Magnetic field and flavor effects on the neutrino fluxes from cosmic accelerators

Particle and astroparticle theory seminar

July 4, 2011

MPIK Heidelberg, Germany

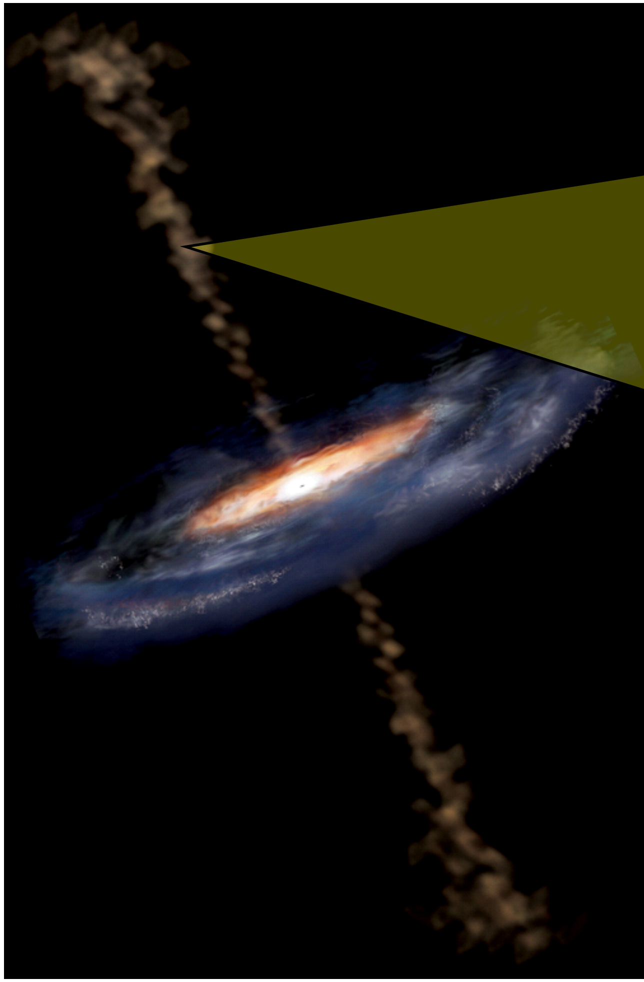
Walter Winter

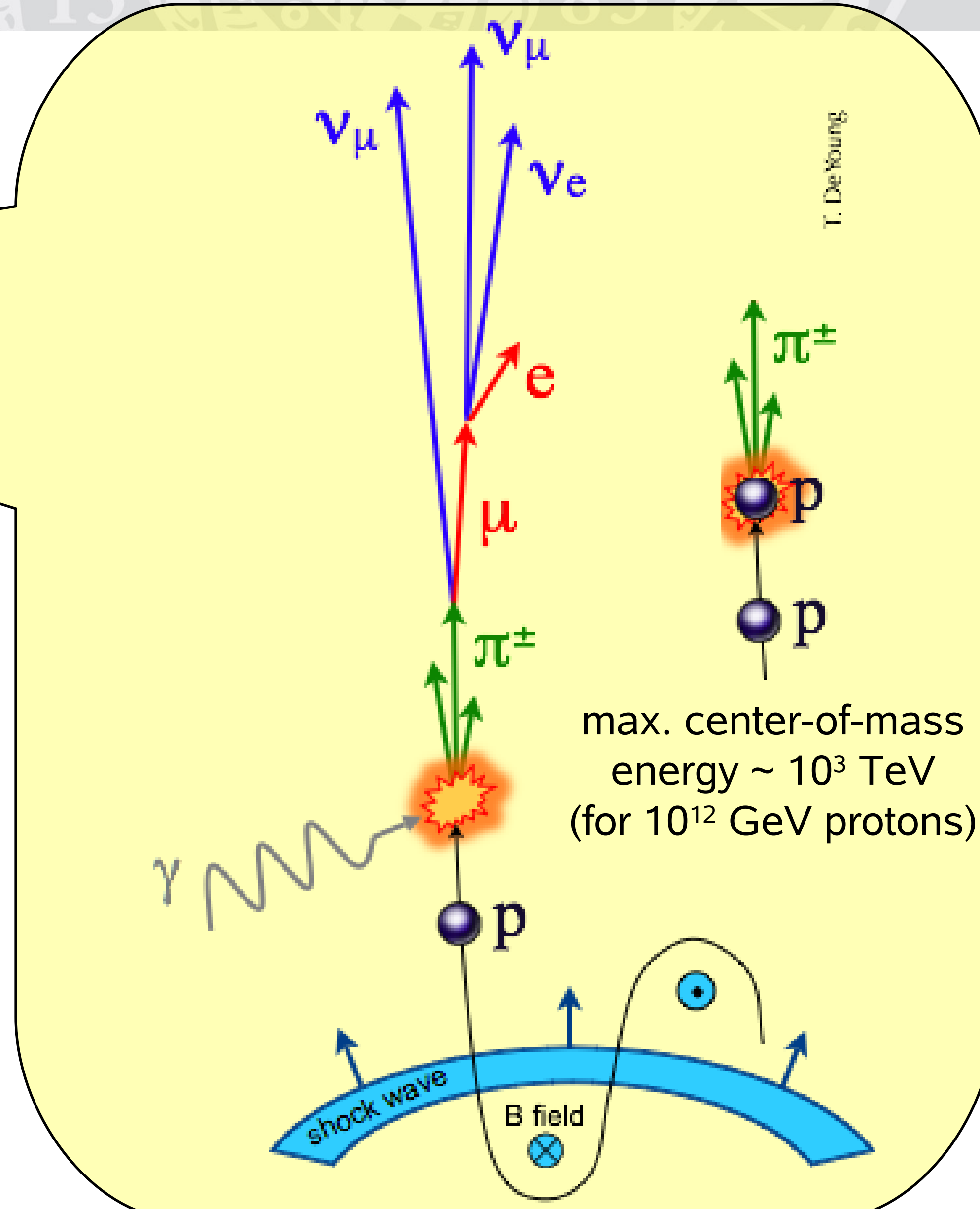
Universität Würzburg

Contents

- Introduction
- Simulation of sources
  A self-consistent approach
  Flavor composition?
- Neutrino propagation and detection:
  Measuring flavor, new physics tests?
- Comments on Glashow resonance
- On gamma-ray burst (GRB) neutrino fluxes
- Summary

Neutrino production in astrophysical sources

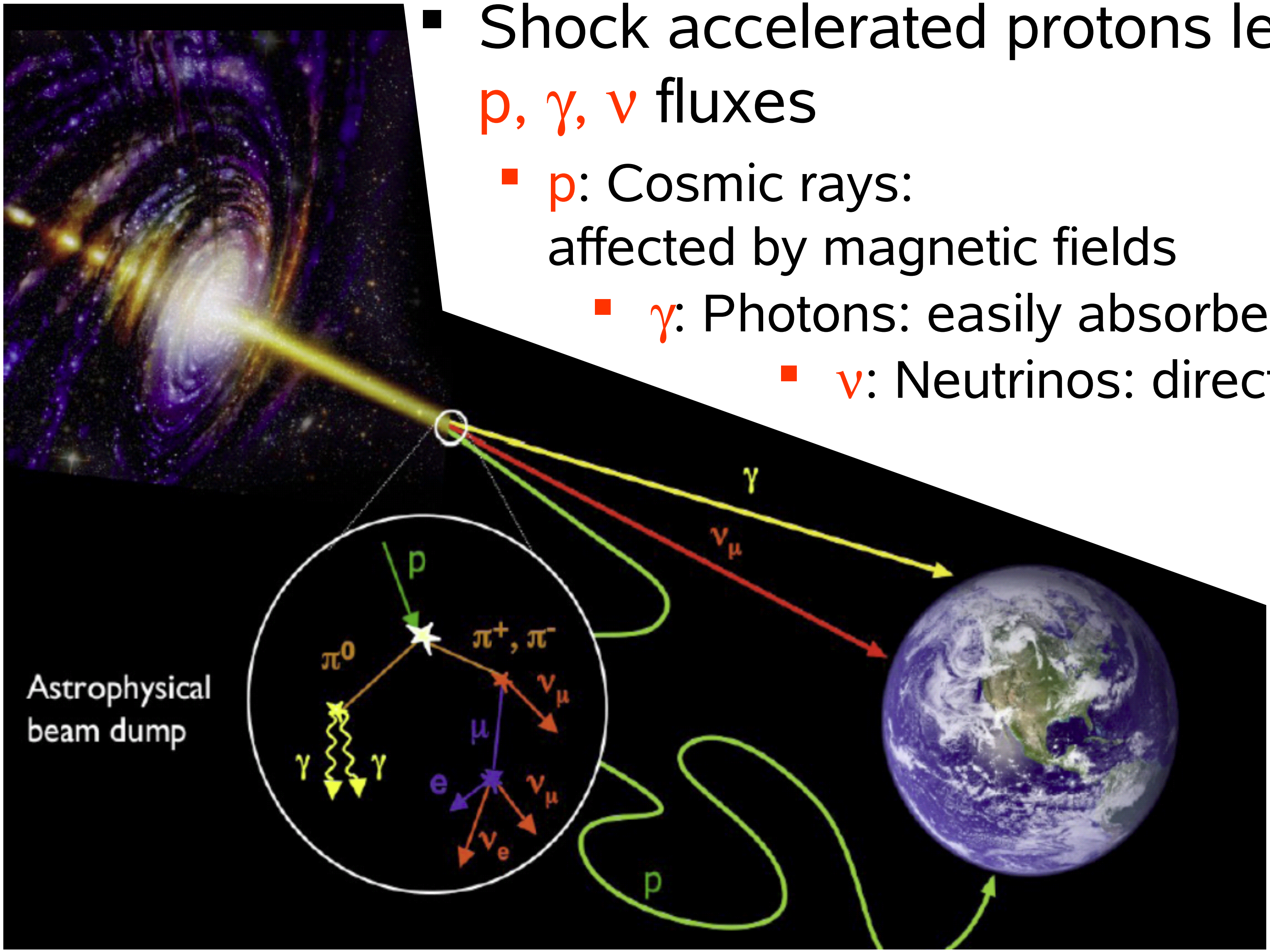




Example: Active galaxy
(Halzen, Venice 2009)

Different messengers

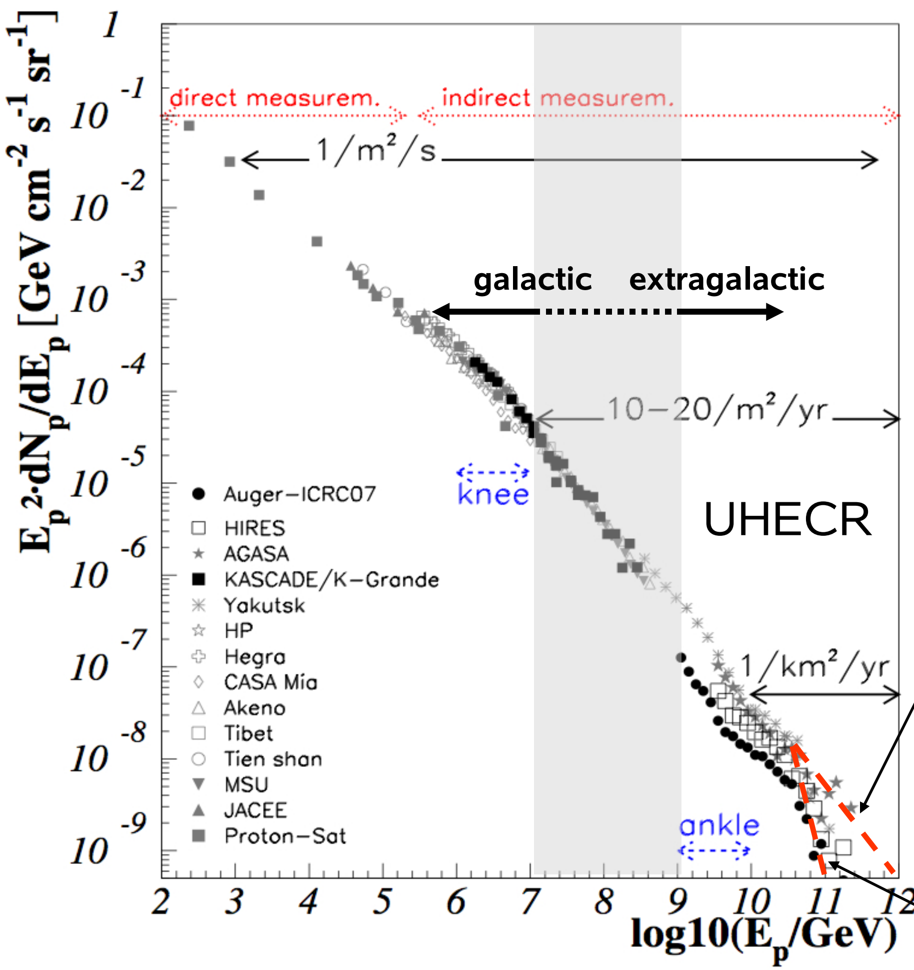
- Shock accelerated protons lead to p, γ, ν fluxes
  - p: Cosmic rays: affected by magnetic fields
    - γ: Photons: easily absorbed/scattered
      - ν: Neutrinos: direct path



(Teresa Montaruli, NOW 2008)

Evidence for proton acceleration, hints for neutrino production

- Observation of cosmic rays: need to accelerate protons/hadrons somewhere
- The same sources should produce neutrinos:
  - in the source (pp, pγ interactions)
  - Proton ($E > 6\ 10^{10}$ GeV) on CMB ⇨ GZK cutoff + cosmogenic neutrino flux



