# Measurement of TeV γ-rays from Markarian 501 during a state of high activity

F. Aharonian[1], A.G. Akhperjanian[2], J. A. Barrio[3,4], A. R. Bazlarkin[2], K. Bernlöhr[1], J. J. G. Beteta[4], S. M. Bradbury[5], J. L. Contreras[4], J. Cortina[6], A. Daum[1], T. Deckers[6], E. Feigl[7], J. Fernandez[3,4], V. Fonseca[4], A. Frass[6], H. Fruck[1], J. C. Gonzalez[4], V. Haustein[1], G. Heinzelmann[6], M. Hemberger[1], G. Hermann[1], M. Hess[1], A. Heusler[1], W. Hofmann[1], I. Holl[6], D. Horns[1], C. Köhler[1], R. Kankanian[2,8], O. Kirstein[6], A. Konopelko[1], I. Kornmeyer[6], D. Kranich[6], H. Krawczynski[1], H. Lampeitl[6], A. Lindner[6], E. Lorenz[7], N. Magnussen[6], H. Meyer[5], R. Mirzoyan[6,9], H. Möller[6], A. Moralejo[4], L. Padilla[4], M. Panter[1], D. Petry[6], R. Plaga[6], J. Prahl[6], C. Prosch[6], G. Pühlhofer[1], G. Rauterberg[6], W. Rhode[5], V. Sahakian[2], M. Samorski[6], J. A. Sanchez[4], D. Schmele[5], W. Stamm[6], M. Ulrich[1], H. J. Völk[1], S. Wagner[6], H. Wetteskind[6], O. A. Wiedner[1], M. Willmer[6], and H. Wirth[1]

(HEGRA Collaboration)

[1] Max-Planck-Institut für Kernphysik, P.O. Box 103980, D-69029 Heidelberg, Germany
[2] Yerevan Physics Institute, Yerevan, Armenia
[3] Max-Planck-Institut für Physik, Föhringer Ring 6, D-80805 München, Germany
[4] Facultad de Ciencias Físicas, Universidad Complutense, E-28040 Madrid, Spain
[5] Universität Kiel, Inst. für Kernphysik, Olshausenstraße 40, D-24118 Kiel, Germany
[6] ETH Hönggerberg, Institut für Physik, Gloriastr.35, D-8093 Hönggerberg, Germany
[7] Universität Hamburg, II. Inst. für Experimentalphysik, Luruper Chaussee 149, D-22761 Hamburg, Germany

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**Abstract**

From March 15, 1997 to April 14, 1997, a high count rate level of about 25 mHz γ-rays above 1 TeV was observed from Mrk 501, using the HEGRA system of four imaging Cherenkov telescopes. The average count rate level is about twice the count rate of γ-rays observed from the Crab Nebula. The energy spectrum of the radiation was measured from 1 TeV to beyond 10 TeV, and follows a power law. For Mrk 501, the value of the differential spectral index of 2.49 ± 0.11 ± 0.05 is comparable to that of the Crab Nebula, 2.71 ± 0.17 ± 0.05. The time variation of the γ-ray flux was studied on time scales between 5 min. and days.

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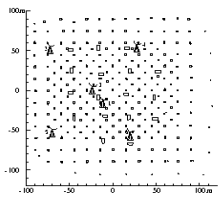
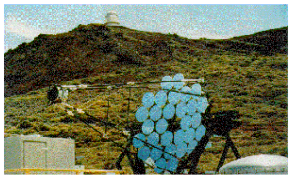



Layout of the HEGRA experiment

Observatorio del Roque de los Muchachos
Canary Island La Palma (2200 m a.s.l)
28.75° N, 17.89° W

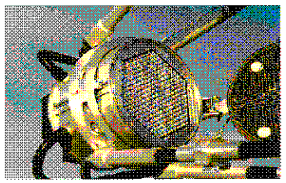
## The HEGRA Telescope System

Central Telescope: CT3



The HEGRA System of IACTs: 5 identical telescopes

- reflector: 50 (front aluminized & quartz coated) 60 cm diameter spherical glass mirrors; $S_R \sim 8.5\,m^2$, $F \sim 5\,m$; $\sigma \approx 5$ arcmin
- mount: alt-azimuth; pointing accuracy ≈ 20 arcsec
- camera: close-packed cluster of 271 PMTs (EMI 9083) fixed with hollow hexagonal light guide; pixel size ~0.25°, FoV ≈ 4.6°
- electronics: in VME standard; 120 MHz FADCs
- triggers: local ~2 pixels from 271 ($\vartheta_{thr} \sim 10$ ph. ~ e-); system ~ "n IACTs from 5" (n ≥ 2)



The HEGRA Telescope System: Camera

### Performance of the HEGRA 5-IACT array

- Energy Threshold: $E_{th} \sim 500$ GeV
- Angular Resolution: $\vartheta\vartheta < 0.1°$
- Energy resolution: $\Delta E/E \leq 25\%$
- Flux sensitivity:

*Point Sources*

STROBO mode: "background-free" ~ 10 γ's in effective FoV ~ 1.5° × 1.5°

$I_{min} \sim 0.5\,Crab\,(1/1h)^{-1}$ for $t \leq 1h$

TASTIGHT mode: "5σ" detection

$I_{min} \sim 0.04\,Crab\,(1/25h)^{-1/2}$ for $t \geq 1h$

*Extended Sources*

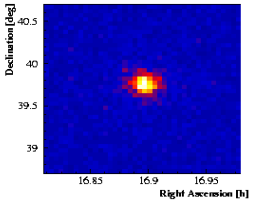
$I_{min} \sim 0.15\,Crab\,(\vartheta/1°)\,(1/100\,h)^{-1/2}$

