Open Questions in Cosmic Ray Physics: From Astrophysics to Particle Physics

Günter Sigl

1. Introduction and Overview
2. Particle Physics at High Energies
3. Astrophysics

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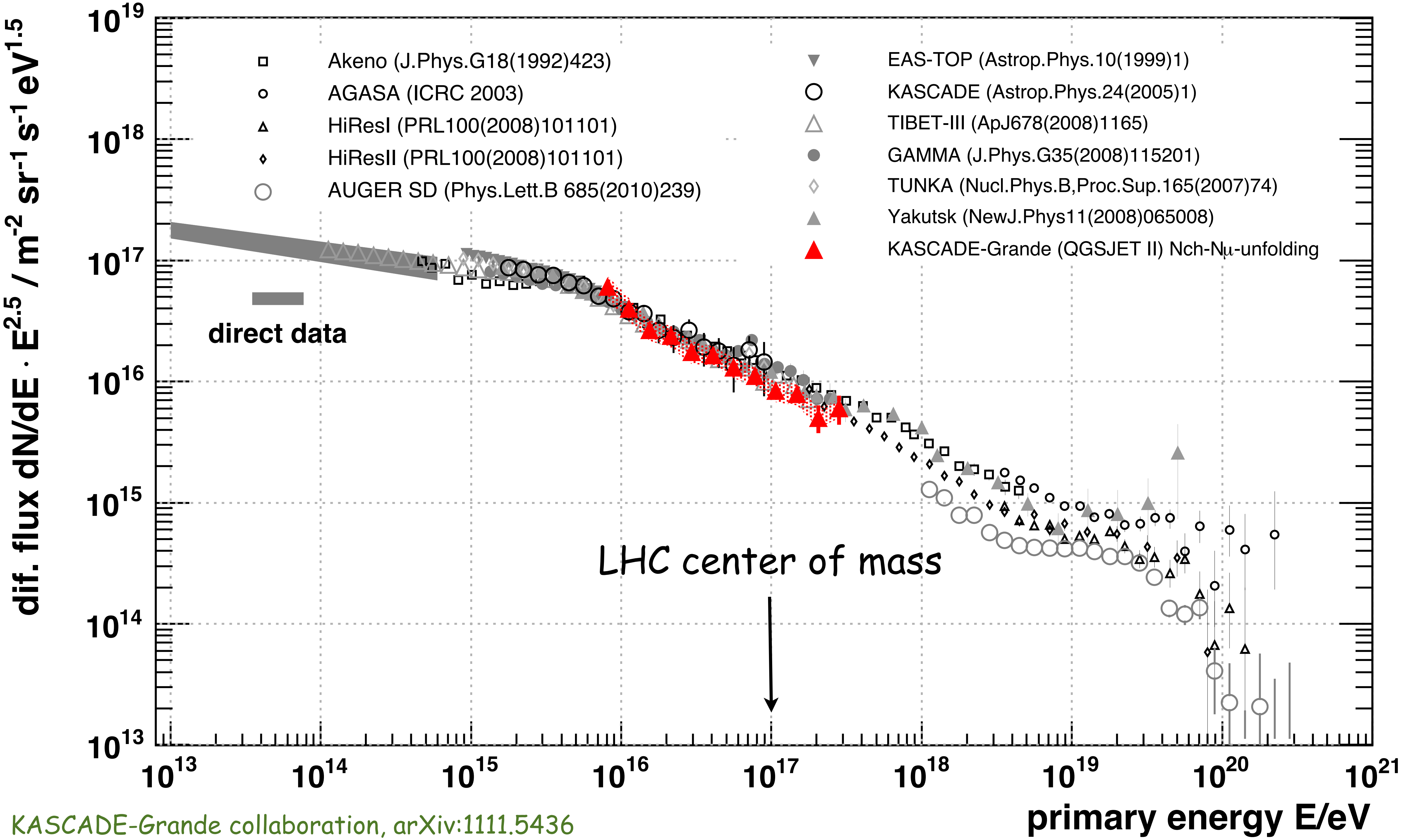
The All Particle Cosmic Ray Spectrum

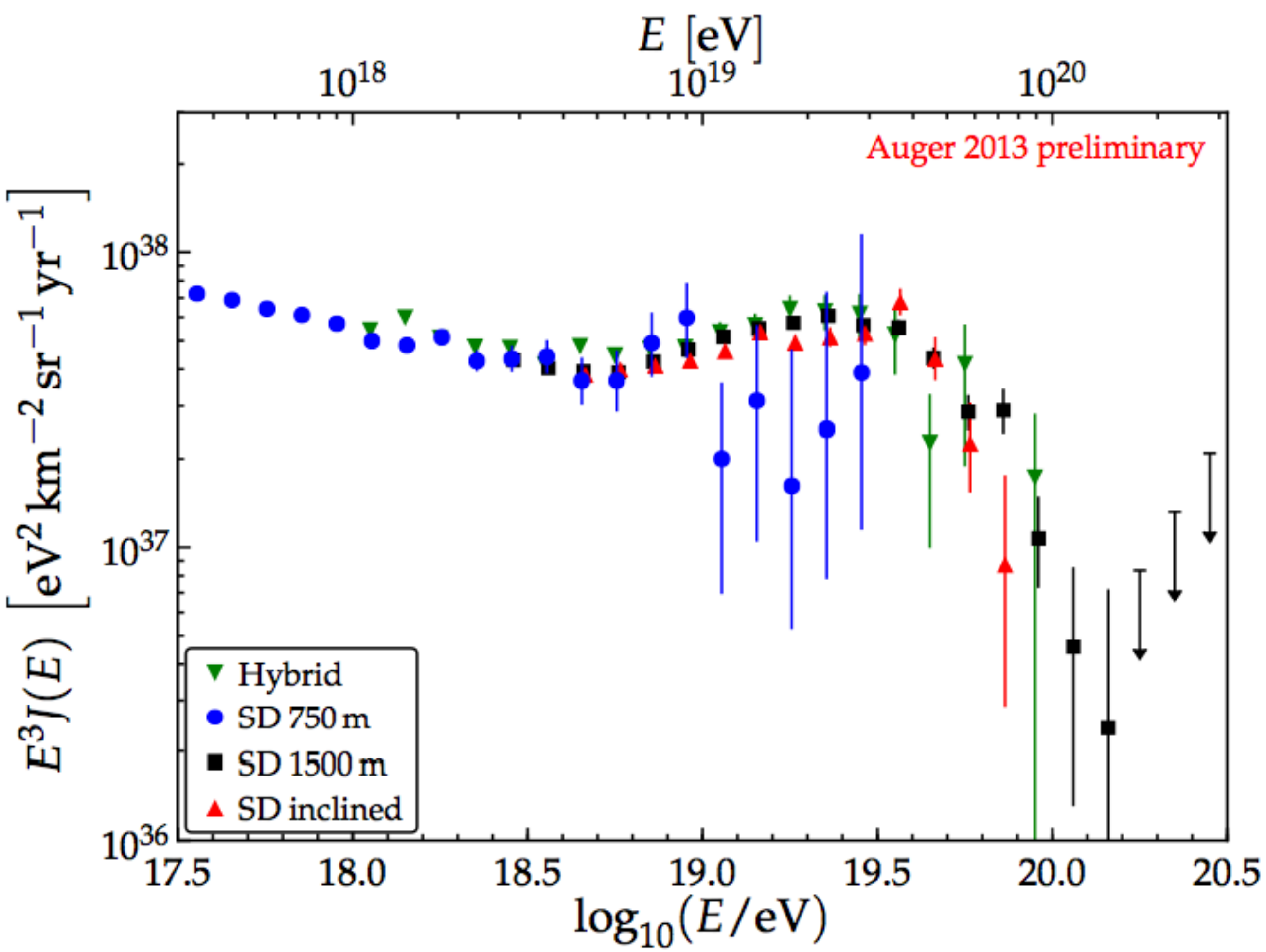


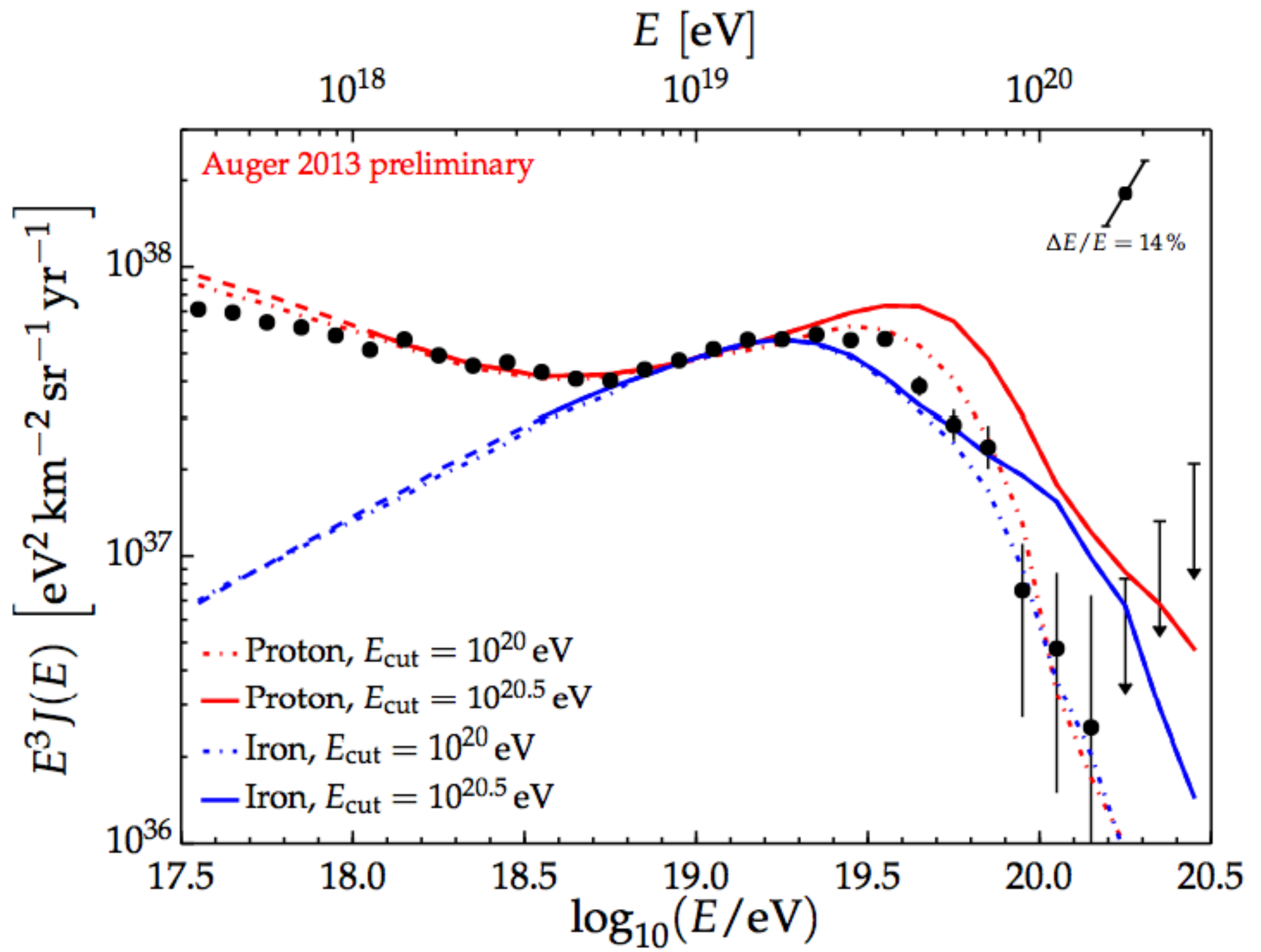
KASCADE-Grande collaboration, arXiv:1111.5436

Pierre Auger Spectra

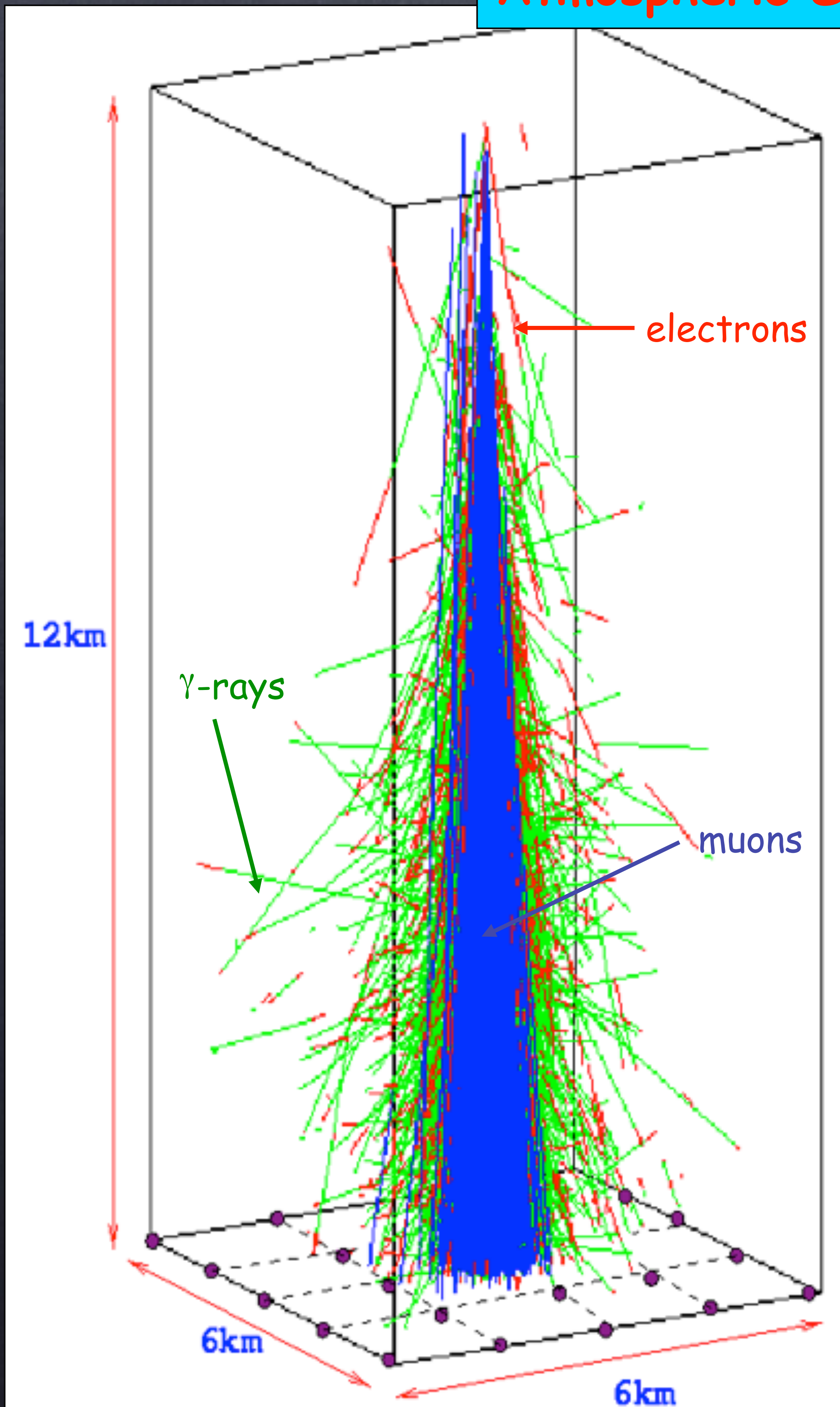
Auger exposure = 31645 km^2 sr yr
up to December 2012

Pierre Auger Collaboration, PRL 101, 061101 (2008)
and Phys.Lett.B 685 (2010) 239
and ICRC 2013, arXiv:1307.5059, higlight talk Letessier-Selvon





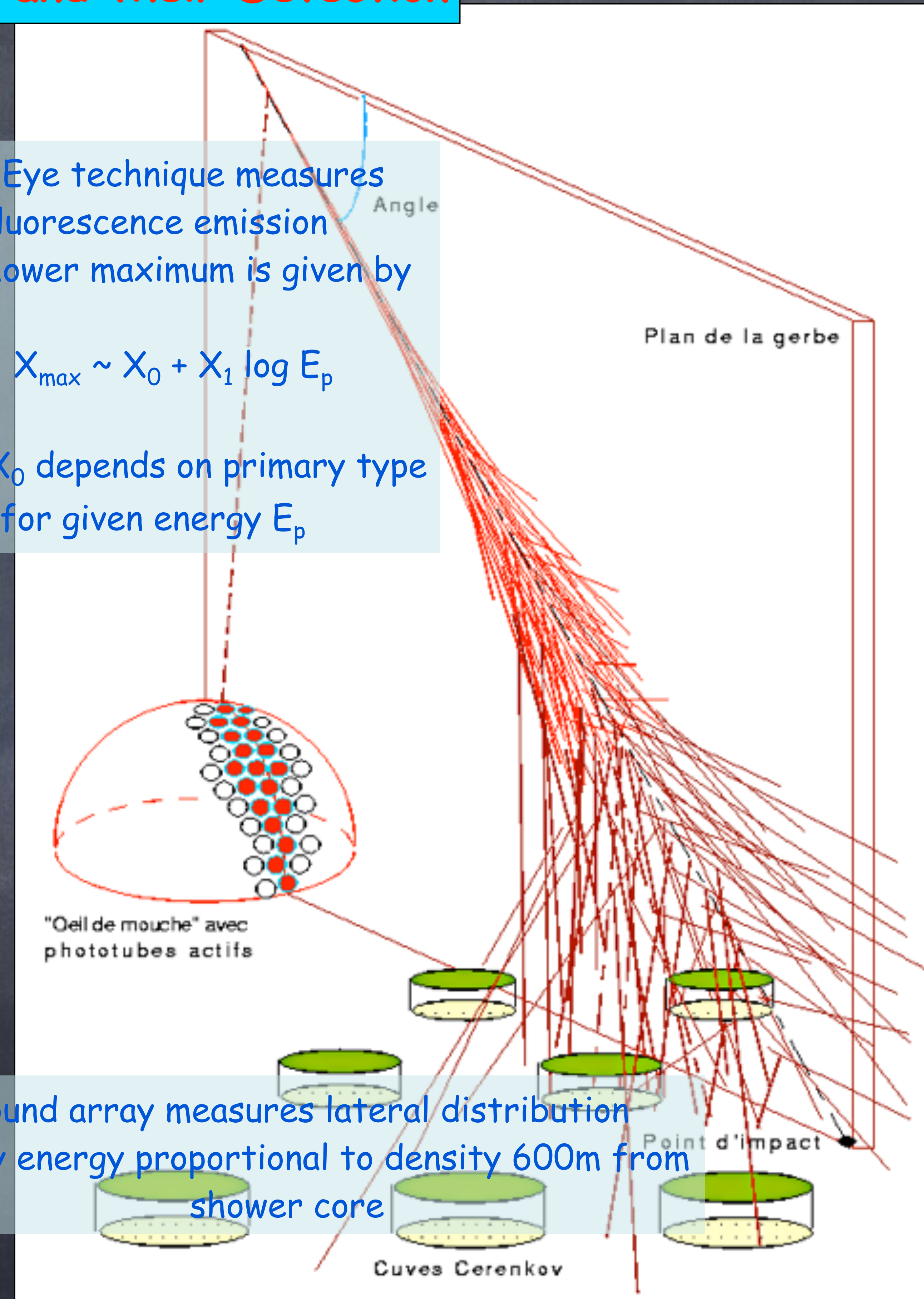
Atmospheric Showers and their Detection



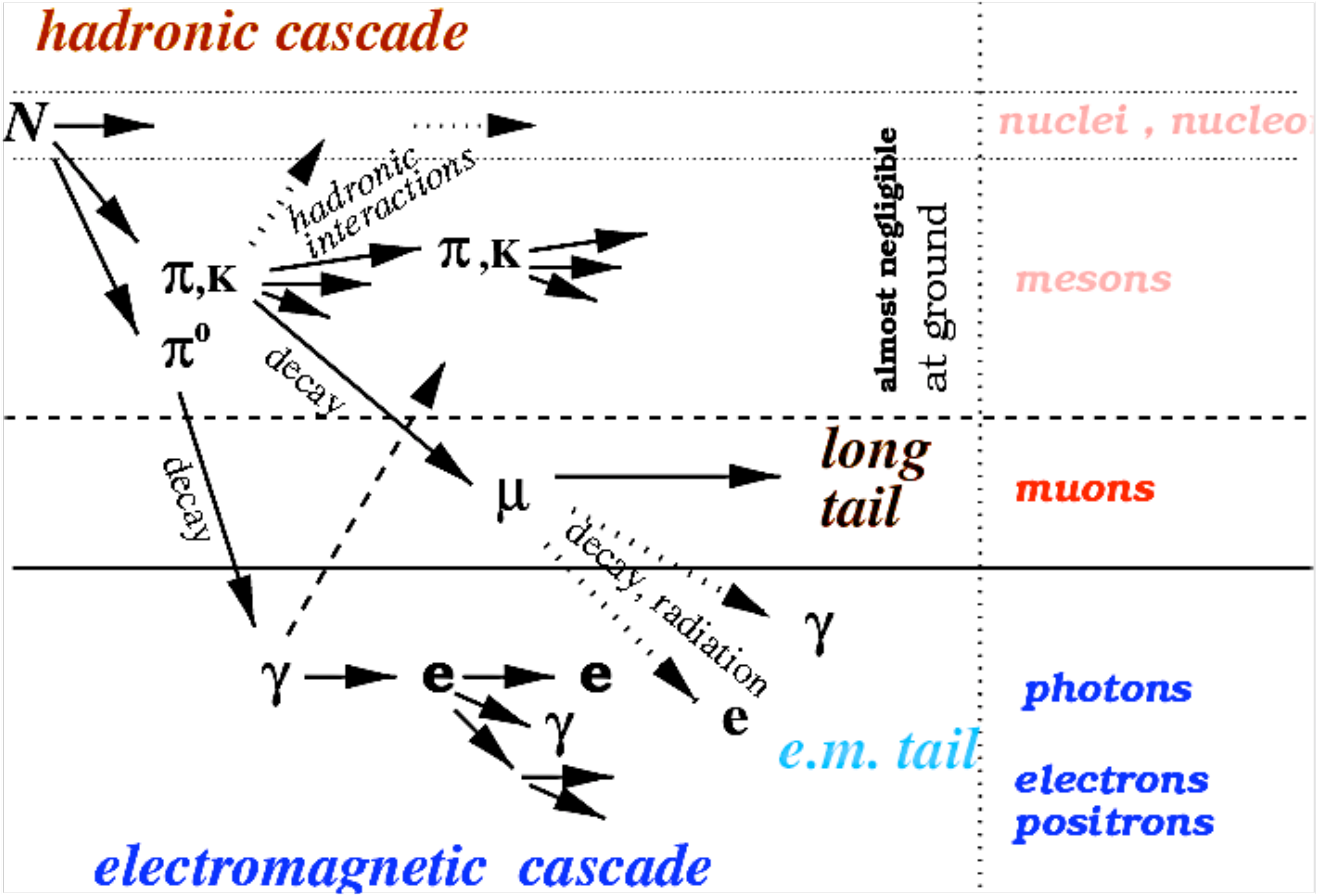
Fly's Eye technique measures fluorescence emission
The shower maximum is given by

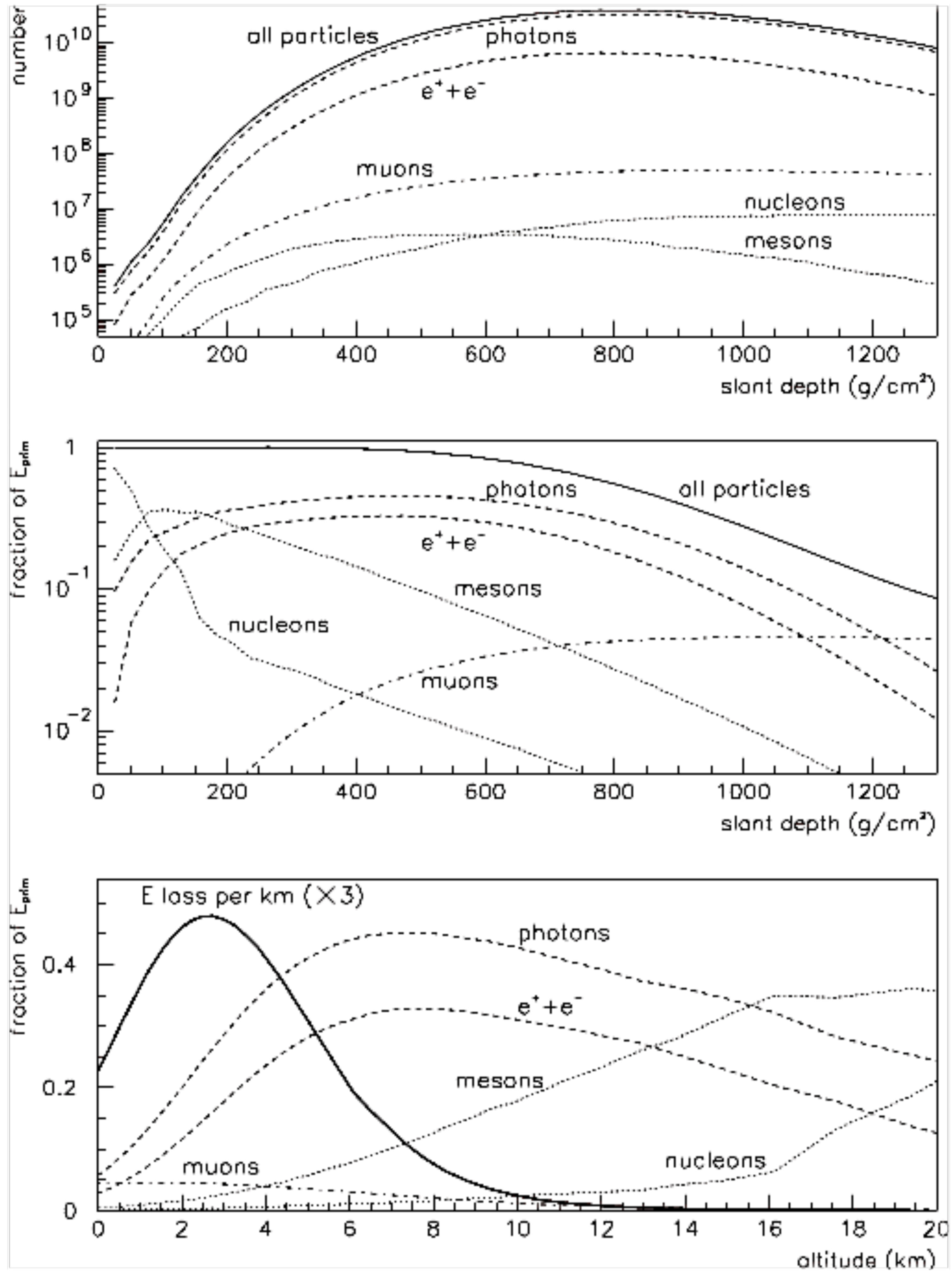
$$X_{max} \sim X_0 + X_1 \log E_p$$

where X_0 depends on primary type for given energy E_p



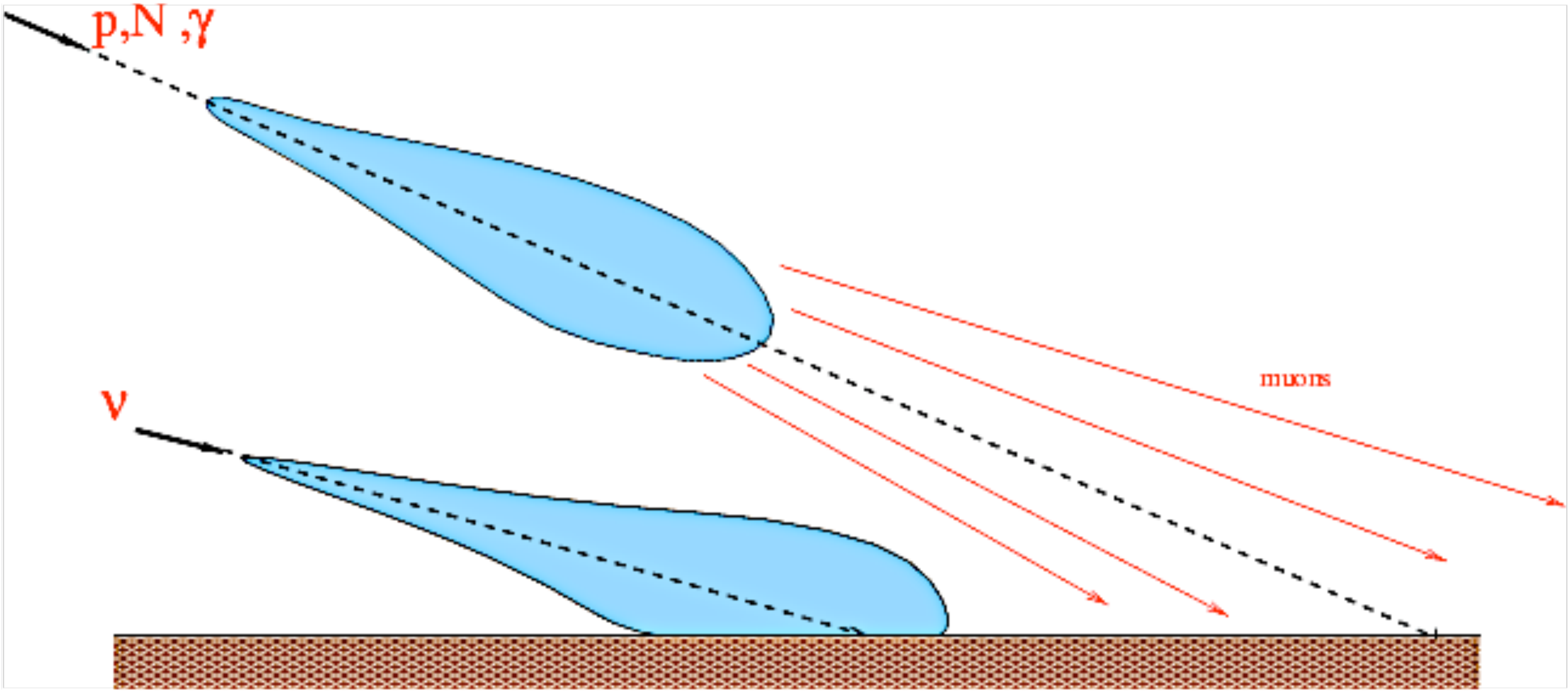
Ground array measures lateral distribution
Primary energy proportional to density 600m from shower core

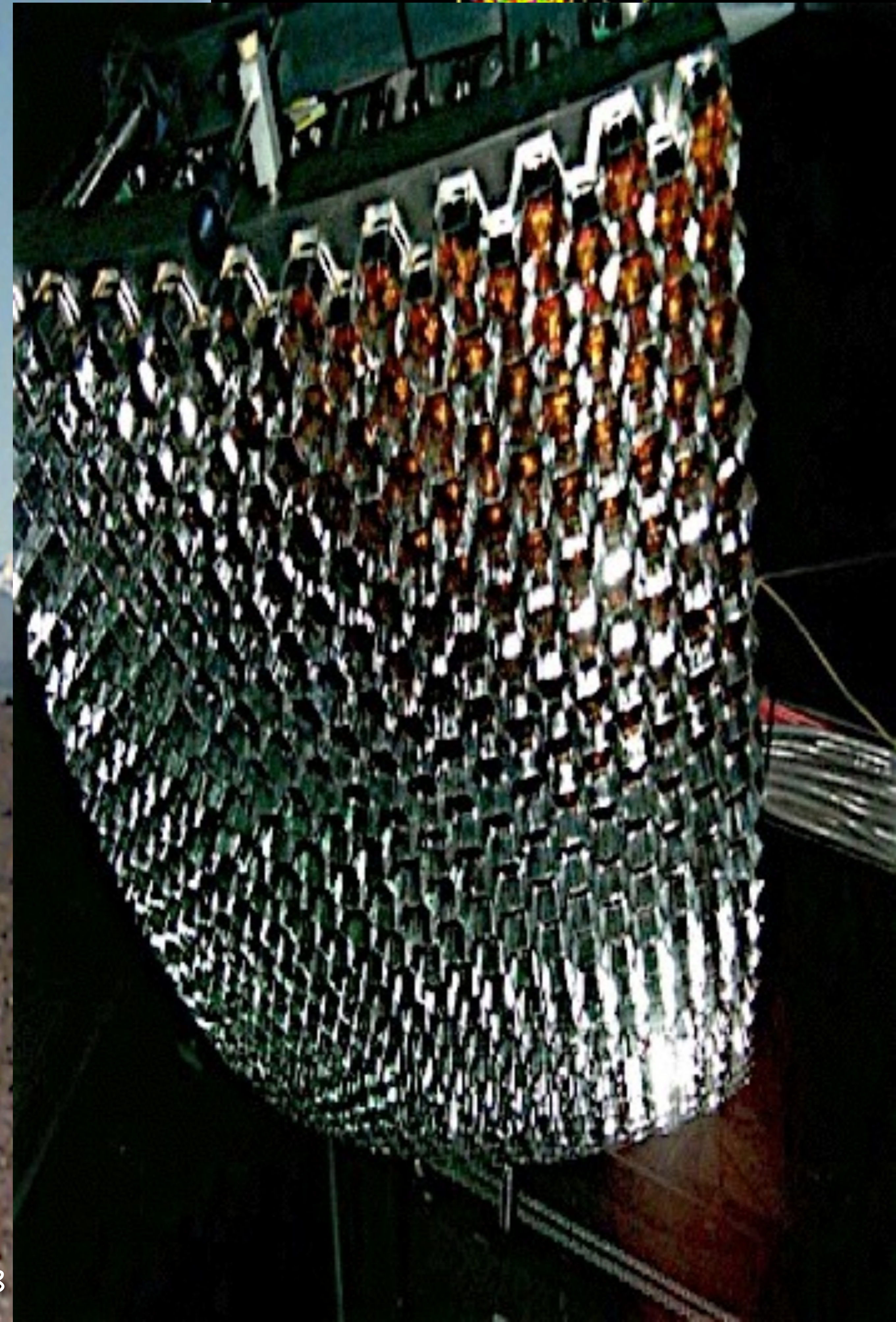

hadronic cascade
N
nuclei , nucleo
hadronic interactions
π,K
π ,K
π⁰
mesons
almost negligible at ground
decay
decay
μ
long tail
muons
decay, radiation
γ
γ
e
e
γ
e
e.m. tail
photons
electrons positrons
electromagnetic cascade



number
all particles
photons
e⁺+e⁻
muons
nucleons
mesons
slant depth (g/cm²)
fraction of E_prim
all particles
photons
e⁺+e⁻
mesons
nucleons
muons
slant depth (g/cm²)
fraction of E_prim
E loss per km (X3)
photons
e⁺+e⁻
mesons
muons
nucleons
altitude (km)

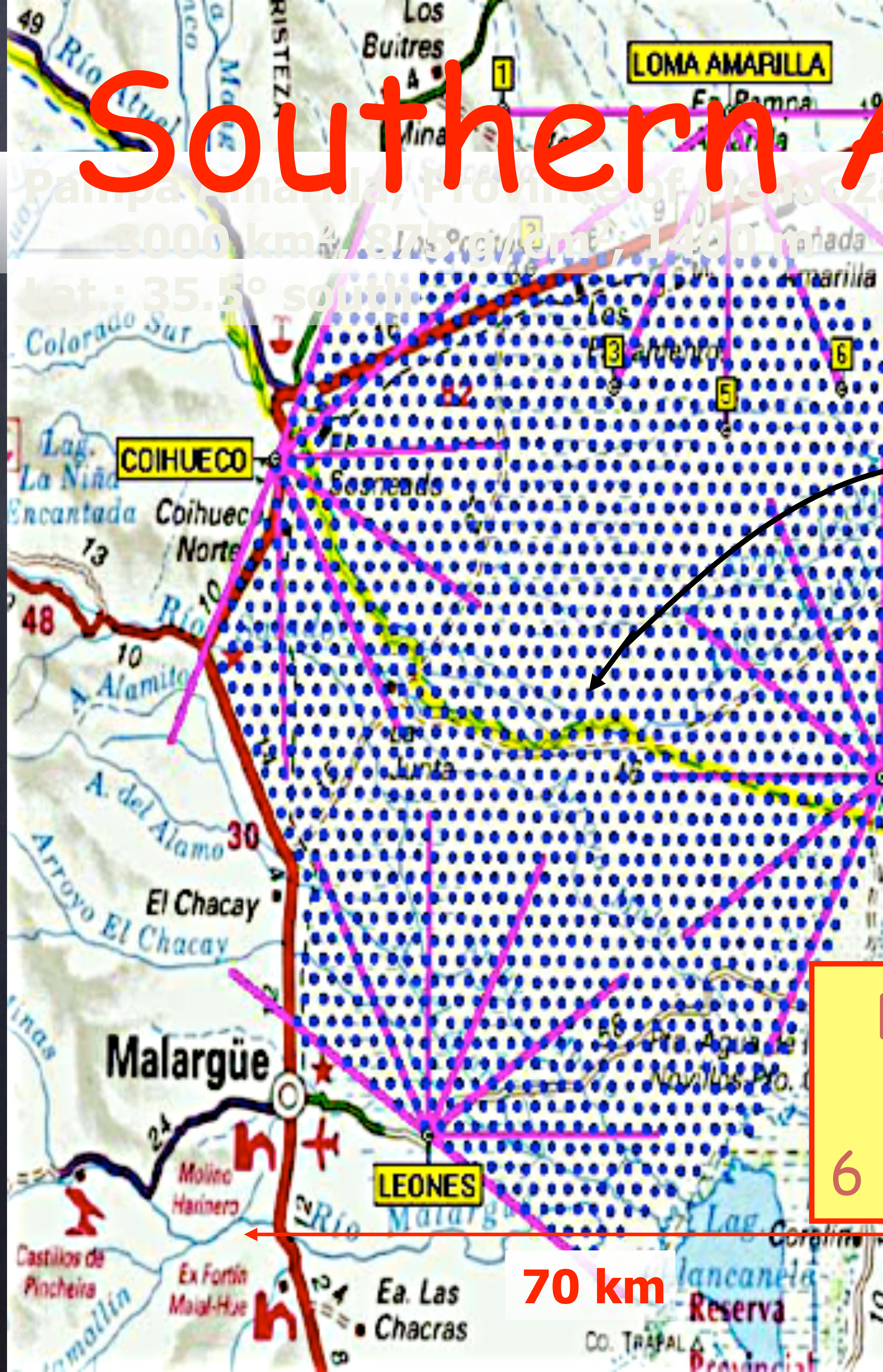
Cosmic ray versus neutrino induced air showers