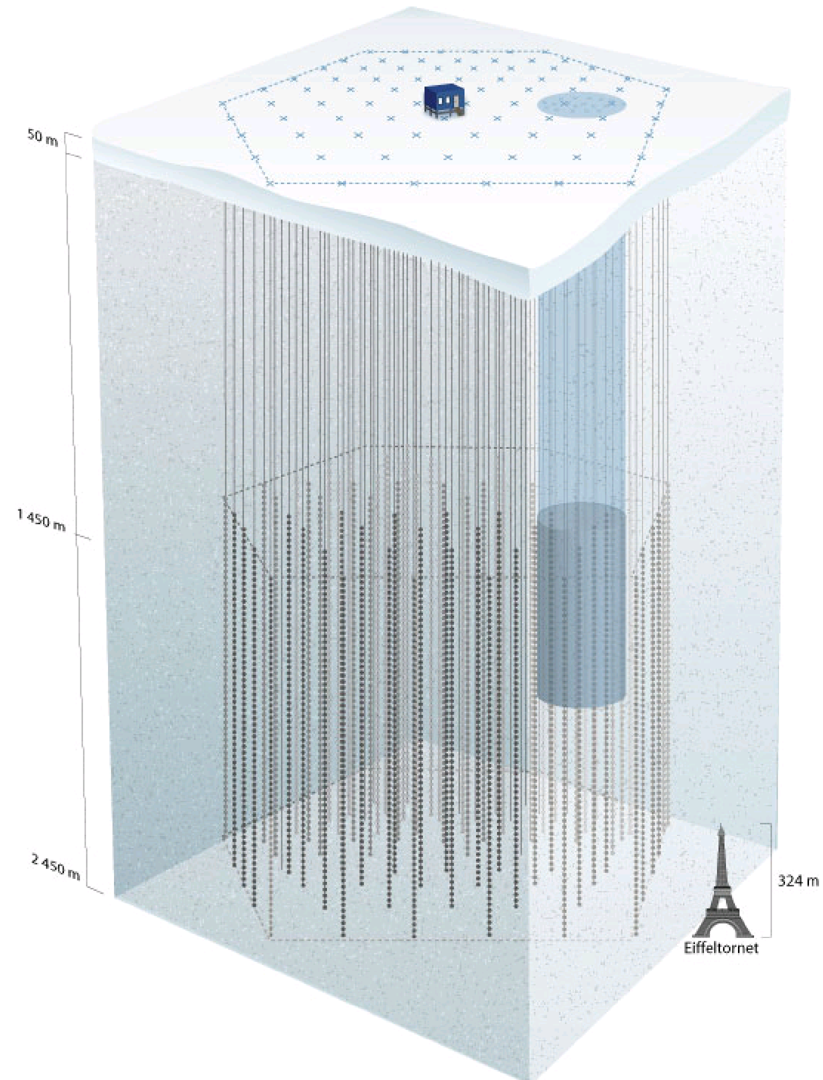
In the source: $E_{p,max}$ up to 10^{12} GeV?

GZK cutoff?

(Source: F. Halzen, Venice 2009)

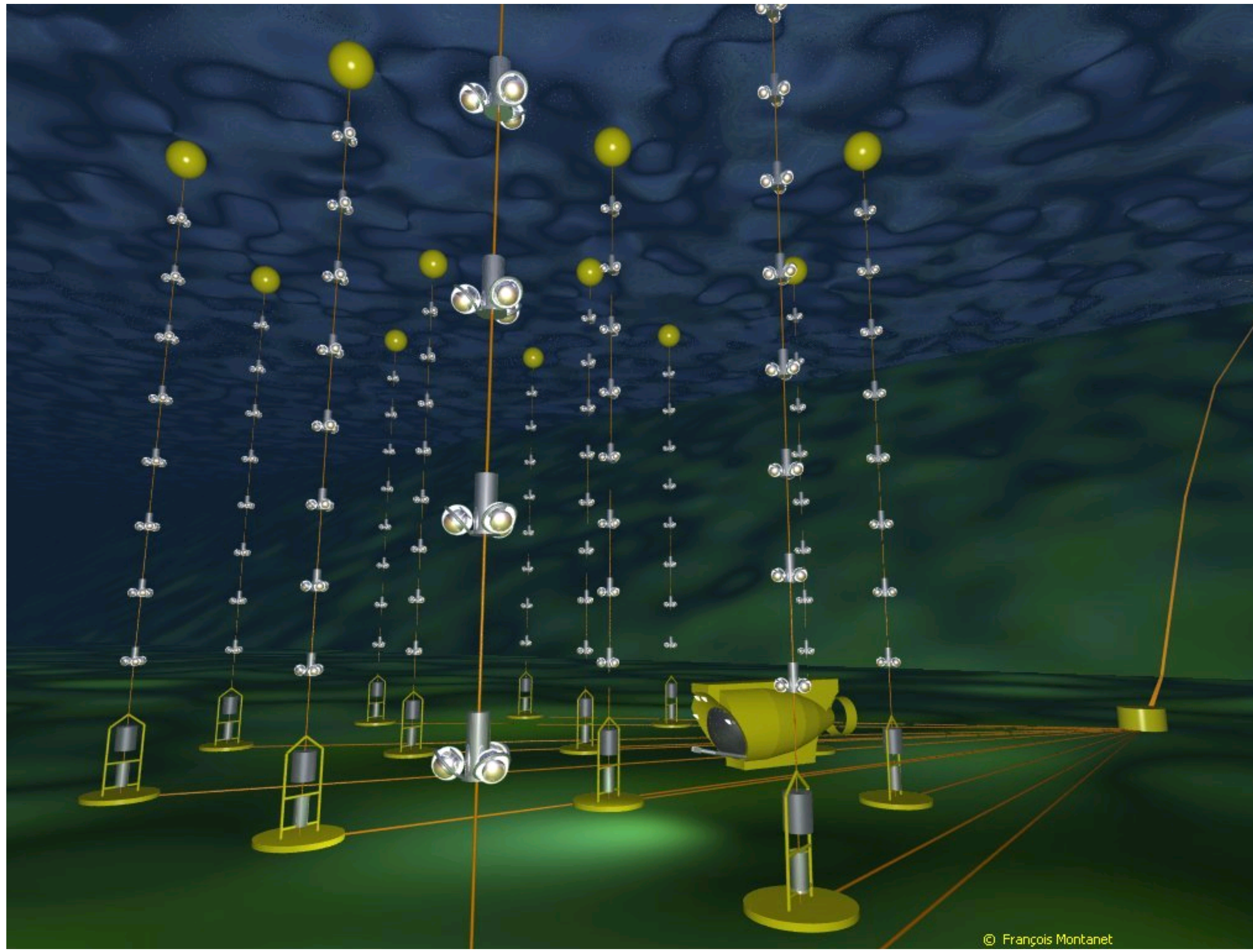
Neutrino detection: IceCube

- Example:
  IceCube at South Pole
  Detector material: ~ 1 km^3 antarctic ice
- Completed 2010/11 (86 strings)
- Recent data releases, based on parts of the detector:
  - Point sources IC-40 [IC-22]
    arXiv:1012.2137, arXiv:1104.0075
  - GRB stacking analysis IC-40
    arXiv:1101.1448
  - Cascade detection IC-22
    arXiv:1101.1692
- Have not seen anything (yet)
  - ➢ What does that mean?
  - ➢ Are the models wrong?
  - ➢ Which parts of the parameter space does IceCube actually test?



http://icecube.wisc.edu/

Neutrino astronomy in the Mediterranean: Example ANTARES



http://antares.in2p3.fr/

When do we expect a ν signal?

[some personal comments]

- Unclear if specific sources lead to neutrino production; spectral energy distribution can be often described by other processes as well (e.g. inverse Compton scattering, proton synchrotron, …)
- However: whereever cosmic rays are produced, neutrinos should be produced to some degree
- There are a number of additional candidates, e.g.
  - „Hidden“ sources (e.g. „slow jet supernovae“ without gamma-ray counterpart)
    (Razzaque, Meszaros, Waxman, 2004; Ando, Beacom, 2005; Razzaque, Meszaros, 2005; Razzaque, Smirnov, 2009)
  - Large fraction of Fermi-LAT unidentified sources?

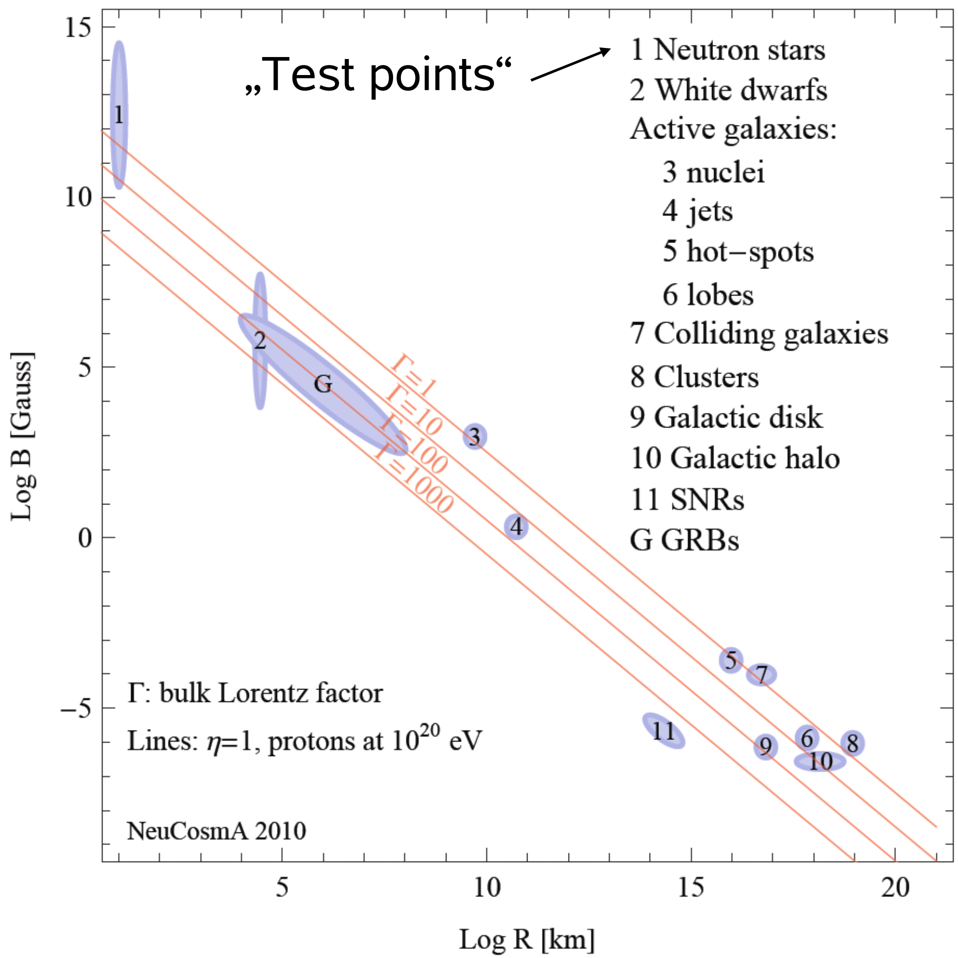
➢ From the neutrino point of view: „Fishing in the dark blue sea“? Looking at the wrong places?

➢ Need for tailor-made neutrino-specific approaches? [unbiased by gamma-ray and cosmic ray observations]

➢ Also: huge astrophysical uncertainties; try to describe at least the particle physics as accurate as possible!

Where to look for sources?

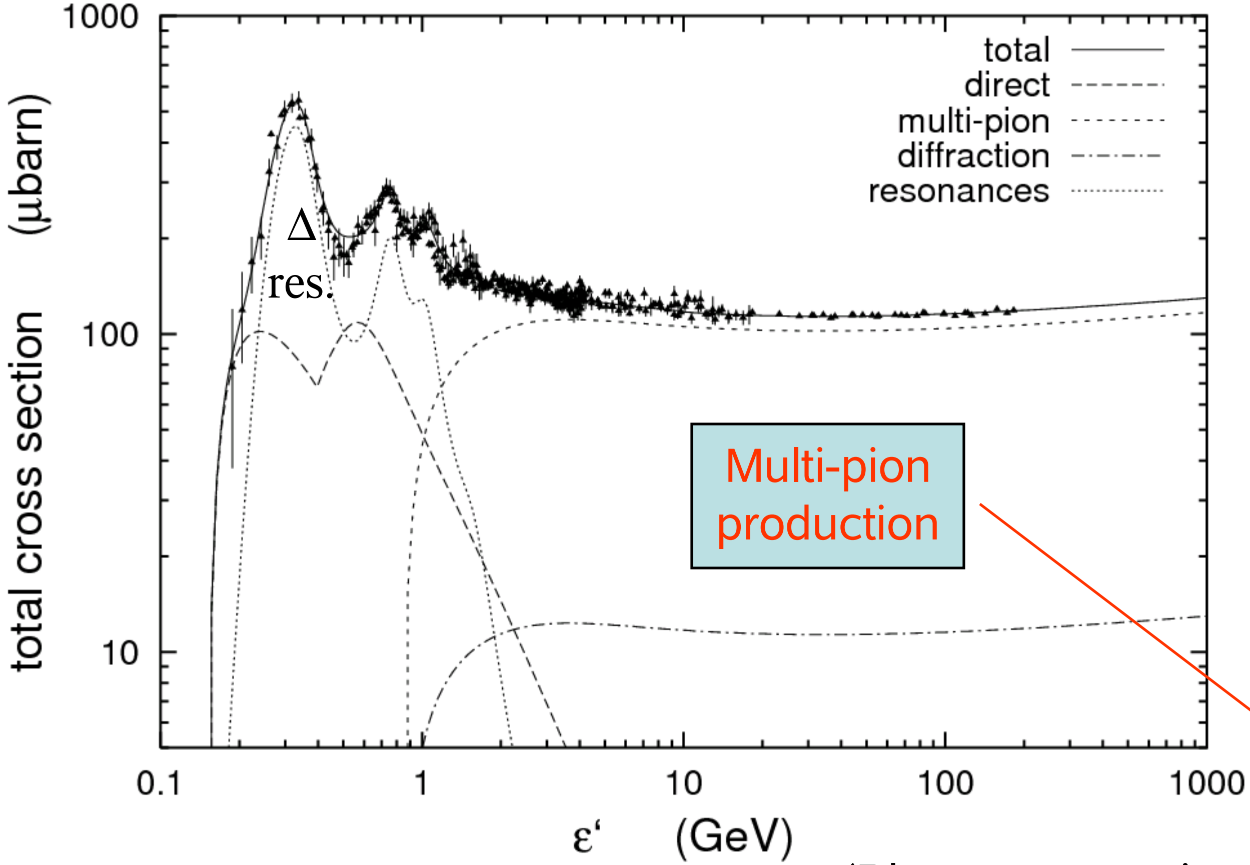
- Model-independent (necessary) condition:
  $E_{max} \sim Z\, e\, B\, R$
  (Larmor-Radius < size of source)
  - ➢ Particles confined to within accelerator!
- Sometimes: define acceleration rate
  $t^{-1}_{acc} = \eta\, Z\, e\, B/E$
  (η: acceleration efficiency)
- Caveat: condition relaxed if source heavily Lorentz-boosted (e.g. GRBs)

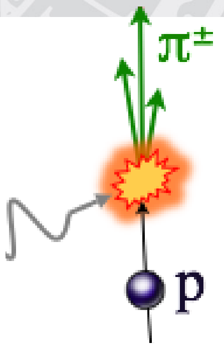


(Hillas, 1984; version adopted from M. Boratav)

Simulation of sources

Pion photoproduction





Power law injection spectrum from Fermi shock acc.