*Transient Sources*

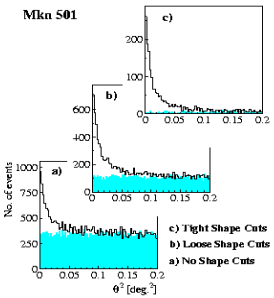
Fluence: $\geq 3 \times 10^{-8}\,erg/cm^2$ for $\Delta t \leq 1h$

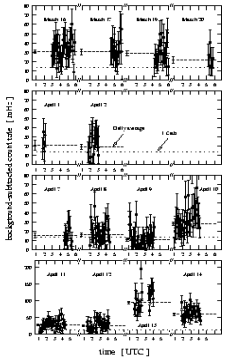
HEGRA CT-System

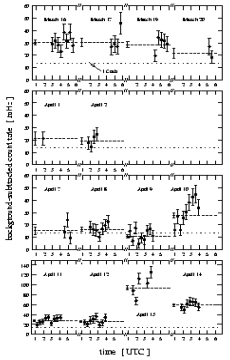
Markarian 501

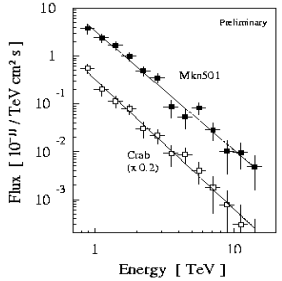



Mkn 501

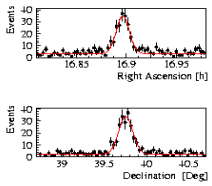



Markarian 501: Source Position




Ra = 16.8978 h , Dec = 39.757 [Deg]

The distribution in the angle $\vartheta$ between the shower axis and the source location is plotted; shown is $dN/d\vartheta^2$.

- For the uniform background from charged cosmic rays one expects a flat distribution in $\vartheta^2$. A γ-ray source causes an excess around $\vartheta \approx 0$. The background is shown as a shaded histogram; it is flat in $\vartheta^2$. In the region out to $\vartheta \sim 0.7°$ on average about 50 excess events per 1 h are counted. This already in the raw data, before selection cuts, a clear signal is visible (a).
- The shapes of Cherenkov images can be used to suppress cosmic-ray background relative to γ-ray showers; γ-rays generate narrower and more compact images. As selection parameter, the width of each image in a given event is scaled to the Monte-Carlo expected width of γ-ray images as a function of image amplitude and distance to the shower core, and an average scaled width is calculated for all telescopes participating in an event. To maintain high efficiency and to minimize corrections, a very loose cut is applied (b). While the background is reduced by a factor of about 3 the number of events in the peak is almost unchanged.
- At the expense of signal statistics the background can be reduced much further. One (c) illustrates the effect of tight cuts which almost completely eliminate the background.

The source position is determined from the two-dimensional distribution of the reconstructed shower directions relative to the direction to Mrk 501, after application of the loose cuts. The accuracy achieved is about 0.01°.

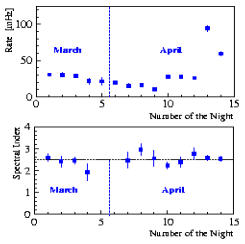
To investigate time variabilities, data were grouped in different time bins, ranging from a hour 5 min., to entire nights. We required that at least two telescopes triggered, and applied a loose angular cut at $\vartheta \sim 0.22°$ as well as a loose shape cut of images.

Shown in the left panel is the detection rate as a function of time in bins of 5 min., over the observation period of 14 days. The errors given are statistical only; the dashed line show daily averages.

The right panel shows bins of 20 min. over the entire observation period.

The HEGRA IACT system with its energy resolution of about 25% allows detailed studies of the spectrum of γ-ray sources. To derive the flux $F(E)$ only events where at least three telescopes triggered were used, to guarantee good stereoscopic reconstruction. Events were counted within a radius of 0.22° from the source, and the background determined from the average yield of events around a virtual source on the opposite side of the source were subtracted. In the determination of spectra, only runs with zenith angles below 30° were included. The resulting Mrk 501 fluxes for the first four days of observations are shown together with the spectrum of the Crab Nebula analyzed in identical fashion. Both spectra are compatible with pure power laws up to at least of 12 TeV, with a differential spectral index of 2.71 ± 0.17 in case of the Crab Nebula, and 2.49 ± 0.11 for Mrk 501 (statistical errors only; at this moment the systematic errors are estimated within ±0.25; the integral flux above 1 TeV is $1.0 \pm 0.1 \cdot 10^{-11}/cm^2 s$ and $2.25 \pm 0.2 \cdot 10^{-11}/cm^2 s$, respectively (systematic errors ±25%).

For the first twelve nights of the observation period, the average daily excess count rates were around 25 mHz, corresponding to about 90 γ-rays/h, whereas on night 14 the rate had increased by about a factor of 4, coming down to about 200 γ-rays/hour the last night, still about 2 times higher than the average rate.

The average daily differential spectral indices have an average value of 2.5 (±0.11 ± 0.25) over the observation period.

Temporal Behaviour of Flux and Spectral Index