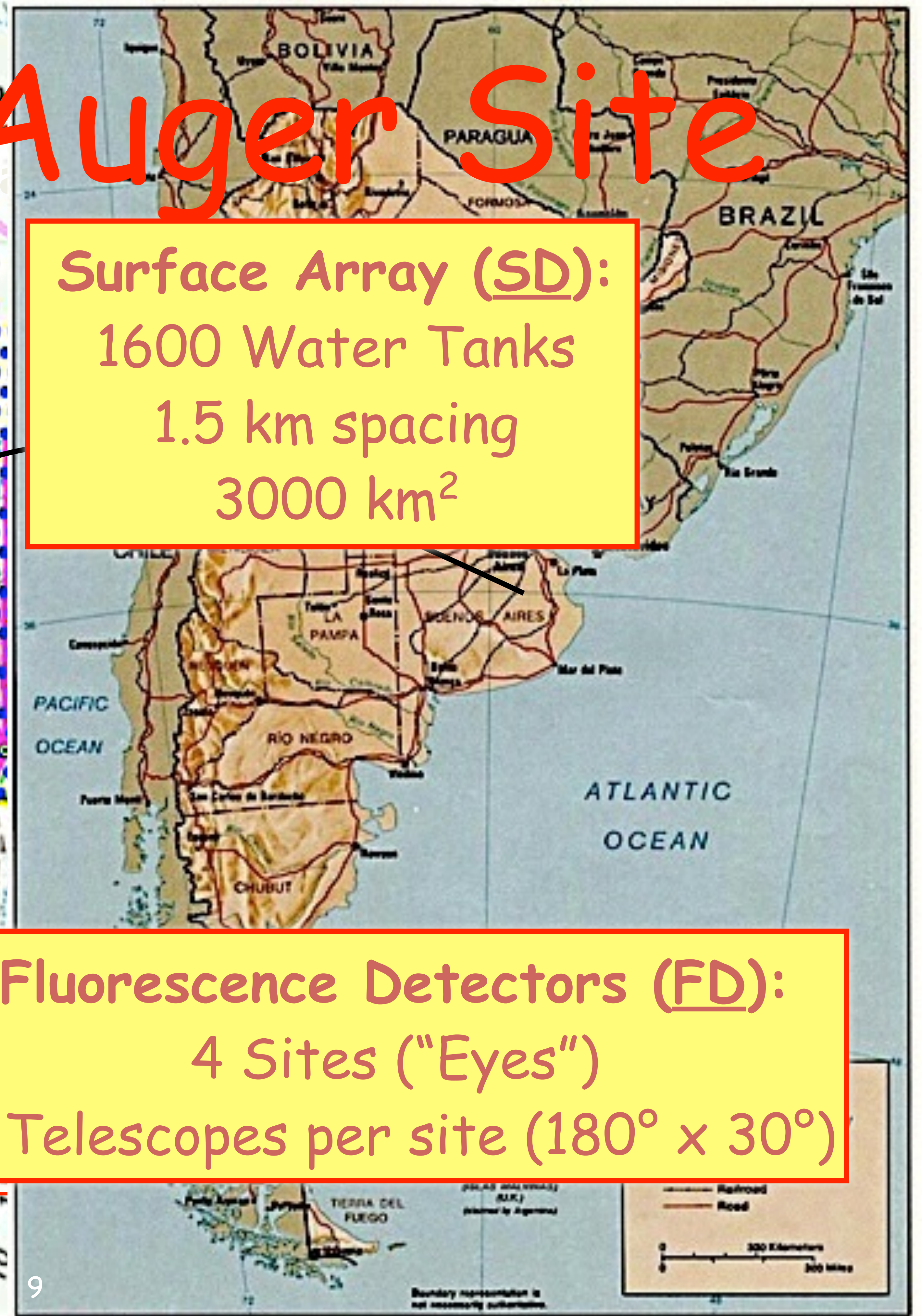




Southern Auger Site

Pampa Amarilla, Province of Mendoza
3000 km², 875 g/cm², 1400 m
Lat.: 35.5° south





Surface Array (SD):
1600 Water Tanks
1.5 km spacing
3000 km^2

Fluorescence Detectors (FD):
4 Sites ("Eyes")
6 Telescopes per site (180° x 30°)

The Ultra-High Energy Cosmic Ray Mystery consists of (at least) Four Interrelated Challenges

1.) electromagnetically or strongly interacting particles above 10^{20} eV loose energy within less than about 50 Mpc.

2.) in most conventional scenarios exceptionally powerful acceleration sources within that distance are needed.

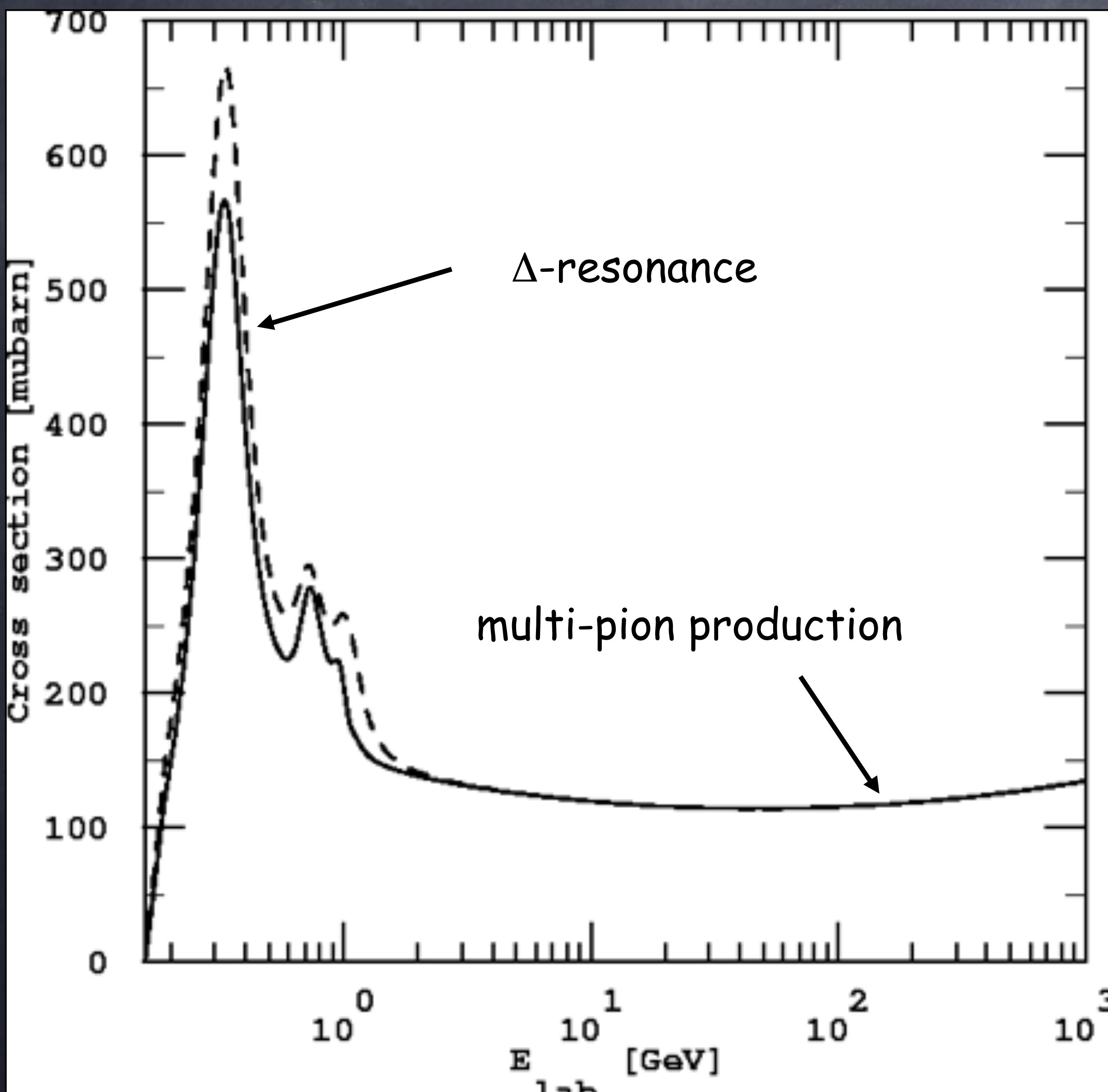
3.) The observed distribution does not yet reveal unambiguously the sources, although there are hints of correlations with local large scale structure

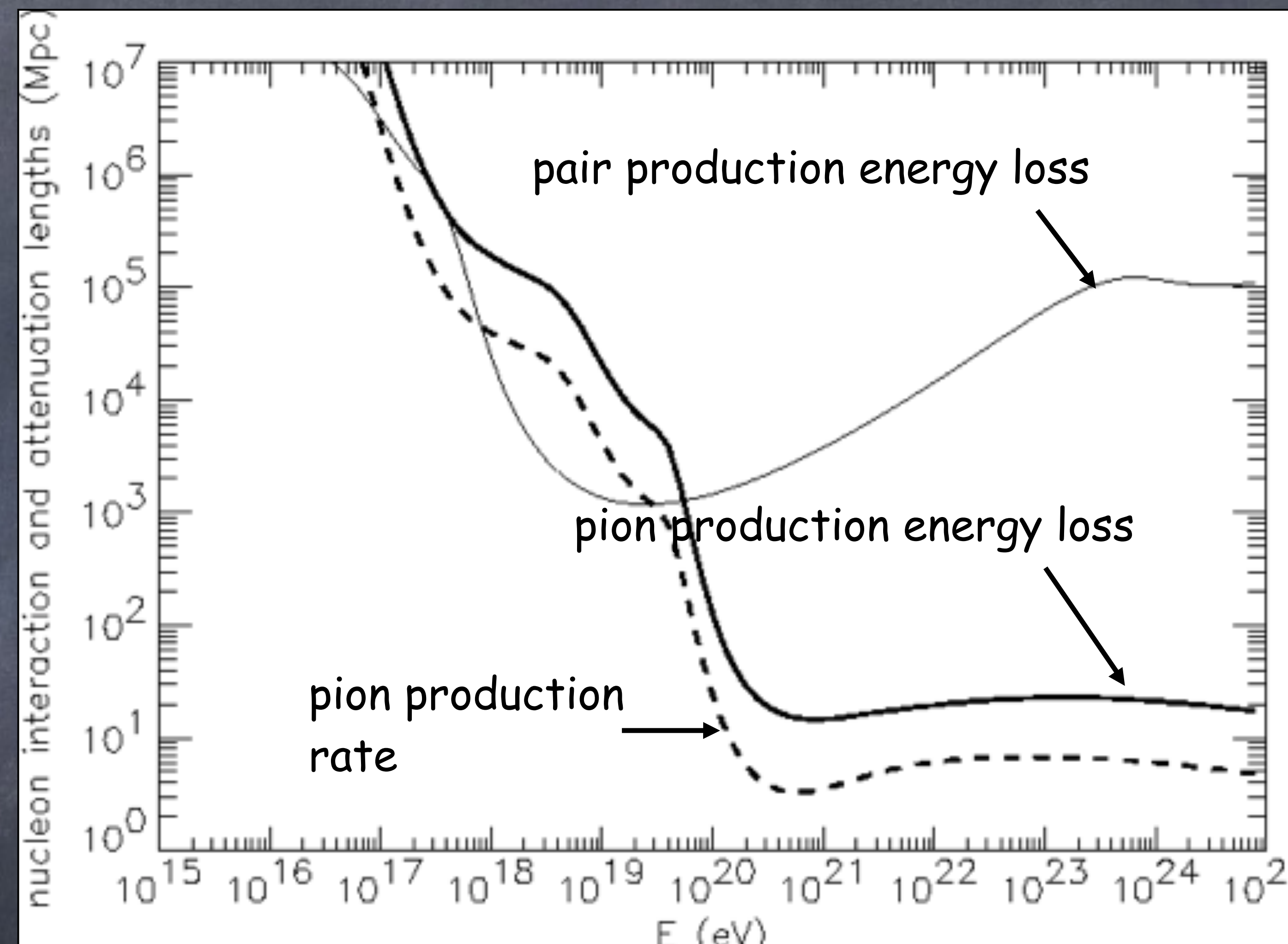
4.) The observed mass composition may become heavy toward highest energies, but no completely clear picture yet between experiments and air shower models

The Greisen-Zatsepin-Kuzmin (GZK) effect

Nucleons can produce pions on the cosmic microwave background

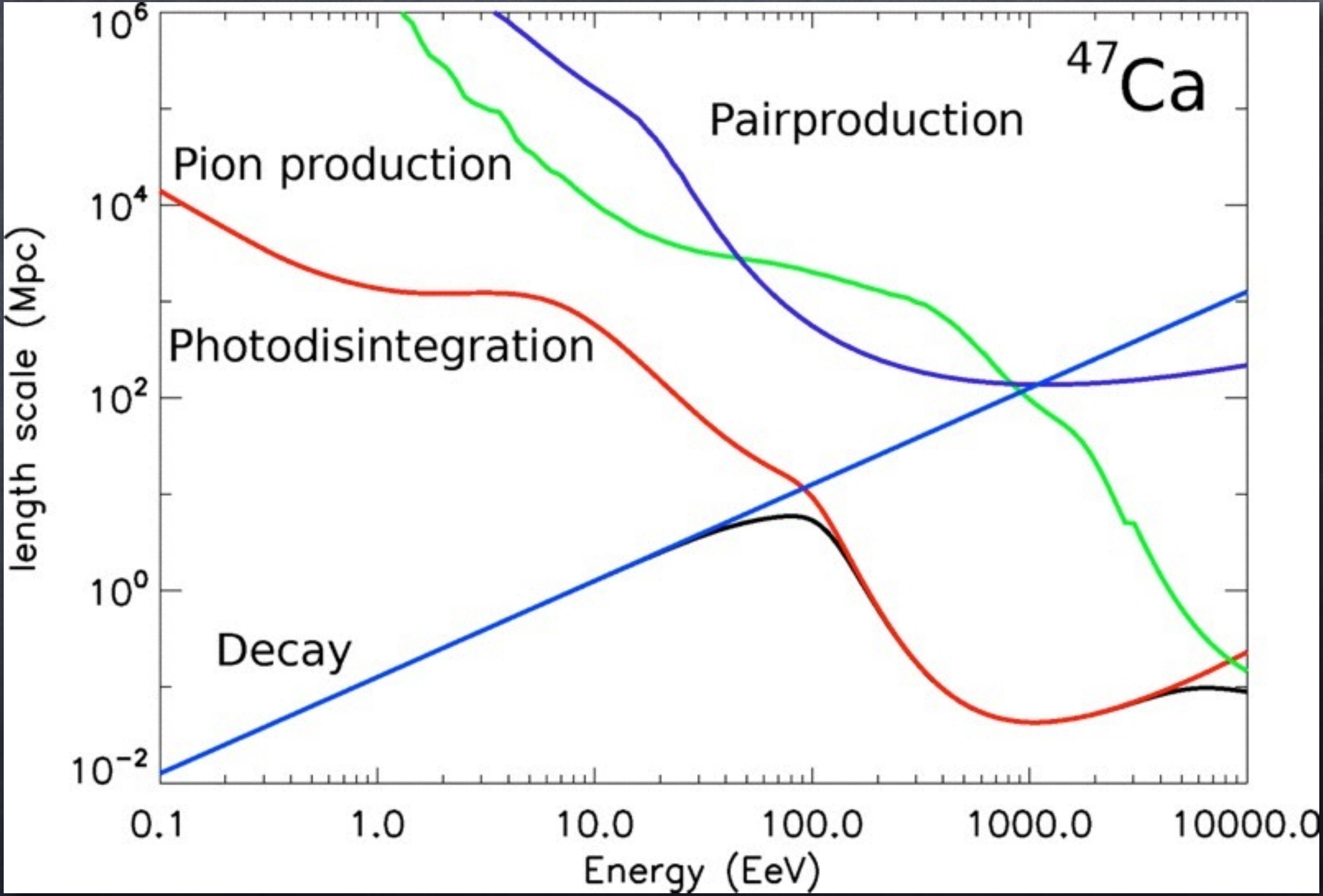
$$E_{\rm th} = \frac{2 m_N m_\pi + m_\pi^2}{4\varepsilon} \simeq 4 \times 10^{19}\,{\rm eV}$$



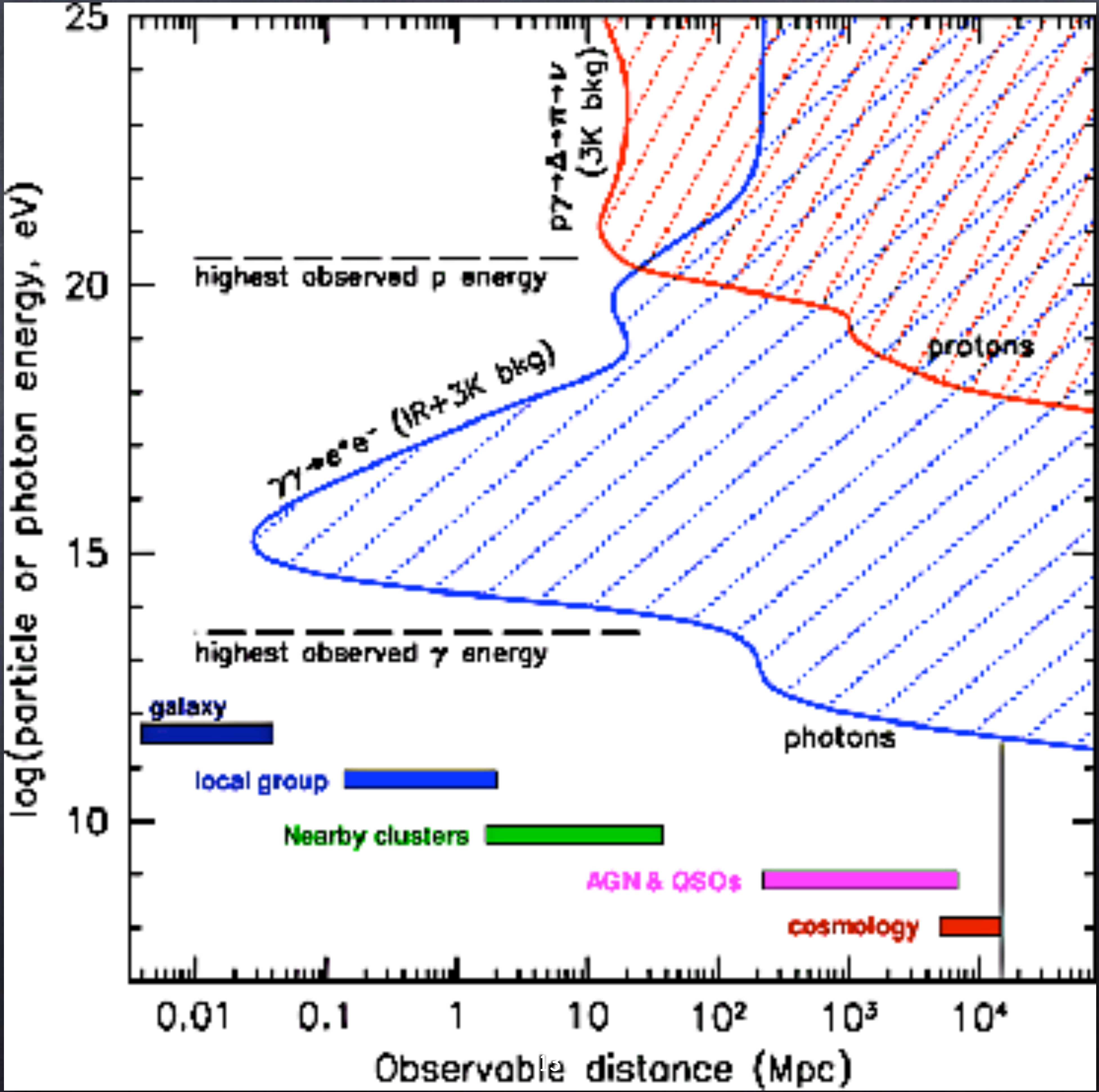


sources must be in cosmological backyard
Only Lorentz symmetry breaking at $\Gamma > 10^{11}$ could avoid this conclusion.

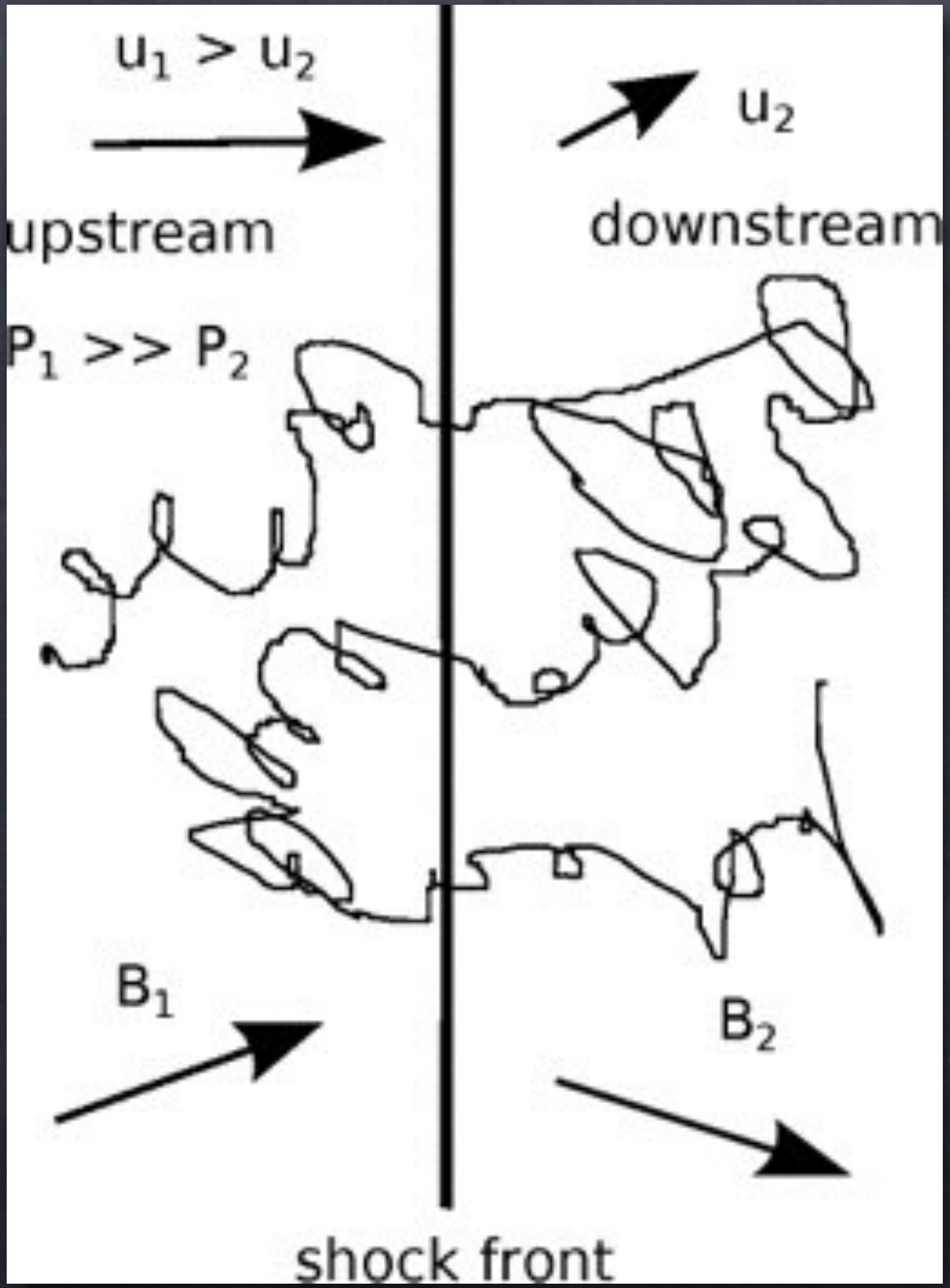
Length scales for relevant processes of a typical heavy nucleus

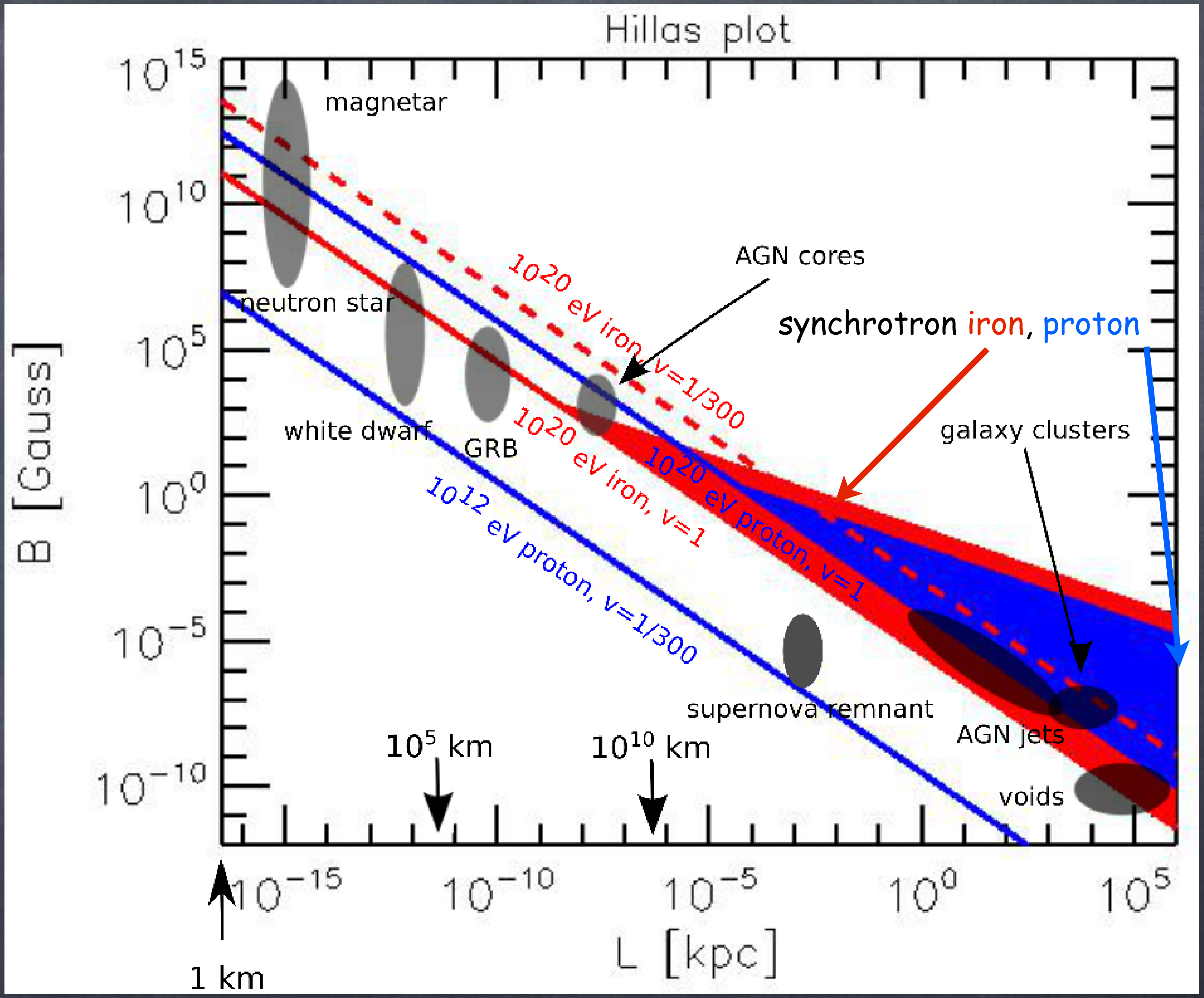


Interaction Horizons



1^{st} Order Fermi Shock Acceleration





Fractional energy gain per shock crossing ~ $u_1 - u_2$ on a time scale r_L/u_2 .

Together with downstream losses this leads to a spectrum E^{-q} with $q > 2$ typically.

Confinement, gyroradius < shock size, and energy loss times define maximal energy

Some general Requirements for Sources

Accelerating particles of charge eZ to energy E_{max} requires induction $\varepsilon > E_{max}/eZ$. With $Z_0 \sim 100\Omega$ the vacuum impedance, this requires dissipation of minimum power of

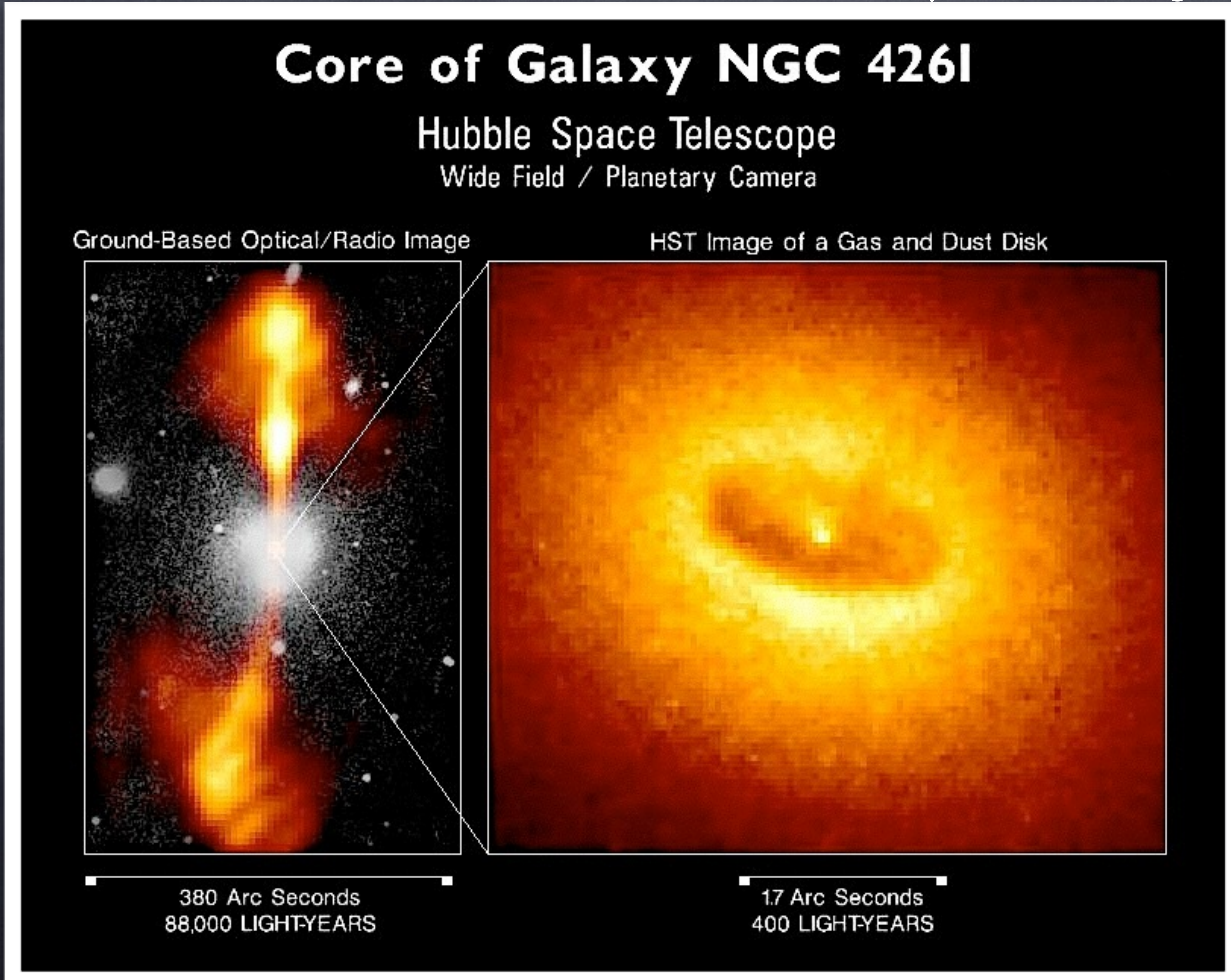
$$L_{\min} \sim \frac{\epsilon^2}{Z_0} \simeq 10^{45}\, Z^{-2} \left(\frac{E_{\max}}{10^{20}\,\mathrm{eV}}\right)^2 \,\mathrm{erg\,s^{-1}}$$

This „Poynting" luminosity can also be obtained from $L_{min} \sim (BR)^2$ where BR is given by the „Hillas criterium":

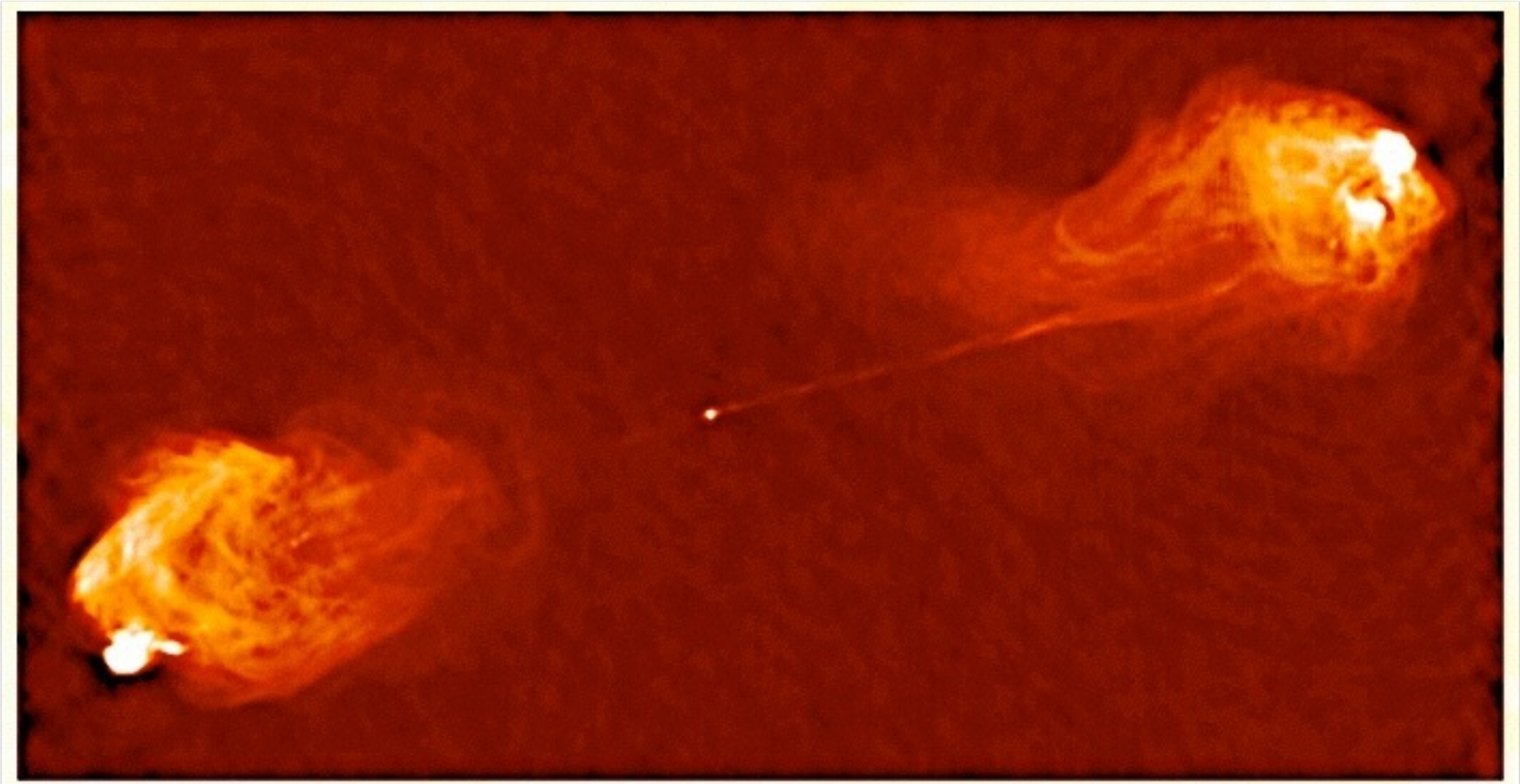
$$BR > 3 \times 10^{17}\,\Gamma^{-1} \left(\frac{E_{\max}/Z}{10^{20}\,\mathrm{eV}}\right) \mathrm{Gauss\,cm}$$

where Γ is a possible beaming factor.
If most of this goes into electromagnetic channel, only AGNs and maybe gamma-ray bursts could be consistent with this.

