Self-Organised Localisation

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Personal view: Do model builders need to think about Eternal Inflation… again…

Inflation

The stochastic approach to inflation developed by Starobinsky and others has been a useful guide.

QFT calculations have confirmed many aspects of this approach as a leading order picture, in the pre-eternal regime.

(See recent papers by Gorbenko, Senatore, and Cohen, Green, Premkumar, Ridgway!)

Still, however, much work remaining...

Inflation

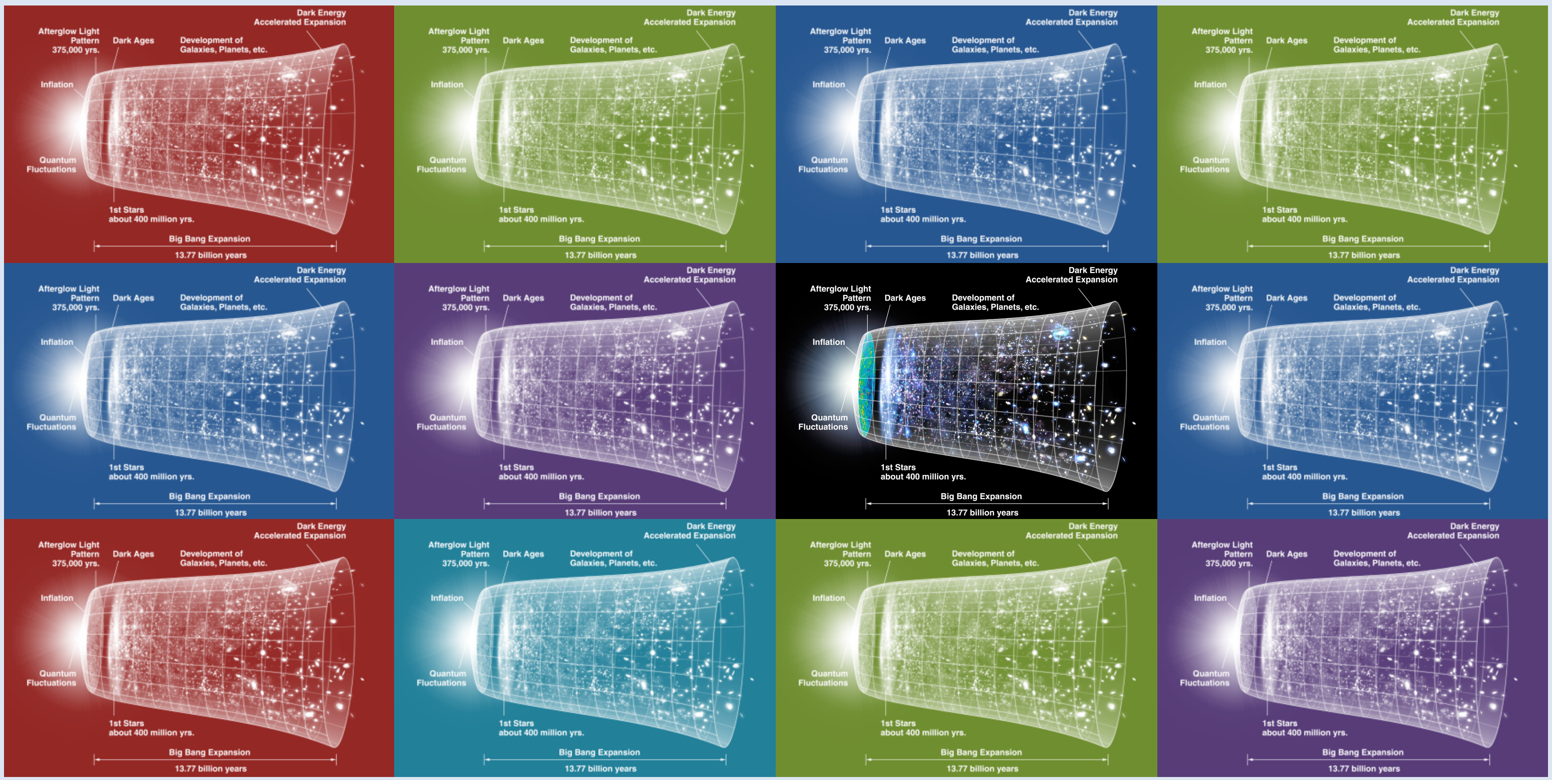
However the stochastic approach has also raised some very deep puzzles.

Eternal inflation appears to be a relatively generic phenomenon…

Why should particle physicists care?

Inflation

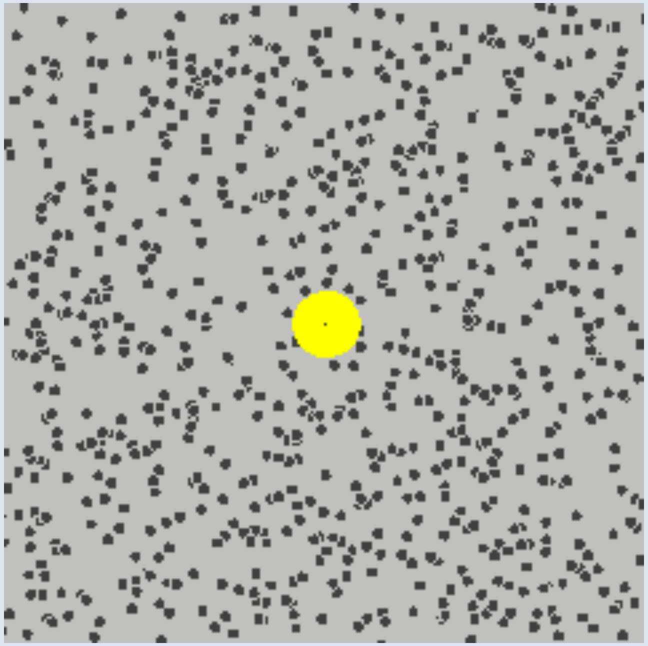
Eternal inflation leads to a multiverse of different Universes. We are but one...



In each one, different parameters, forces...