Multi-pion production

(Photon energy in nucleon rest frame)

Resonant production, direct production

Different characteristics (energy loss of protons; energy dep. cross sec.)

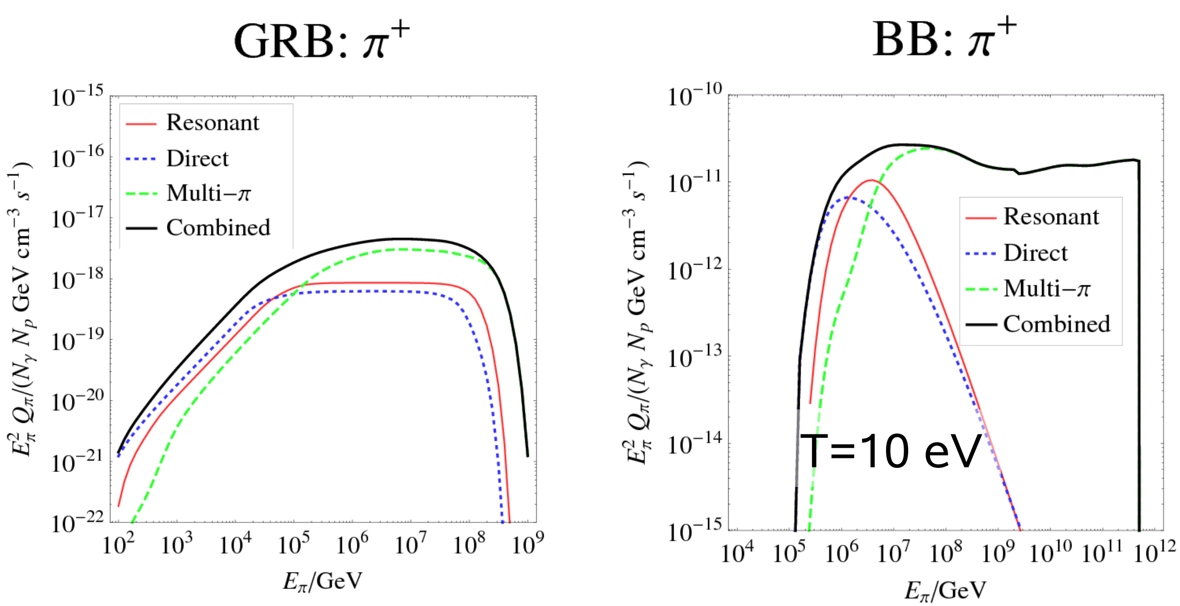
(Mücke, Rachen, Engel, Protheroe, Stanev, 2008; **SOPHIA**)

Meson photoproduction

- Often used: Δ(1232)-resonance approximation

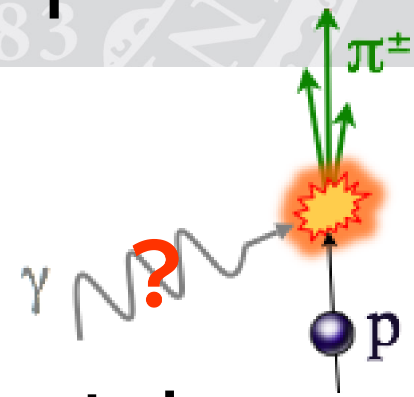
$$p+\gamma \rightarrow \Delta^+ \rightarrow \begin{cases} n+\pi^+ & 1/3 \text{ of all cases} \\ p+\pi^0 & 2/3 \text{ of all cases} \end{cases}$$

- Limitations:
  - No π^- production; cannot predict π^+/π^- ratio (affects neutrino/antineutrino)
  - High energy processes affect spectral shape (X-sec. dependence!)
  - Low energy processes (t-channel) enhance charged pion production
  - ➢ Charged pion production underestimated compared to π^0 production by factor of 2.4 (independent of input spectra!)



from: Hümmer, Rüger, Spanier, Winter, ApJ 721 (2010) 630

A self-consistent approach



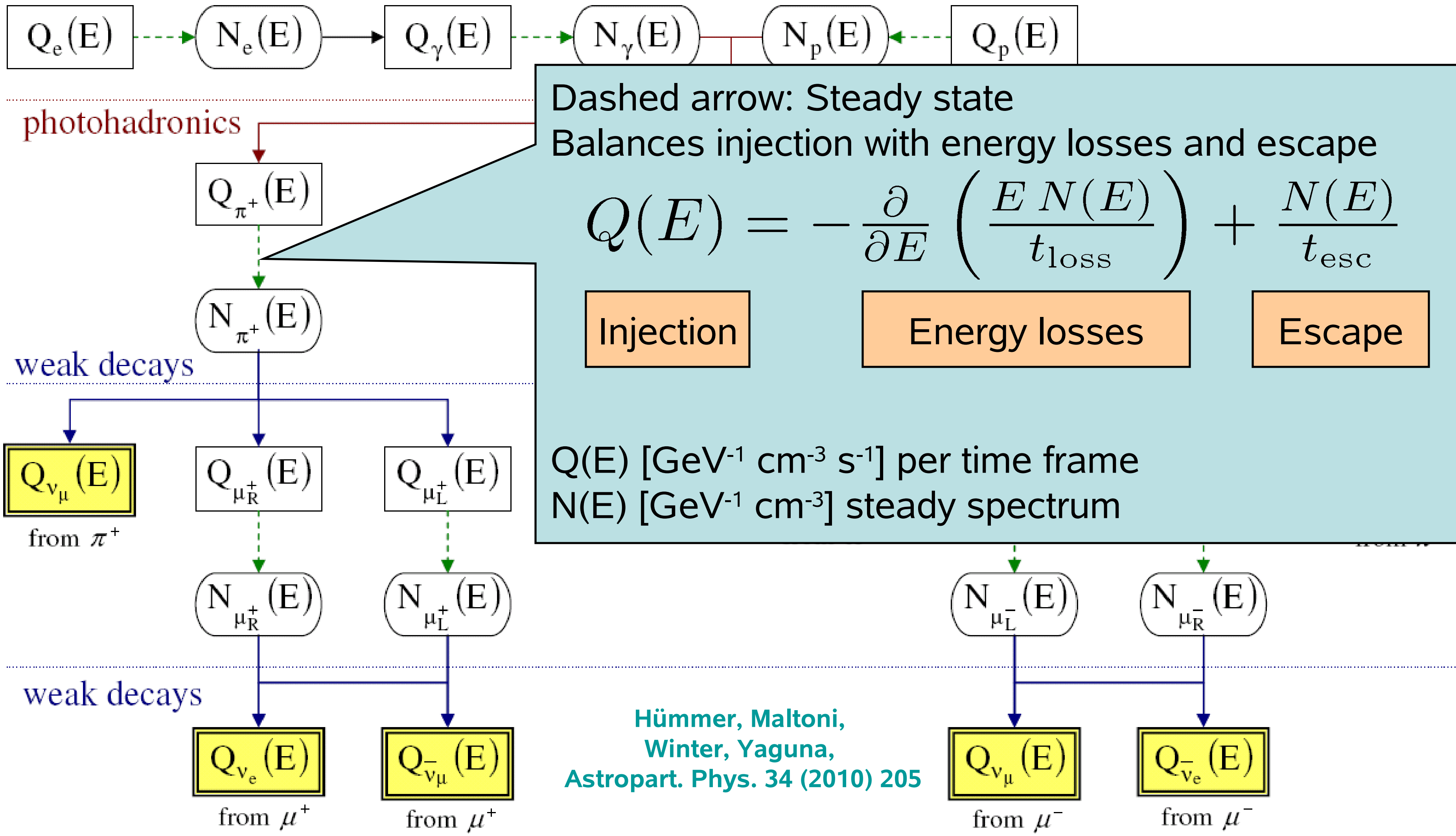
- Target photon field typically:
  - Put in by hand (e.g. obs. spectrum: GRBs)
  - Thermal target photon field
  - **From synchrotron radiation of co-accelerated electrons/positrons (AGN-like)**
- Requires few model parameters, mainly

Parameter	Units	Description	Typical values used
R	km (kilometers)	Size of acceleration region	10^{1} km ... 10^{21} km
B	G (Gauss)	Magnetic field strength	10^{-9} G ... 10^{15} G
α	1	Universal injection index	1.5 ... 4

- Purpose: describe wide parameter ranges with a simple model; minimal set of assumptions for ν!?

Model summary

Dashed arrows: include cooling and escape



Dashed arrow: Steady state

Balances injection with energy losses and escape

$$Q(E) = -\frac{\partial}{\partial E}\left(\frac{E\,N(E)}{t_{\mathrm{loss}}}\right) + \frac{N(E)}{t_{\mathrm{esc}}}$$

Injection | Energy losses | Escape

Q(E) [$\mathrm{GeV^{-1}\,cm^{-3}\,s^{-1}}$] per time frame

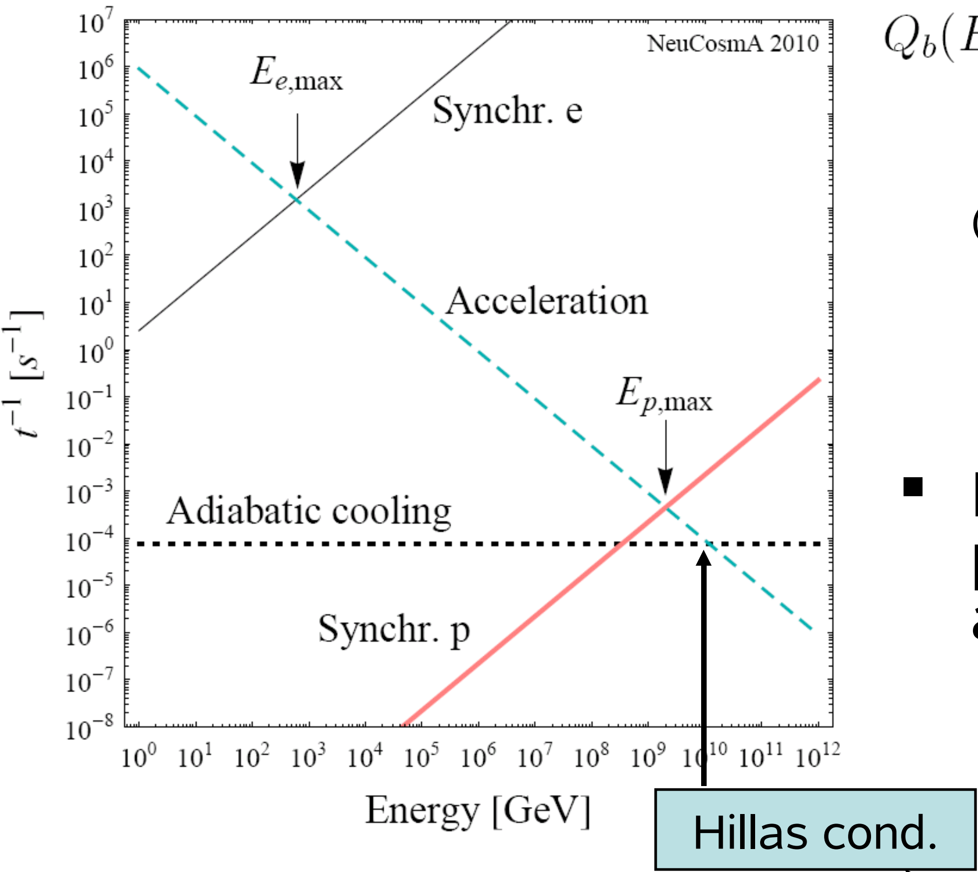
N(E) [$\mathrm{GeV^{-1}\,cm^{-3}}$] steady spectrum

Hümmer, Maltoni, Winter, Yaguna, Astropart. Phys. 34 (2010) 205

An example: Primaries

TP 3: α=2, B=10^3 G, R=$10^{9.6}$ km

Maximum energy: e, p



Hümmer, Maltoni, Winter, Yaguna, 2010

- Meson production described by

$$Q_b(E_b) = \int \frac{dE_p}{E_p} N_p(E_p) \int d\varepsilon \, N_\gamma(\varepsilon) \, R_b(x, y)$$

$$x = E_b/E_p$$

$$y \equiv (E_p \varepsilon)/m_p$$

(summed over a number of interaction types)
  - ➢ Only product normalization enters in pion spectra as long as synchrotron or adiabatic cooling dominate
- Maximal energy of primaries (e, p) by balancing energy loss and acceleration rate