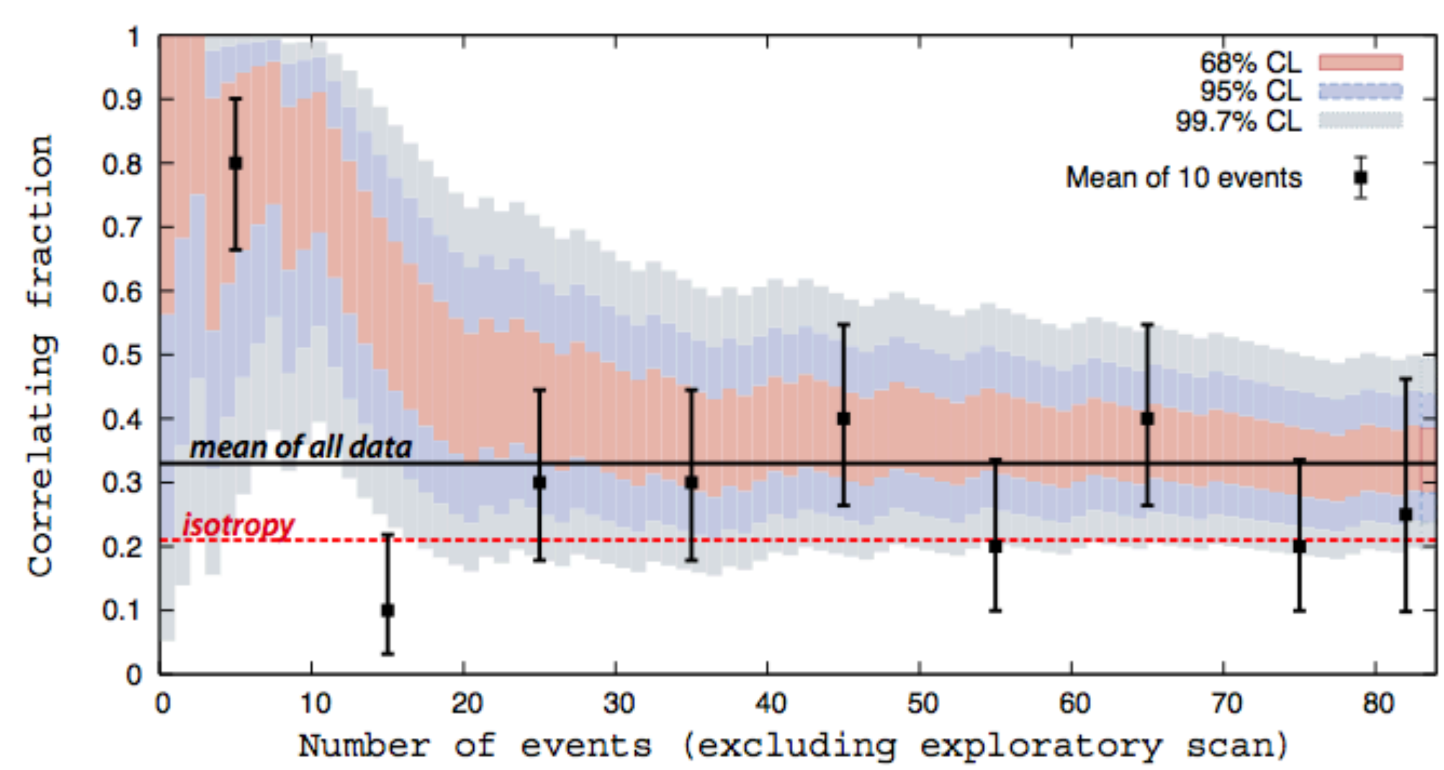
A possible acceleration site associated with shocks in hot spots of active galaxies

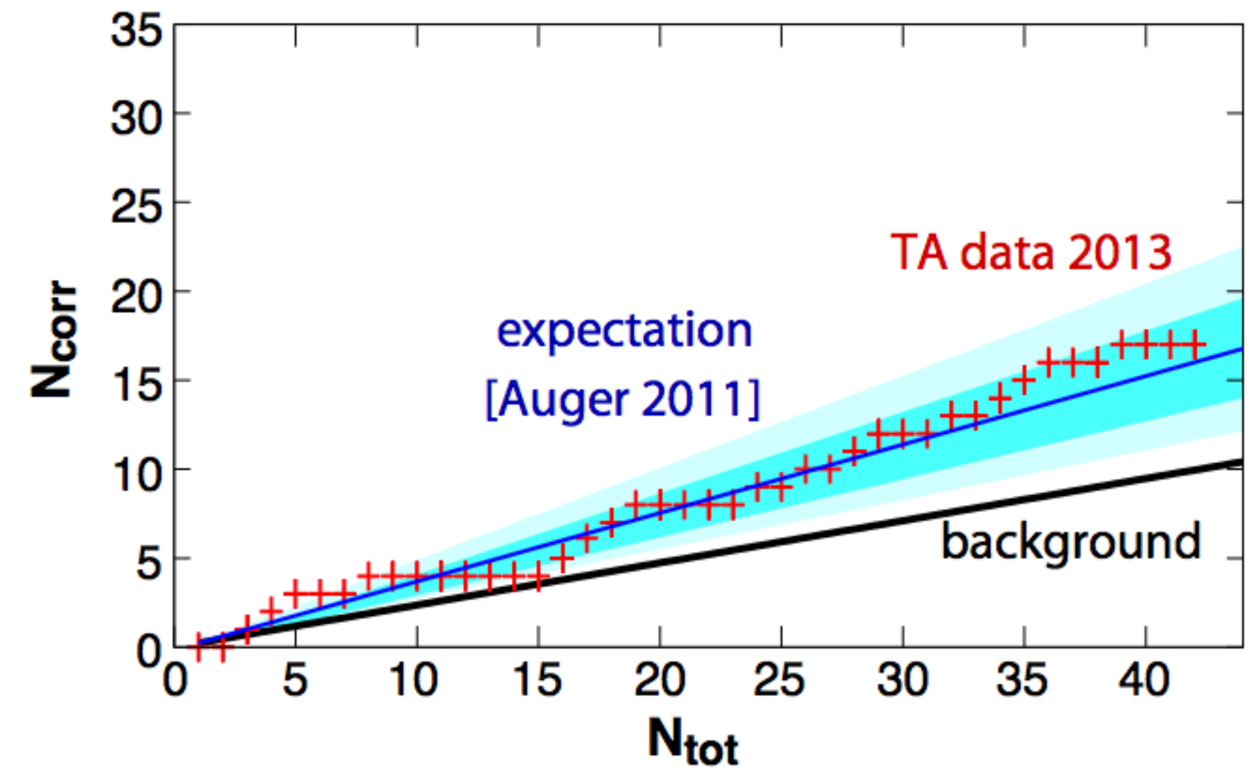


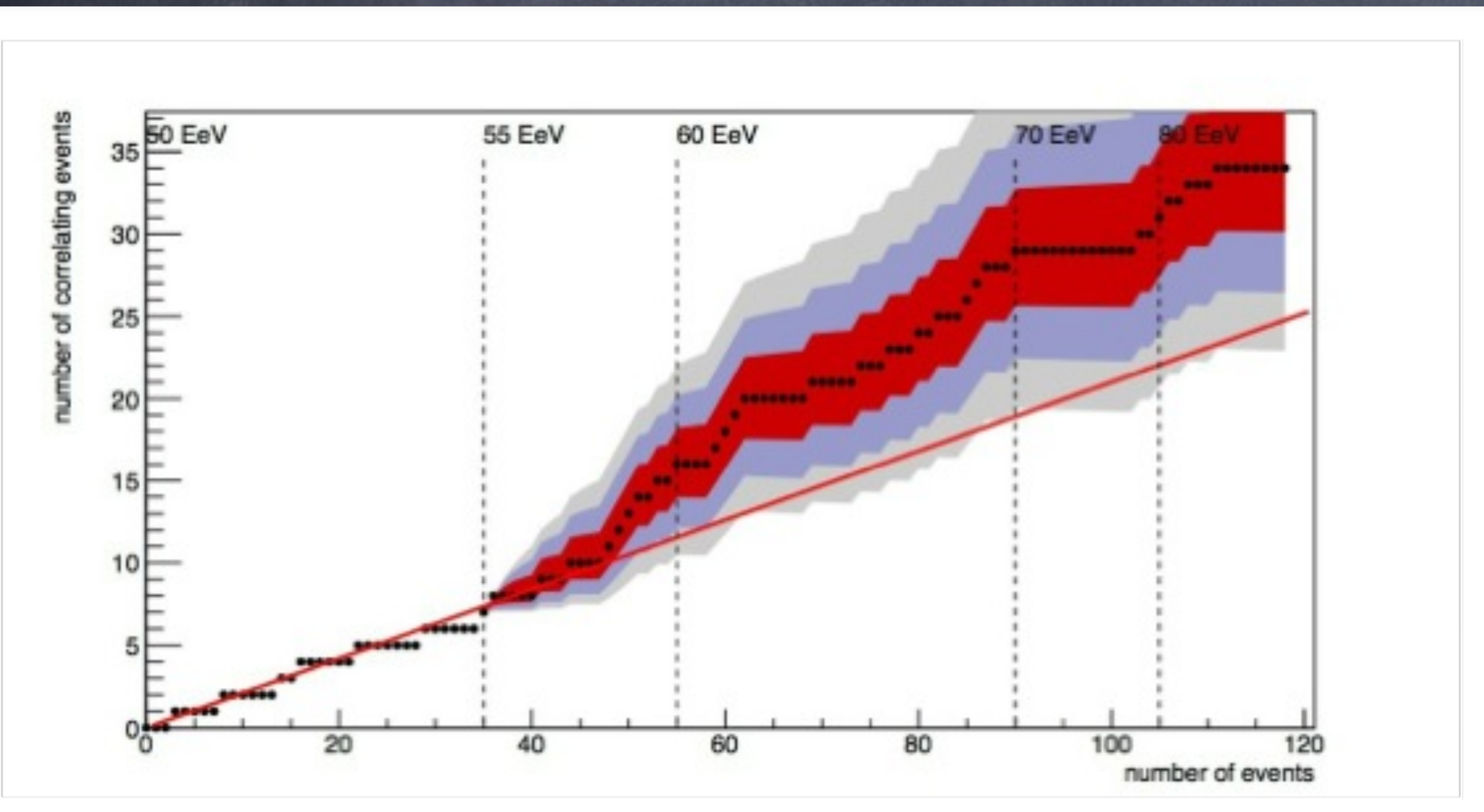
Or Cygnus A



Status of Large Scale UHECR Anisotropy

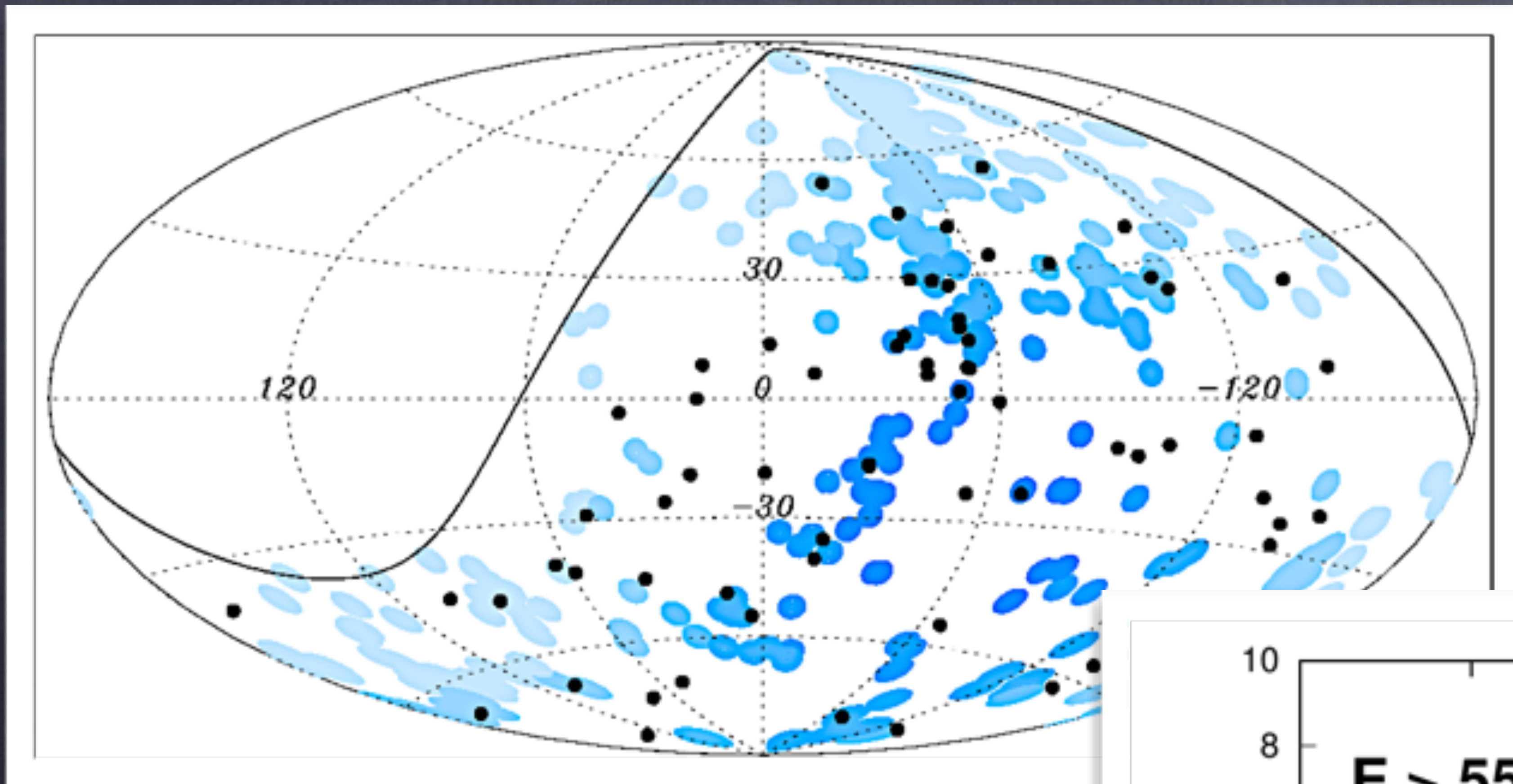






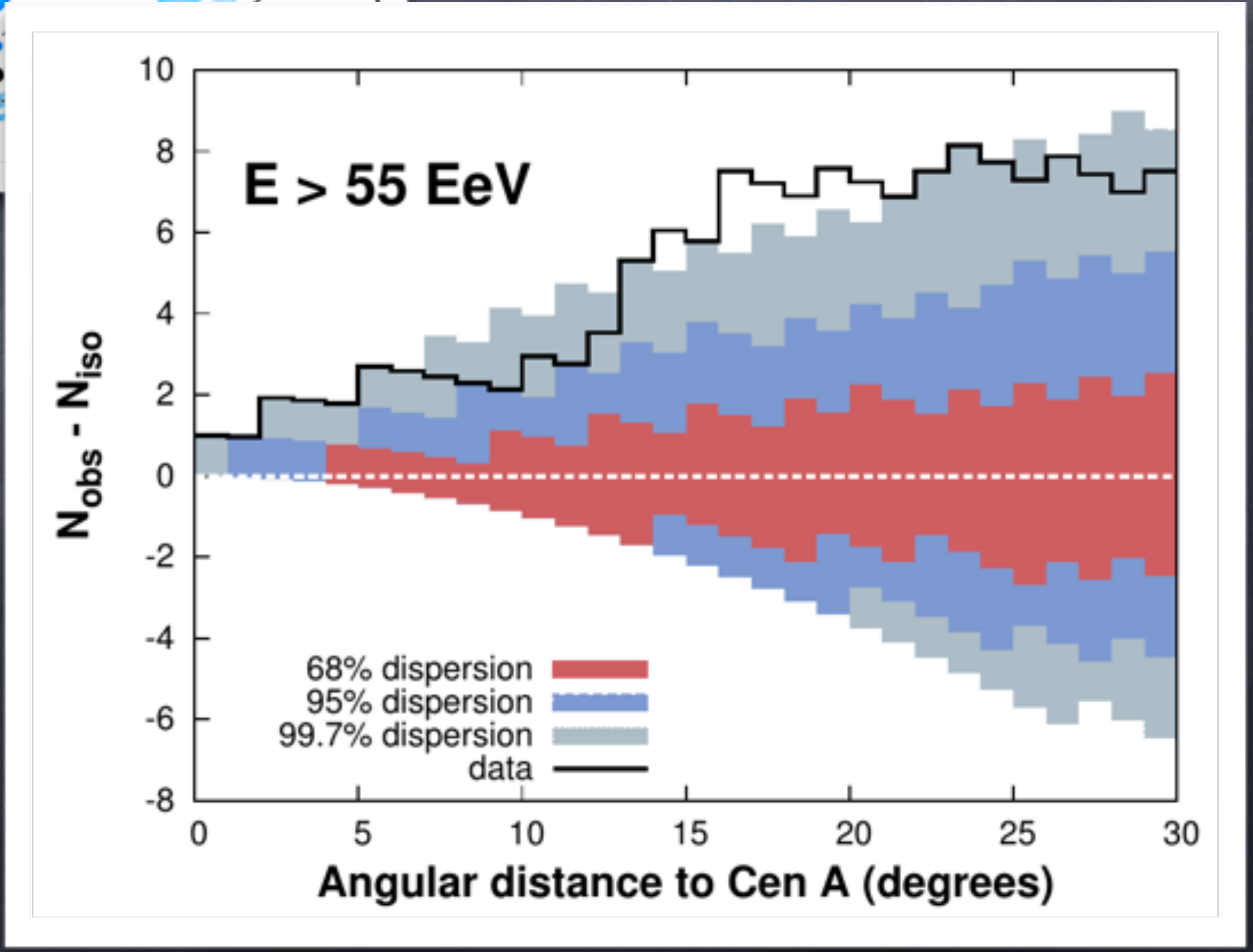
Kampert and Tinyakov,arXiv:1405.0575

Centaurus A is a UHECR source candidate



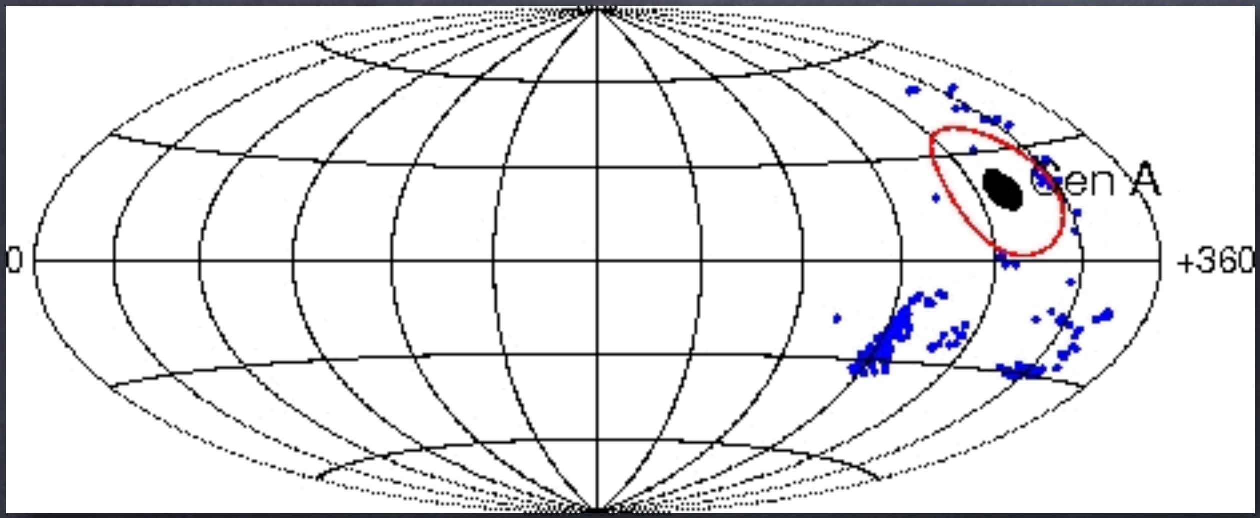
Pierre Auger Collaboration, Astropart.Phys. 34 (2010) 314

Pierre Auger sees an excess in the direction of Centaurus A above 55 EeV

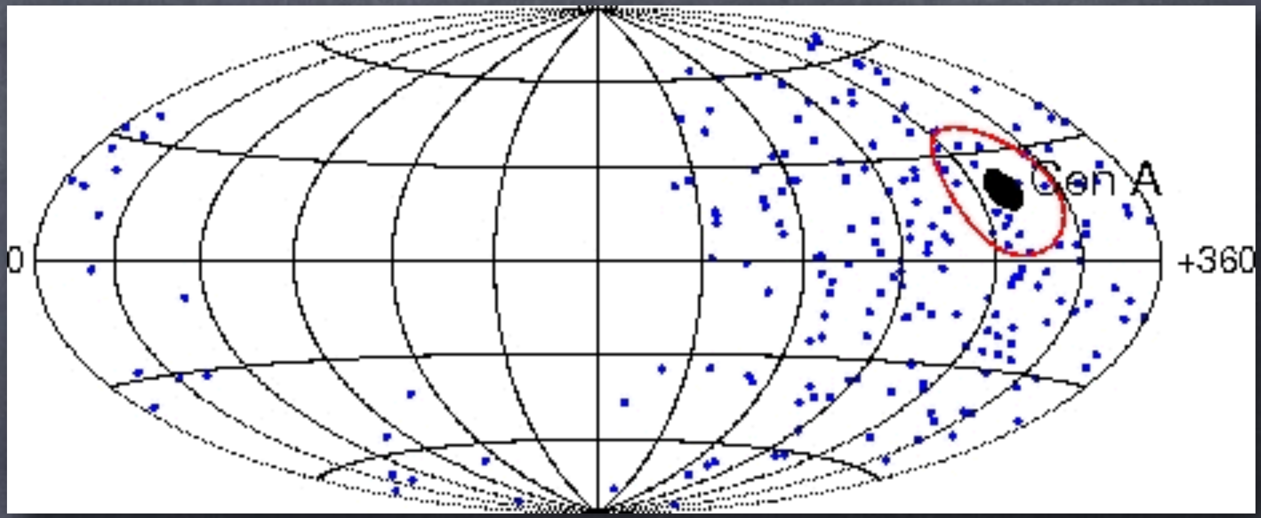


Kampert and Tinyakov,arXiv:1405.0575

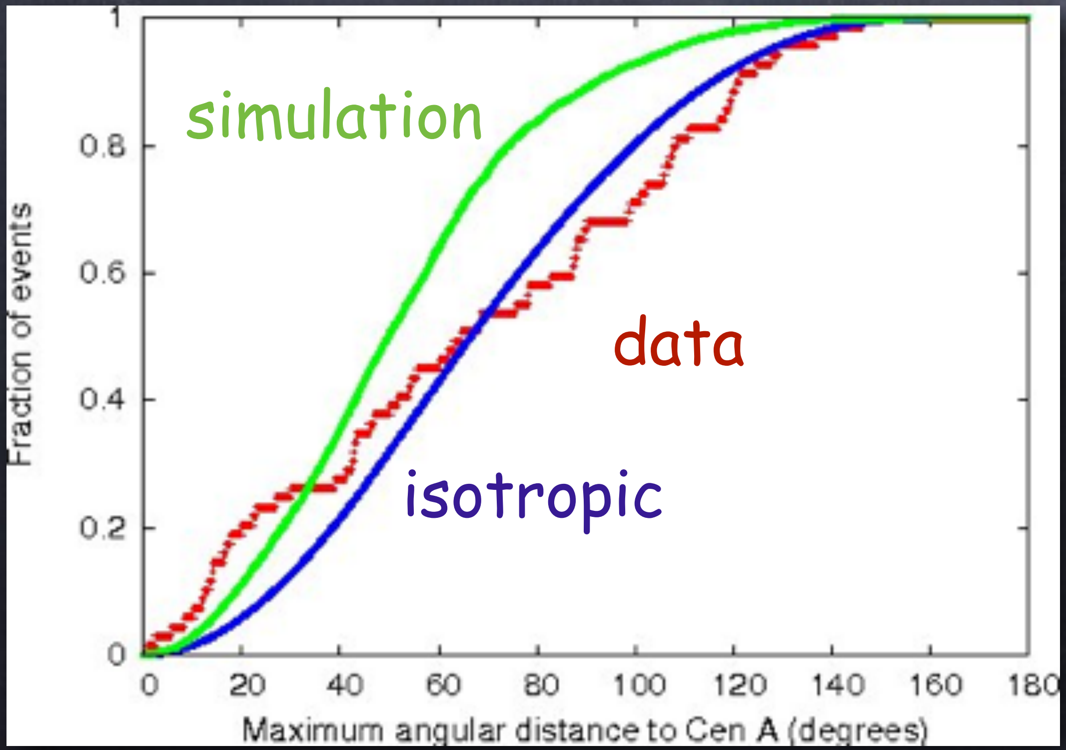
But even for iron primaries Centaurus A can not be the only UHECR source



Iron Image of Cen A in the Prouza-Smida Galactic magnetic field model

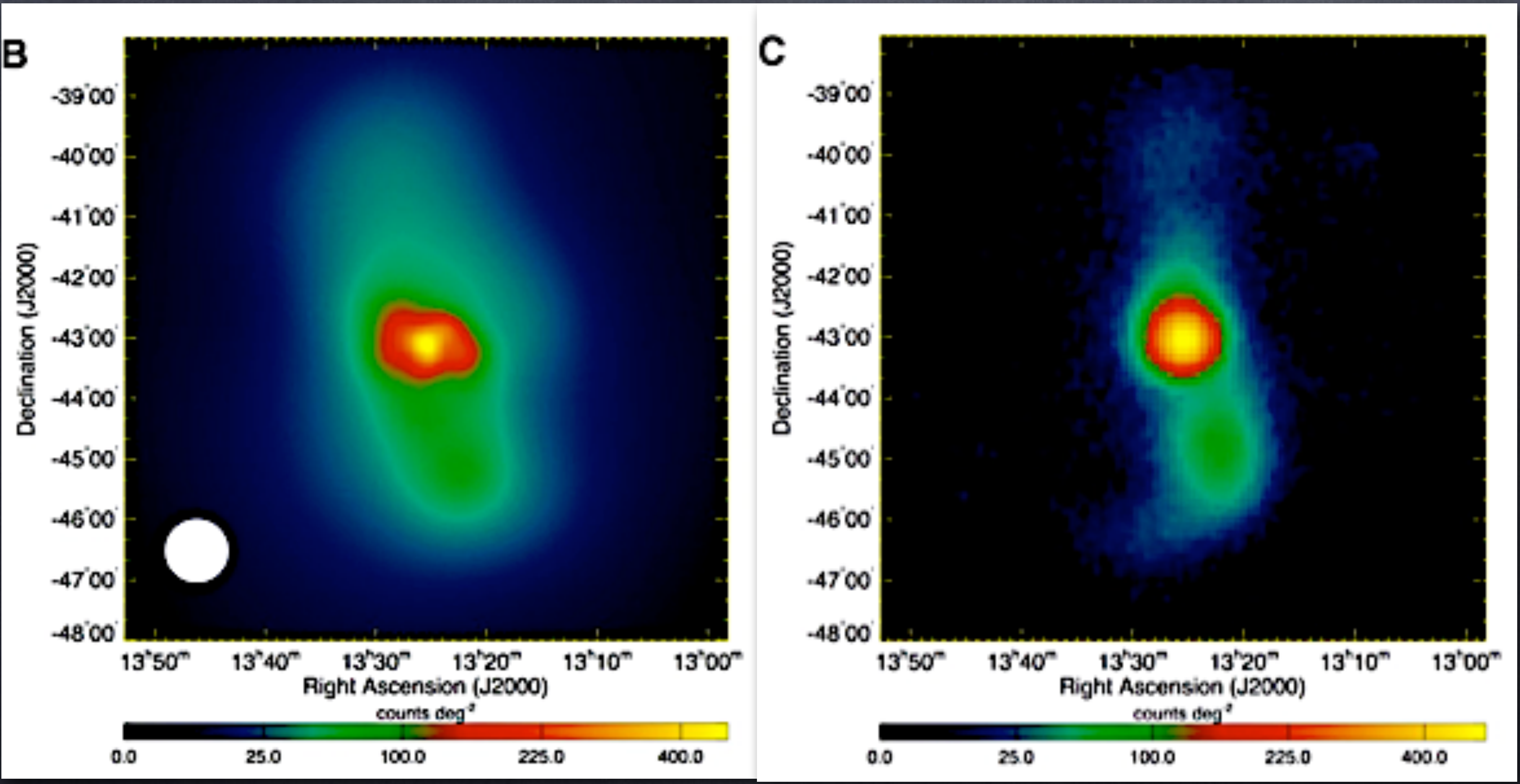


Including an extreme choice for the turbulent Galactic field component with strength 10 μG, coherence length 50 pc, 10 kpc halo extension



Giacinti, Kachelriess, Semikoz, Sigl, Astropart.Phys. 35 (2011) 192

Lobes of Centaurus A seen by Fermi-LAT

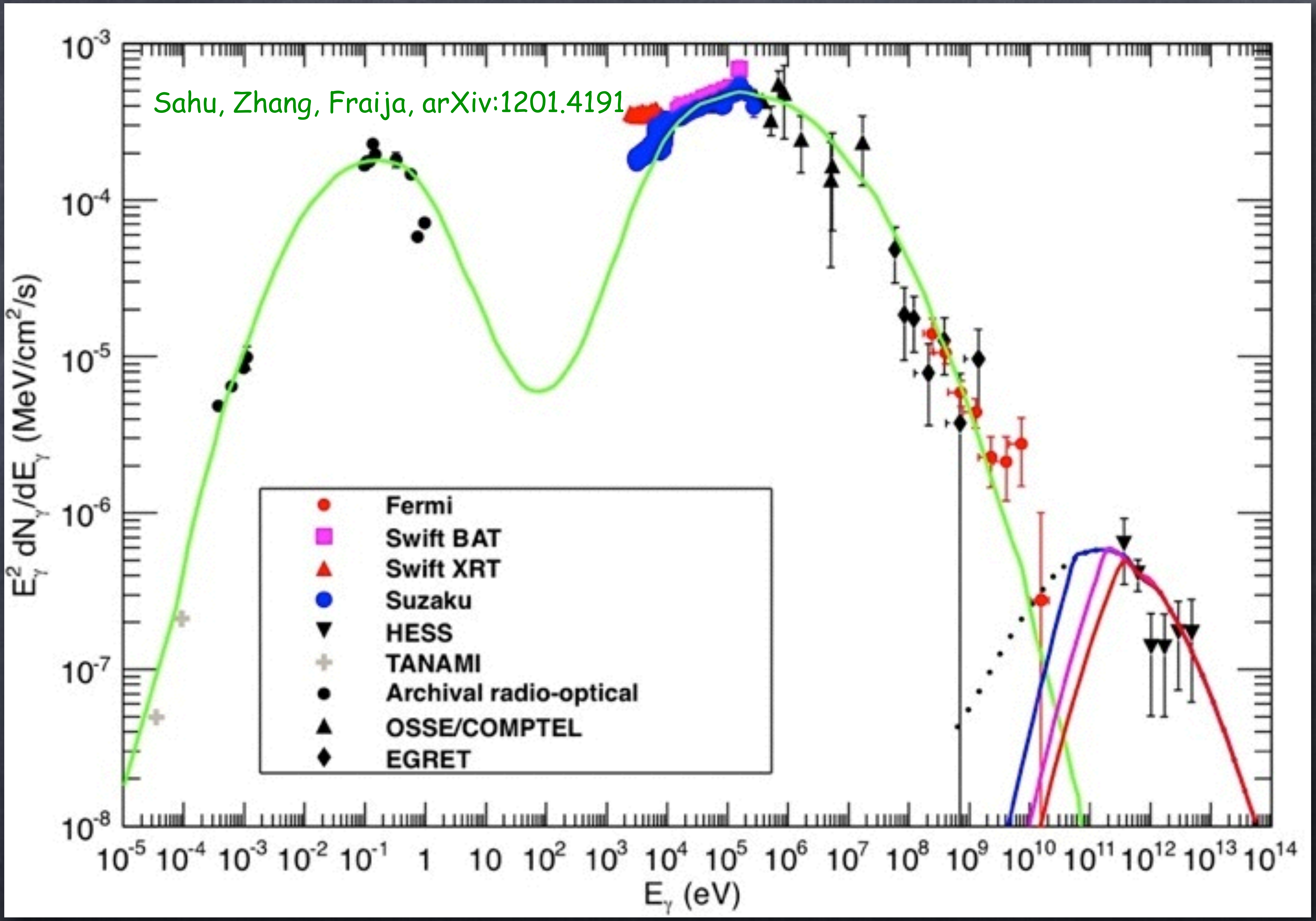


> 200 MeV γ-rays

Radio observations

Abdo et al., Science Express 1184656, April 1, 2010

Centaurus A as Multimessenger Source: A Mixed hadronic+leptonic Model



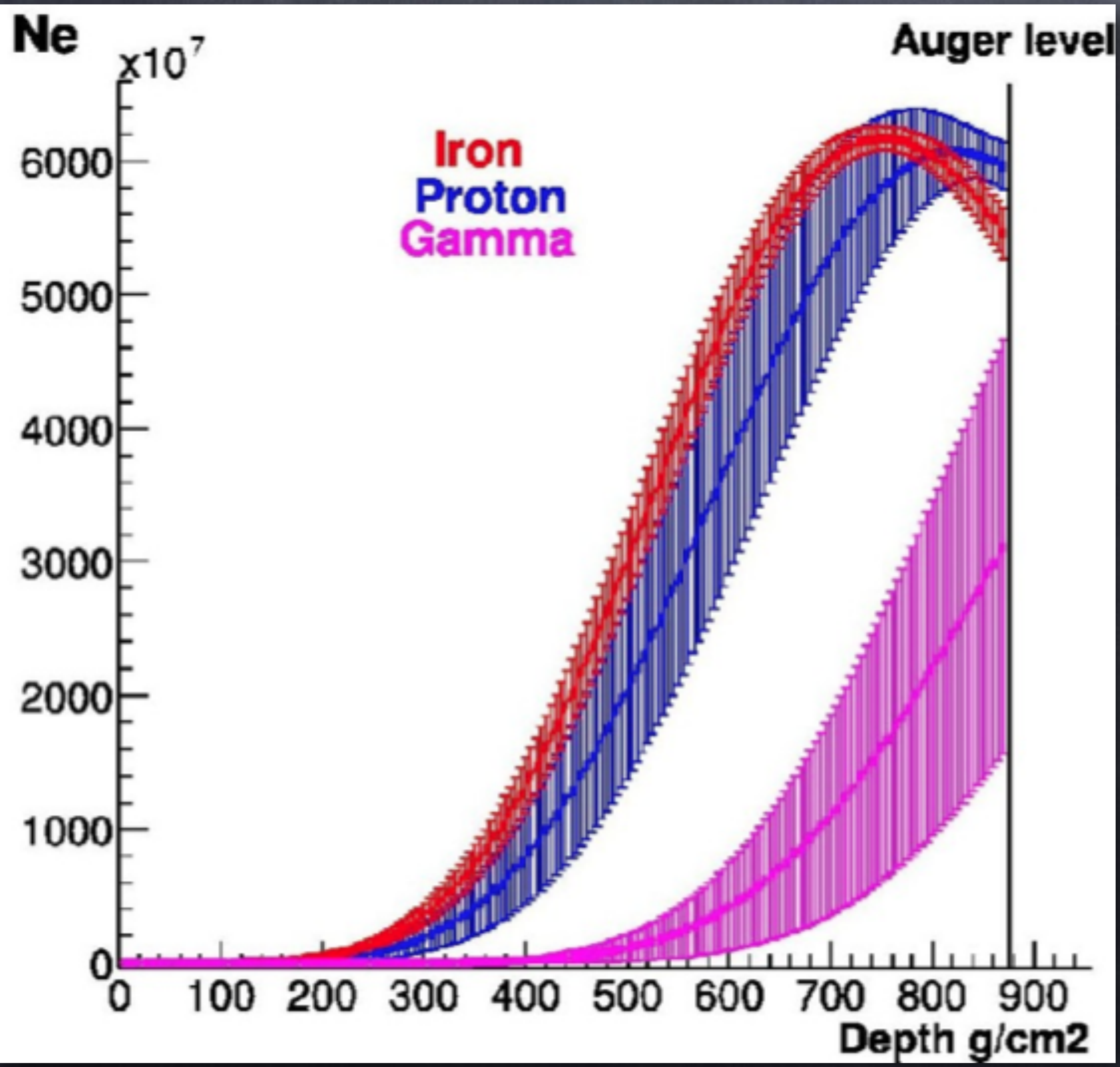
Low energy bump = synchrotron

high energy bump = synchrotron self-Compton

TeV-γ-rays: pγ interactions of shock-accelerated protons

Mass Composition

Depth of shower maximum and its distribution contain information on primary mass composition