Inflation

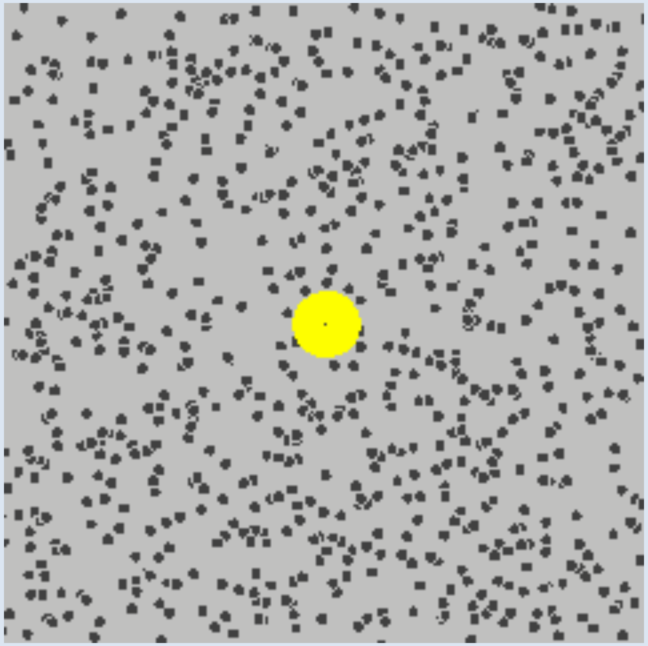
Suppose you have a box of gas and you measure the velocity of one atom, once.



Is that value of velocity likely, or unlikely?

Inflation

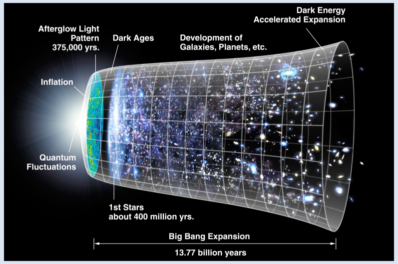
Suppose you have a box of gas and you measure the velocity of one atom, once.



If we know the properties of the statistical ensemble at equilibrium, we have context.

Inflation

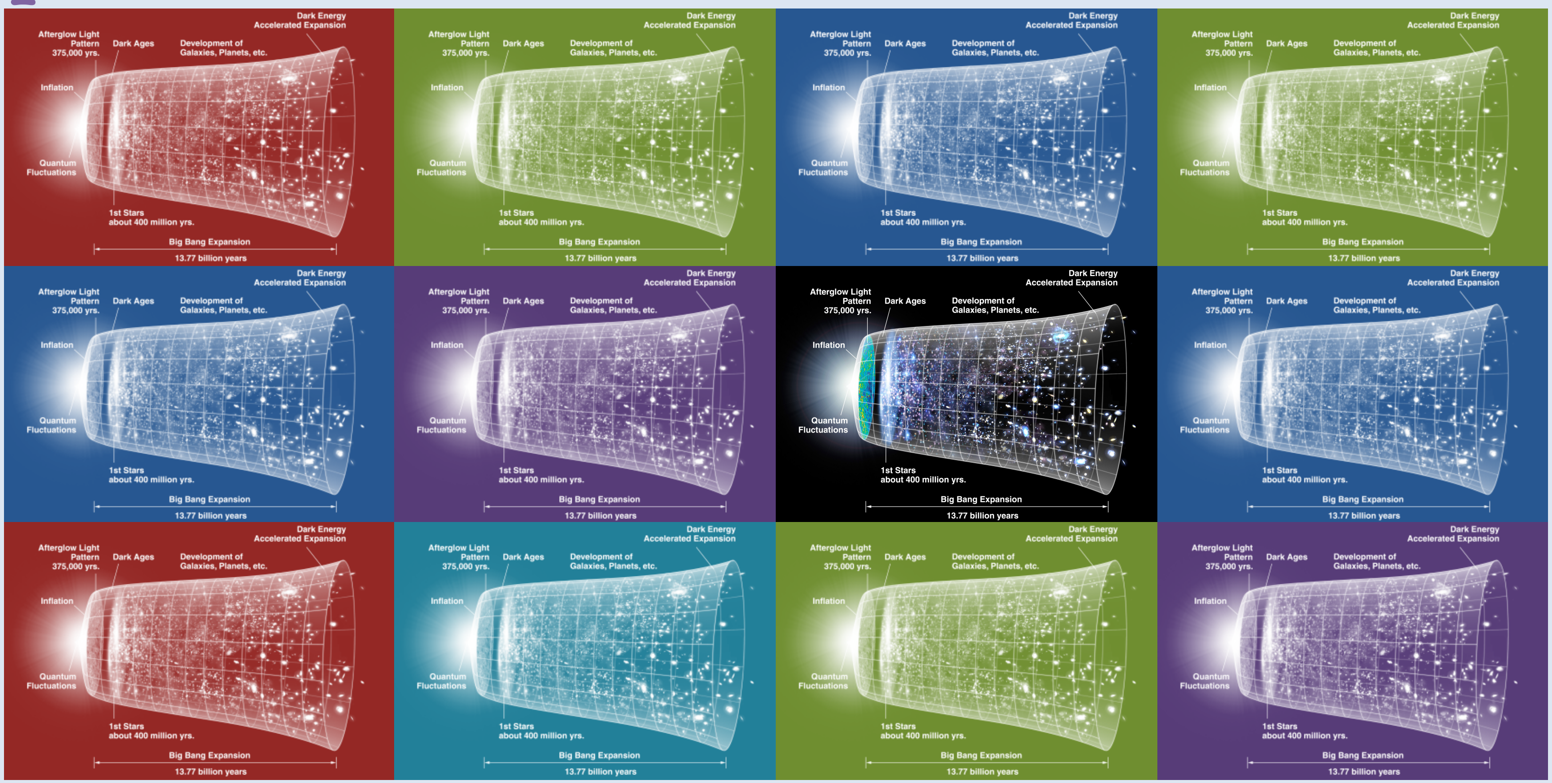
If there is a multiverse in which parameters, forces, etc are scanned then by measuring SM parameters...



how can we know if they are likely, unlikely, tuned, etc? Anthropics...?

Inflation

Instead, we need to know the macroscopic properties of the statistical ensemble



to assign context to parameters.

Towards quasi-statistics?

The Good

The stochastic approach offers possibility of estimating macroscopic parameters.

$$\frac{\partial}{\partial\phi}\left[\frac{\hbar}{8\pi^2}\frac{\partial(H^3\,P_{\rm FP})}{\partial\phi}+\frac{V'\,P_{\rm FP}}{3H}\right]=\frac{\partial P_{\rm FP}}{\partial t}$$

"Quantum"
Gaussian Noise

EOM
Drift

Time
Evolution

Light scalar fields follow a Langevin-like trajectory. Average of trajectories described by a Fokker-Planck equation.

The Good

We're interested in the volume distribution

$$\frac{\partial}{\partial\phi}\left[\frac{\hbar}{8\pi^2}\frac{\partial(H^3\,P)}{\partial\phi}+\frac{V'\,P}{3H}\right]+3H\,P=\frac{\partial P}{\partial t}$$

"Quantum" Gaussian Noise — EOM Drift — Volume expansion — Time Evolution

Light scalar fields follow a Langevin-like trajectory. Volume average of trajectories described by a Fokker-Planck-like equation.