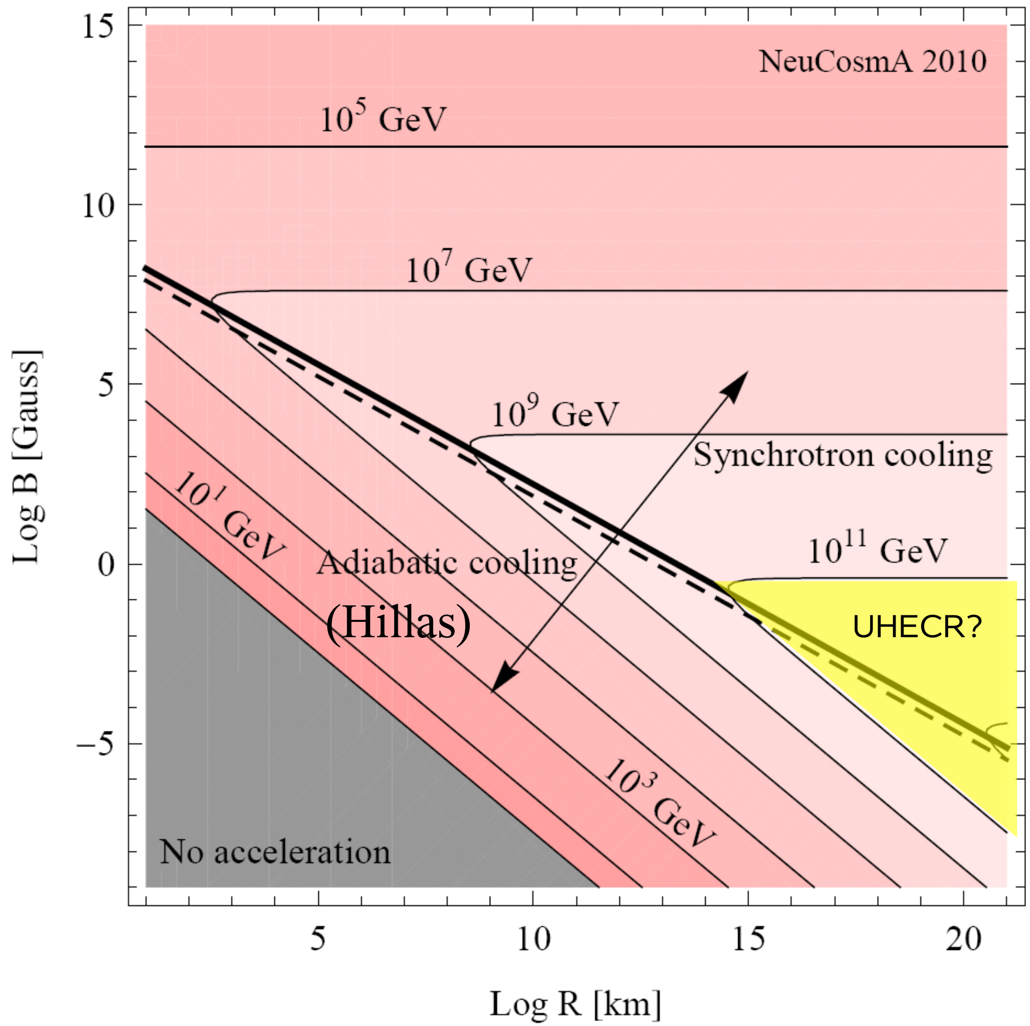
$$t_{\mathrm{acc}}^{-1} = \eta \frac{c^2 e B}{E}$$

➢ Hillas condition often necessary, but not sufficient!

Maximal proton energy (general)

- Maximal proton energy (⇨ UHECR) often constrained by proton synchrotron losses

- Sources of UHECR in lower right corner of Hillas plot?



Hümmer, Maltoni, Winter, Yaguna, 2010

An example: Secondaries

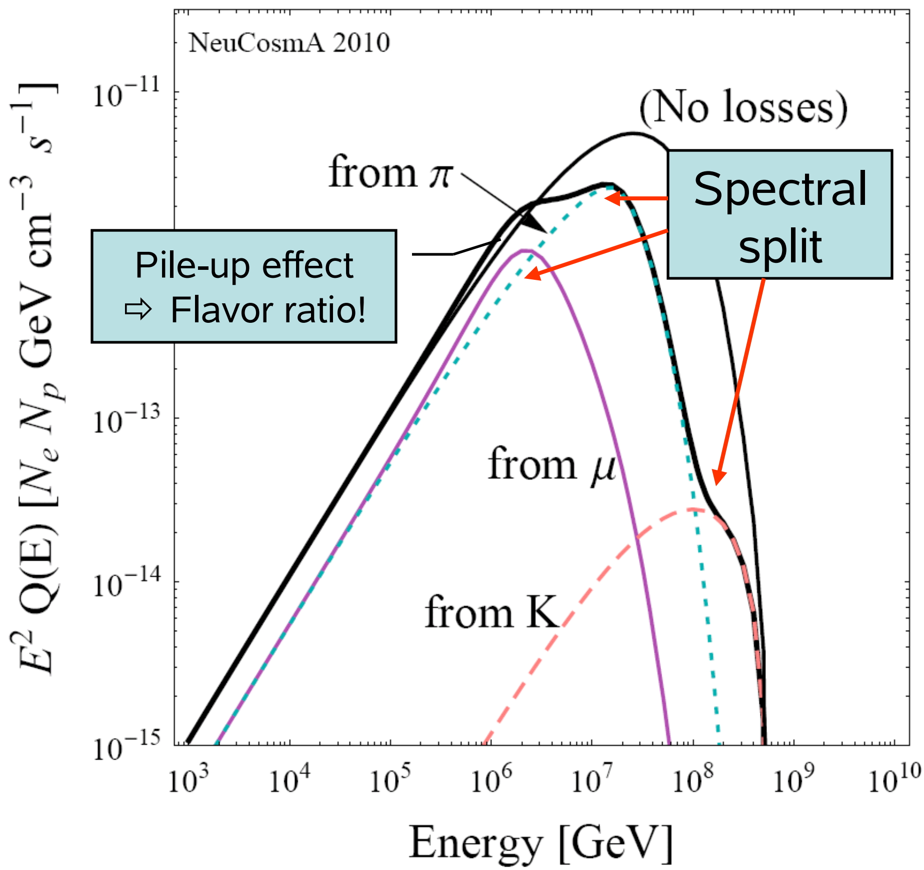
α=2, B=10^3 G, R=$10^{9.6}$ km

- Secondary spectra (μ, π, K) become loss-steepend above a critical energy

$$E_c = \sqrt{\frac{9\pi\epsilon_0 m^5 c^5}{\tau_0 e^4 B^2}}$$

  - E_c depends on particle physics only (m, τ_0), and **B**
  - Leads to characteristic flavor composition
  - Any additional cooling processes mainly affecting the primaries will not affect the flavor composition
  - Flavor ratios most robust predicition for sources?
  - The only way to directly measure B?

Injection: ν_μ

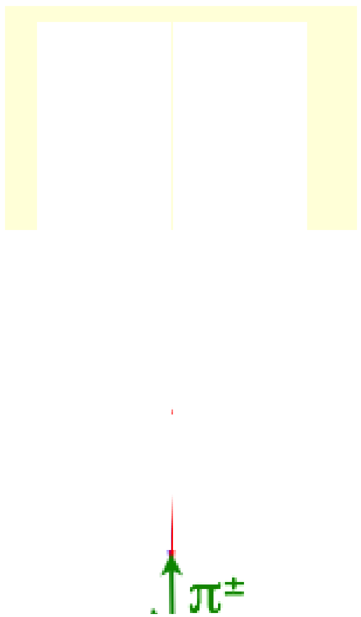


Hümmer et al,
Astropart. Phys. 34 (2010) 205

Flavor composition at the source

Flavor composition at the source
(Idealized – **energy independent**)

- Astrophysical neutrino sources produce certain flavor ratios of neutrinos (ν_e:ν_μ:ν_τ):
- Pion beam source (1:2:0)
  Standard in generic models
- Muon damped source (0:1:0)
  at high E: Muons loose energy before they decay
- Muon beam source (1:1:0)
  Cooled muons pile up at lower energies (also: heavy flavor decays)
- Neutron beam source (1:0:0)
  Neutron decays from pγ
  (also possible: photo-dissociation of heavy nuclei)



$$n \longrightarrow p + e^- + \bar{\nu}_e$$

➢ At the source: Use ratio ν_e/ν_μ (nus+antinus added)

However: flavor composition is energy dependent!

Muon beam
⇨ muon damped

Energy window with large flux for classification

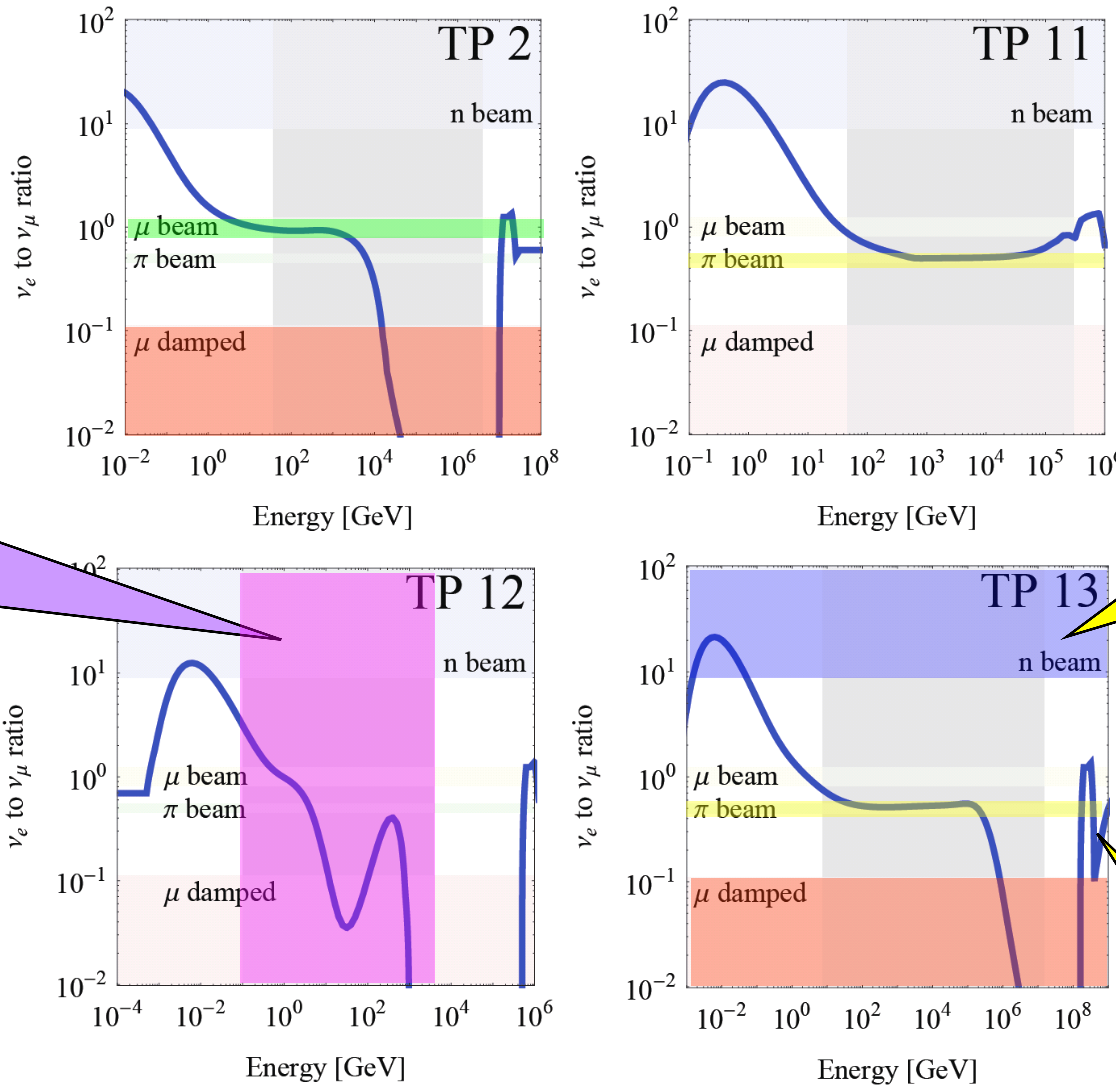
Undefined (mixed source)

Pion beam

Typically n beam for low E (from pγ)