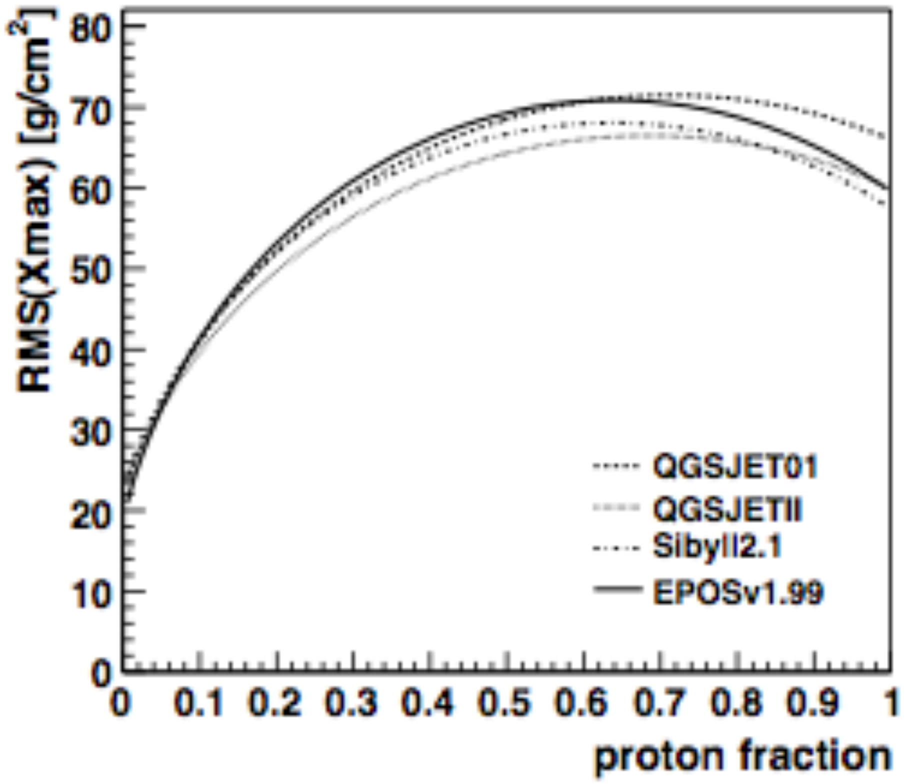
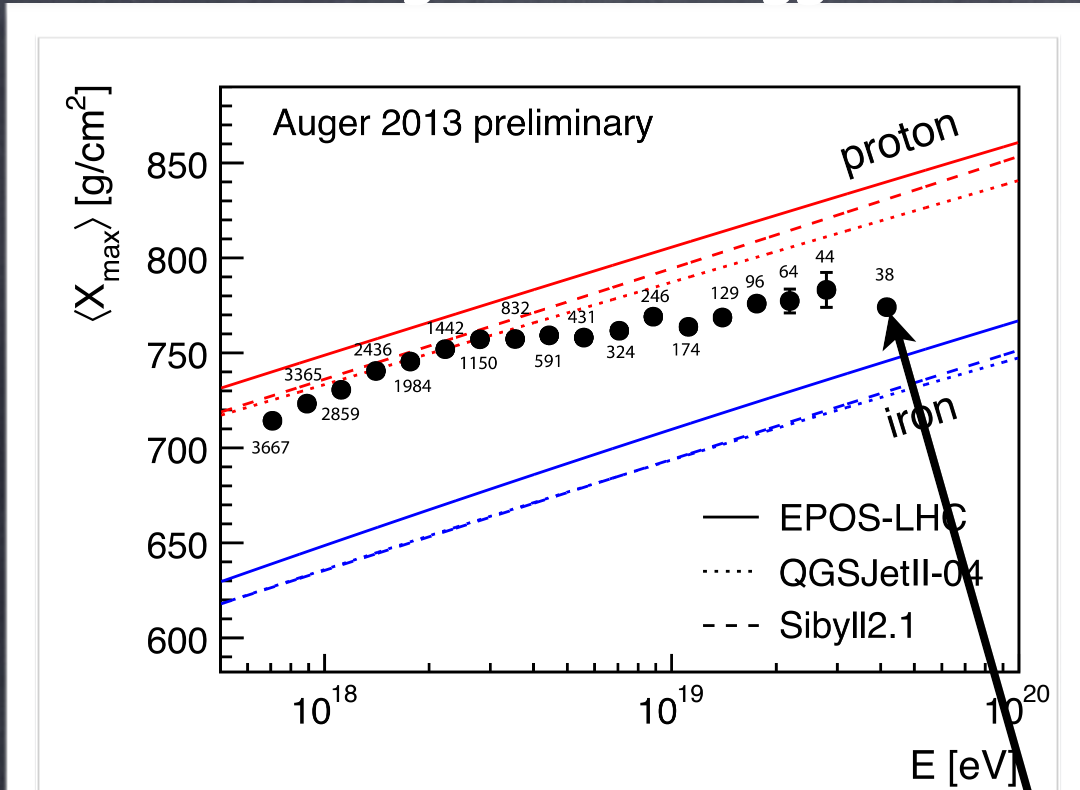
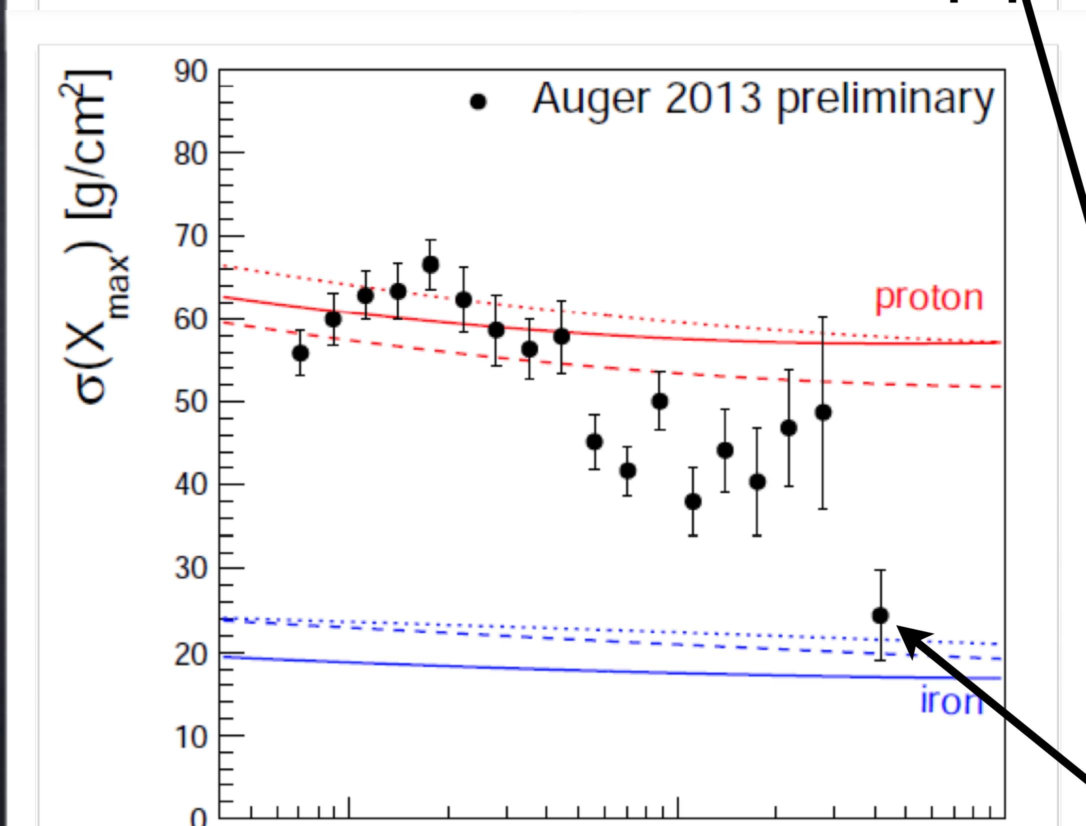


FIGURE 1. RMS(X_{max}) from different hadronic interaction models [23] and a two-component p/Fe composition model ($E = 10^{18}$ eV).

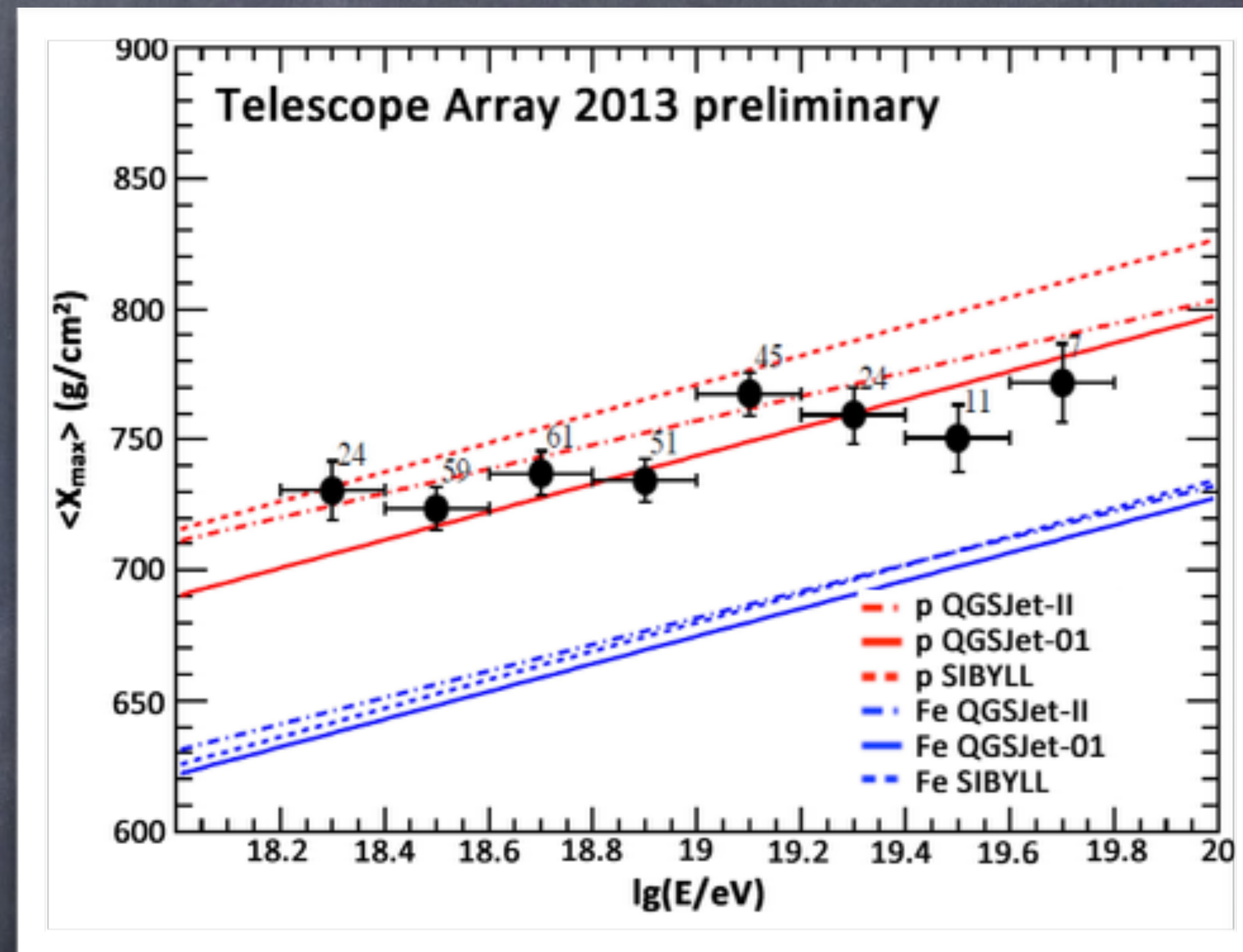
Pierre Auger data suggest a heavier composition toward highest energies:





Pierre Auger Collaboration, Phys.Rev.Lett., 104 (2010) 091101, and ICRC 2013, arXiv:1307.5059

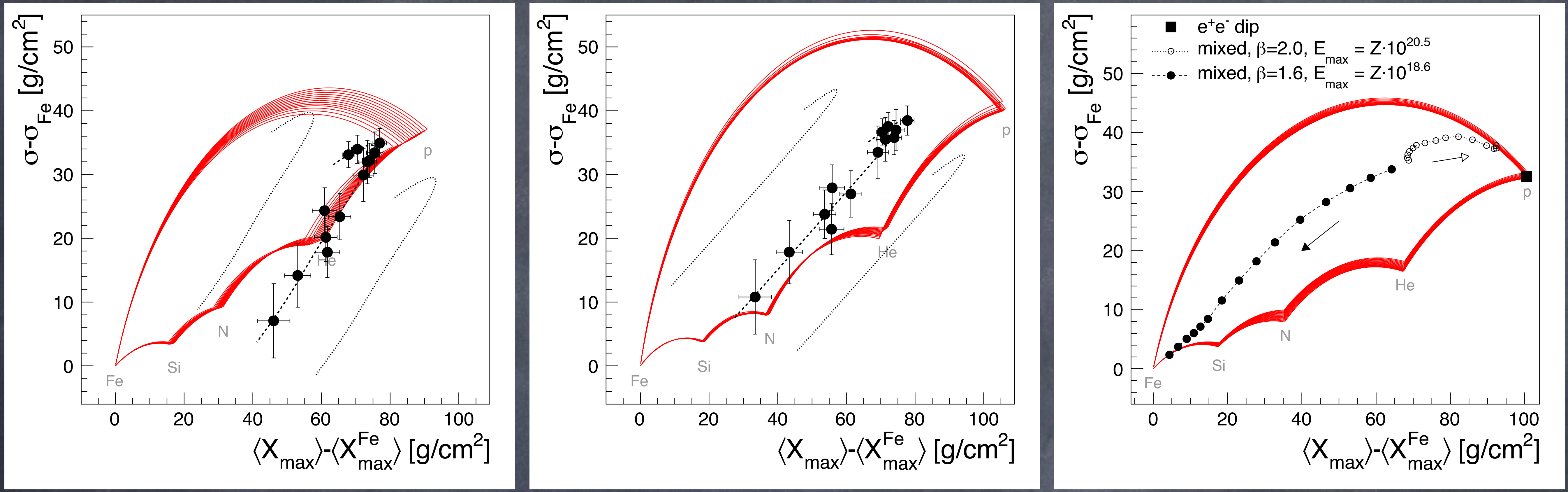
but not confirmed on the northern hemisphere by HiRes and Telescope Array which are consistent with protons

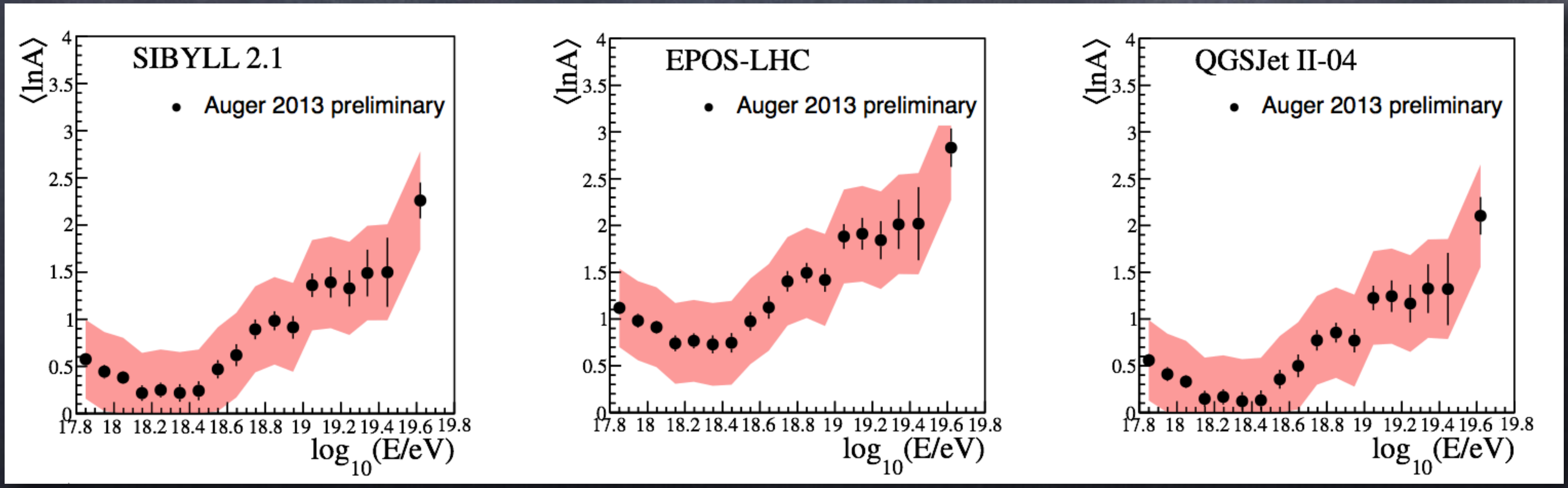


potential tension with air shower simulations and some hadronic interaction models because a mixed composition would predict larger RMS(X_{max})

combined measurement of X_{max} and its fluctuation σ constrains composition

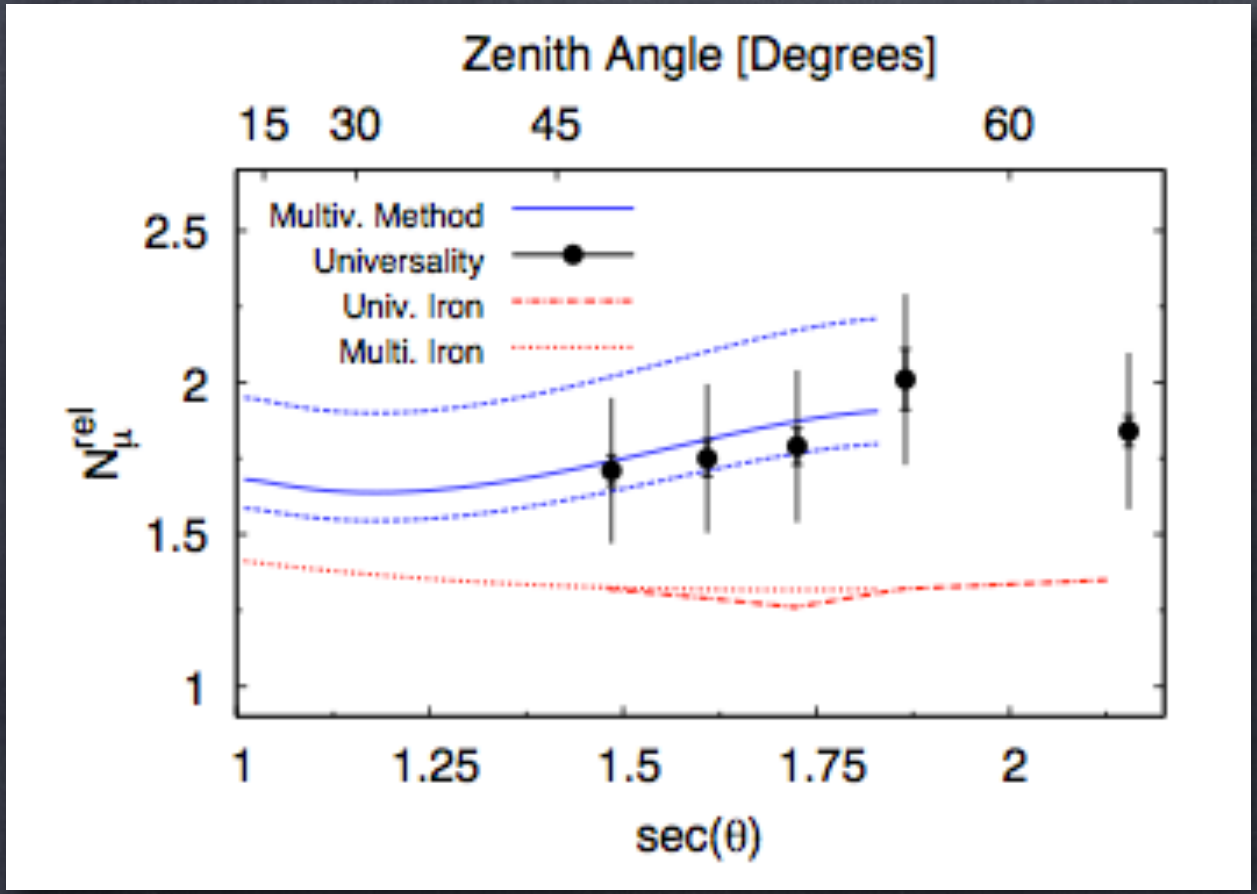
within a given hadronic interaction model



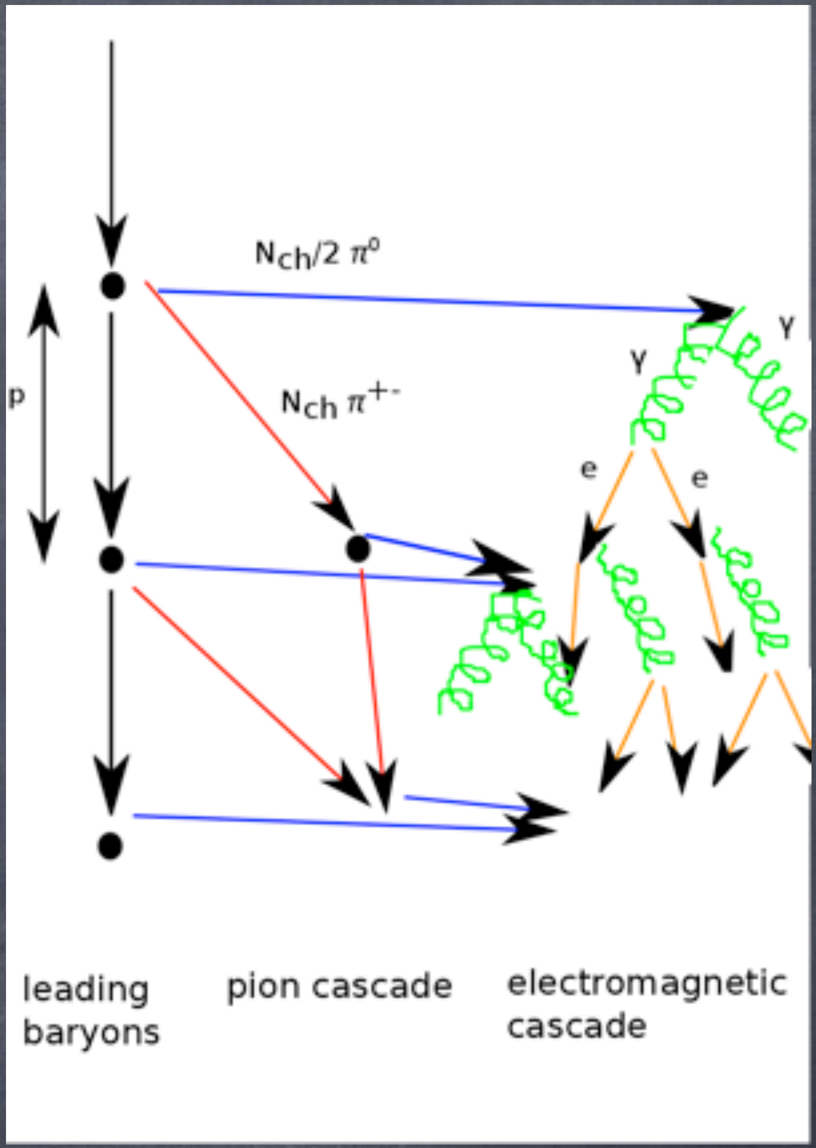


Kampert and Unger, arXiv:1201.0018, M. Roth at TeVPA 2013 and ICRC 2013

Muon number measured at 1000 m from shower core a factor ~2 higher than predicted



Pierre Auger Collaboration, ICRC 2011, Allen et al., arXiv:1107.4804



The muon number scales as

$$N_\mu \propto E_{\rm had} \propto (1 - f_{\pi^0})^N \ ,$$

with the fraction going into the electromagnetic channel $f_{\pi^0} \simeq \frac{1}{3}$ and the number of generations N strongly constrained by $X_{\rm max}$. Larger N_μ thus requires smaller f_{π^0} !

KASCADE data suggest a heavy composition below ~10^{18} eV possibly becoming lighter around 10^{18} eV

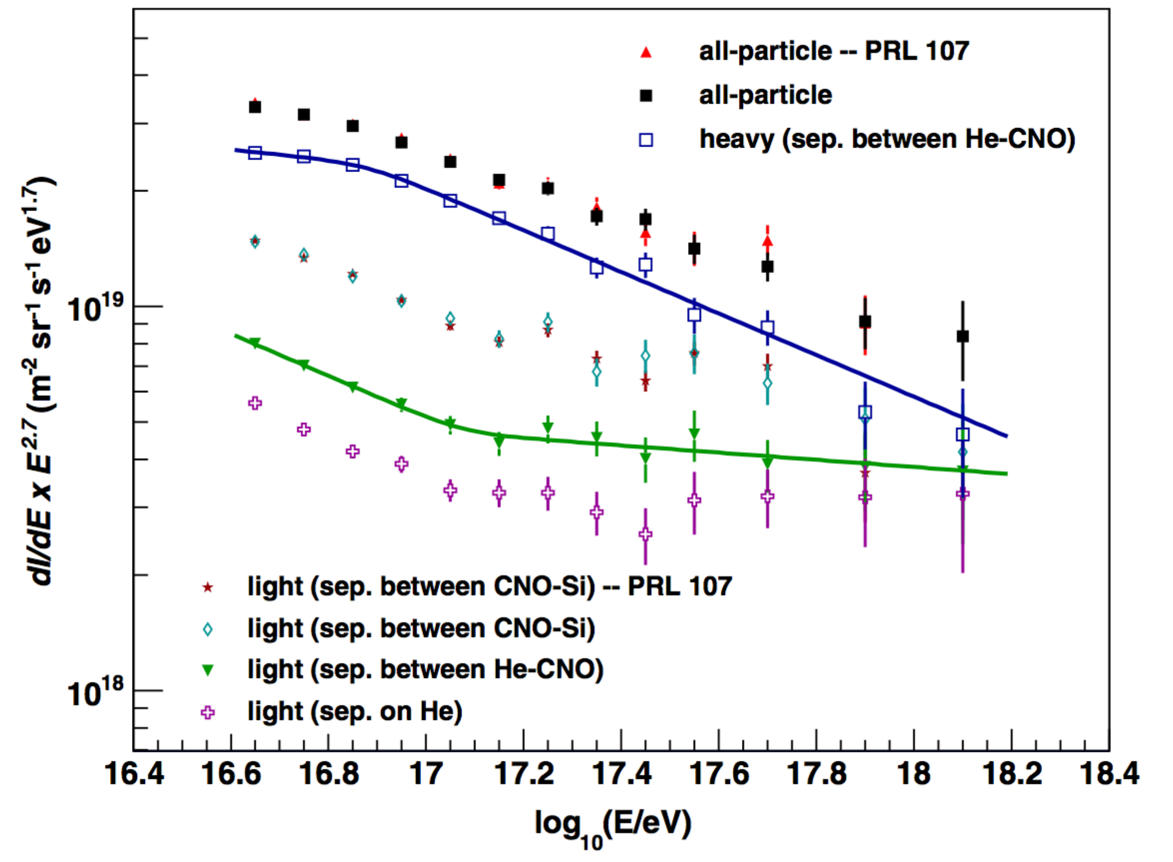
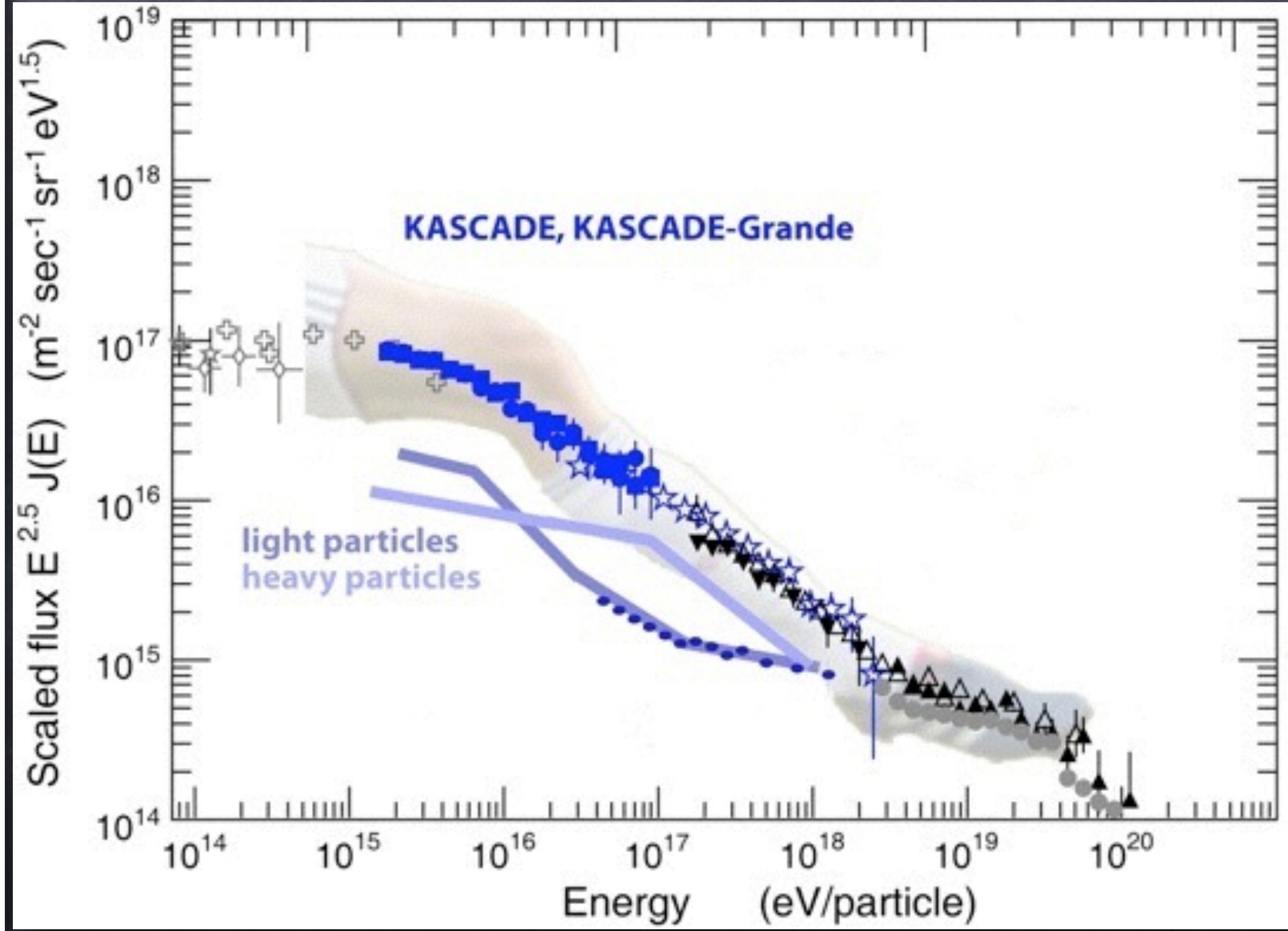
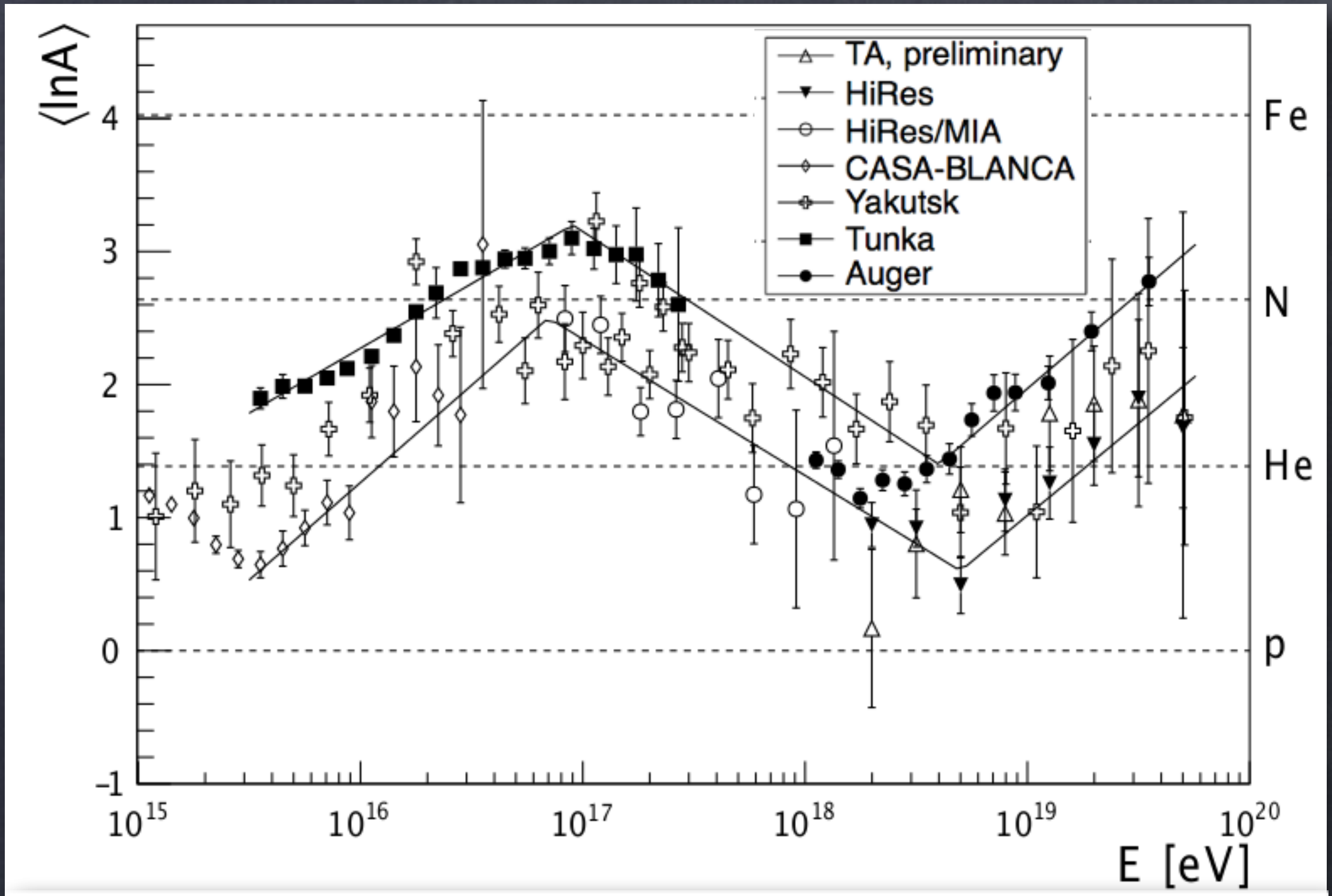


FIG. 4 (color online). The all-particle and electron-rich spectra from the analysis [8] in comparison to the results of this analysis with higher statistics. In addition to the light and heavy spectrum based on the separation between He and CNO, the light spectrum based on the separation on He is also shown. The error bars show the statistical uncertainties.