The Good

We're interested in the volume distribution

$$\frac{\partial}{\partial\ldots}\left[\ldots\frac{\partial(H^3 P)}{\ldots} + \frac{V' P}{\ldots H}\right] + 3HP = \frac{\partial P}{\partial t}$$

We're interested in the equilibrium solutions, where time-dependence is trivial $\dot{P} \propto P$, since these are the final states, independent of initial conditions.

Scalar field fo... trajectory. Volume average... described by a Fokker-Planck-like equation.

The Good

We're interested in the volume distribution

$$\ldots\partial(H^3 P) \ldots + \frac{V' P}{\ldots}\Big] + 3HP = \frac{\partial P}{\partial t}$$

Boundary conditions are just as much a part of the theory as the equation itself. I won't always explicitly state them, but ask at any point about sensitivity to BCs.

Scalar field follo... trajectory. Volume average ... described by a Fokker-Planck-like equation.

Questions / Issues

The Good?

In our applications to reach steady state we have to wait for

$$N > S_{\rm dS} = \frac{8\pi^2 M_P^2}{\hbar\, H^2}$$

e-foldings.

This necessarily implies eternal inflation. (See e.g. Arkani-Hamed, Dubovsky, Senatore, Villadoro).

The Bad

We have to choose a clock:

$$dt_\xi/dt = (H/H_0)^{1-\xi}$$

Because Hubble depends on the scalar field, the clock can involve field dependence. For example, evolution w.r.t. proper time is not the same as scale factor.

$$\frac{\partial}{\partial\phi}\left[\frac{\hbar}{8\pi^2}\frac{\partial(H^{2+\xi}\,P)}{\partial\phi} + \frac{V'\,P}{3H^{2-\xi}}\right] + 3H^\xi P = H_0^{\xi-1}\,\frac{\partial P}{\partial t_\xi}$$

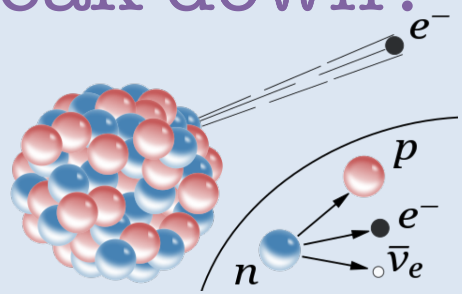
Worse

Quantum corrections accumulate to large effect on classical system, essentially because field climbs the potential. Does semi-classics break down?

Similar to the page time for BHs. For BHs it seems semi-classics remains valid, but new field configurations may be important.

Worse

Quantum corrections accumulate to large effect on classical system, essentially because field climbs the potential. Does semi-classics break down?



Take a large quantity of radioactive material that beta-decays. EFT valid for each decay, yet after long enough system is completely changed.

Even Worse

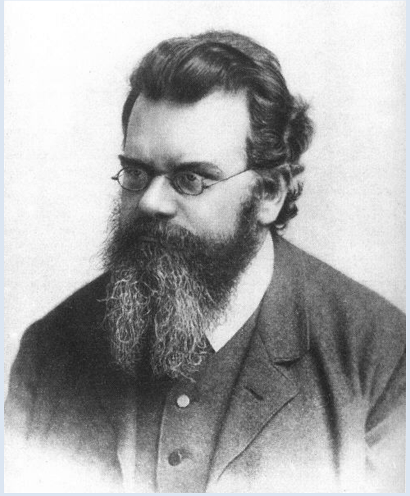
In inflation reheating occurs when the inflaton passes the “reheating surface”. Can perform statistics by studying this surface.

$$\frac{\infty}{\infty}$$

But in eternal inflation reheating surface is infinite. Must be regulated somehow. Results depend on measure...

The Worst

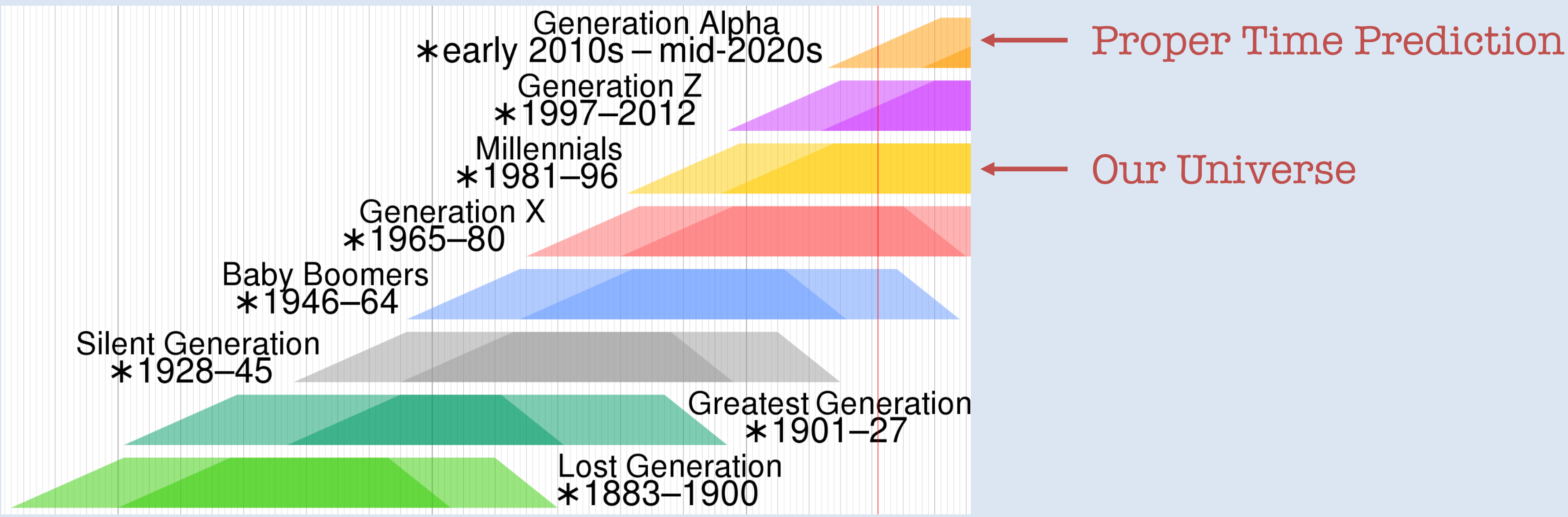
Since we don't commit to a specific inflationary model, we take proper time as our time-slicing measure.