Pion beam
⇨ muon damped

Behavior for small fluxes undefined

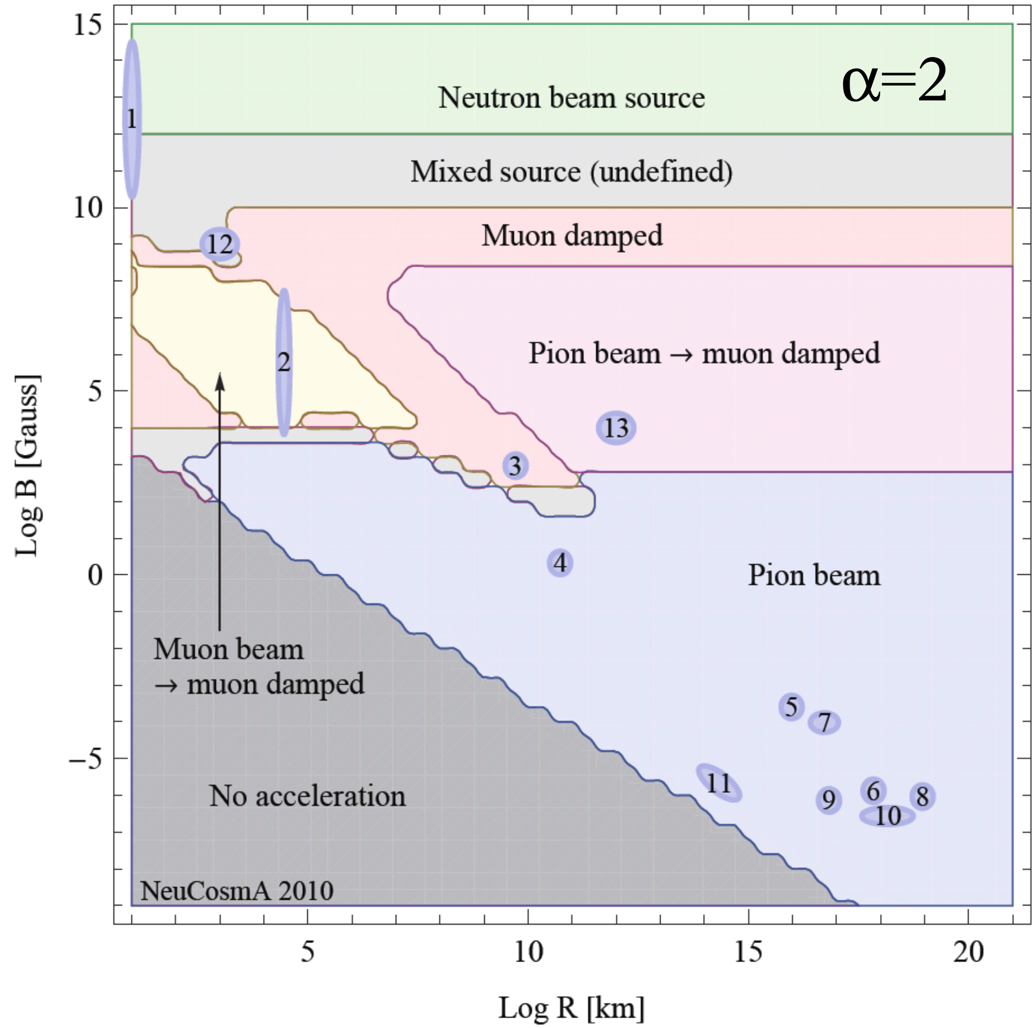


(from Hümmer, Maltoni, Winter, Yaguna, 2010;
see also: Kashti, Waxman, 2005; Kachelriess, Tomas, 2006, 2007; Lipari et al, 2007)

Parameter space scan

- All relevant regions recovered
- GRBs: in our model α=4 to reproduce pion spectra; pion beam ⇨ muon damped

  (confirms Kashti, Waxman, 2005)
- Some dependence on injection index



Hümmer, Maltoni, Winter, Yaguna, 2010

Neutrino propagation and detection

Neutrino propagation

- Key assumption: Incoherent propagation of neutrinos
- Flavor mixing: $P_{\alpha\beta} = \sum_{i=1}^{3} |U_{\alpha i}|^2 |U_{\beta i}|^2$
- Example: For $\theta_{13}=0$, $\theta_{23}=\pi/4$:

(see Pakvasa review, arXiv:0803.1701, and references therein)

$$\begin{pmatrix} \nu_e^{source} \\ \nu_\mu^{source} \\ \nu_\tau^{source} \end{pmatrix} = \begin{pmatrix} 1 \\ 2 \\ 0 \end{pmatrix} \quad \Rightarrow \quad \begin{pmatrix} \nu_e^{Earth} \\ \nu_\mu^{Earth} \\ \nu_\tau^{Earth} \end{pmatrix} = \begin{pmatrix} 1 \\ 1 \\ 1 \end{pmatrix}$$

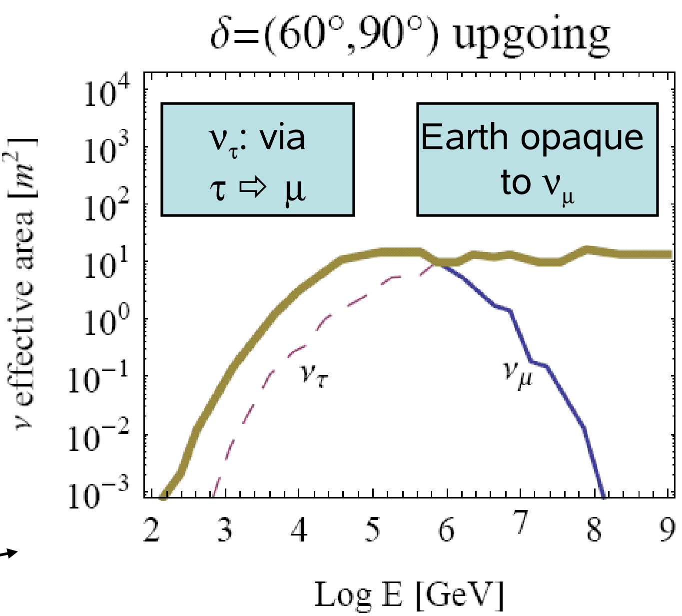
- NB: No CPV in flavor mixing only!
  But: In principle, sensitive to Re exp(-i δ) ~ cosδ

Neutrino detection: Muon tracks

- Number of events depends on neutrino effective area and observ. time t_{exp}:

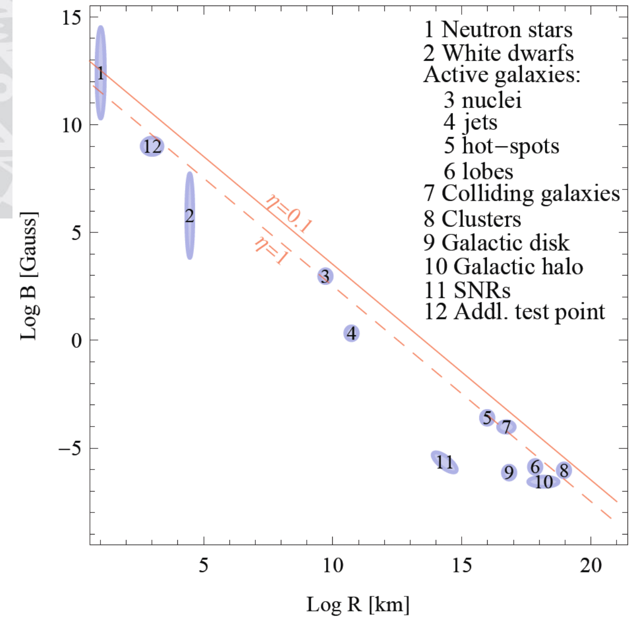
$$N = \int dE \, \mathrm{Exp}(E,\delta) \frac{dN(E)}{dE} = \int dE \, A_\nu^{\mathrm{eff}}(E,\delta) \, t_{\mathrm{exp}} \frac{dN(E)}{dE}$$

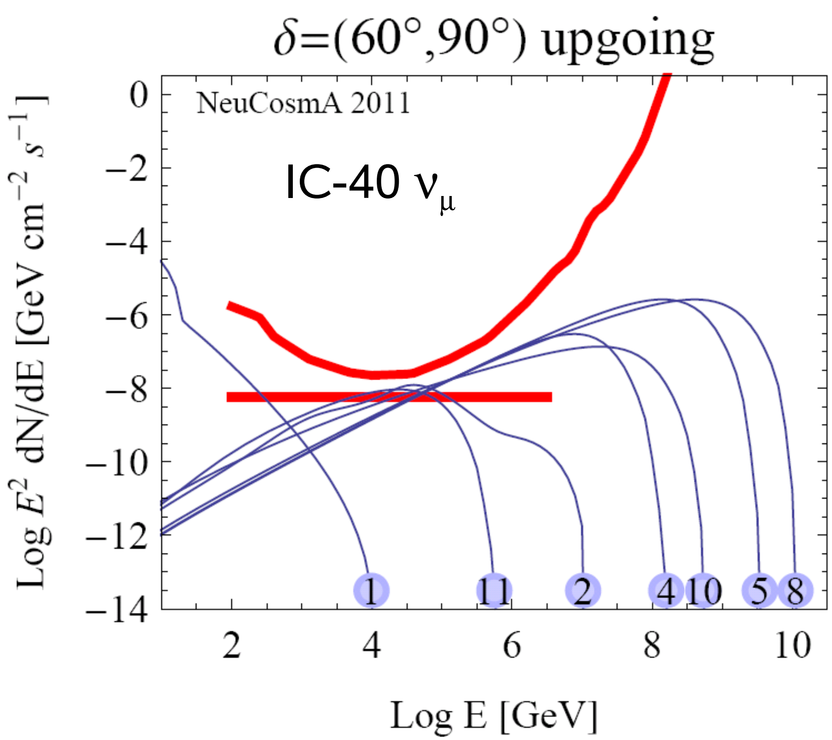
- Neutrino effective area ~ detector area x muon range (E); but: cuts, uncontained events, …
- Time-integrated point source search, IC-40 (arXiv:1012.2137)

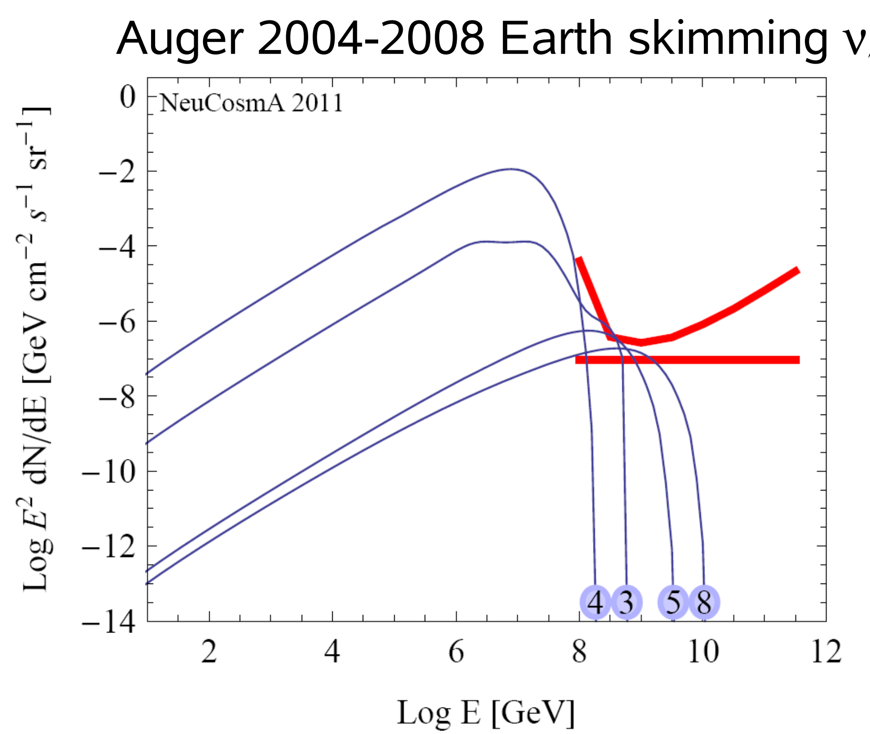


Individual spectra

- Differential limit $2.3\ E/(A_{eff}\ t_{exp})$ illustrates what spectra the data limit best



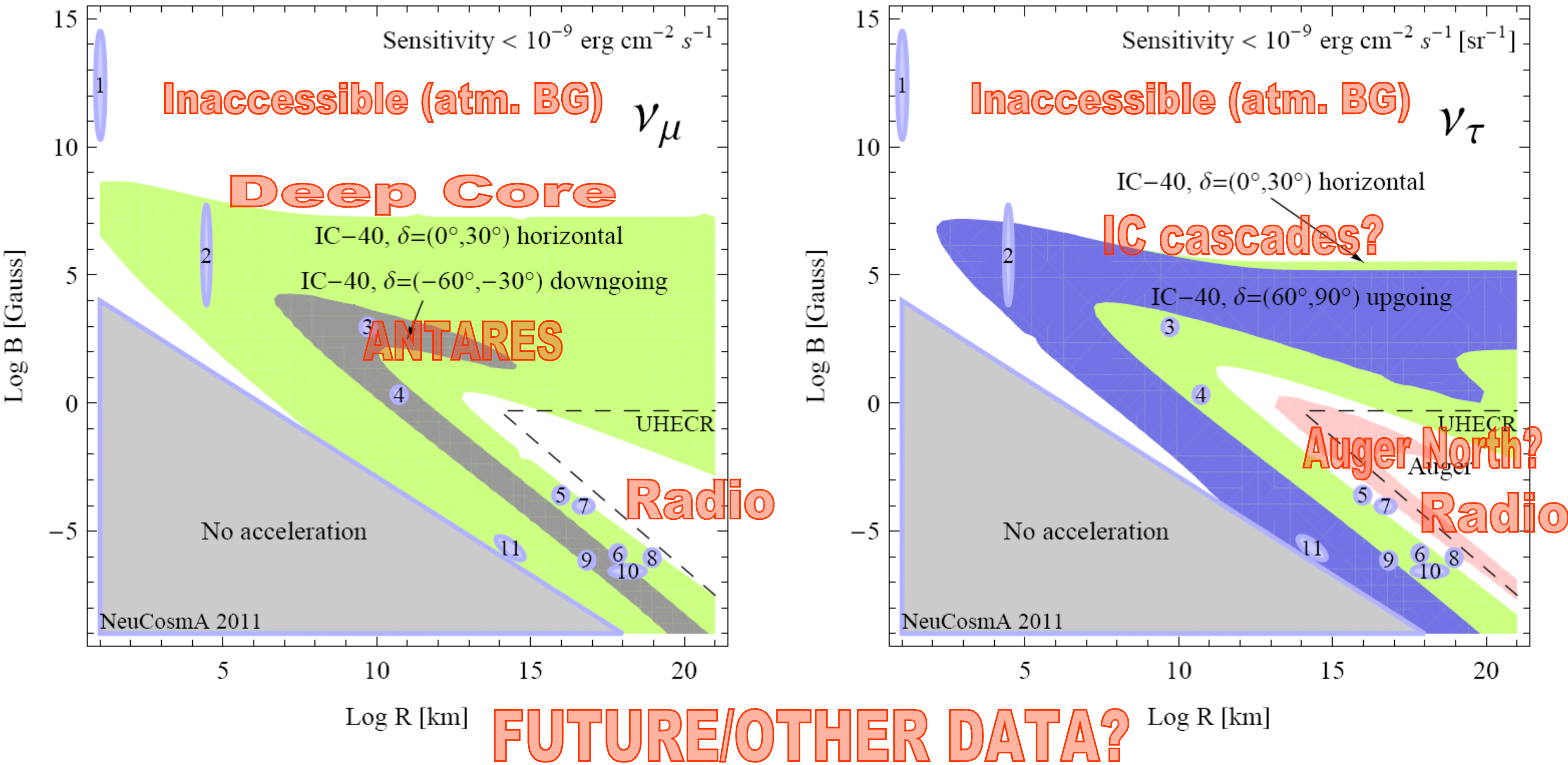




(Winter, arXiv:1103.4266)

Which point sources can specific data constrain best?