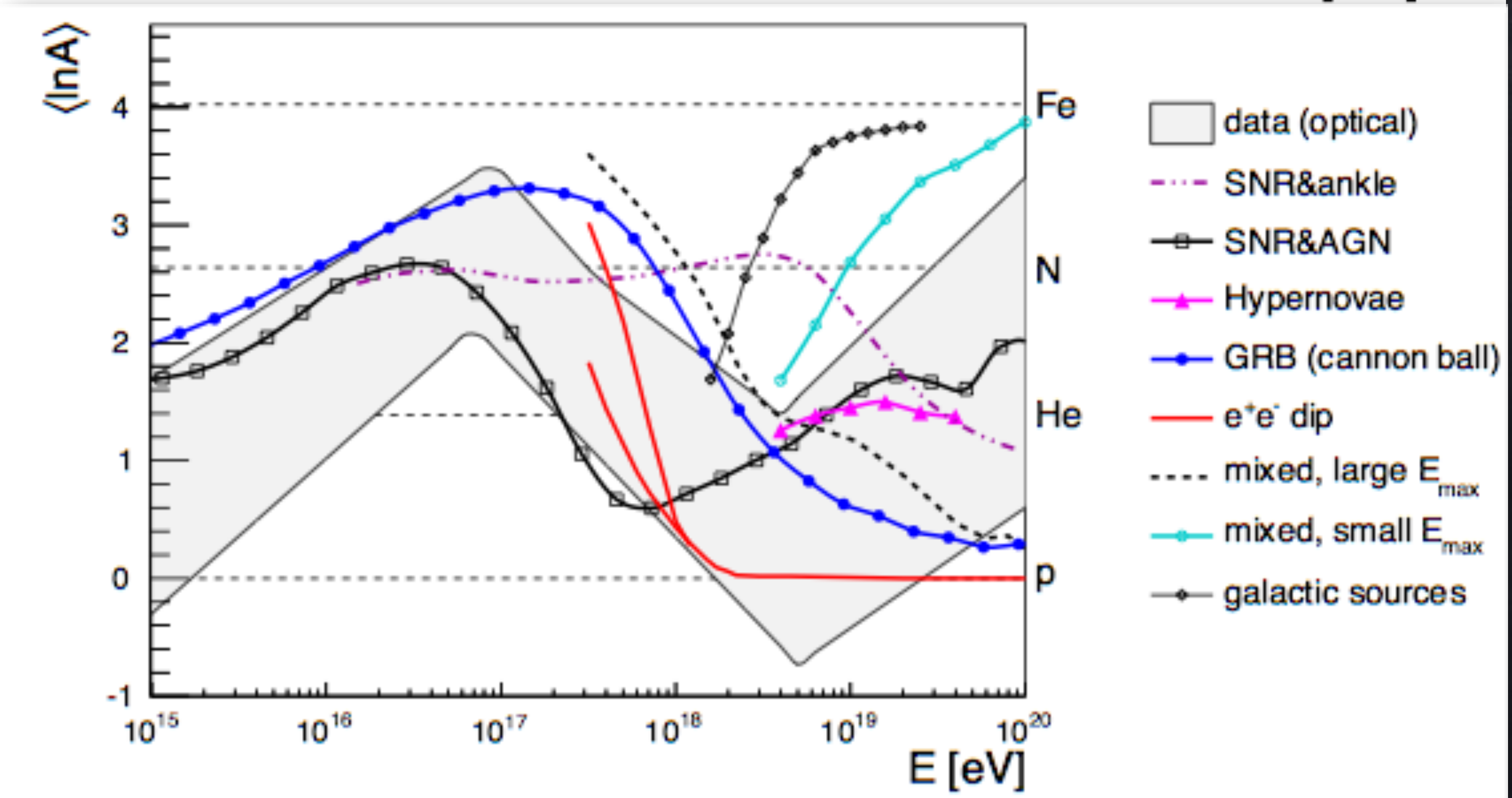
KASCADE Collaboration, Phys.Rev. D87 (2013) 081101,

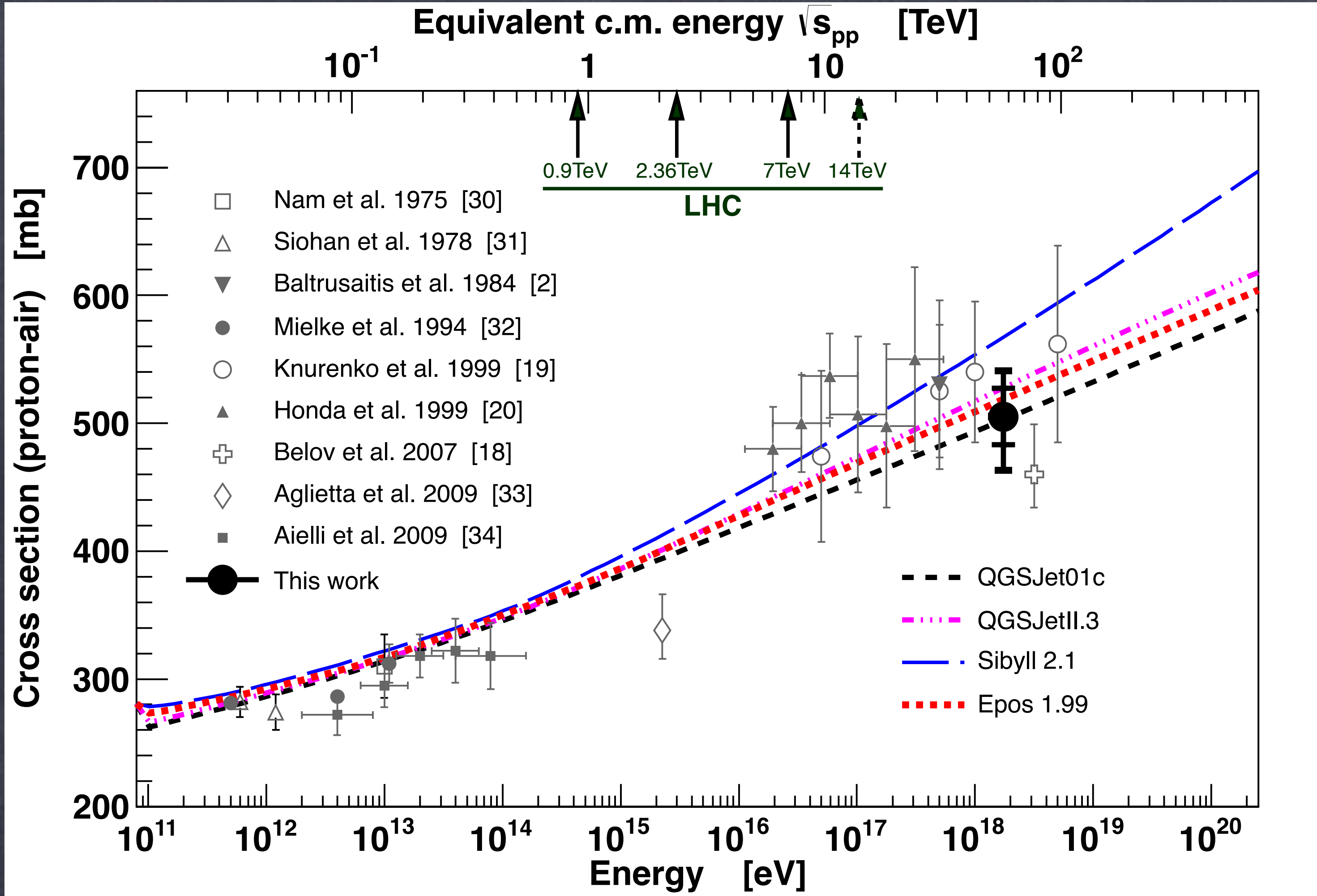


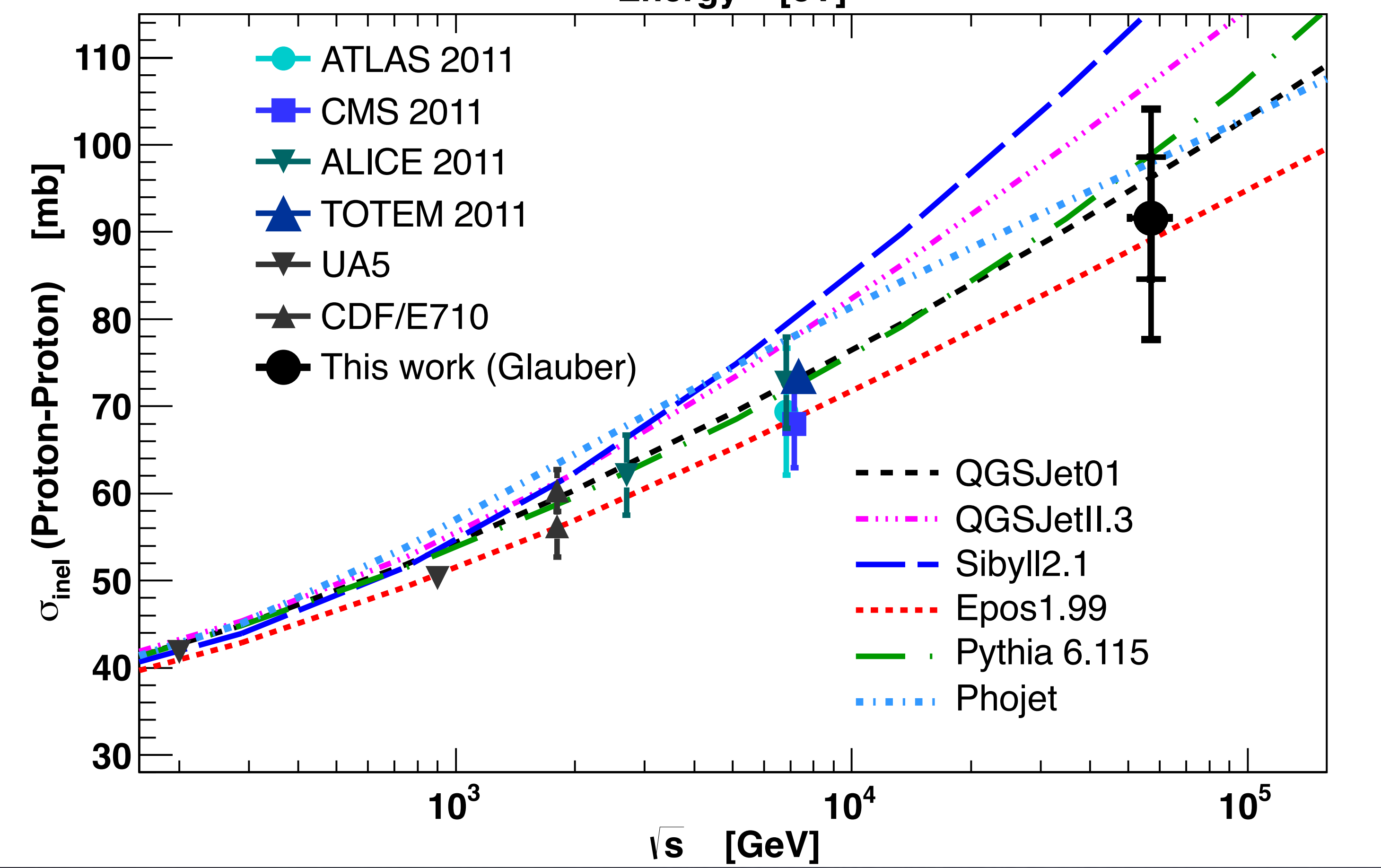
The global picture for the mass composition



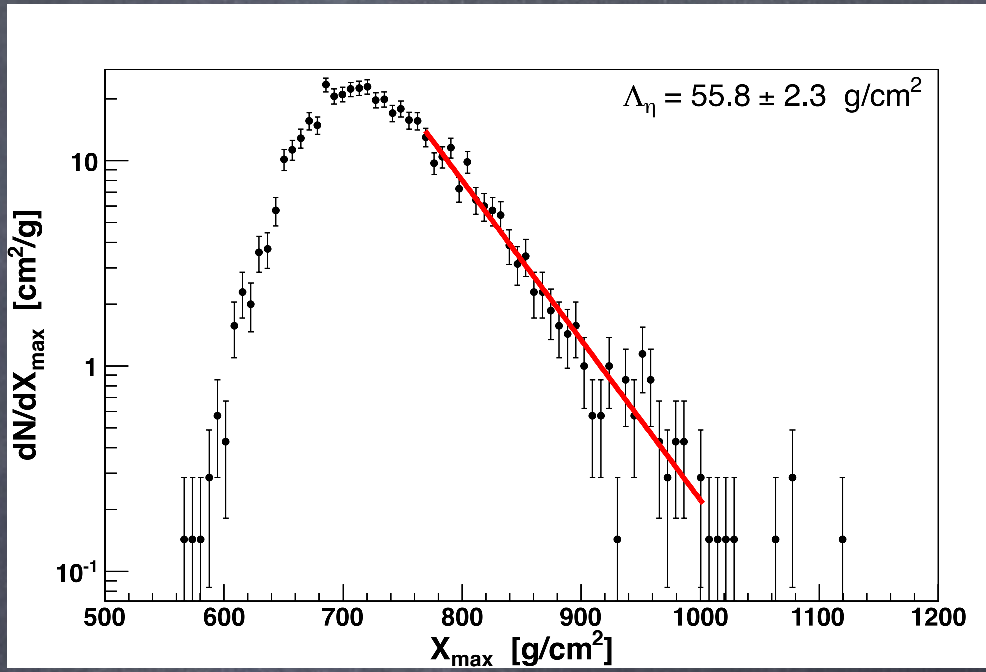
K.-H.Kampert and M.Unger, Astropart.Phys. 35 (2012) 660







p-air cross section derived from exponential tail of depth of shower maxima

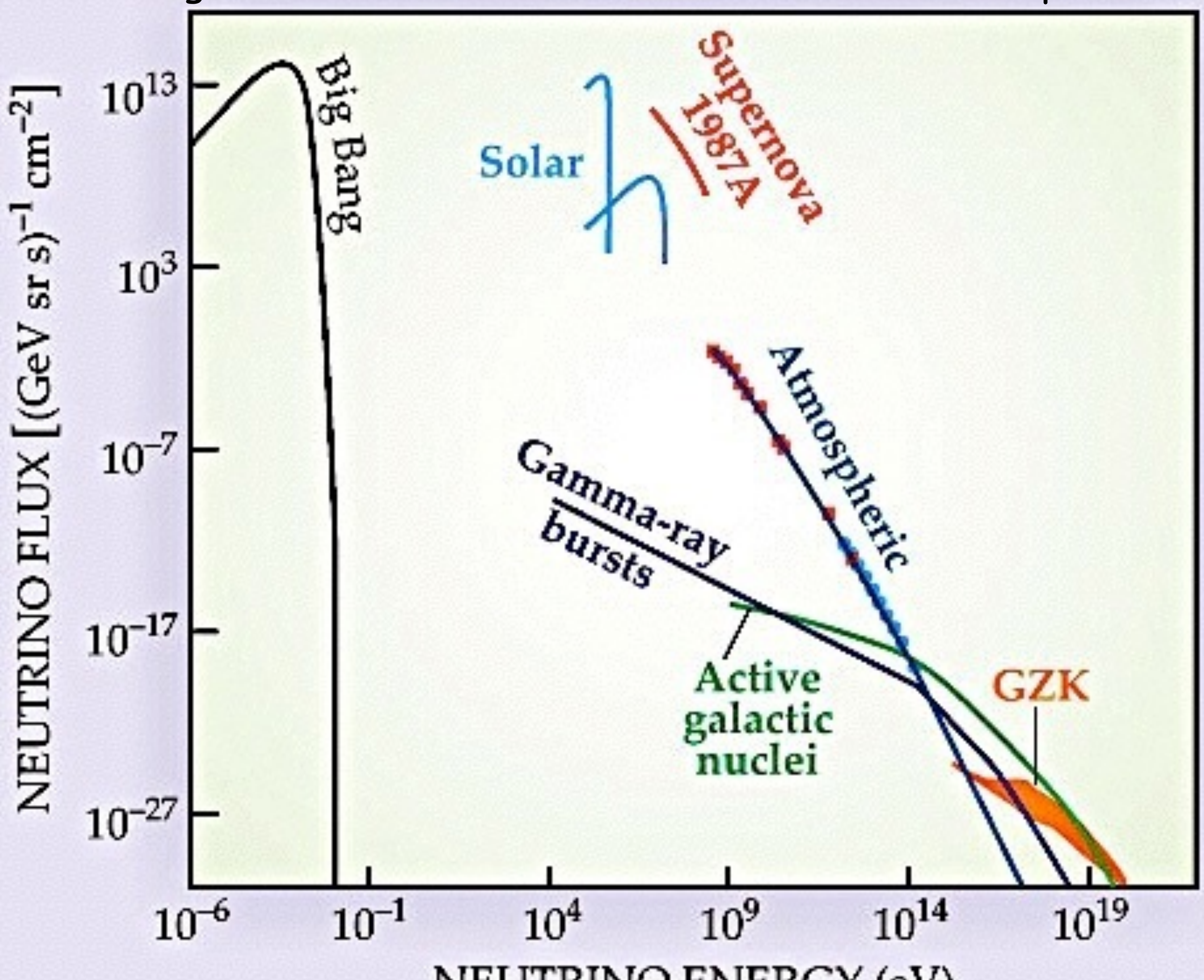


pp cross section derived from Glauber model

Pierre Auger Collaboration, PRL 109, 062002 (2012)

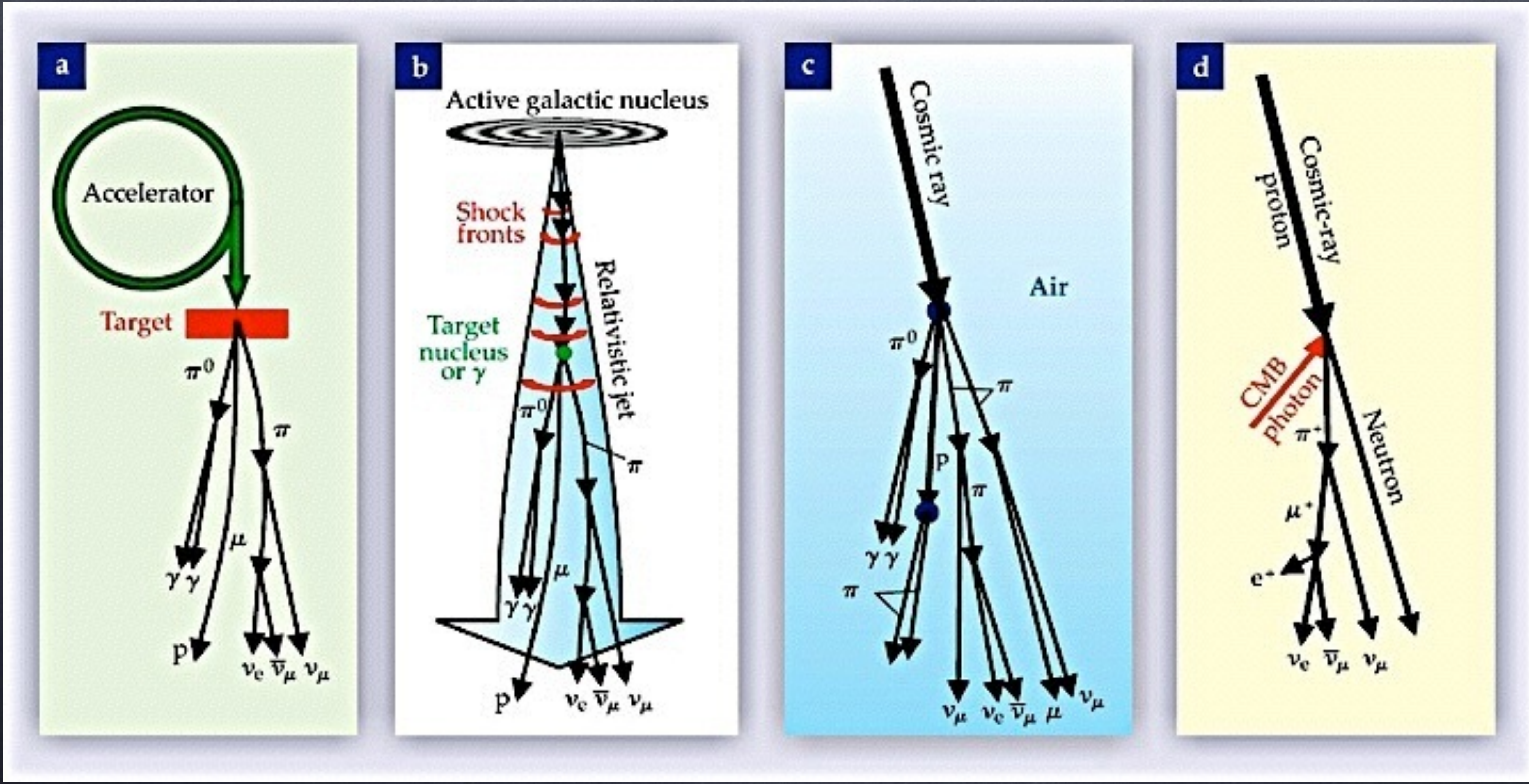
Very High High Energy Neutrinos

The „grand unified“ differential neutrino number spectrum

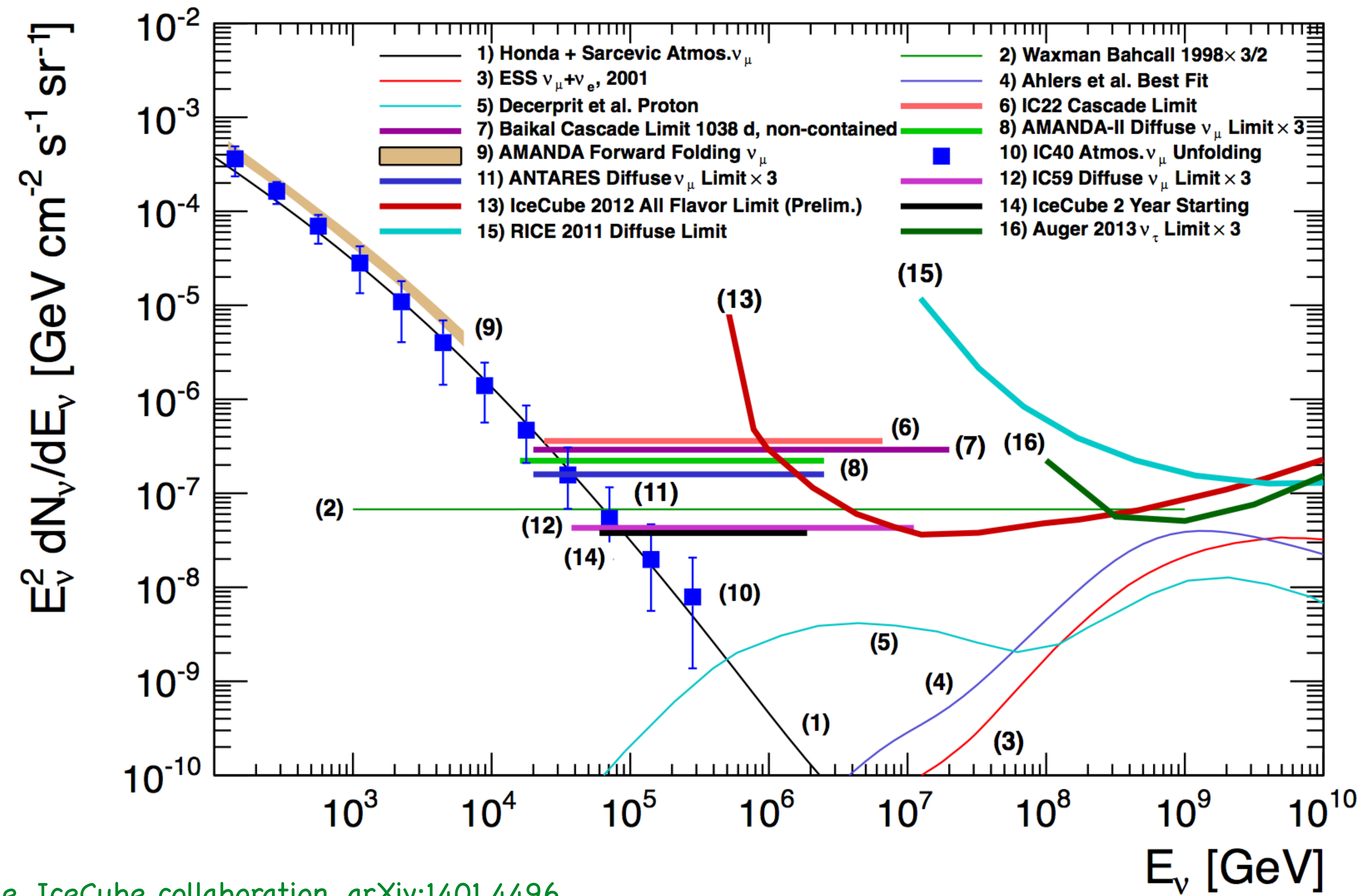


From Physics Today

Summary of neutrino production modes



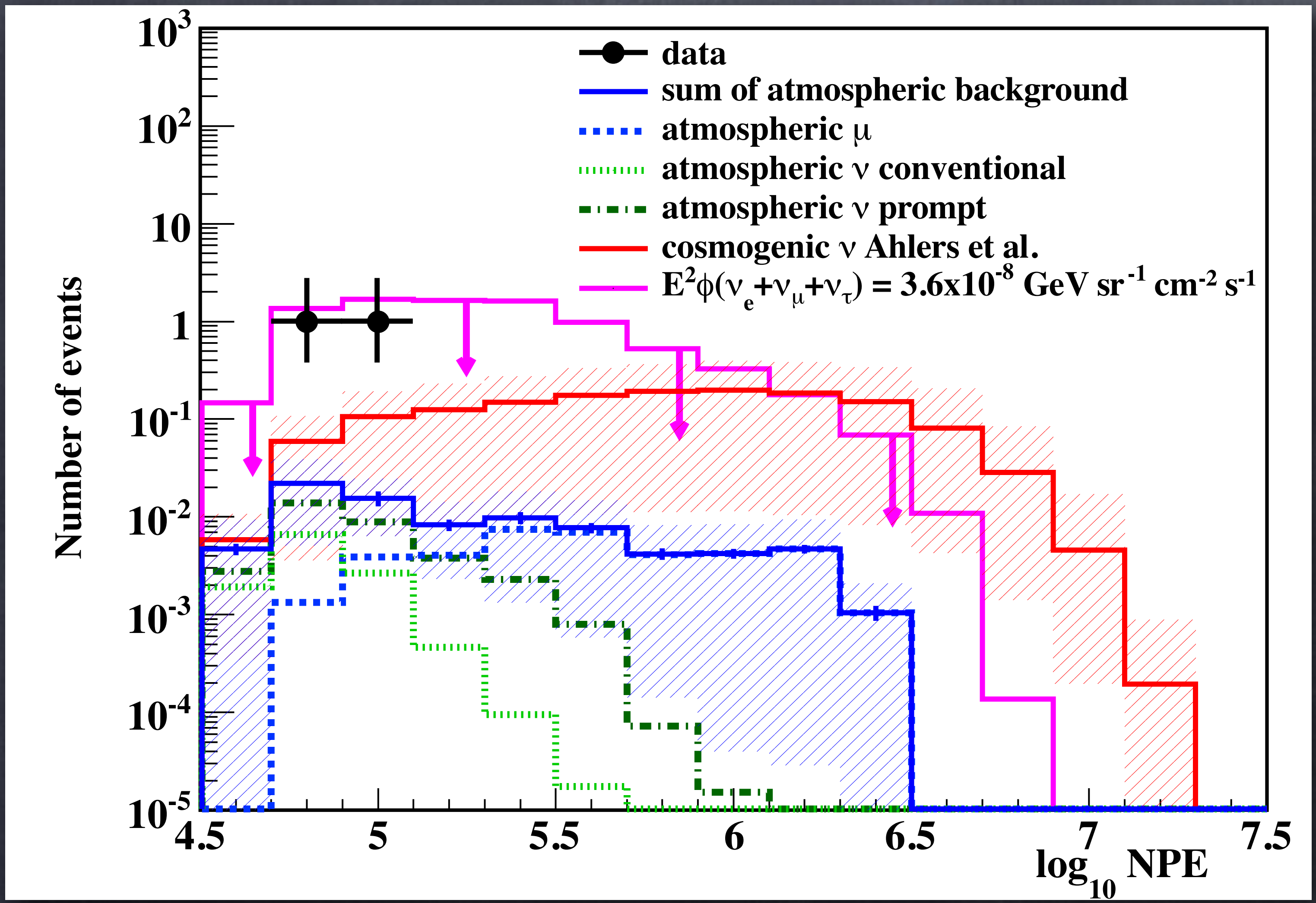
From Physics Today

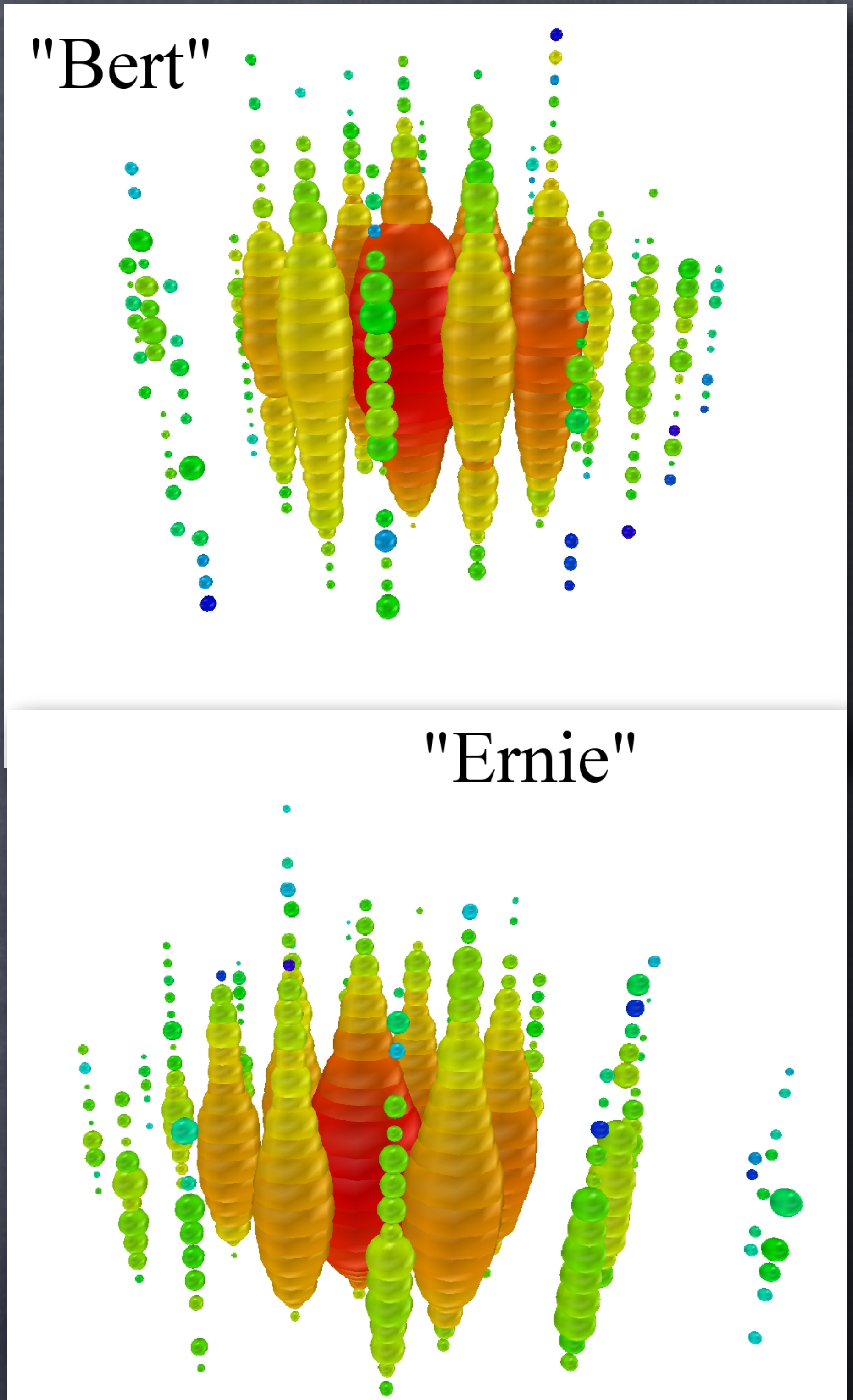


A.Karle, IceCube collaboration, arXiv:1401.4496

Figure 7: An overview is presented of observed atmospheric neutrino fluxes, upper limits to diffuse fluxes and models. The IceCube 2012 differential upper limit (11) turn up sharply at 1PeV because of observed PeV events. The best fit diffuse flux using starting events in IceCube (12) forms evidence for a diffuse astrophysical flux up to PeV energies above the atmospheric neutrino spectrum extending to a few 100 TeV.

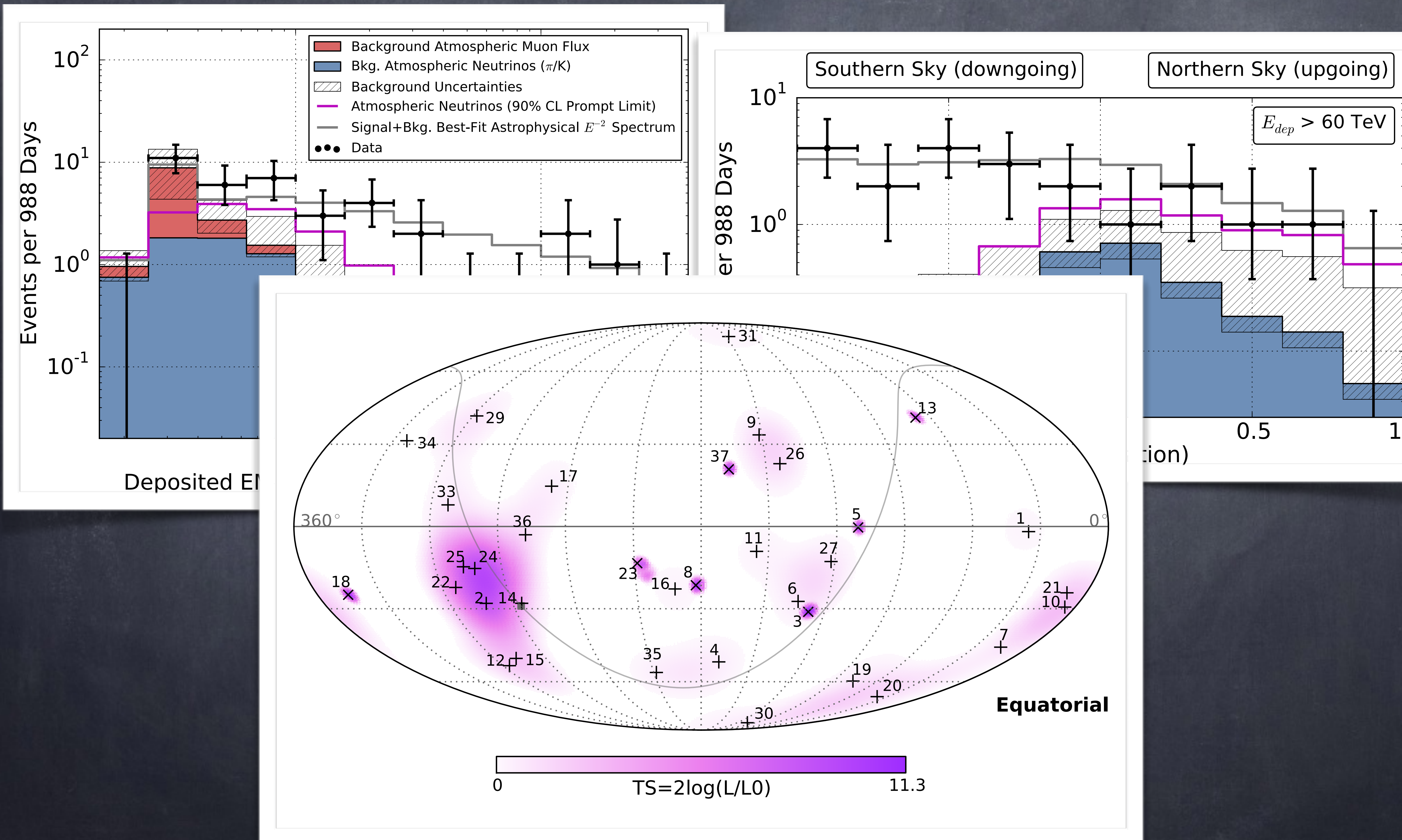
But now two PeV energy candidate neutrinos observed by IceCube





IceCube collaboration, arXiv:1304.5356

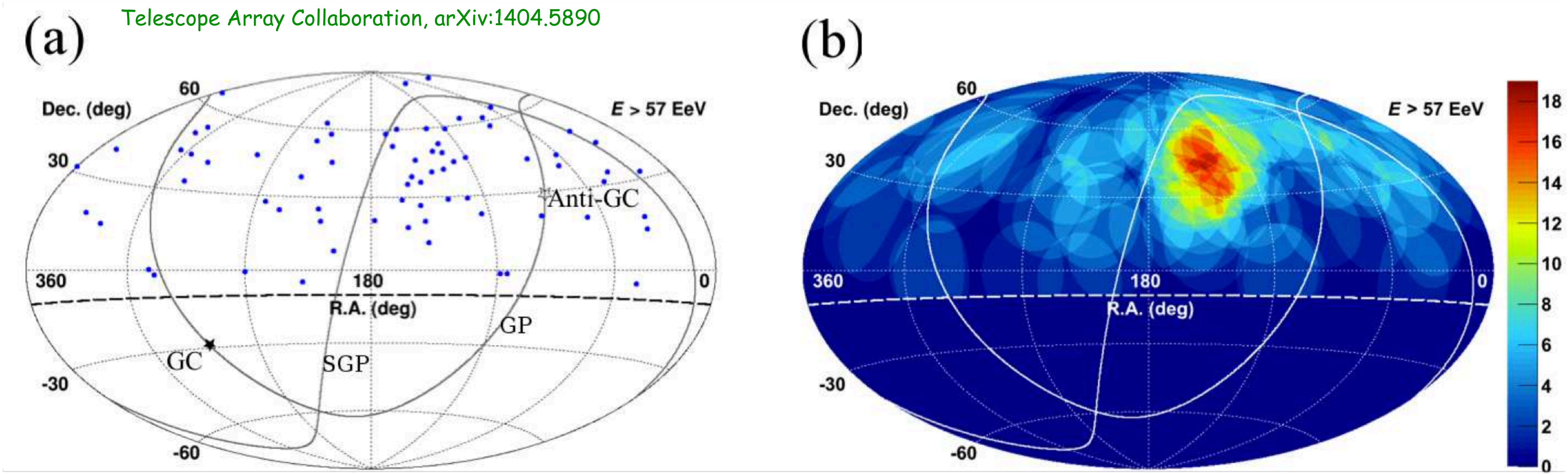
and a total of 37 events above 30 TeV deposited energy:



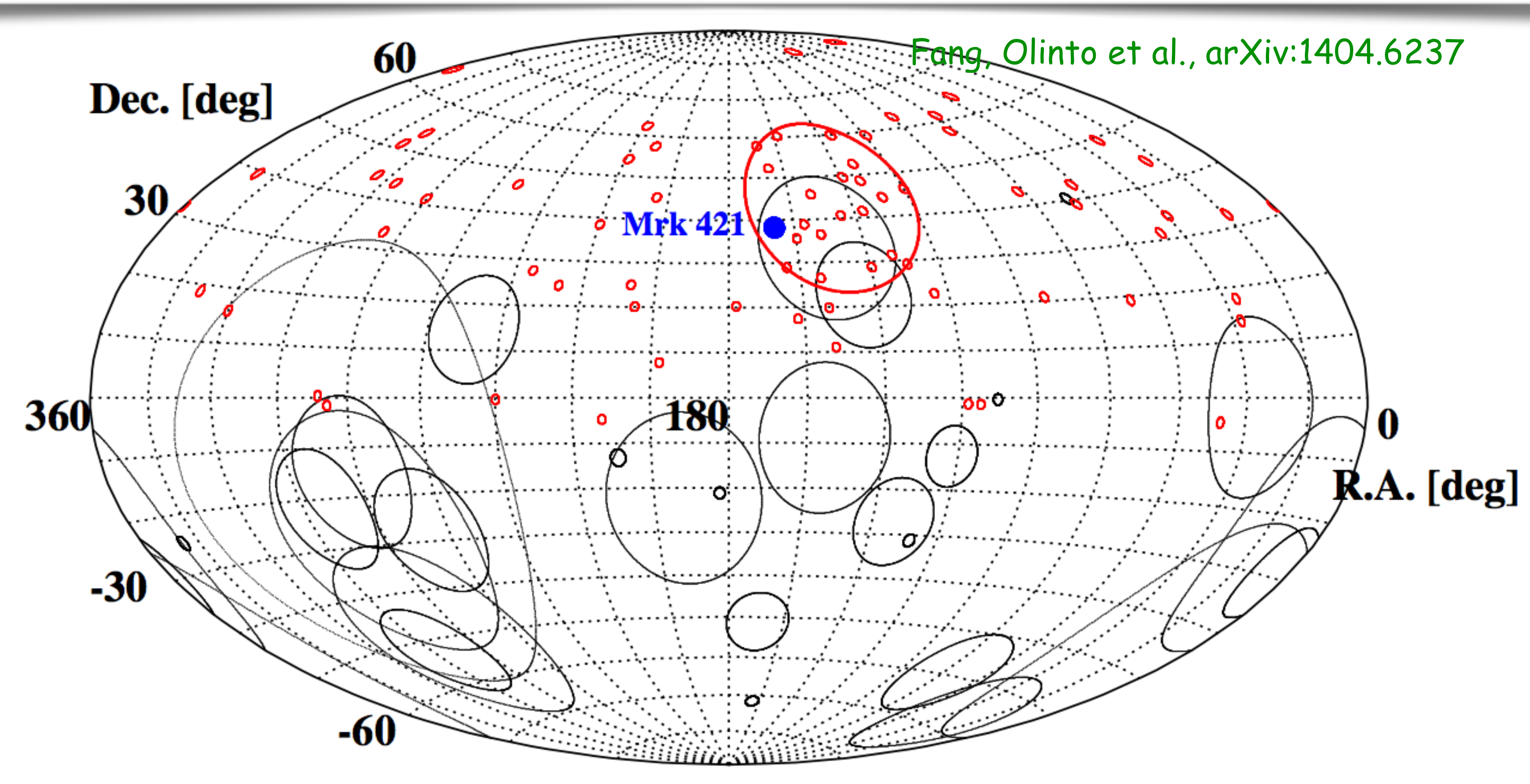
IceCube collaboration, arXiv:1405.5303

A possible Correlation of IceCube Neutrinos with the Cosmic Ray Excess seen by Telescope Array ?