Issues such as Boltzmann brains arise in this case. May not apply to our applications though, as all our Universes are unstable...

The Worst

The youngness paradox is much more severe. Emphasised to us by Andrei Linde.



Universe should be much younger and hotter if proper time cutoff naively extrapolated to our time.

The Ugly

Our scenario (and most of inflationary parameter space) is in The Swampland.

De Sitter conjecture is:

$$|\nabla V| > c\frac{V}{M_P}$$

Slow-roll parameter is:

$$\epsilon = \frac{M_P^2\, V'^2}{2\, V^2}$$

Clear tension...

The Ugly

Our scenario (and most of inflationary parameter space) is in The Swampland.
D conjecture is:

V

Also, in all our models, the field range is super-Planckian, violating distance conjecture too.

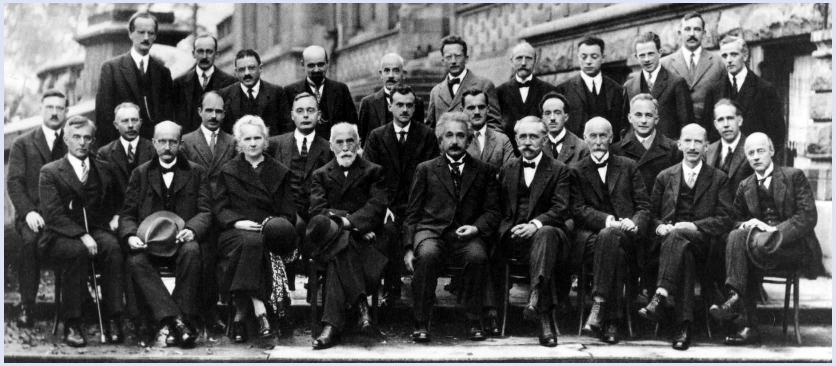
Slow-roll parame

$$\epsilon = \frac{M_P^2\, V'^2}{2\, V^2}$$

Clear tension…

Personal View

Past breakthroughs were made by venturing into incomplete frameworks, often even involving unregulated infinities. QM, QFT...



Given the success of symmetry-based approaches in taking us beyond the SM, perhaps we need to spend more time in uncharted territory?

Nature 😍 Criticality

Puzzle 1

A Stable Universe?

It is as if there is some additional piece in the potential for the entire Universe that knew in advance, before the electroweak phase had even happened, to precisely cancel the Higgs contribution

$$V = V_H + V_0$$

where $$\tau \sim \frac{\sqrt{3}M_P}{\sqrt{V_H + V_0}} \sim 10^{18}\,\mathrm{s}$$

The Universe is delicately and calmly balanced between two violent phases. Why?

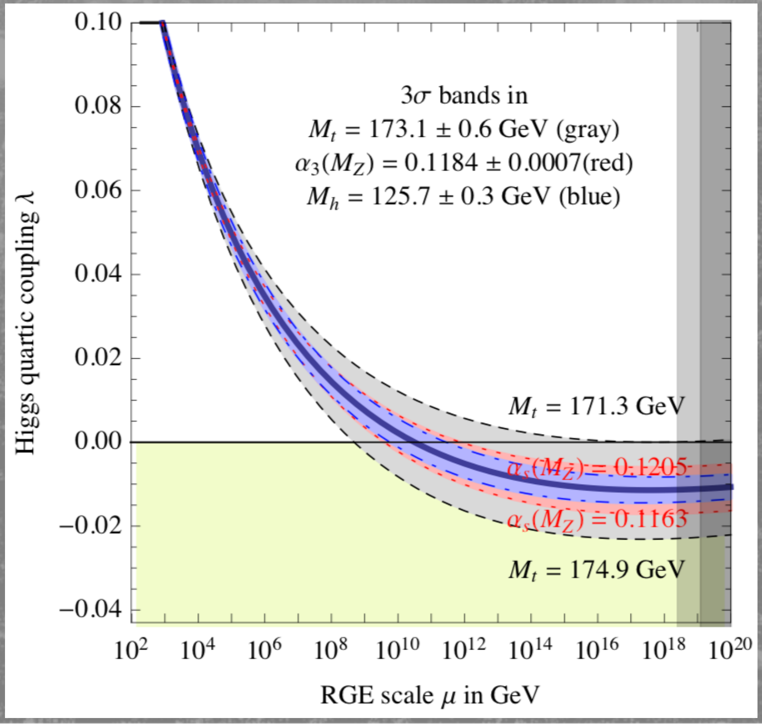
Puzzle 2

A Metastable Universe?

The Higgs potential also depends on the Higgs itself, due to quantum mechanics:

$$V_H = M^2(H)|H|^2 + \lambda(H)|H|^4$$

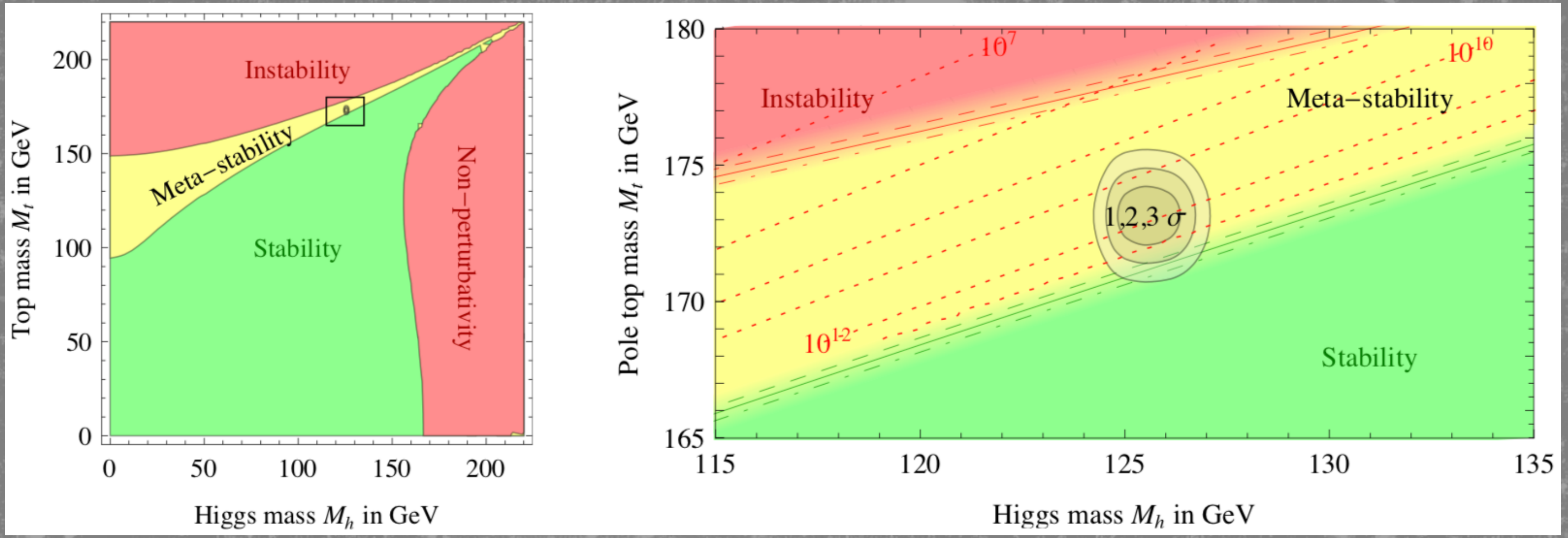
The Higgs quartic interaction effectively turns negative at large field values:



Taken from 1205.6497

A Metastable Universe?

This means that the vacuum we are in, as in the Mexican hat pictures, is just local, but there is a deeper one out at large field values.



The fundamental parameters we have measured in our "room" imply that nature is delicately balanced at a critical point where two Higgs phases may coexist.

Puzzle 3

Critical Higgs

Consider pions:

$$\pi \longleftrightarrow \langle \bar{q} q \rangle$$

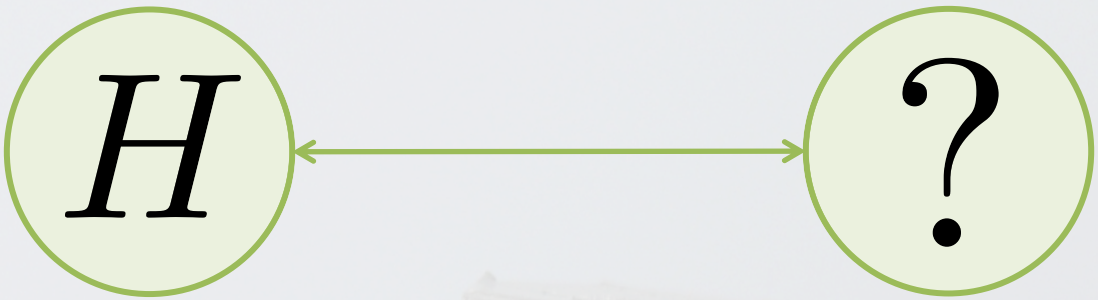
The order parameter for the condensate and the pion mass are both calculable in terms of microscopic theory

$$f_\pi \sim \frac{\Lambda_{\rm QCD}}{g_\star} \qquad\qquad m_\pi^2 \sim m_q \Lambda_{\rm QCD}$$

and both follow typical symmetries + scales.

Critical Higgs

What about the Higgs?

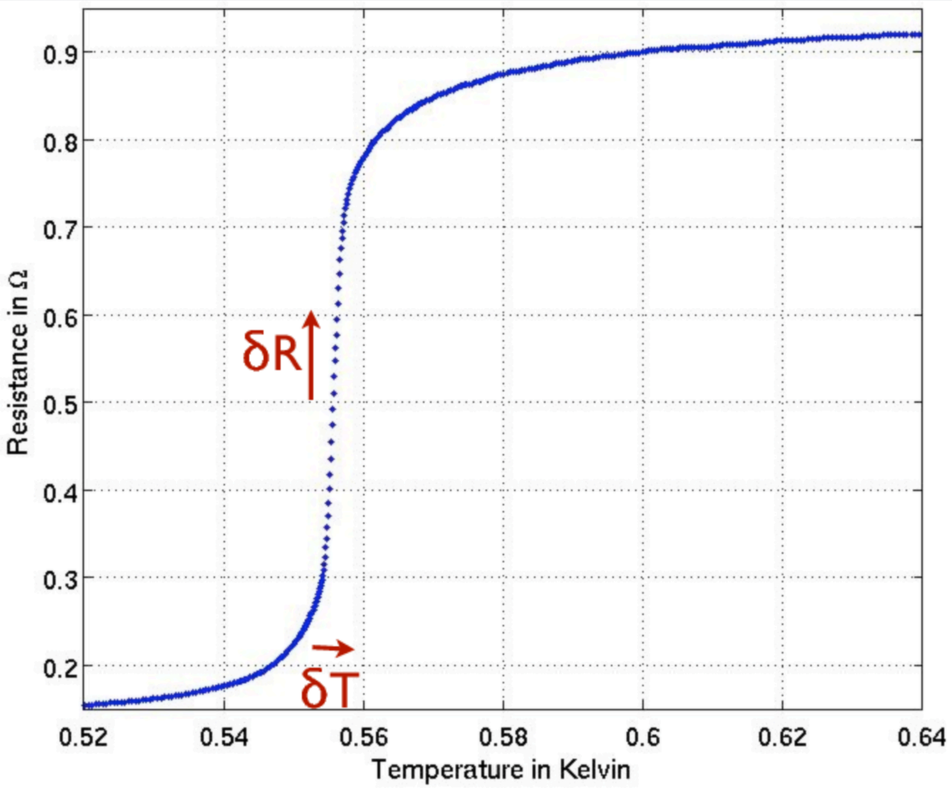


If there is some scale at which the electroweak scale (order parameter) and Higgs mass become calculable in terms of the microscopic theory then the LHC is telling us that:

$$v \ll \frac{\Lambda}{g_\star} \qquad\qquad m_h^2 \ll \hbar \lambda^2 \Lambda^2$$

Criticality and Tuning: TES

In a transition edge sensor the temperature is fine-tuned, through a feedback loop, to sit at precisely the critical point…



Taken from 1309.5383

A big fluctuation gives a big change!