Constraints to energy flux density $\phi = \int E \frac{dN(E)}{dE} dE$

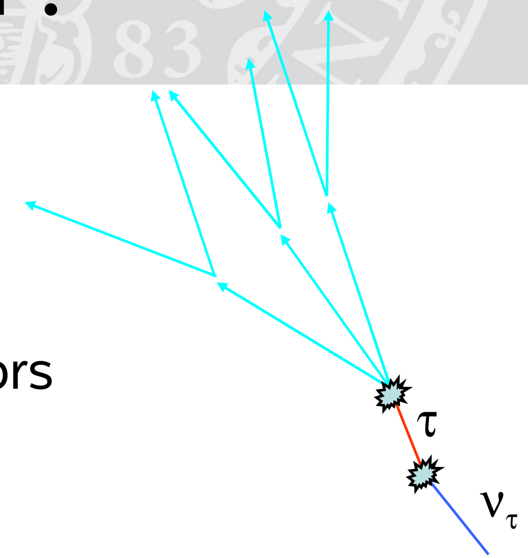


(Winter, arXiv:1103.4266)

Measuring flavor?

- In principle, flavor information can be obtained from different event topologies:
  - Muon tracks - ν_μ
  - Cascades (showers) – CC: ν_e, ν_τ, NC: all flavors
  - Glashow resonance (6.3 PeV): $\bar{\nu}_e$
  - Double bang/lollipop: ν_τ (sep. tau track) →

  (Learned, Pakvasa, 1995; Beacom et al, 2003)



- In practice, the first (?) IceCube „flavor“ analysis appeared recently – IC-22 cascades **(arXiv:1101.1692)**

  Flavor contributions to cascades for E^{-2} extragalatic test flux (after cuts):
    - Electron neutrinos 40%
    - Tau neutrinos 45%
    - Muon neutrinos 15%

  - ➢ Electron and tau neutrinos detected with comparable efficiencies
  - ➢ Neutral current showers are a moderate background

Flavor ratios at detector

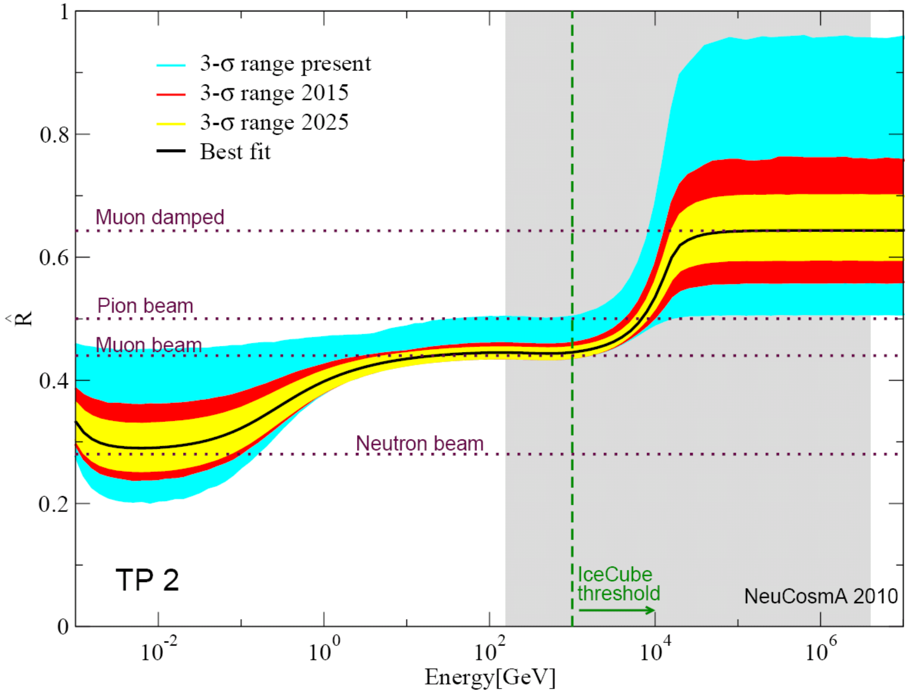
- At the detector: define observables which
  - take into account the unknown flux normalization
  - take into account the detector properties
- Example: Muon tracks to showers
  Do not need to differentiate between electromagnetic and hadronic showers!

$$\hat{R} = \frac{\phi_\mu^{\mathrm{Det}}}{\phi_e^{\mathrm{Det}} + \phi_\tau^{\mathrm{Det}}}$$

- Flavor ratios have recently been discussed for many particle physics applications

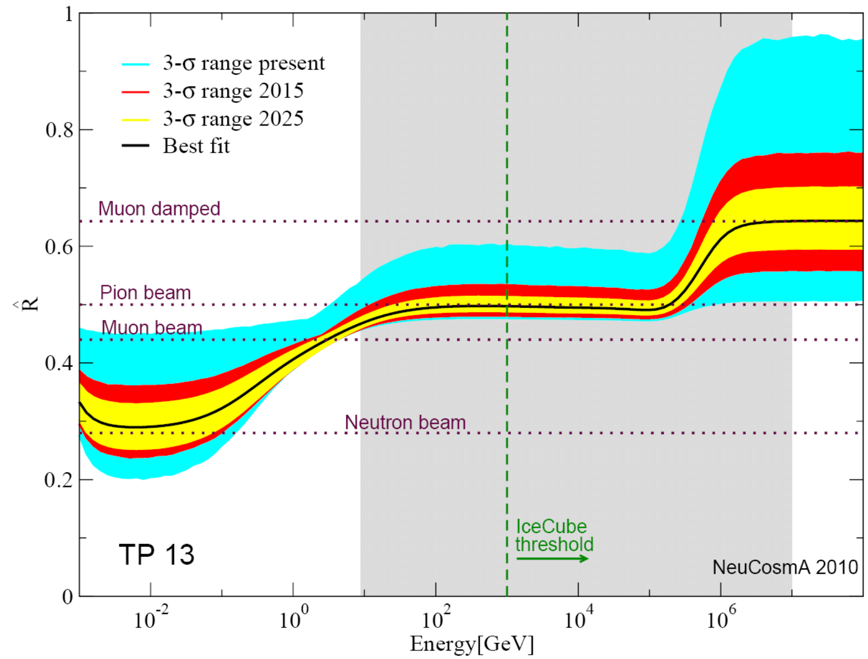
(for flavor mixing and decay: Beacom et al 2002+2003; Farzan and Smirnov, 2002; Kachelriess, Serpico, 2005; Bhattacharjee, Gupta, 2005; Serpico, 2006; Winter, 2006; Majumar and Ghosal, 2006; Rodejohann, 2006; Xing, 2006; Meloni, Ohlsson, 2006; Blum, Nir, Waxman, 2007; Majumar, 2007; Awasthi, Choubey, 2007; Hwang, Siyeon,2007; Lipari, Lusignoli, Meloni, 2007; Pakvasa, Rodejohann, Weiler, 2007; Quigg, 2008; Maltoni, Winter, 2008; Donini, Yasuda, 2008; Choubey, Niro, Rodejohann, 2008; Xing, Zhou, 2008; Choubey, Rodejohann, 2009; Esmaili, Farzan, 2009; Bustamante, Gago, Pena-Garay, 2010; Mehta, Winter, 2011…)

Parameter uncertainties



- Basic dependence recovered after flavor mixing

- However: mixing parameter knowledge ~ 2015 (Daya Bay, T2K, etc) required



Hümmer, Maltoni, Winter, Yaguna, 2010

Julius-Maximilians-
UNIVERSITÄT
WÜRZBURG

New physics in R?

$$\widehat{X}(E) = \frac{\Phi_e^0(E)}{\Phi_\mu^0(E)}$$

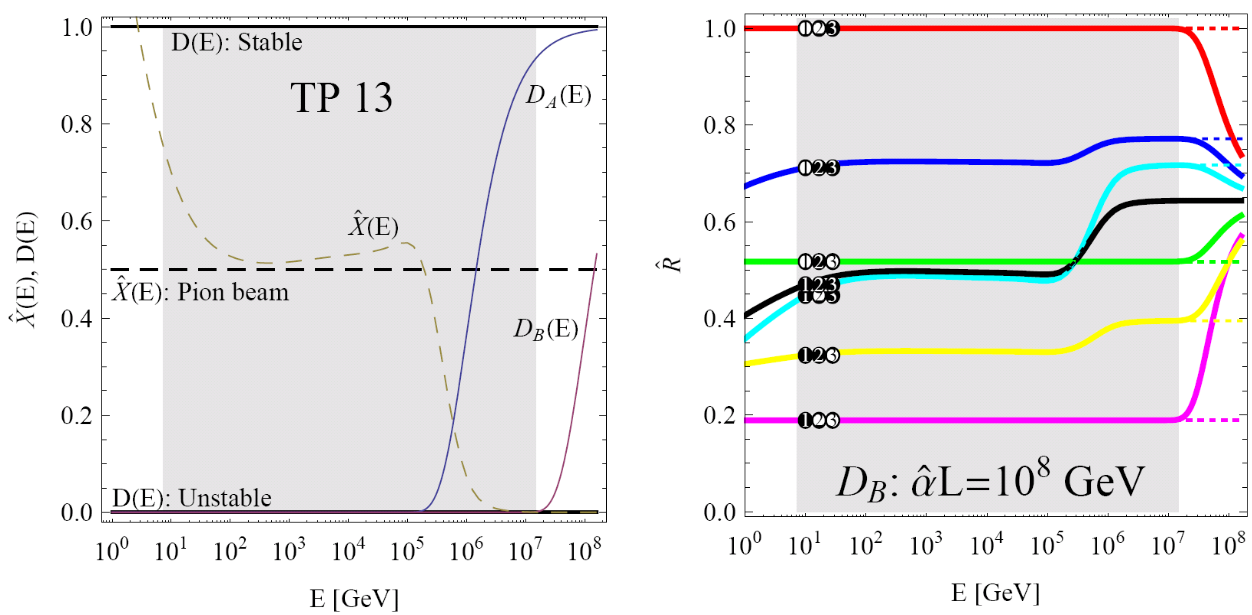
Energy dependence flavor comp. source

$$\widehat{R} \equiv \frac{\Phi_\mu^{\mathrm{Det}}}{\Phi_e^{\mathrm{Det}} + \Phi_\tau^{\mathrm{Det}}} = \frac{P_{e\mu}(E)\,\widehat{X}(E) + P_{\mu\mu}(E)}{[P_{ee}(E) + P_{e\tau}(E)]\,\widehat{X}(E) + [P_{\mu e}(E) + P_{\mu\tau}(E)]}$$

$$P_{\alpha\beta} = \sum_{i=1}^{3} |U_{\beta i}|^2 \, |U_{\alpha i}|^2 \; D_i(E) \quad \text{with} \quad D_i(E) = \exp\left(-\hat{\alpha}_i \frac{L}{E}\right)$$

Energy dep. new physics

(Example: [invisible] neutrino decay)



1 Stable state

1 Unstable state

Mehta, Winter, JCAP 03 (2011) 041; see also Bhattacharya, Choubey, Gandhi, Watanabe, 2009/2010

Glashow resonance?

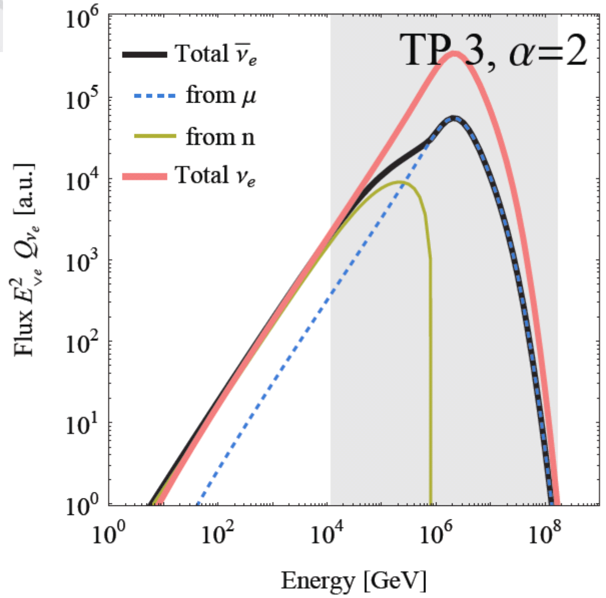
Glashow resonance

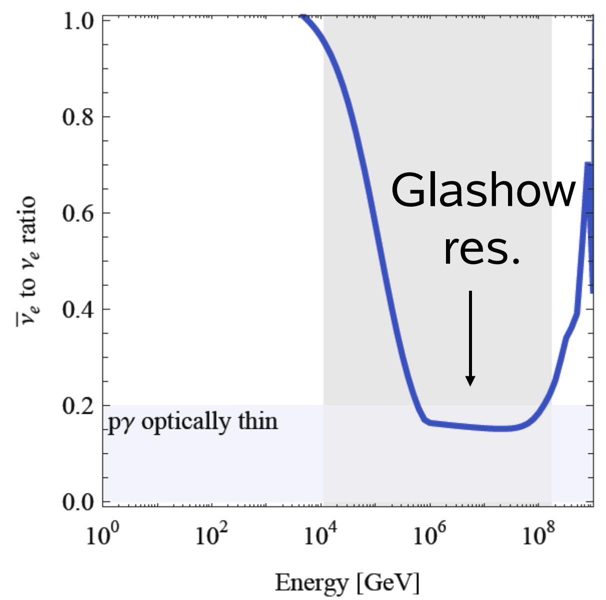
… at source

- pp: Produce π^+ and π^- in roughly equal ratio
- pγ: Produce mostly π^+
  - ➢ Glashow resonance (6.3 PeV, electron antineutrinos) as source discriminator?