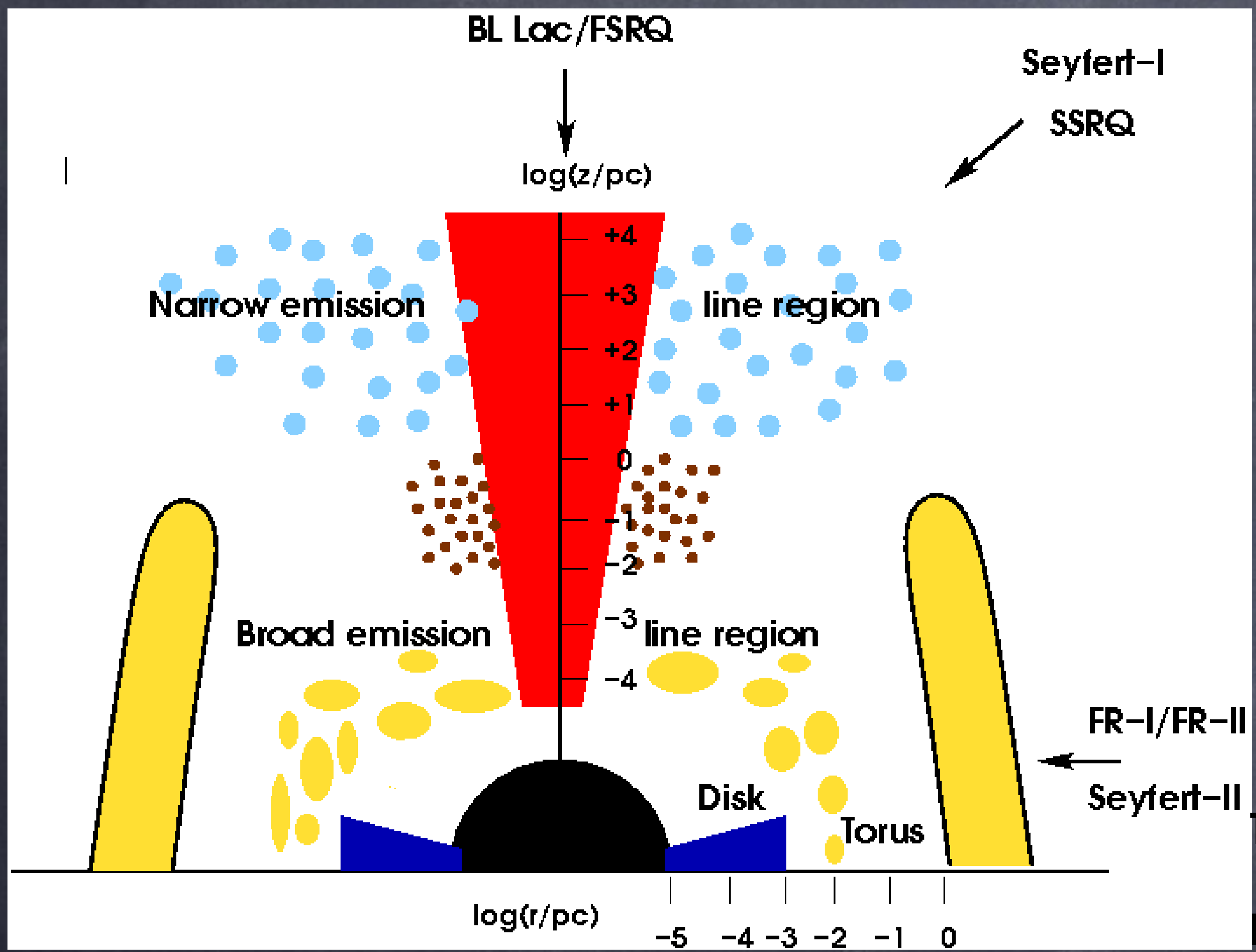
Telescope Array Collaboration, arXiv:1404.5890

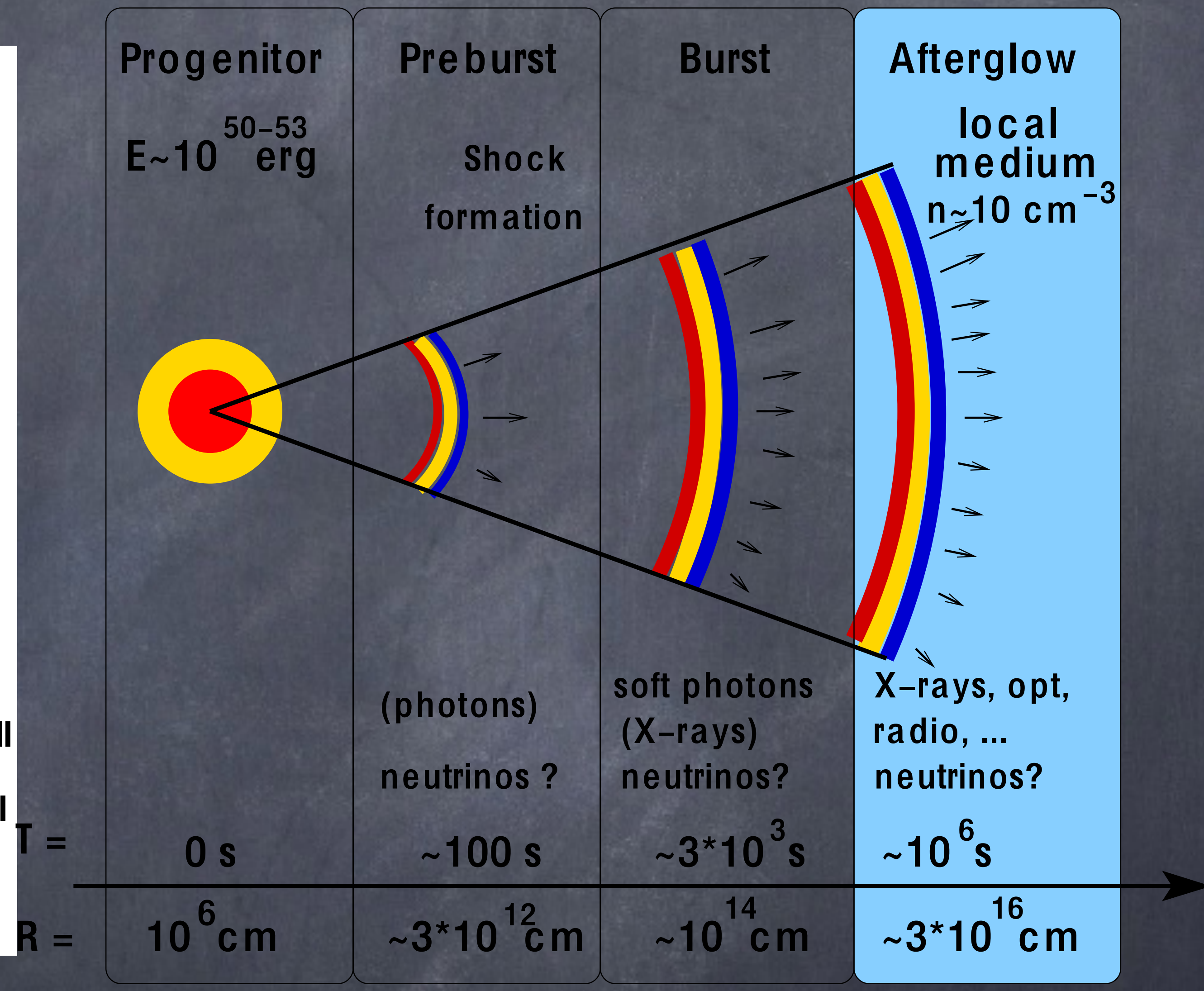


Fang, Olinto et al., arXiv:1404.6237



Discrete Extragalactic High Energy Neutrino Sources





active galaxies **gamma ray bursts**

Figures from J. Becker, Phys.Rep. 458 (2008) 173

Neutrino Fluxes from Gamma-Ray Bursts

GRBs are optically thick to charged cosmic rays and nuclei are disintegrated => only neutrons escape and contribute to the UHECR flux by decaying back into protons

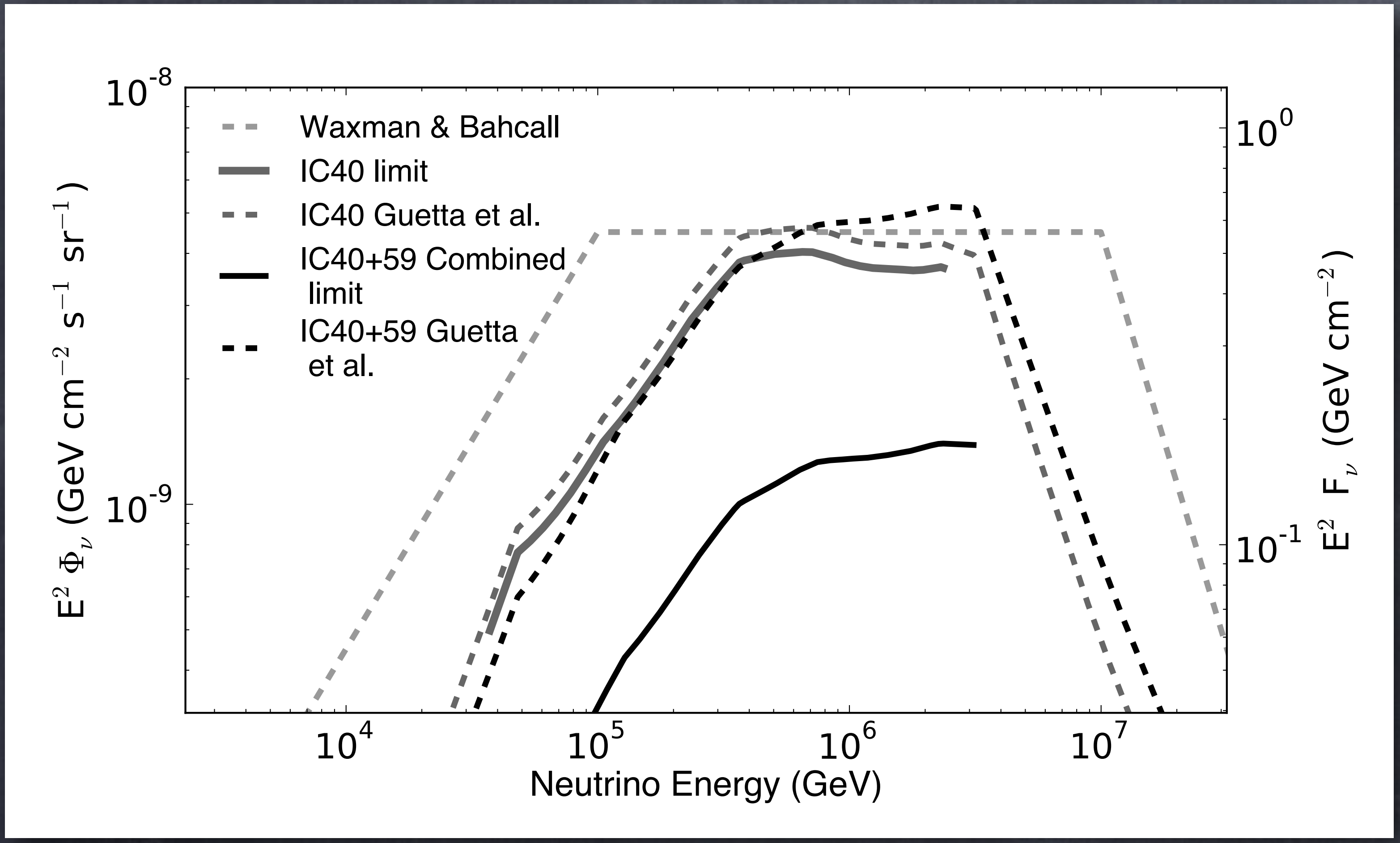
Diffuse neutrino flux from GRBs can thus be linked to UHECR flux (if it is dominantly produced by GRBs)

$$\Phi_\nu(E_\nu) \sim \frac{1}{\eta_\nu} \Phi_p \left(\frac{E}{\eta_\nu} \right) ,$$

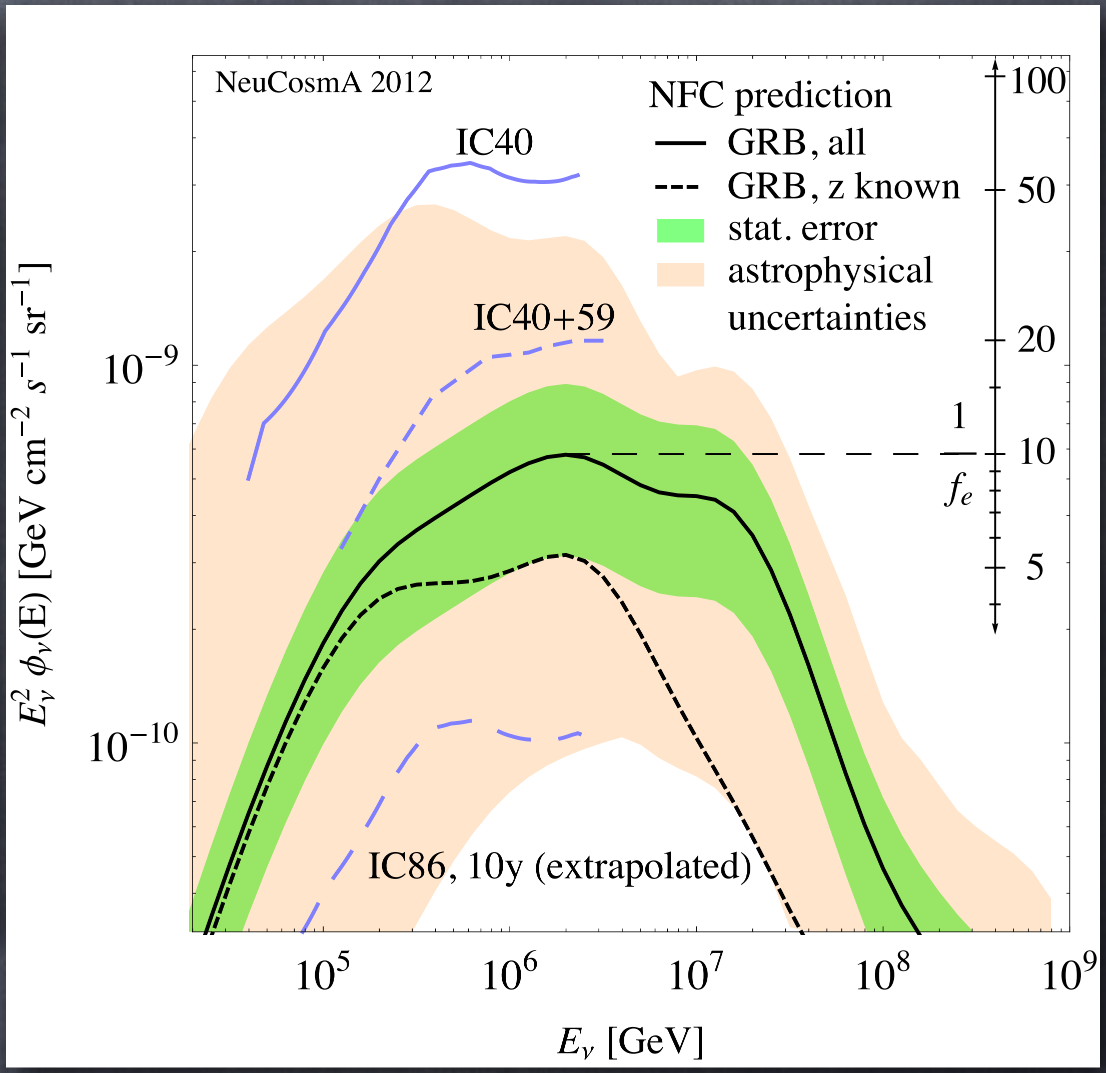
where $\eta_\nu \simeq 0.1$ is average neutrino energy in units of the parent proton energy.

Above ~ 10^{17} eV neutrino spectrum is steepened by one power of E_ν because pions/muons interact before decaying

GRBs as UHECR sources now strongly constrained by neutrino fluxes observed by IceCube



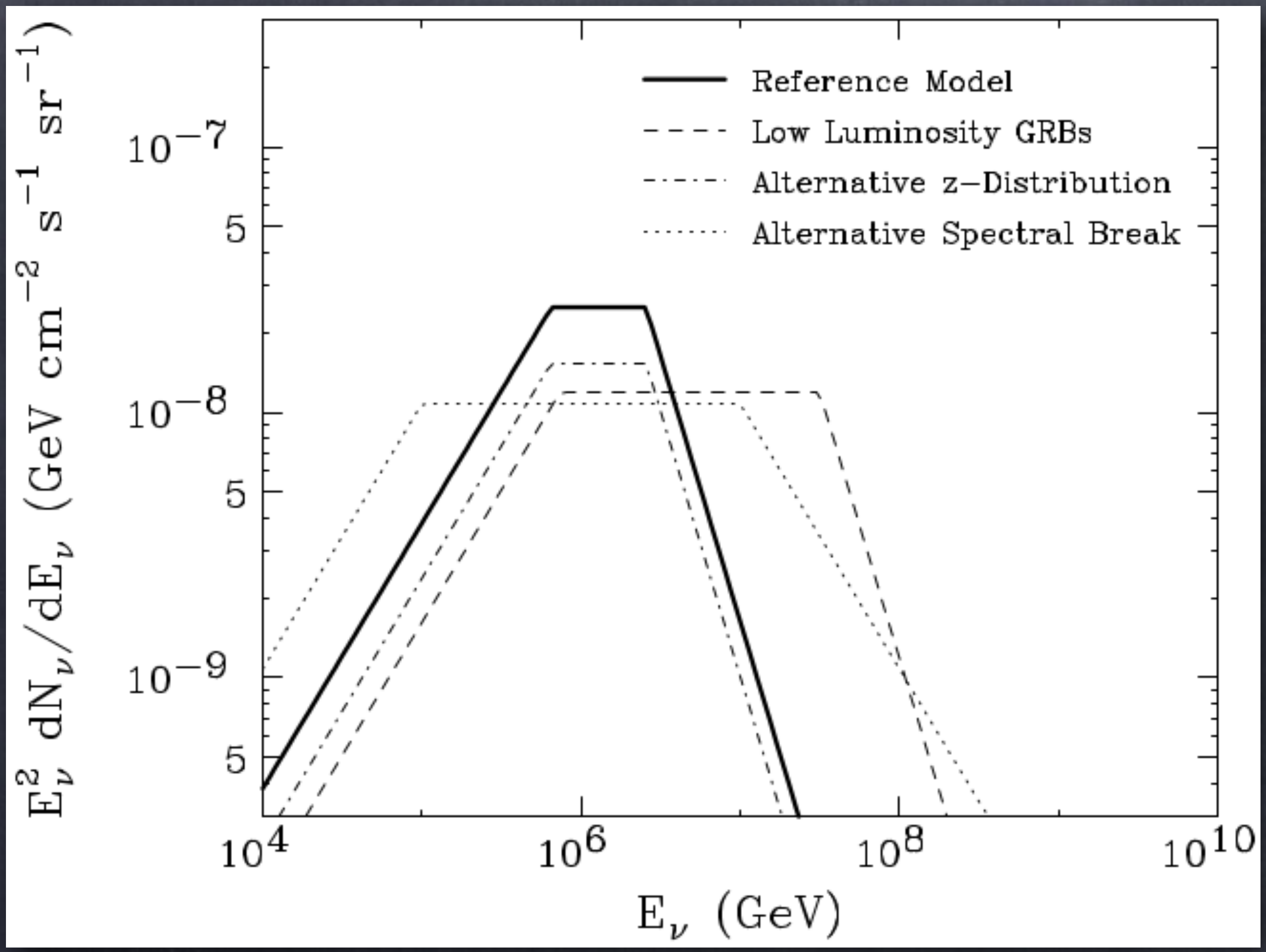
IceCube collaboration, Nature 484 (2012) 351

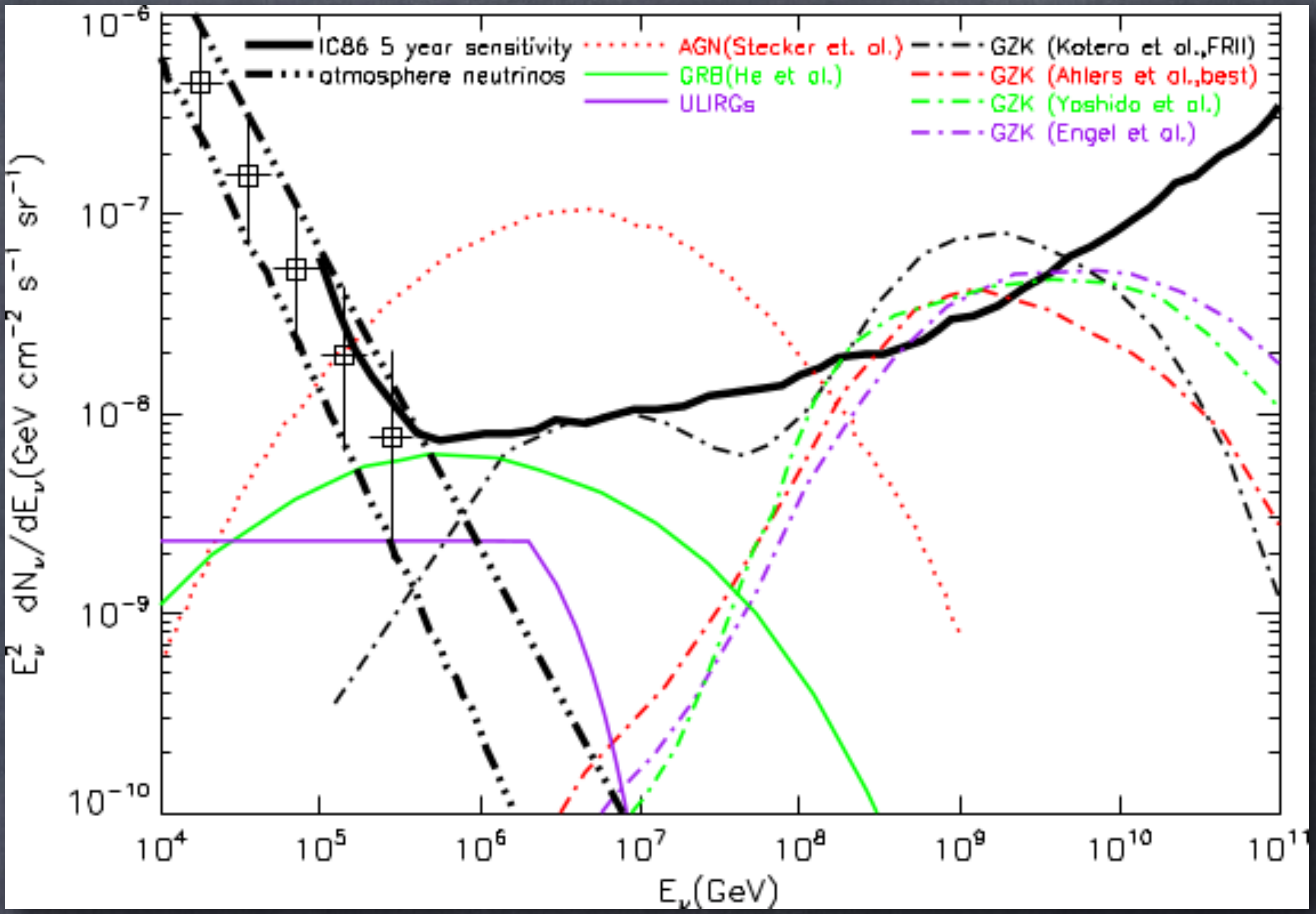


but re-evaluation of diffuse neutrino flux from GRBs gave factor ~10 smaller fluxes

Hümmer, Baerwald, Winter, PRL 108 (2012) 231101

But GRB models can still be tweaked to explain the IceCube events



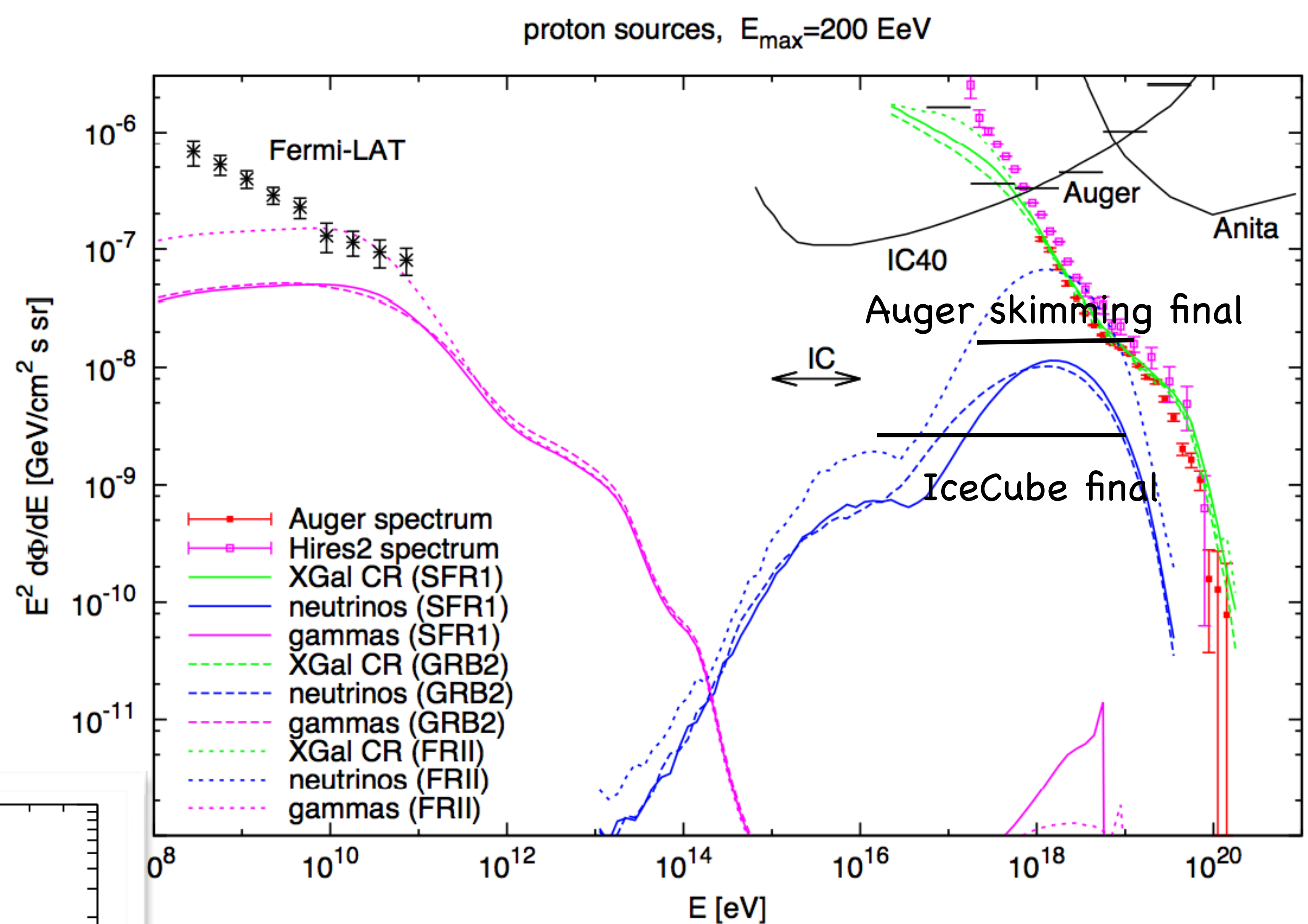


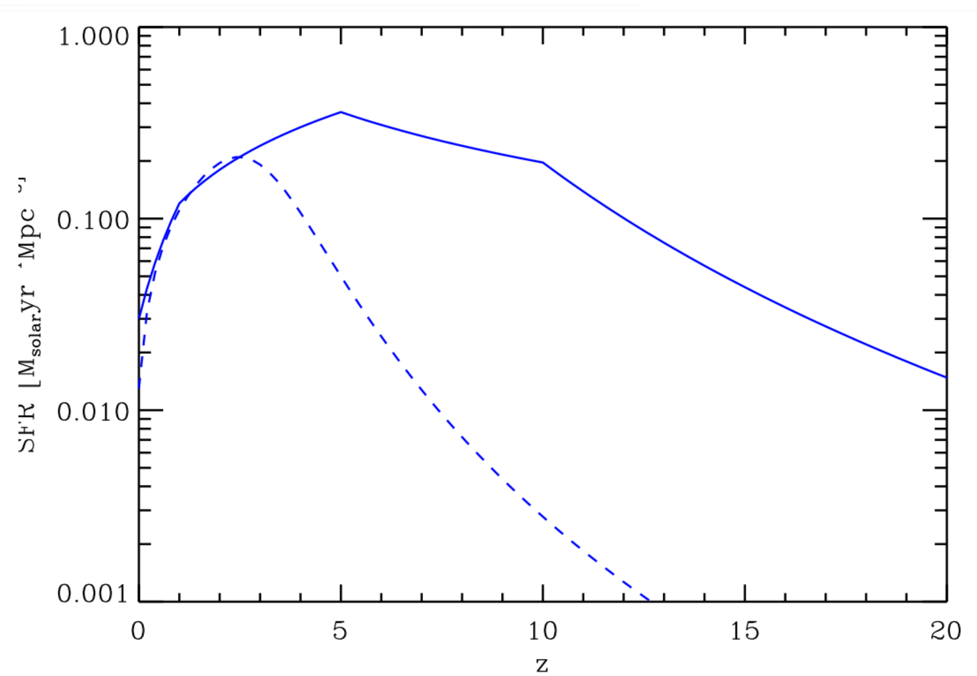
Cholis and Hooper, arXiv:1211.1974

He et al., arXiv:1303.1253

Cosmogenic Neutrinos: Maximal Fluxes for Pure Proton Injection

- Including secondary photons
- strong source evolution is here constrained by Fermi-LAT results





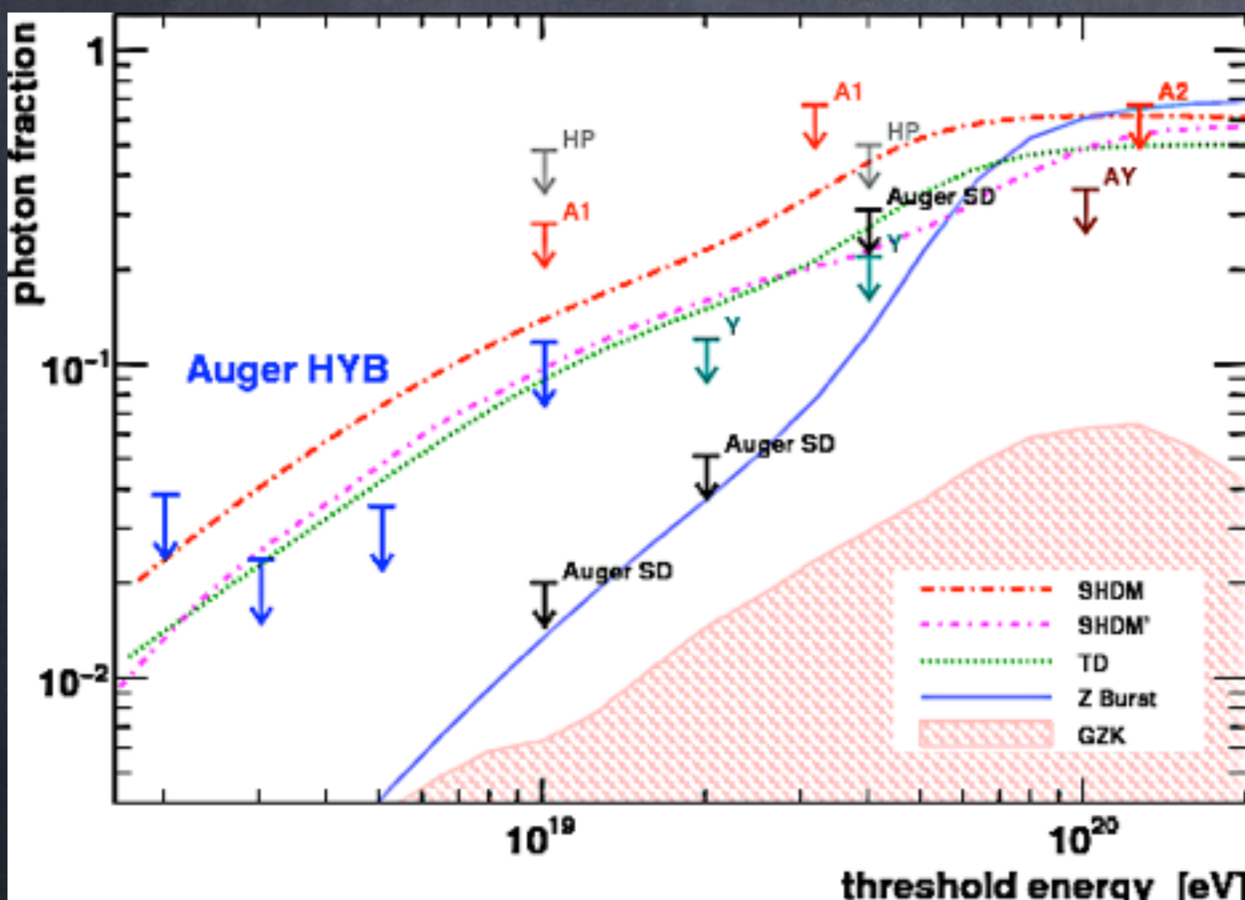
...n scenario with $E_{\max} = 200\,\mathrm{EeV}$ for different source evolution models (SFR1, GRB2 ... source spectral index is $\alpha = 2.4$ for the SFR1 and GRB2 models, while $\alpha = 2.2$ for ... Indicated are the propagated proton spectrum, the resulting (all flavor) neutrino ...luxes. The photon background measured by Fermi-LAT [10] is indicated, besides the ... ν bounds included in figure 1.