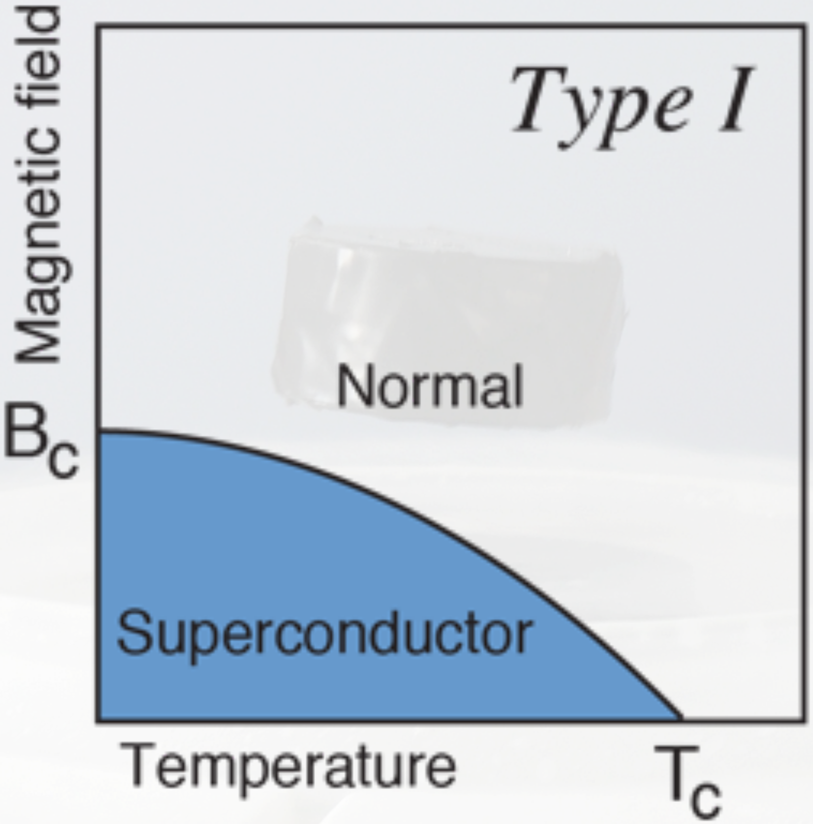
Proximity to the critical point is signaled by:

Macroscopic discontinuities of derivatives of free energy across the critical point (Ehrenfest)...

...which follow from a microscopic coexistence of two phases...

Other Background Parameters

The phase doesn't only depend on the temperature, but other background parameters such as an external magnetic field.



Taken from Hyperphysics

For a range of temperatures could tune B to sit arbitrarily close to the critical point.

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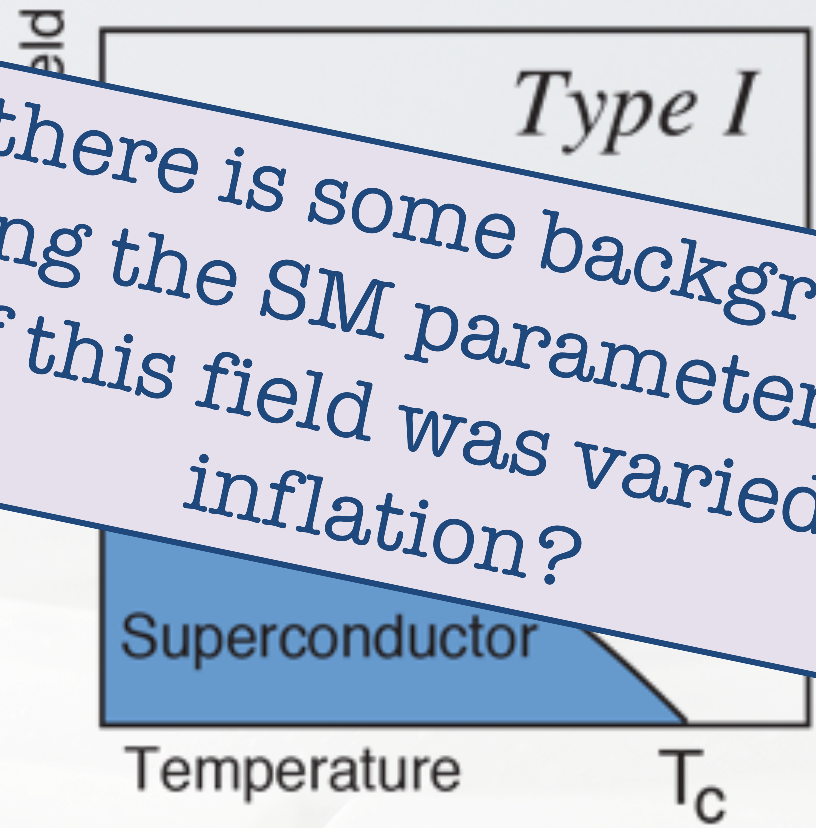

Taken from Hyperphysics

What if there is some background field controlling the SM parameters and the value of this field was varied during inflation?

For a range of temperatures could tune B to sit arbitrarily close to the critical point.

Criticality and Eternal Inflation...

Our Setup

Suppose the background parameters are controlled by some scalar field. In EFT language:

$$V = 3H_0^2 M_P^2 + g_\epsilon^2 f^4 \omega(\varphi) \;, \qquad \omega(\varphi) = \sum_{n=1}^{\infty} \frac{c_n}{n!}\varphi^n \;, \qquad \varphi \equiv \frac{\phi}{f} \;, \qquad \omega(0) = 0$$

When scalar potential is a small perturbation across the field range can expand perturbatively to find:

$$\frac{\alpha}{2}\frac{\partial^2 P}{\partial \varphi^2} + \frac{\partial\left(\omega' P\right)}{\partial \varphi} + \beta \omega P = \frac{\partial P}{\partial T}$$

Where:

$$\alpha = \frac{3\hbar H_0^4}{4\pi^2 g_\epsilon^2 f^4} \;, \qquad \beta = \frac{3\xi f^2}{2M_P^2} \;, \qquad T = \frac{t}{t_R} \;, \qquad t_R = \frac{3H_0}{g_\epsilon^2 f^2}$$

Quantum | Range | Clock | Timescale

In Stasis

In stasis the solution is an eigenstate of time. Subject to BCs, field distribution is a solution of:

$$\frac{\alpha}{2}\,p'' + \omega' p' + (\omega'' + \beta\omega - \lambda)\,p = 0$$

Reminder:

$$\alpha = \frac{3\hbar H_0^4}{4\pi^2 g_\epsilon^2 f^4}\,, \qquad \beta = \frac{3\xi f^2}{2M_P^2}\,, \qquad T = \frac{t}{t_R}\,, \qquad t_R = \frac{3H_0}{g_\epsilon^2 f^2}$$

Quantum Range Clock Timescale

A peak, if it exists, will have position determined by the inflationary rate.