Caveats:

- Multi-pion processes produce π^-
- If some optical thickness, nγ “backreactions“ equilibrate π^+ and π^-
- Neutron decays fake π^- contribution
  - ➢ May identify “pγ optically thin source“ with about 20% contamination from π^-, but cannot establish pp source!



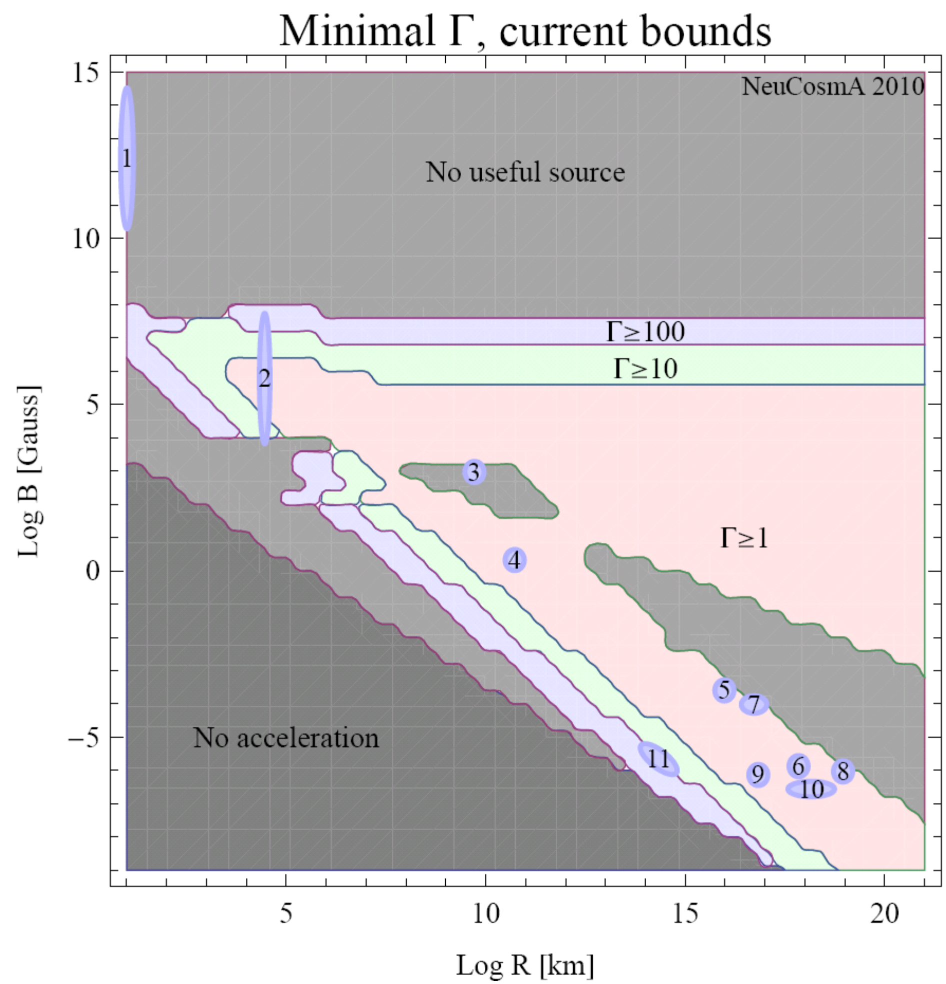


Sec. 3.3 in Hümmer, Maltoni, Winter, Yaguna, 2010; see also Xing, Zhou, 2011

Glashow resonance

… at detector

- Additional complications:
  - Flavor mixing
    (electron antineutrinos from muon antineutrinos produced in μ^+ decays)
  - Have to know flavor composition
    (e.g. a muon damped pp source can be mixed up with a pion beam p γ source)
  - Have to hit a specific energy (6.3 PeV), which may depend on Γ of the source



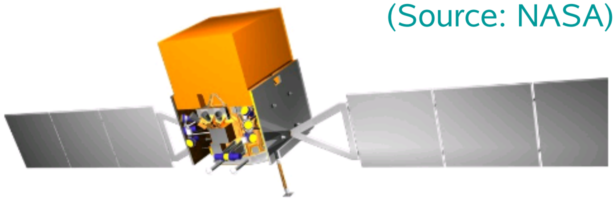
Sec. 4.3 in Hümmer, Maltoni, Winter, Yaguna, 2010

On GRB neutrino fluxes

Example: GRB stacking

- Idea: Use multi-messenger approach

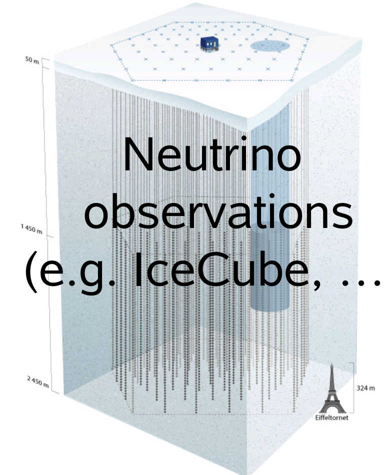
(Source: NASA)



GRB gamma-ray observations
(e.g. Fermi GBM, Swift, etc)

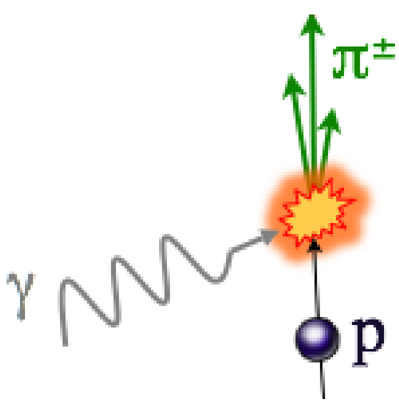
Coincidence!

(Source: IceCube)

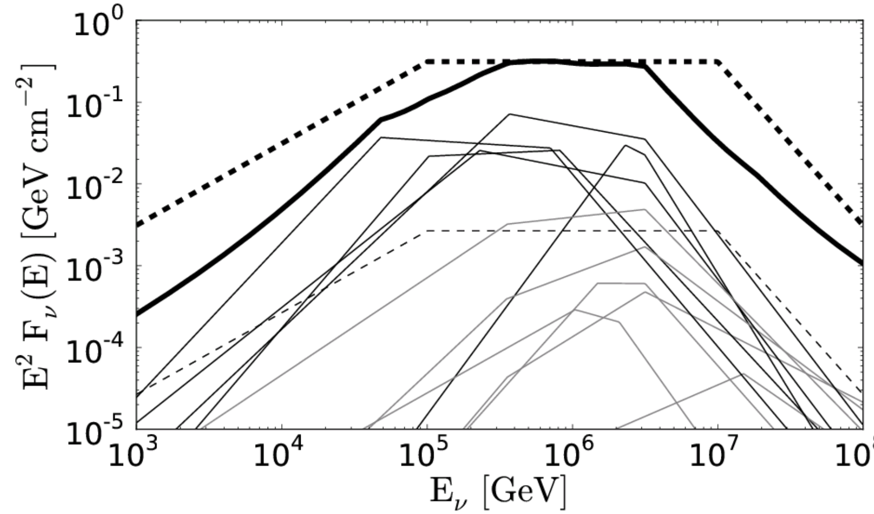


Neutrino observations
(e.g. IceCube, …)

- Predict neutrino flux from observed photon fluxes event by event



Observed:
broken power law
(Band function)

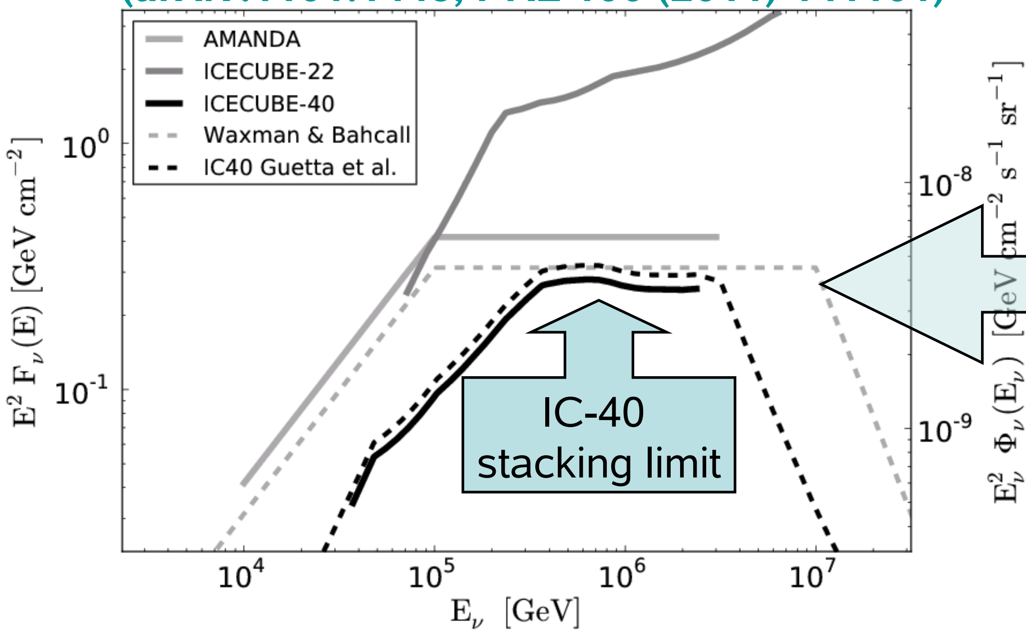


(Example: IceCube, arXiv:1101.1448)

Gamma-ray burst fireball model:

IC-40 data meet generic bounds

(arXiv:1101.1448, PRL 106 (2011) 141101)



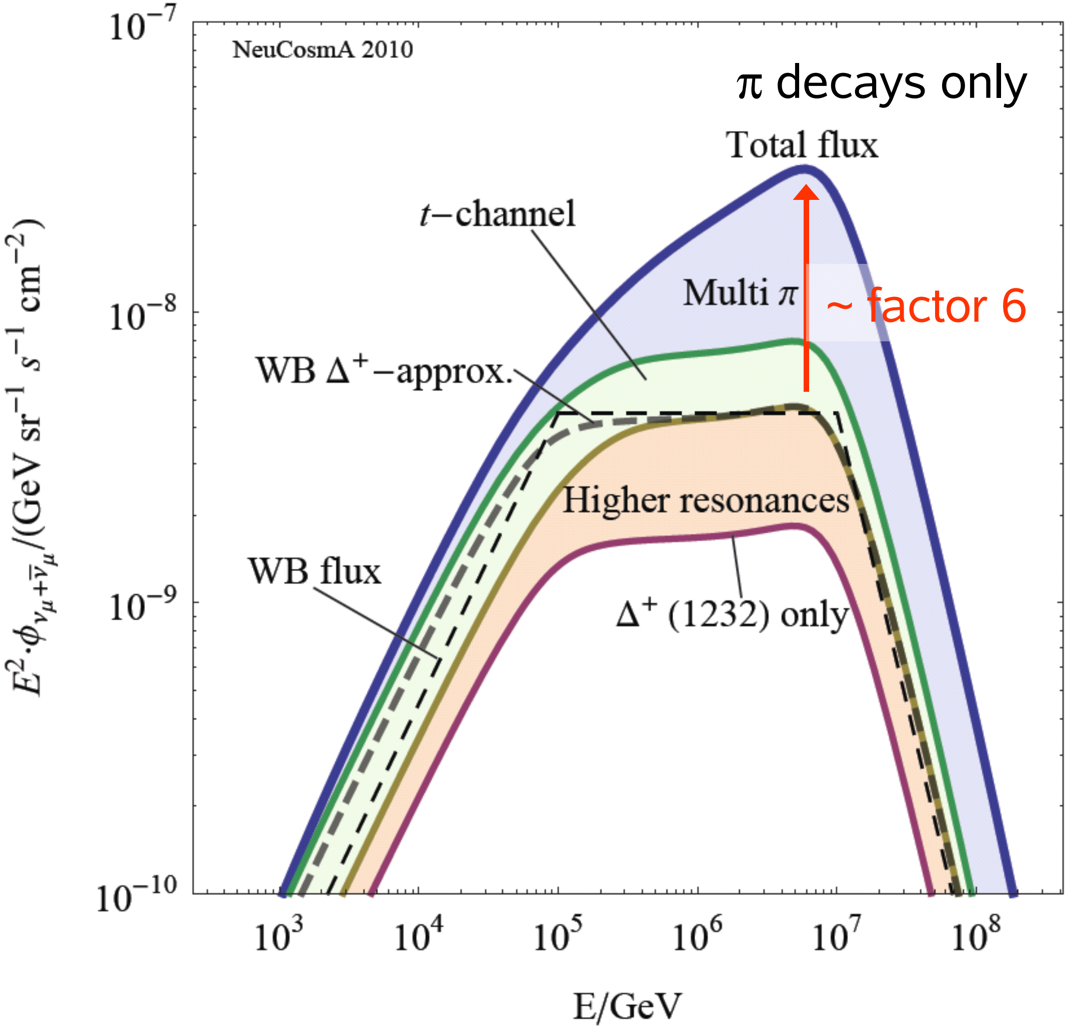
Generic flux based on the assumption that GRBs are the sources of (highest energetic) cosmic rays

(Waxman, Bahcall, 1999; Waxman, 2003; spec. bursts: Guetta et al, 2003)

- Does IceCube really rule out the paradigm that GRBs are the sources of the ultra-high energy cosmic rays?