Roulet, Sigl, van Vliet, Mollerach, JCAP 1301, 028

Lorentz Symmetry Violation in the Electromagnetic Sector

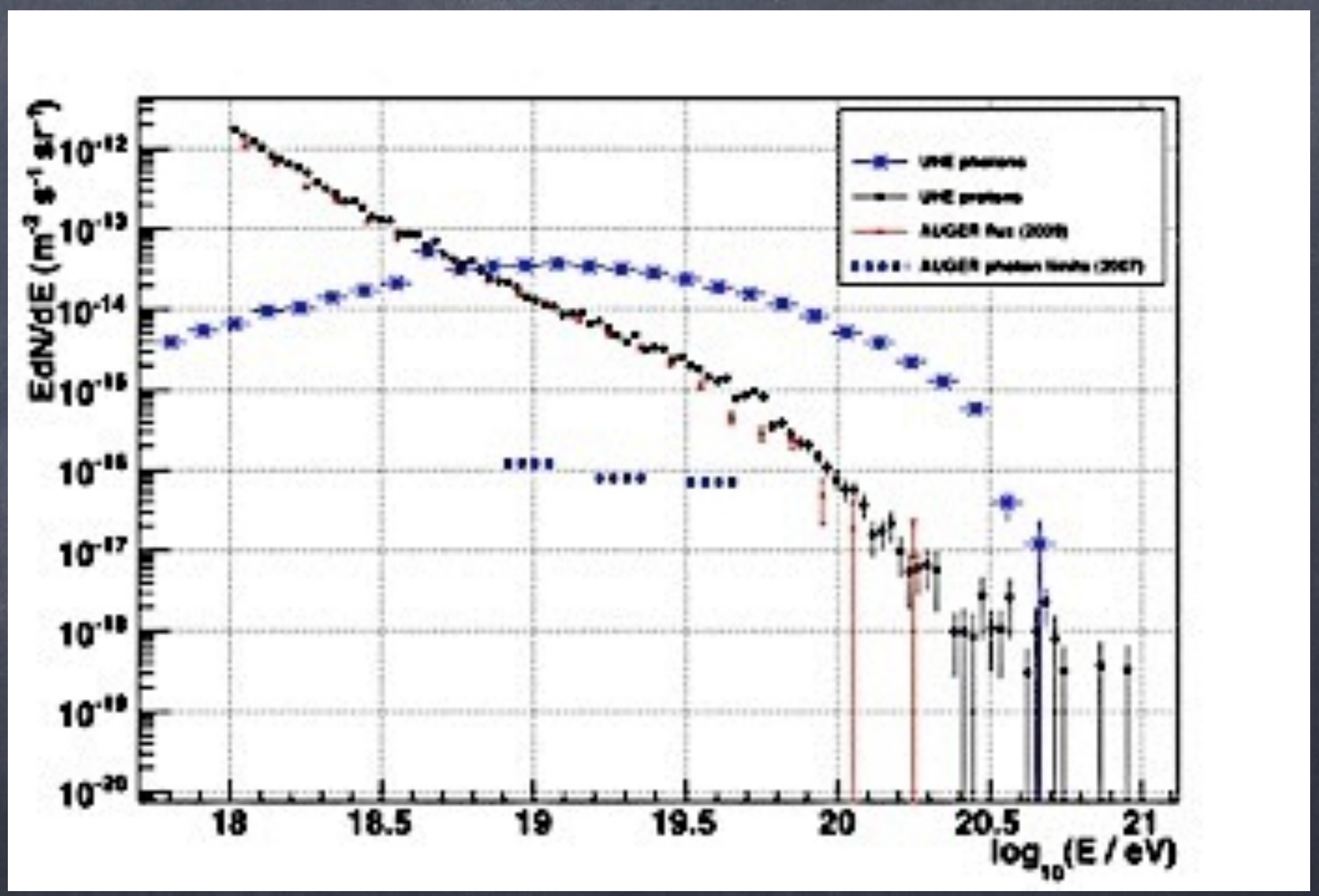
The idea:

Experimental upper limits on UHE photon fraction



Pierre Auger Collaboration, Astropart. Phys. 31 (2009) 399

Contradict predictions if pair production is absent



Maccione, Liberati, Sigl, PRL 105 (2010) 021101

Lorentz Symmetry Violation in the Photon Sector

For a photon dispersion relation

$$\omega_{\pm}^2 = k^2 + \xi_n^{\pm} k^2 \left(\frac{k}{M_{\rm Pl}}\right)^n , n \geq 1 ,$$

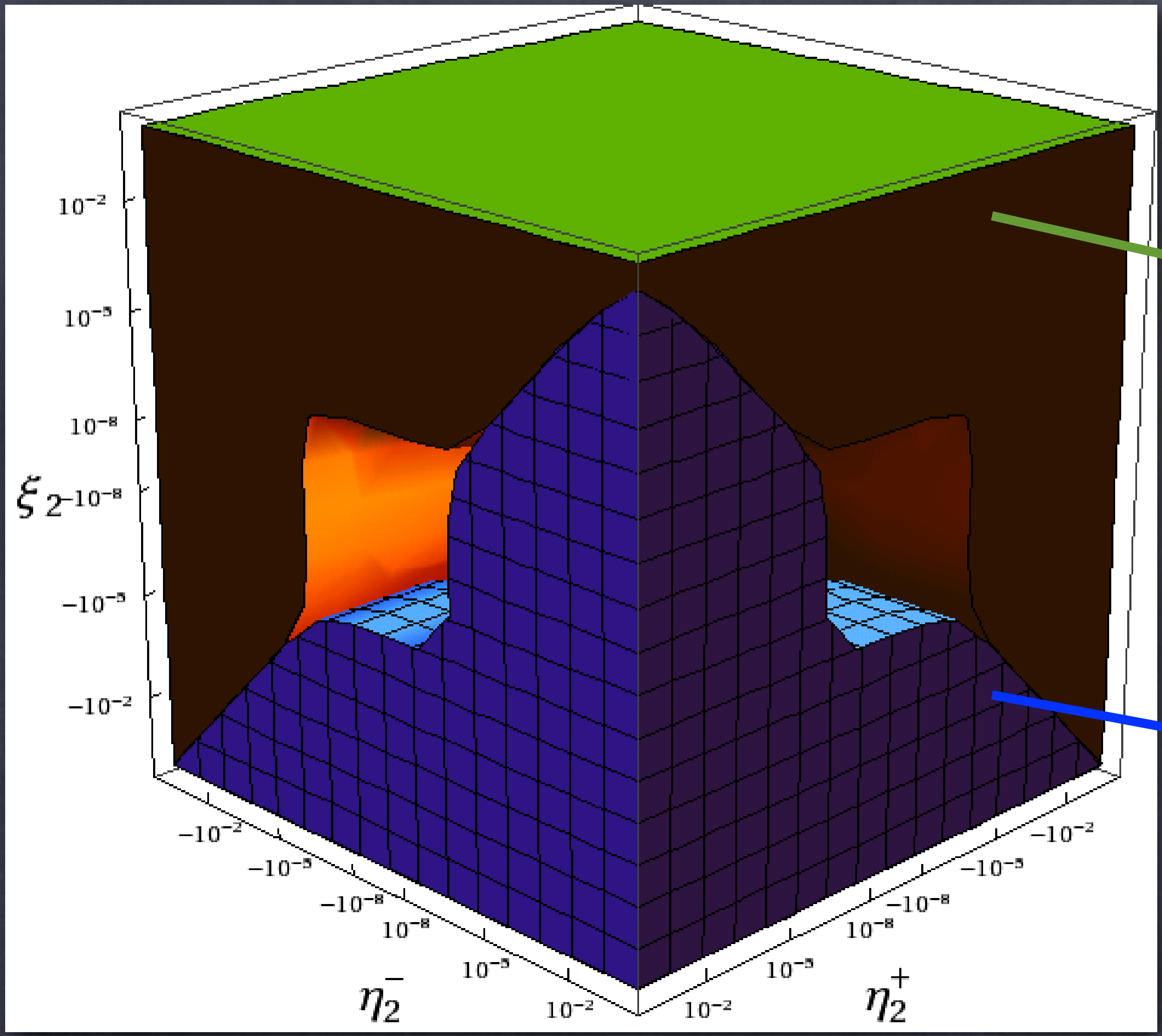
pair production may become inhibited, increasing GZK photon fluxes above observed upper limits: In the absence of LIV for electrons/positrons for n=1 (CPT-odd terms) this yields:

$$\xi_1 \leq 10^{-12}$$

Even for n=2 (CPT-even) one has sensitivity to $\xi_2 \sim 10^{-6}$

Such strong limits may indicate that Lorentz invariance violations are completely absent !

Constraints for n=2



UHE photons are detected

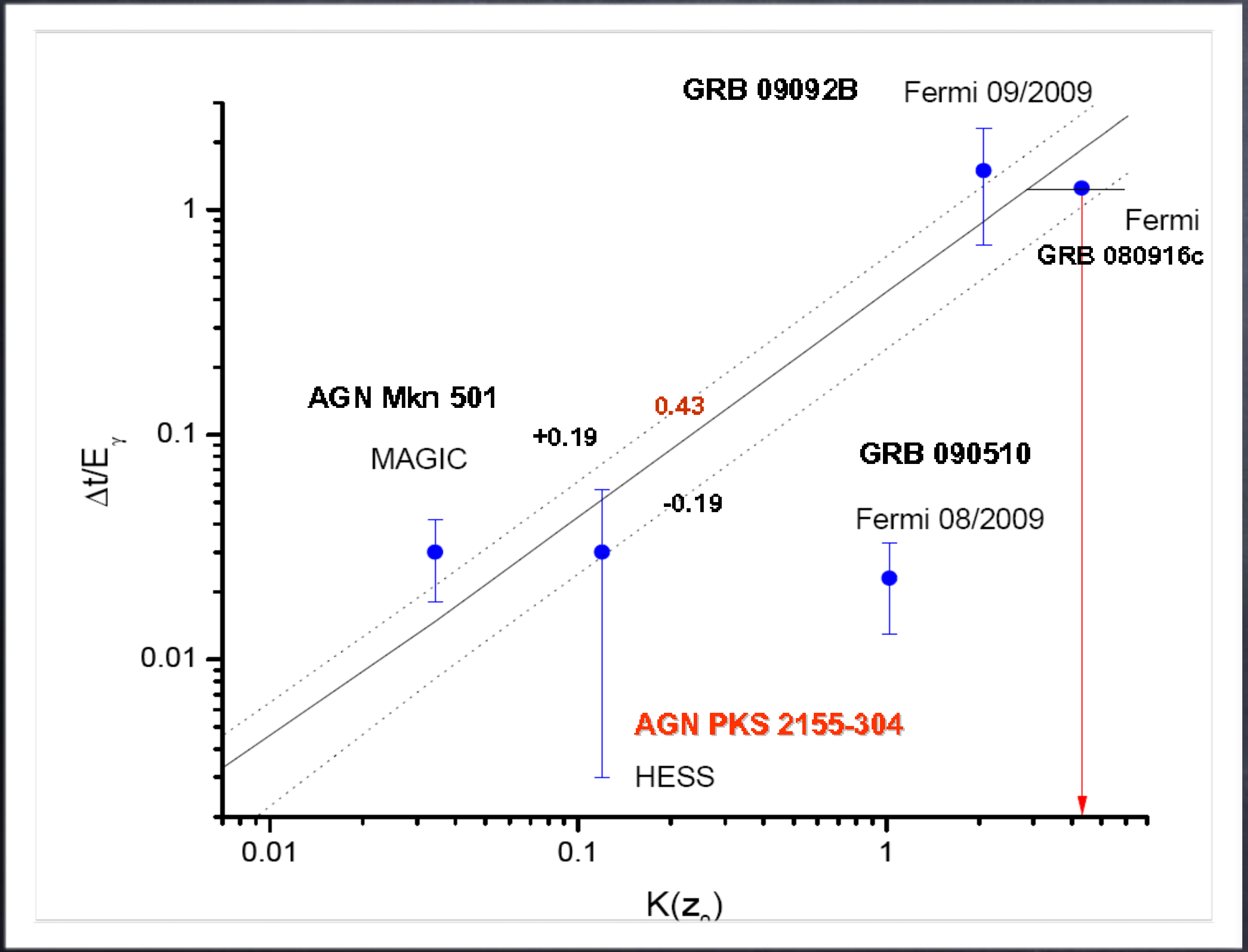
UHE photon absorption takes place

Such strong limits suggest that Lorentz invariance violations are completely absent !

The modified dispersion relation also leads to energy dependent group velocity $V=\partial E/\partial p$ and thus to an energy-dependent time delay over a distance d:

$$\Delta t = -\xi\, d \frac{E}{M_{\rm Pl}} \simeq -\xi \left(\frac{d}{100\,{\rm Mpc}}\right)\left(\frac{E}{\rm TeV}\right)\,{\rm sec}$$

for linearly suppressed terms. GRB observations in TeV γ-rays can therefore probe quantum gravity and may explain that higher energy photons tend to arrive later (Ellis et al.).

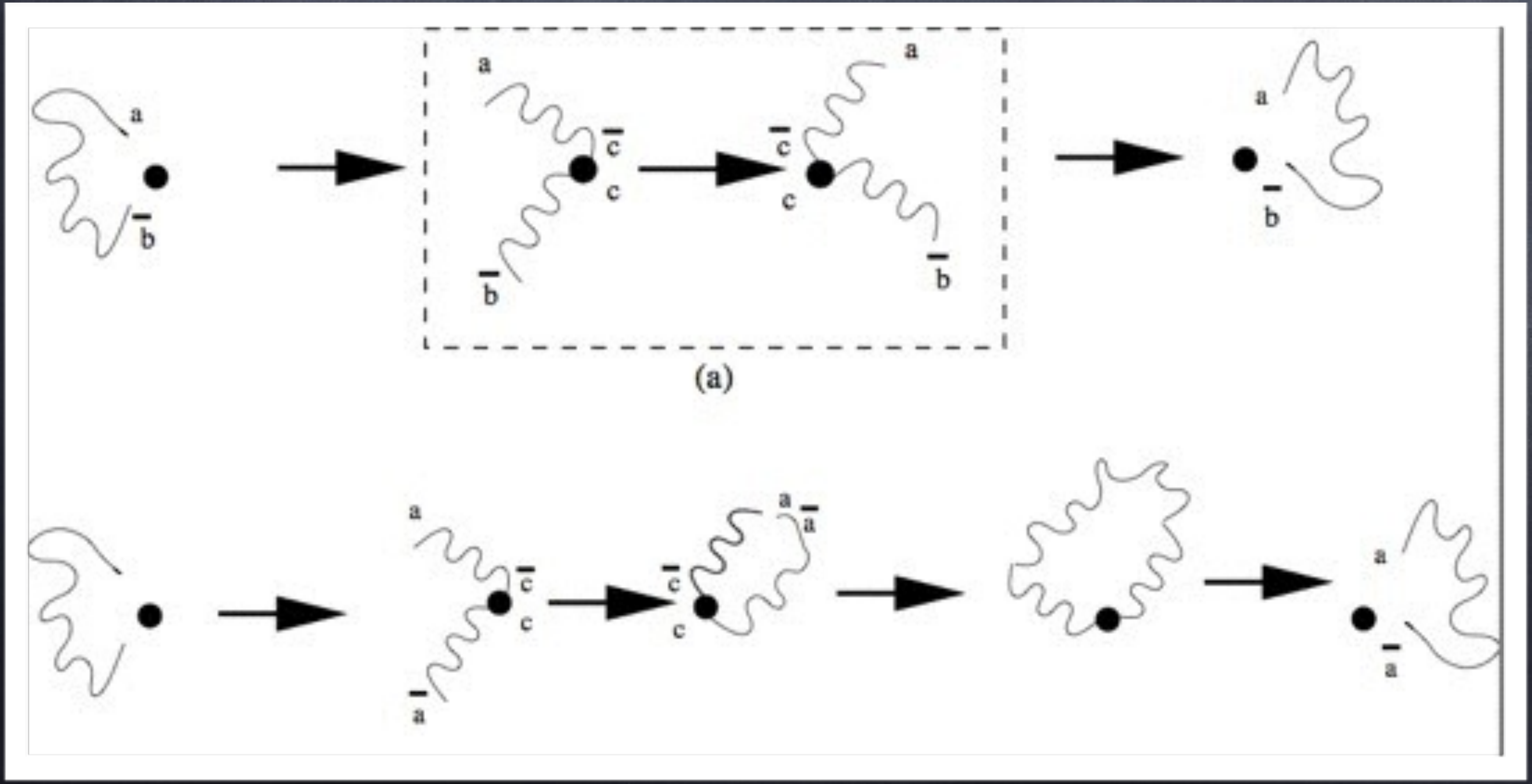


But the UHE photon limits are inconsistent with interpretations of time delays of high energy gamma-rays from GRBs within quantum gravity scenarios based on effective field theory

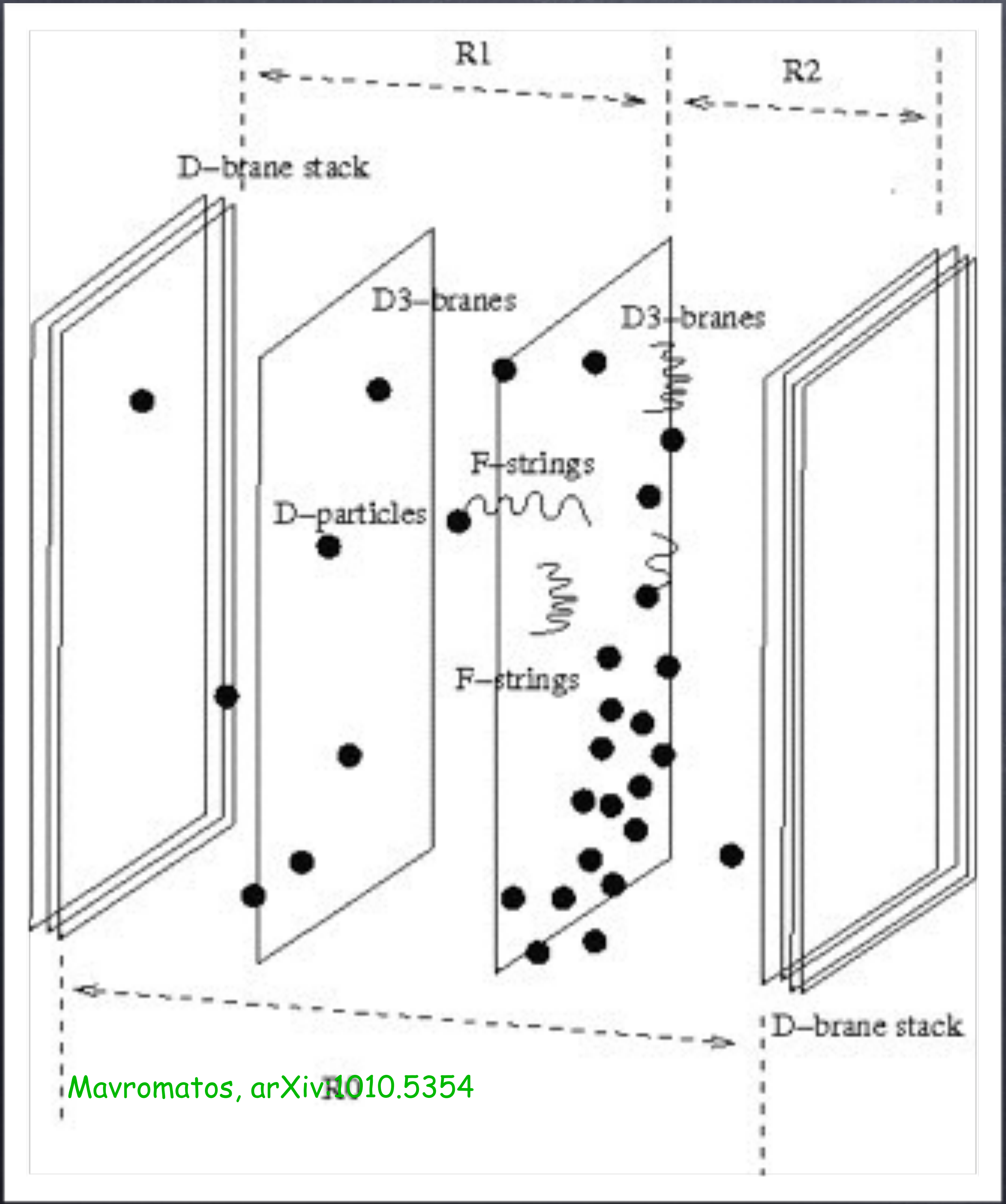
Maccione, Liberati, Sigl, PRL 105 (2010) 021101

Possible exception in space-time foam models,

Ellis, Mavromatos, Nanopoulos, arXiv:1004.4167

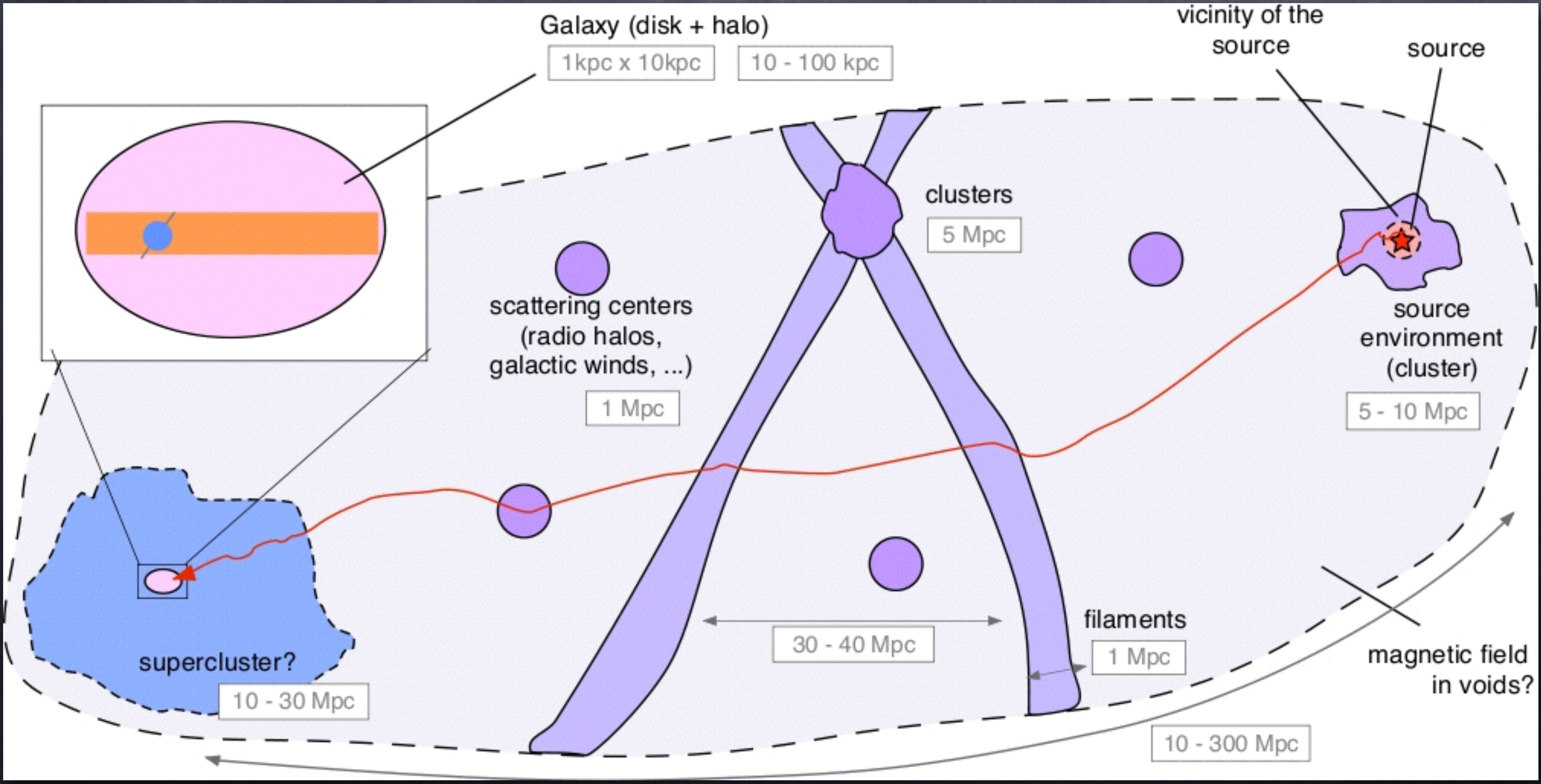


In space-time foam models there may be fluctuating terms in dispersion relation, thus no strict energy-momentum conservation. This could circumvent pair production limits, allowing to interpret time dispersion by quantum gravity effects



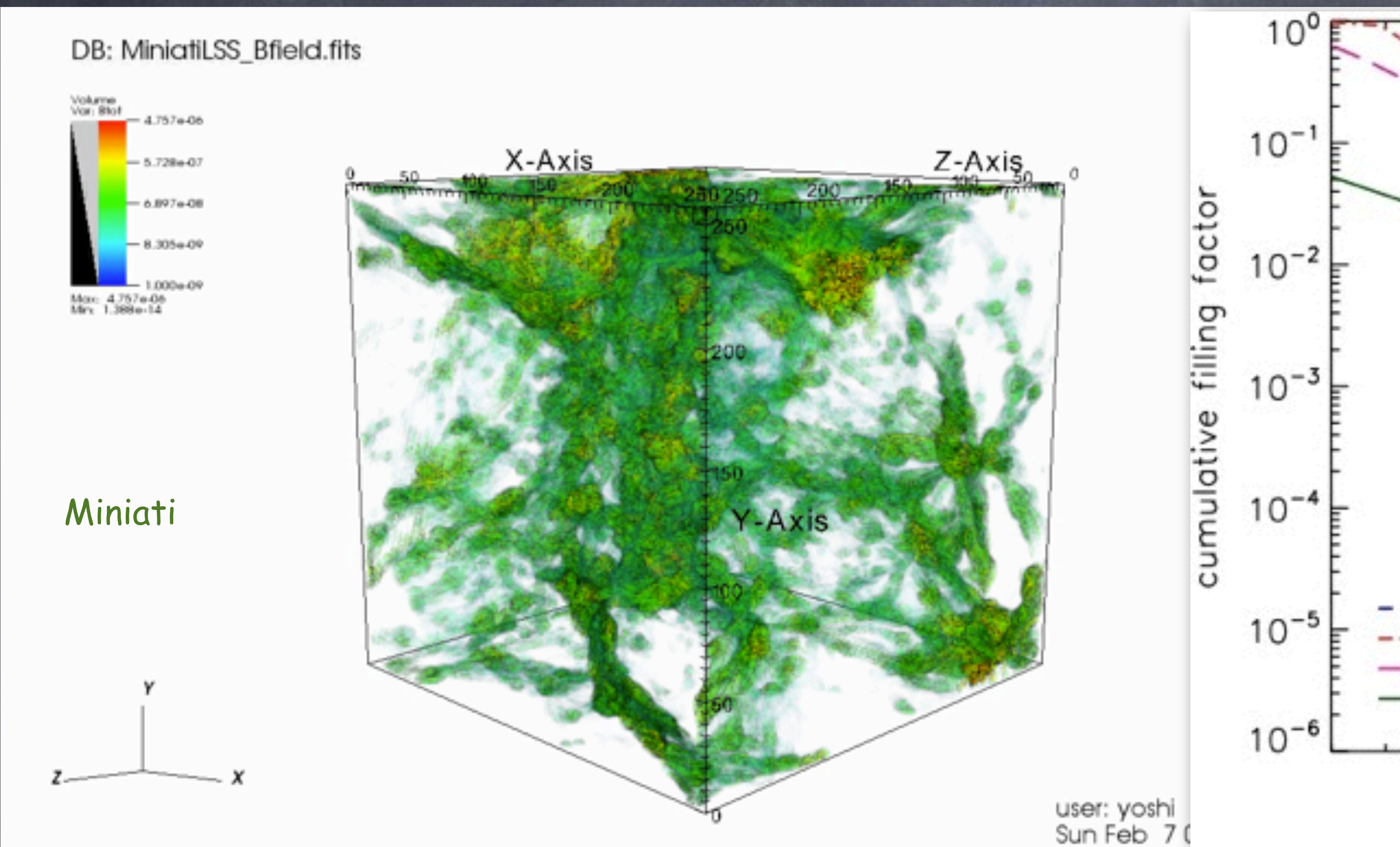
Mavromatos, arXiv:1010.5354

3-Dimensional Effects in Propagation

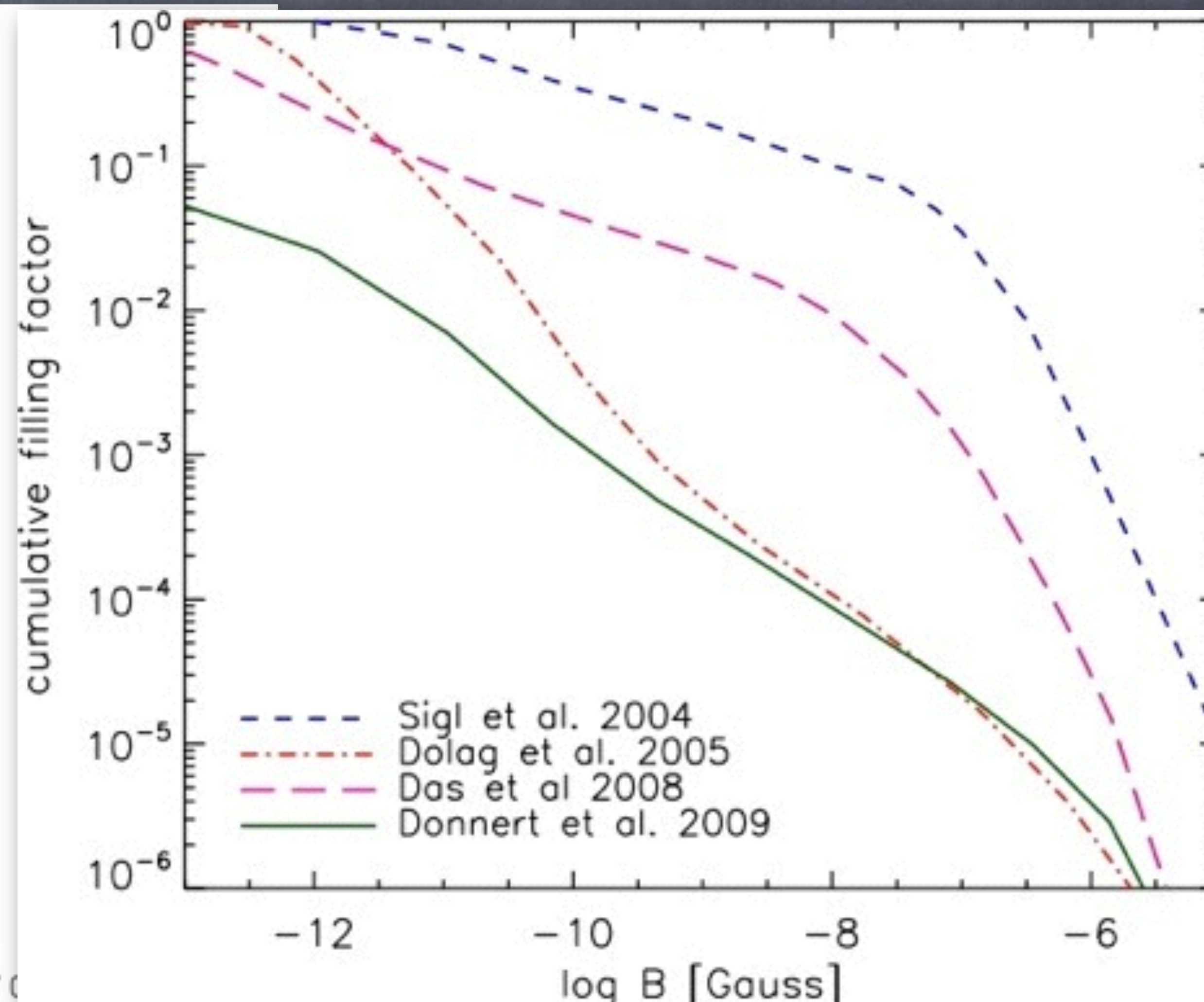


Kotera, Olinto, Ann.Rev.Astron.Astrophys. 49 (2011) 119

Structured Extragalactic Magnetic Fields



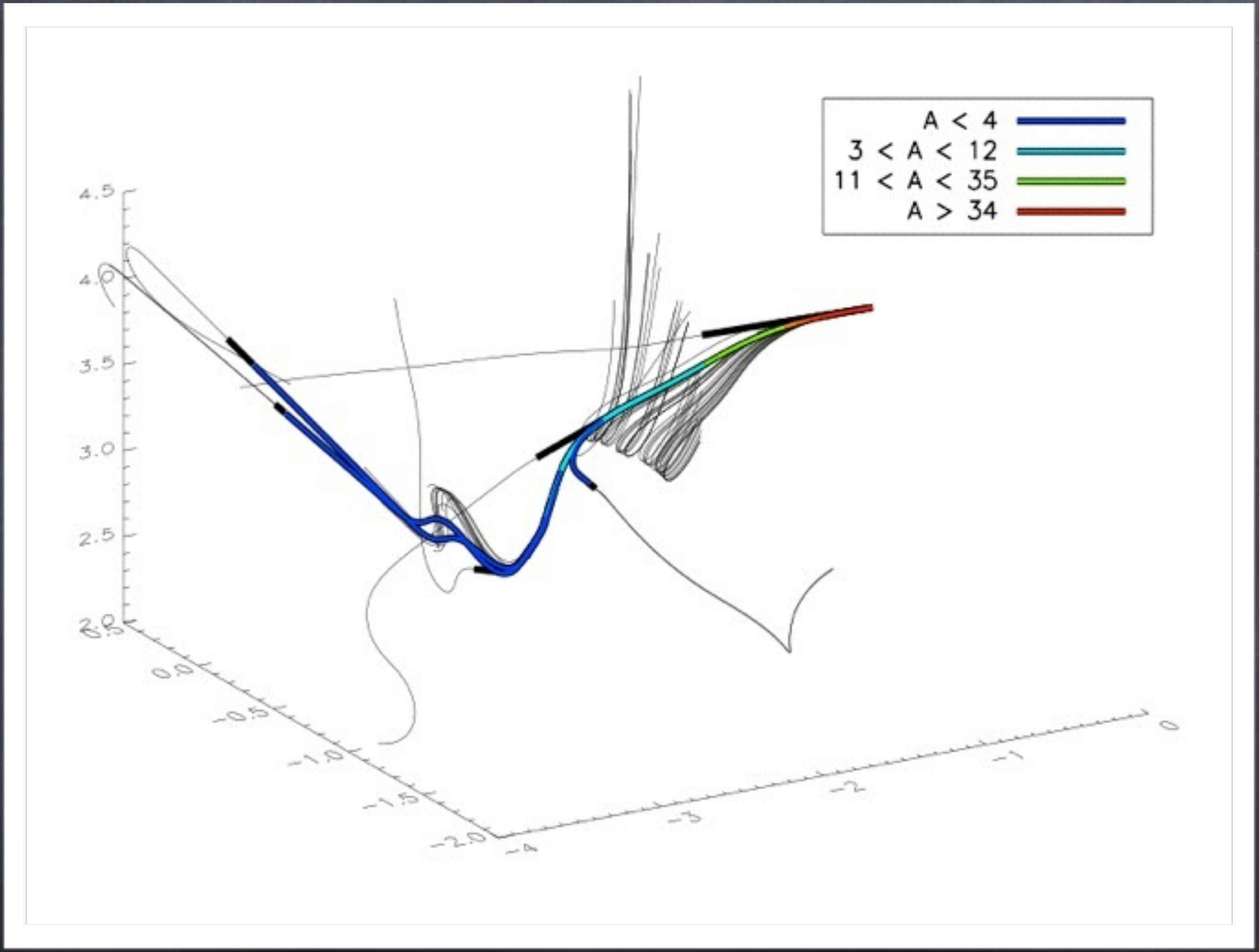
Miniati



Kotera, Olinto, Ann.Rev.Astron.Astrophys. 49 (2011) 119

Filling factors of extragalactic magnetic fields are not well known and come out different in different large scale structure simulations

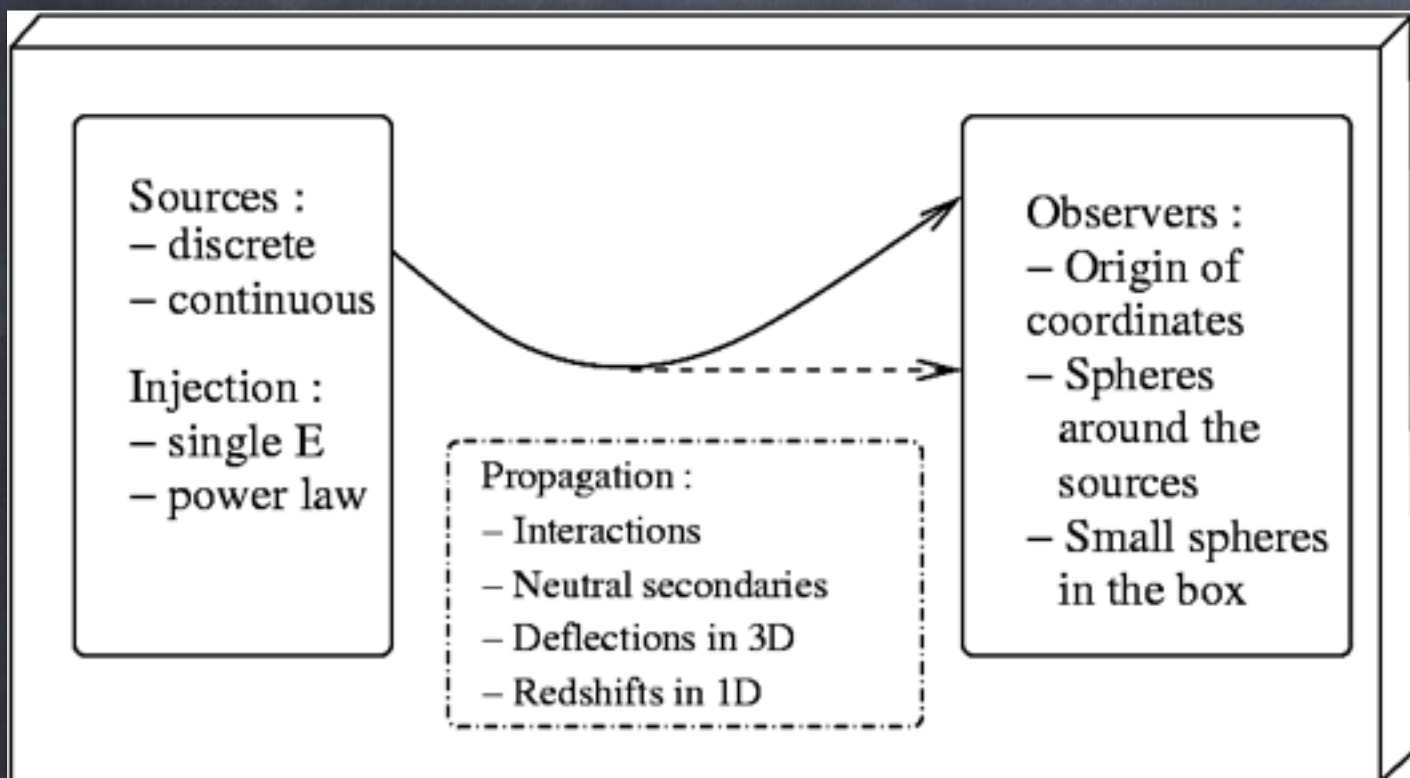
Extragalactic iron propagation produces nuclear cascades in structured magnetic fields:

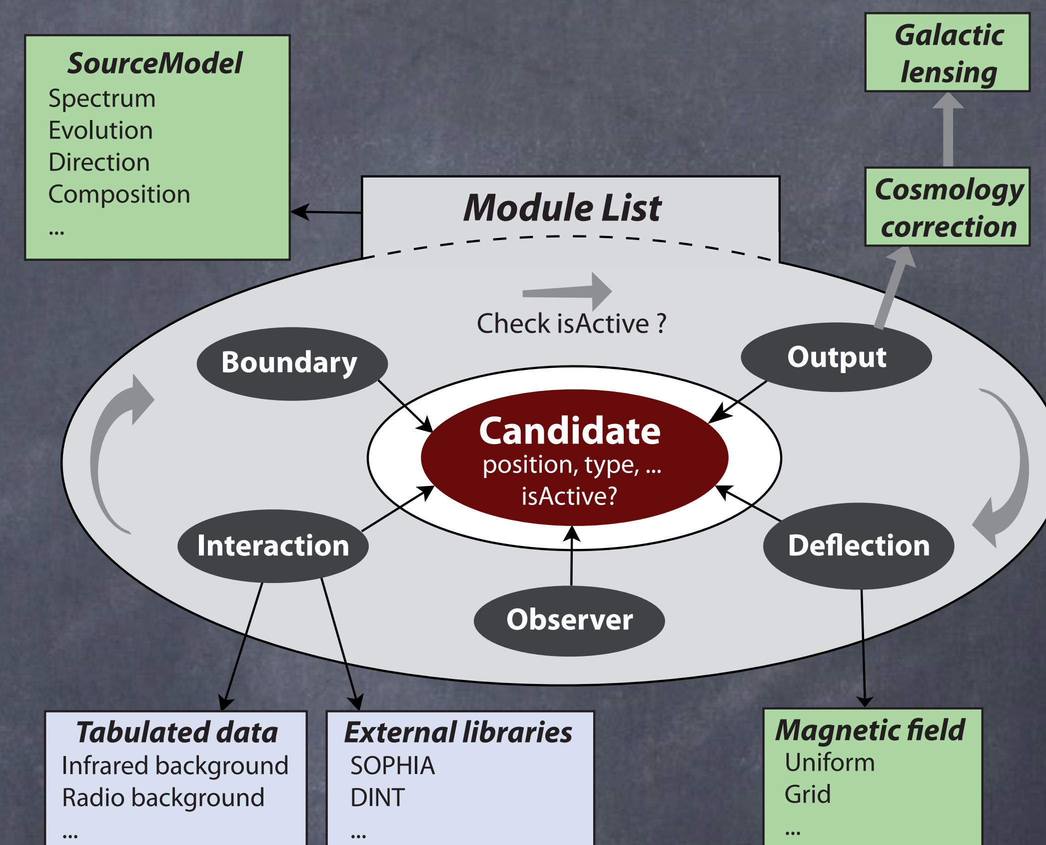


Initial energy 1.2 x 10^{21} eV, magnetic field range 10^{-15} to 10^{-6} G. Color-coded is the mass number of secondary nuclei

CRPropa 2.0/3.0

CRPropa is a public code for UHE cosmic rays, neutrinos and γ-rays being extended to heavy nuclei and hadronic interactions





Version 1.4: Eric Armengaud, Tristan Beau, Günter Sigl, Francesco Miniati, Astropart.Phys.28 (2007) 463.

Version 2.0 at https://crpropa.desy.de/Main_Page

Version 3.0: Luca Maccione, Rafael Alves Batista, David Walz, Gero Müller, Nils Nierstenhoefer, Karl-Heinz Kampert, Peter Schiffer, Arjen van Vliet Astroparticle Physics 42 (2013) 41

The main part of the code is written in C++ and calls some Fortran routines (mainly SOPHIA for interactions photo-pion production of nucleons) nuclear interactions based on TALYS

Electromagnetic cascades are treated by solving one-dimensional transport equations

The set-up (source distributions, environment, magnetic fields, low energy photon backgrounds, injection spectrum, arbitrary composition at fixed energy per nucleon, which interactions/secondaries to take into account) can be provided with xml files.

Output can be in form of whole trajectories or events; possible output formats are ASCII, FITS or ROOT.

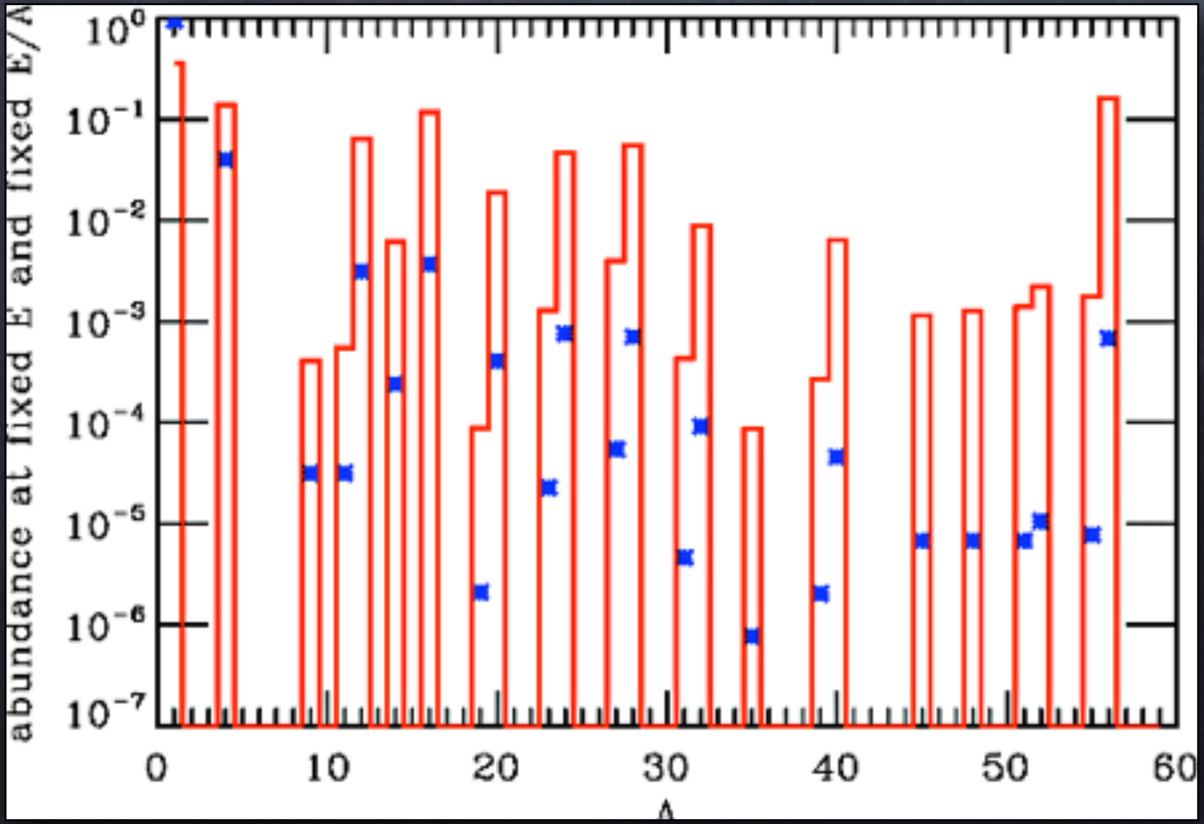
Presented are two examples for 1D and 3D simulations

Mixed mass compositions

For an injection spectrum $E^{-\alpha}$ elemental abundance at given energy E is modified to

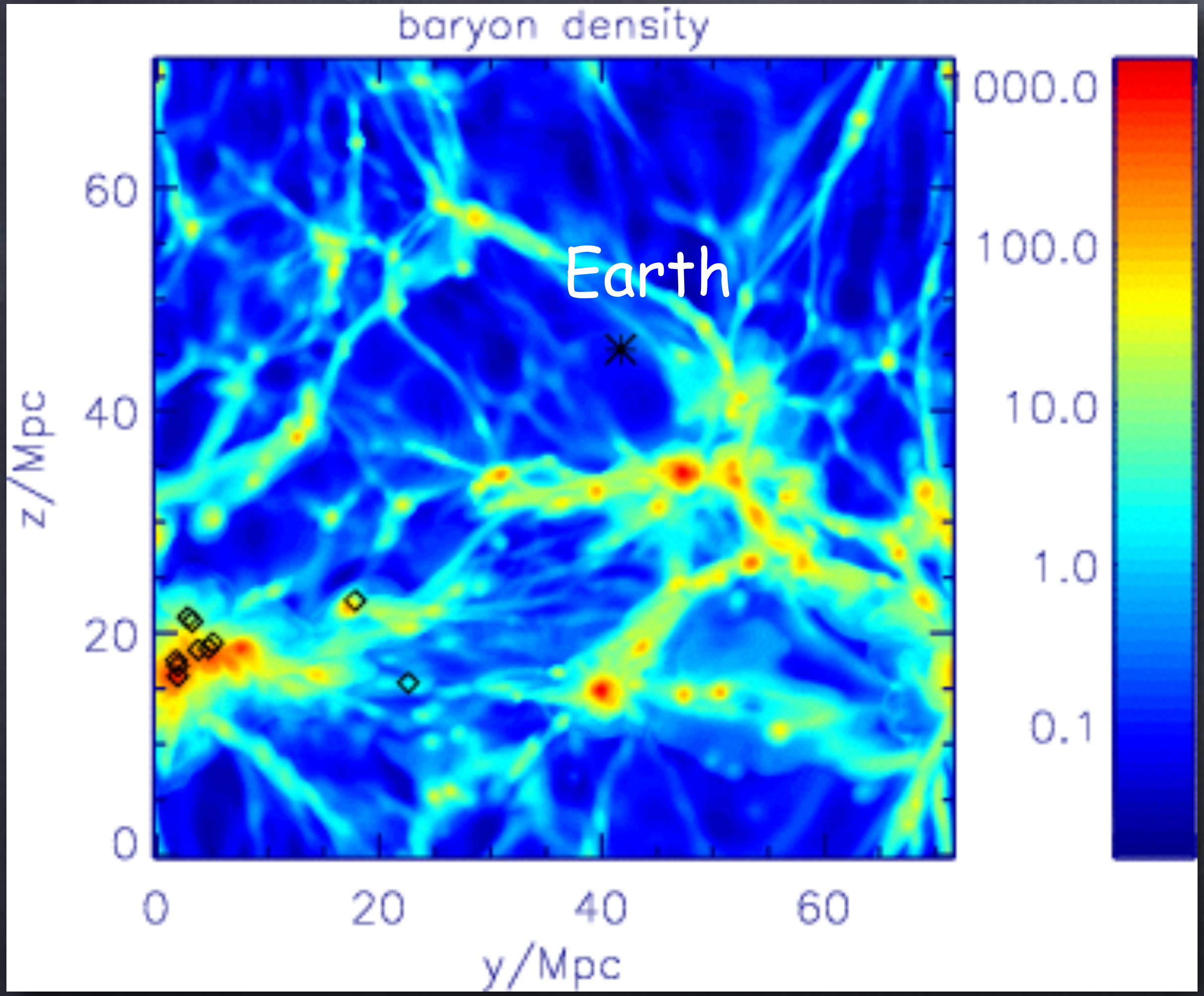
$$\frac{dn_A}{dE}(E) = N x_A \, A^{\alpha-1} \, E^{-\alpha} g(E)$$

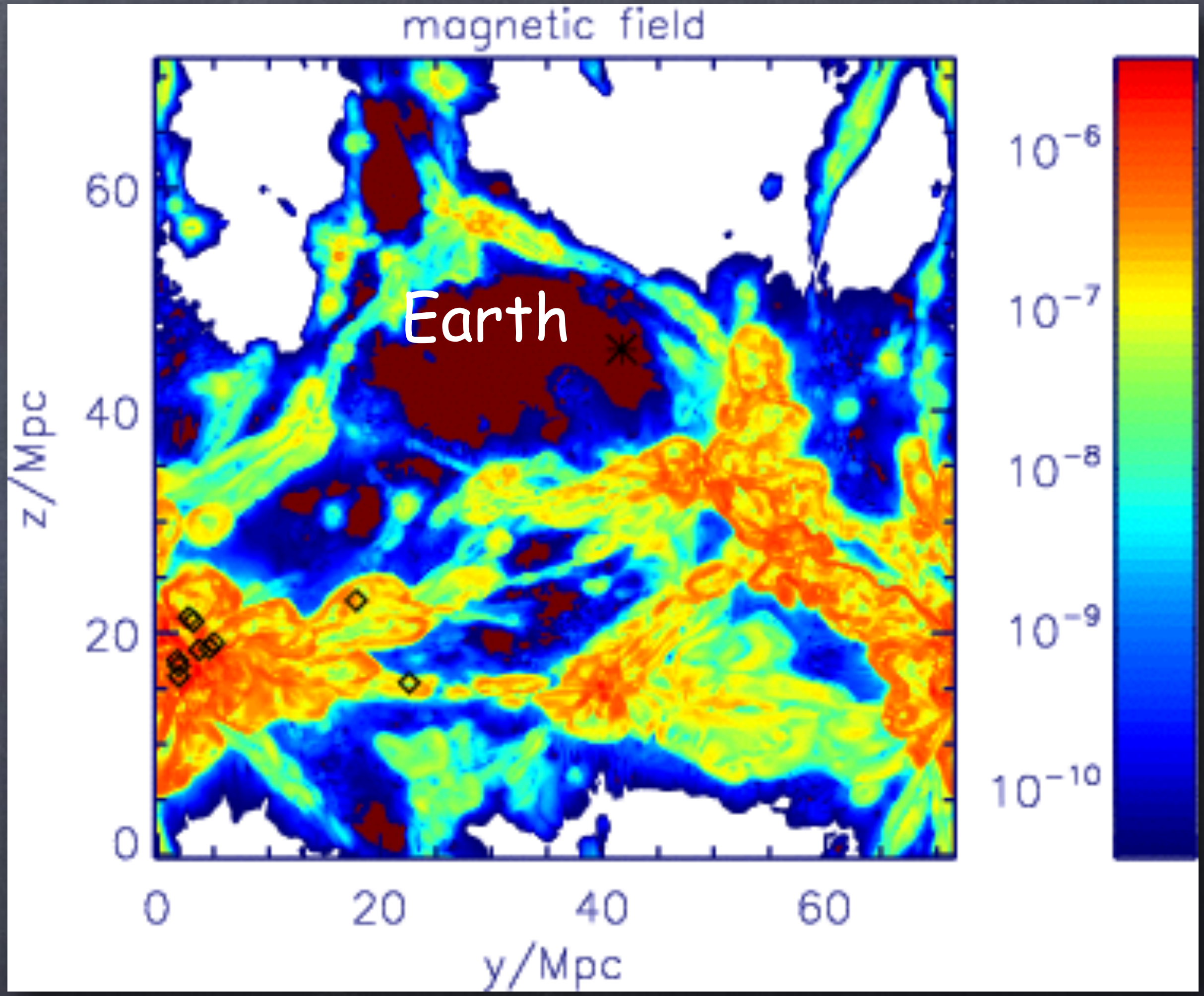
where x_A is the abundance at given energy per nucleon E/A and g(E) is the cut-off shape.



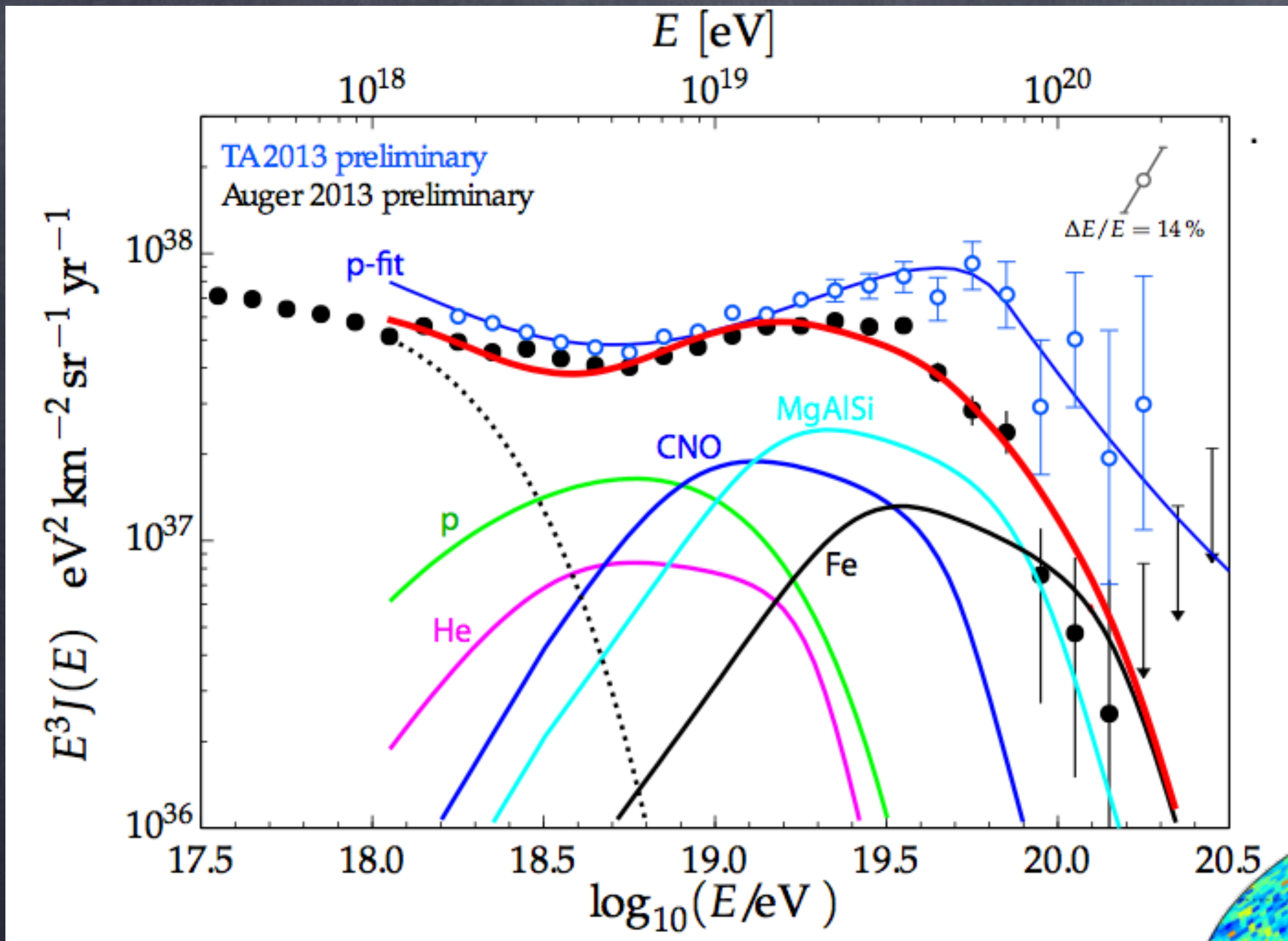
Composition at given E/A (blue) following elemental abundances in the Galaxy
Composition at given E for an $E^{-2.6}$ injection spectrum (red).

Discrete Sources in nearby large scale structure





Building Benchmark Scenarios

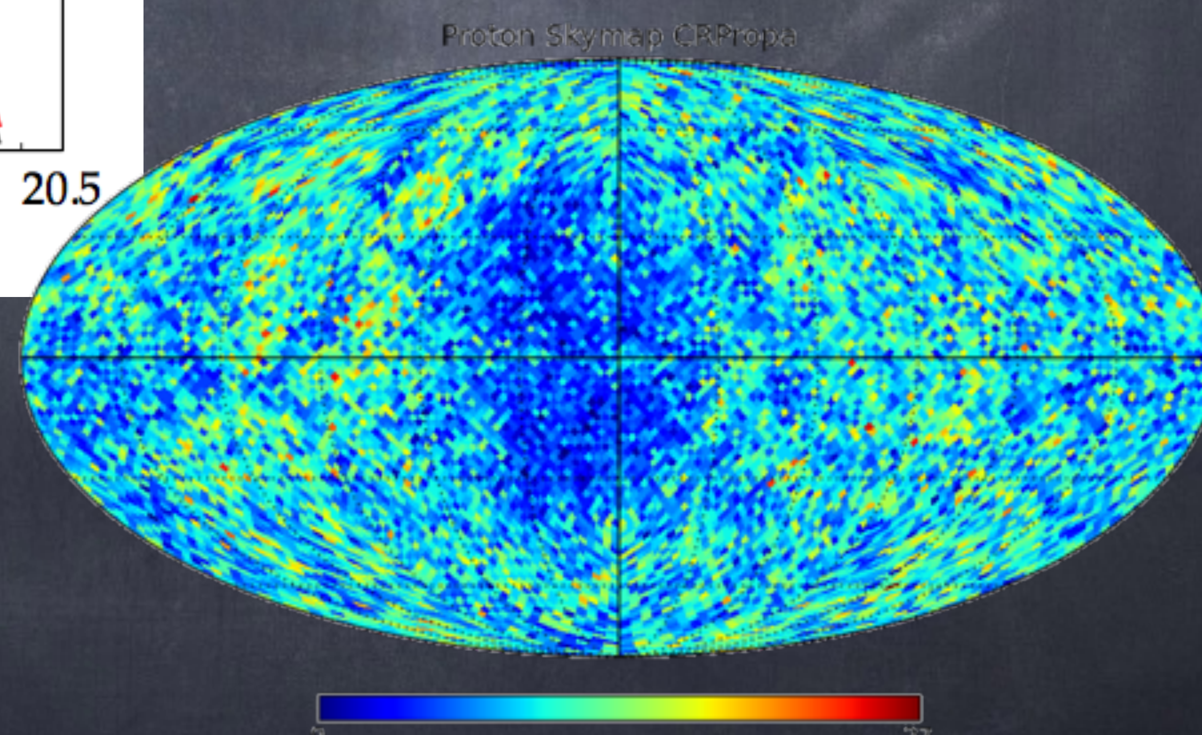


TA fit (red): pure proton injection

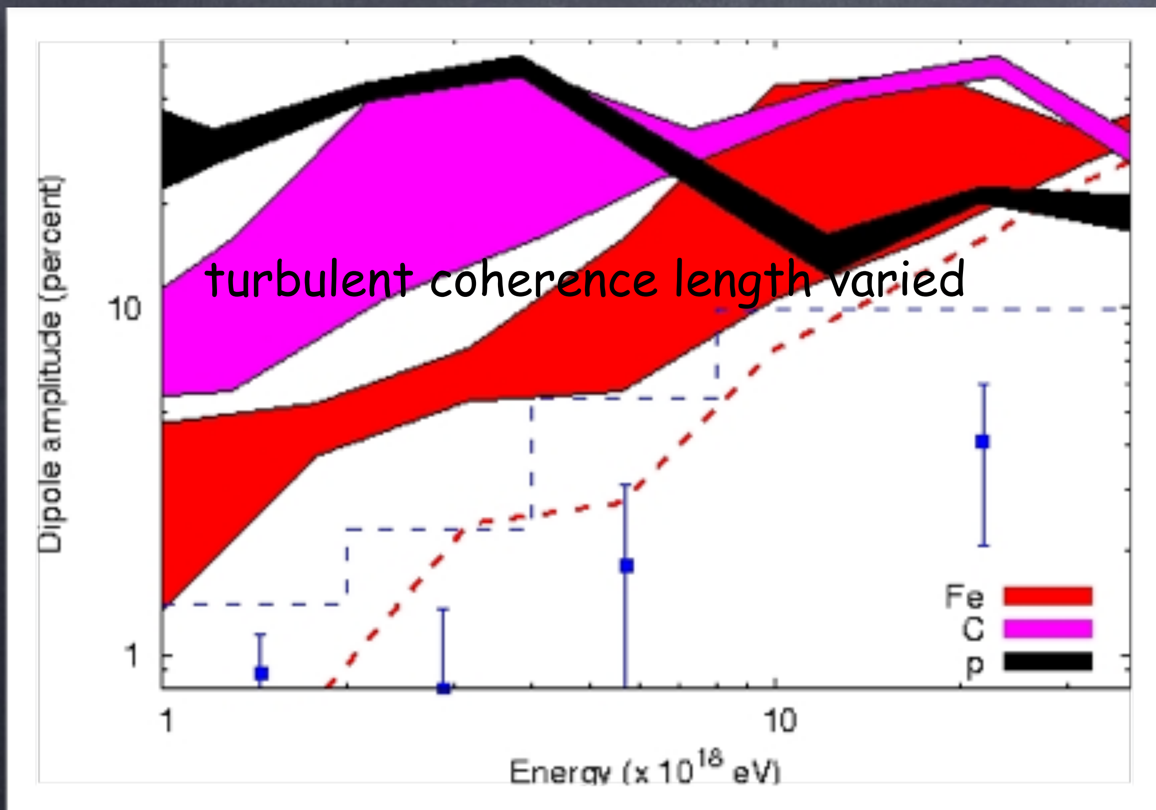
rate $\propto$ $(1+z)^{4.4}$ $E^{-2.36}$

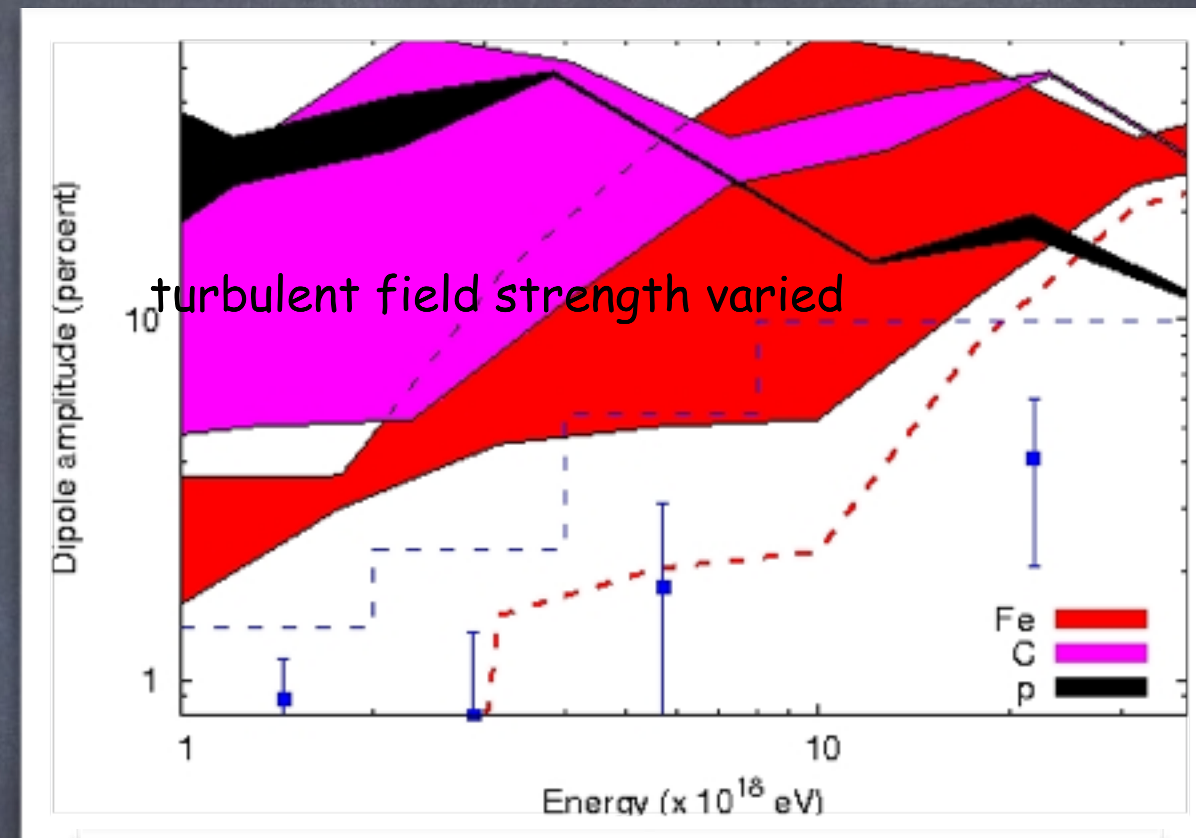
Auger fit: enhanced galactic

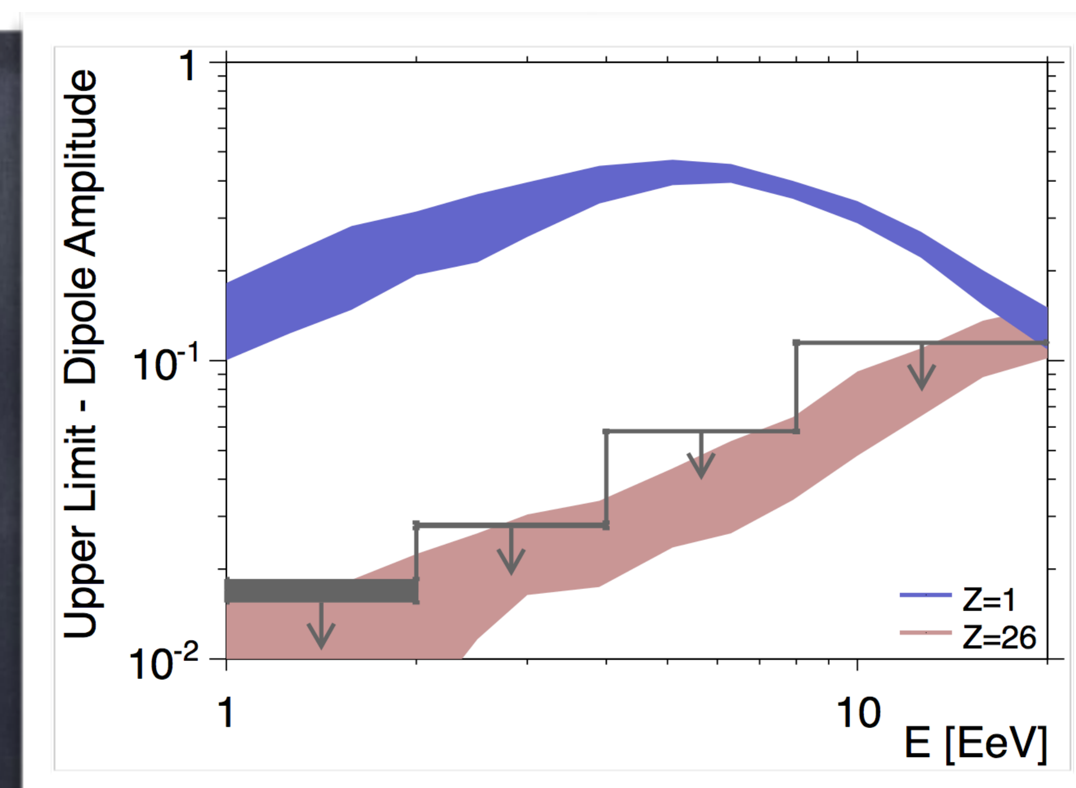
composition $\propto$ $E^{-1.8}$ up to $10^{18.7}$ eV*Z



Composition and the Transition Galactic/Extragalactic Cosmic Rays







Giacinti, Kachelriess, Semikoz, Sigl,
JCAP 07 (2012) 031
and Pierre Auger Collaboration, Astrophys.J. 762 (2012) L13

Light Galactic Nuclei produce too much anisotropy above $\simeq 10^{18}$ eV. This implies:

1.) if composition around 10^{18} eV is light => probably extragalactic (and ankle may be due to pair production by protons)

2.) if composition around 10^{18} eV is heavy => transition could be at the ankle if Galactic nuclei are produced by sufficiently frequent transients, e.g. magnetars

It is surprisingly difficult to construct simple scenarios with structured sources and magnetic fields that reproduce all observations: spectra, energy dependent composition and anisotropy; to explain them separately is quite easy

Relatively hard injection spectra and low maximal rigidities of few times 10^{18} eV seem to be favored

Conclusions

1.) It is surprisingly difficult to construct simple scenarios with structured sources and magnetic fields that reproduce all observations: spectra, energy dependent composition and anisotropy; to explain them separately is quite easy

2.) The observed X_{max} distribution of air showers provides potential constraints on hadronic interaction models: Some models are in tension even when "optimizing" unknown mass composition; however, systematic uncertainties are still high.

Conclusions

3.) Both diffuse cosmogenic neutrino and photon fluxes mostly depend on mass composition, maximal acceleration energy and redshift evolution of sources

4.) Multi-messenger modeling sources including gamma-rays and neutrinos start to constrain the source and acceleration mechanisms

5.) Highest Energy Cosmic Rays, Gamma-rays, and Neutrinos give the strongest constraints on violations of Lorentz symmetry => terms suppressed to first and second order in the Planck mass would have to be unnaturally small