(see also Ahlers, Gonzales-Garcia, Halzen, 2011 for a fit to data)

Waxman-Bahcall, reproduced

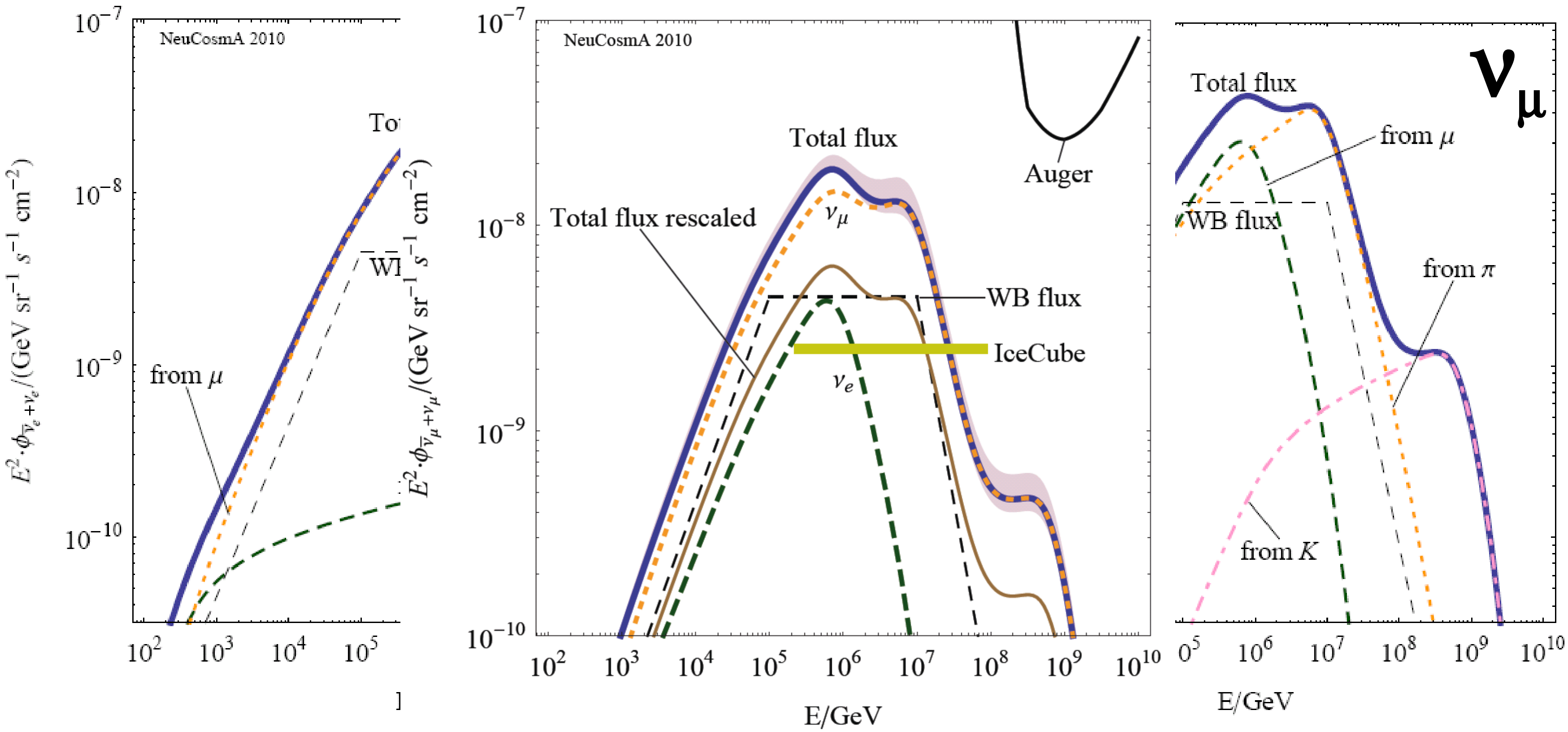


- Reproduced original WB flux with similar assumptions
- Additional charged pion production channels included, also π^-!

Baerwald, Hümmer, Winter, Phys. Rev. D83 (2011) 067303

Fluxes before/after flavor mixing

BEFORE FLAVOR MIXING
AFTER FLAVOR MIXING

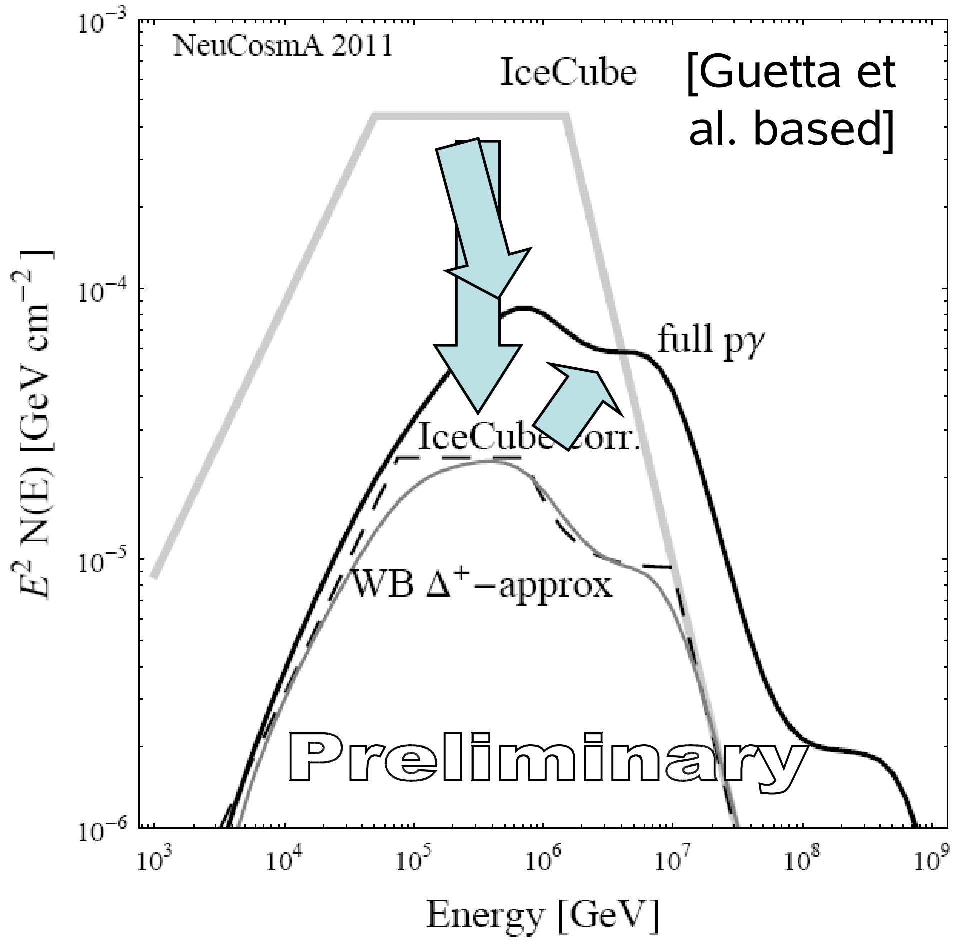


Baerwald, Hümmer, Winter, Phys. Rev. D83 (2011) 067303;
see also: Murase, Nagataki, 2005; Kashti, Waxman, 2005;
Lipari, Lusignoli, Meloni, 2007

Re-analysis of fireball model

- Correction factors from:
  - Cosmological expansion (z)
  - Some crude estimates, e.g. for f_π (frac. of E going pion production)
  - Spectral corrections (compared to choosing the break energy)
  - Neutrinos from pions/muons
- Photohadronics and magnetic field effects change spectral shape
  Baerwald, Hümmer, Winter, PRD83 (2011) 067303
- **Conclusion (preliminary): Fireball flux ~ factor of five lower than expected, with different shape**

(one example)

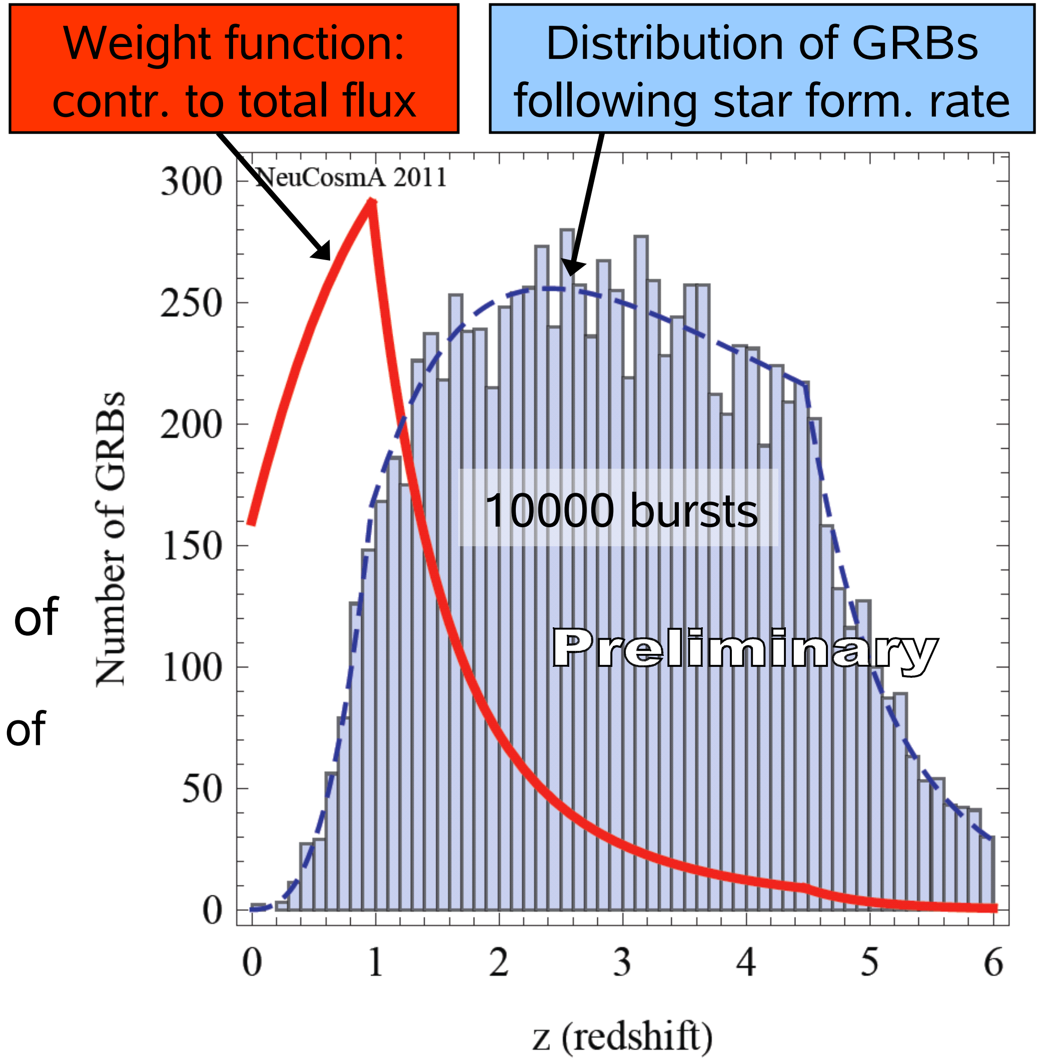


(Hümmer, Baerwald, Winter, in prep.)

Julius-Maximilians-
UNIVERSITÄT
WÜRZBURG

Systematics in aggregated fluxes

- IceCube: Signal from 117 bursts “stacked“ (summed) for current limit **(arXiv:1101.1448)**
  - ➢ Is that sufficient?
- Some (preliminary) results:
  - z ~ 1 “typical“ redshift of a GRB
    - ➢ Flux overestimated if z ~ 2-3 assumed (unless z measured)
  - Peak contribution in a region of low statistics
    - ➢ Probability to be within 20% of the diffuse flux is (roughly)
      - 40% for 100 bursts
      - 50% for 300 bursts
      - 70% for 1000 bursts
      - 95% for 10000 bursts
    - ➢ Need O(1000) bursts for reliable stacking limits!



(Baerwald, Hümmer, Winter, in prep.)

Summary

- Particle production, flavor, and magnetic field effects change the shape of astrophysical neutrino fluxes
  - ➢ Description of the „known“ (particle physics) components should be as accurate as possible for data analysis; e.g. GRBs
- Flavor ratios, though difficult to measure, are interesting because
  - they may be the only way to directly measure B (astrophysics)
  - they are useful for new physics searches (particle physics)
  - they are relatively robust with respect to the cooling and escape processes of the primaries (e, p, γ)
- The flux shape and flavor ratio of a point source can be predicted in a self-consistent way if the astrophysical parameters can be estimated, such as from a multi-messenger observation
  (R: from time variability, B: from energy equipartition, α: from spectral shape)
- Even for point sources searches experiments such as Auger are useful, since they (in principle) test the parameter region relevant for the UHECR in our model