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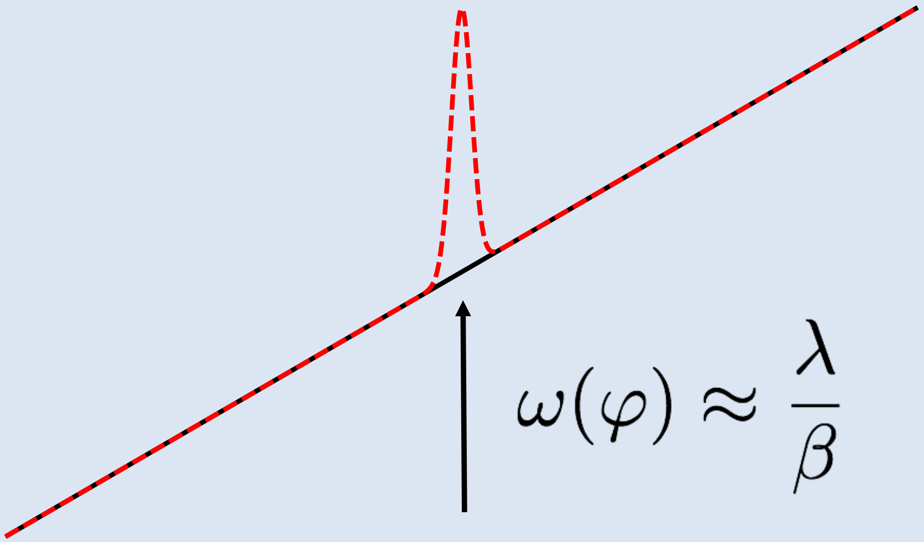
$$\frac{\alpha}{2}\,p'' + \omega' p' + \left(\omega'' + \beta\omega - \lambda\right) p = 0$$

A peak, if it exists, will have position determined by the inflationary rate (eigenvalue):



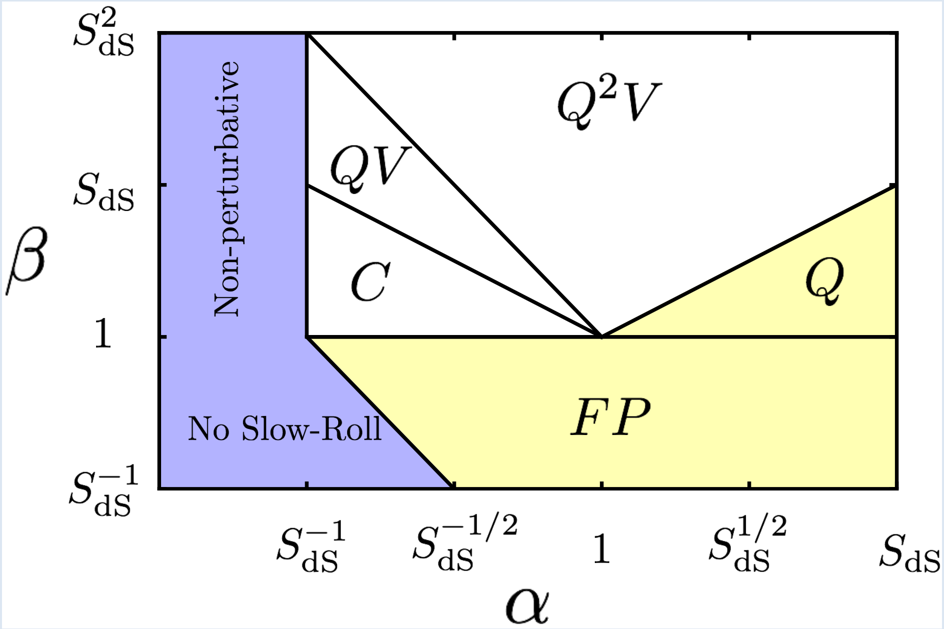
This is very intuitive. If the field is localised at some position, the vacuum energy in slow roll is just the height it is localised at.

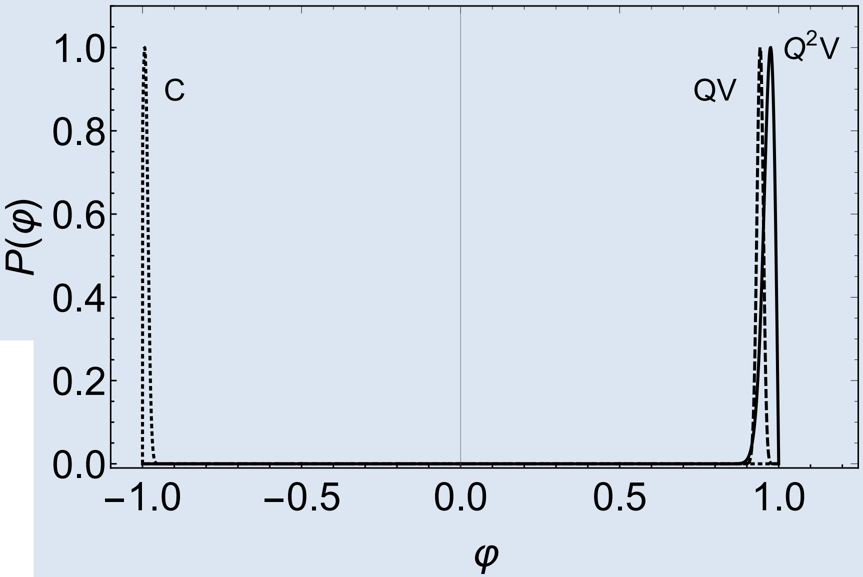
The vacuum energy is the inflationary rate.

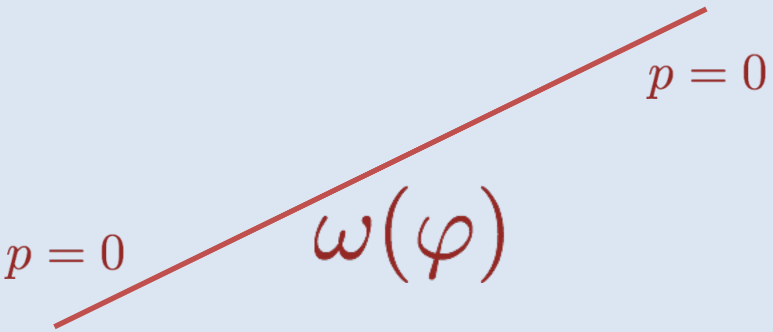
In Stasis

On a linear slope we identify three distinct parameter regimes:

$$C : \alpha\beta < 1$$
$$QV : \alpha\beta > 1$$
$$Q^2V : \alpha^2\beta > 1$$







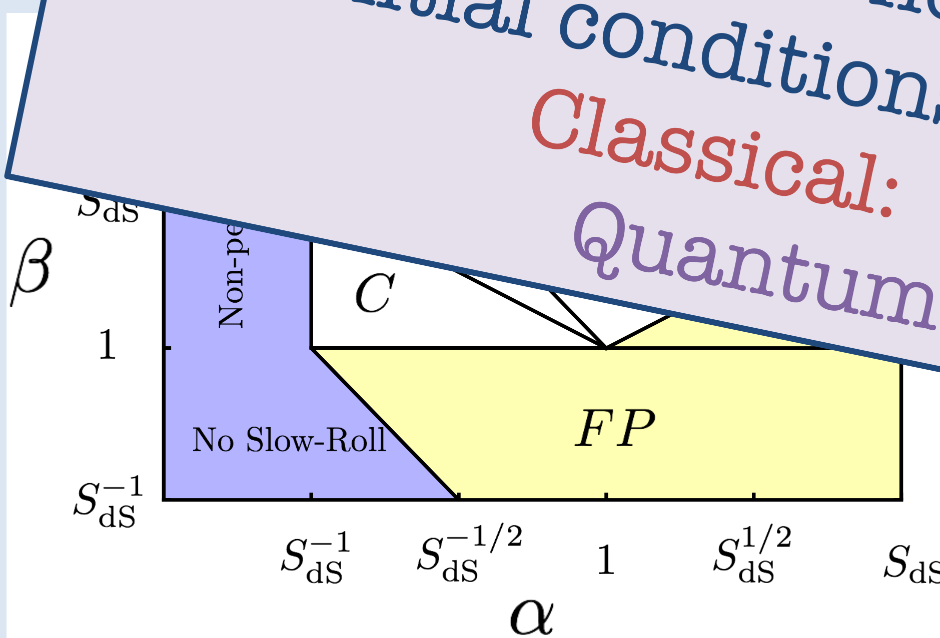
In Stasis

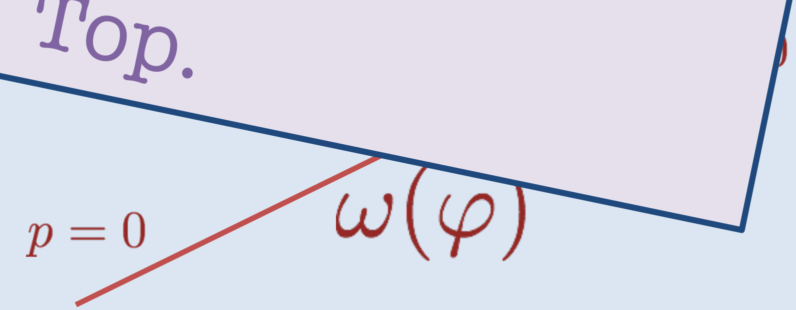
On a linear slope we identify three distinct parameter reg

Punchline: For absorbing boundary conditions on a monotonic curve, for any initial conditions, will settle at:

Classical: Bottom.

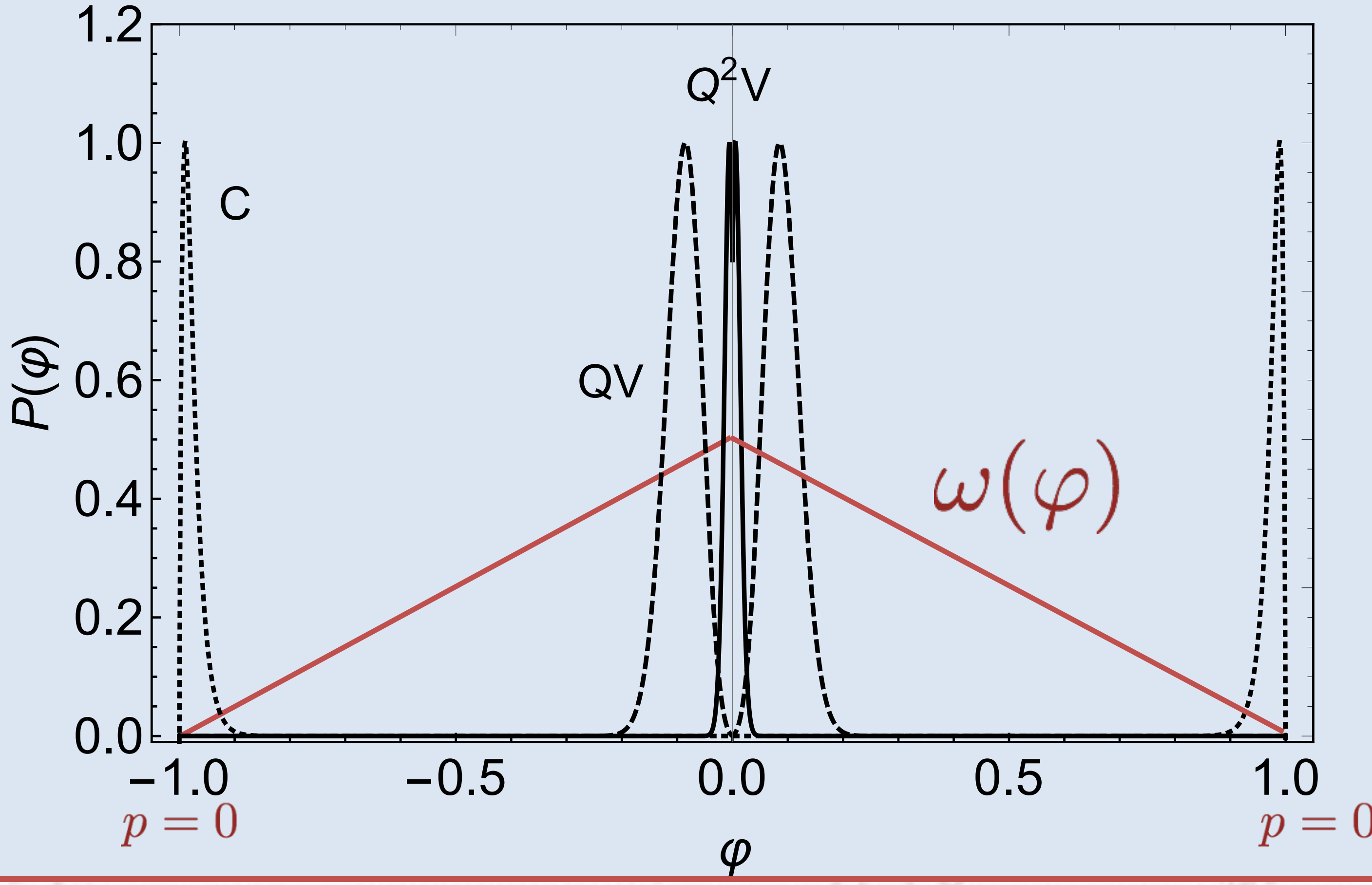
Quantum: Top.





In Stasis

On a pyramid we identify the same three distinct parameter regimes:



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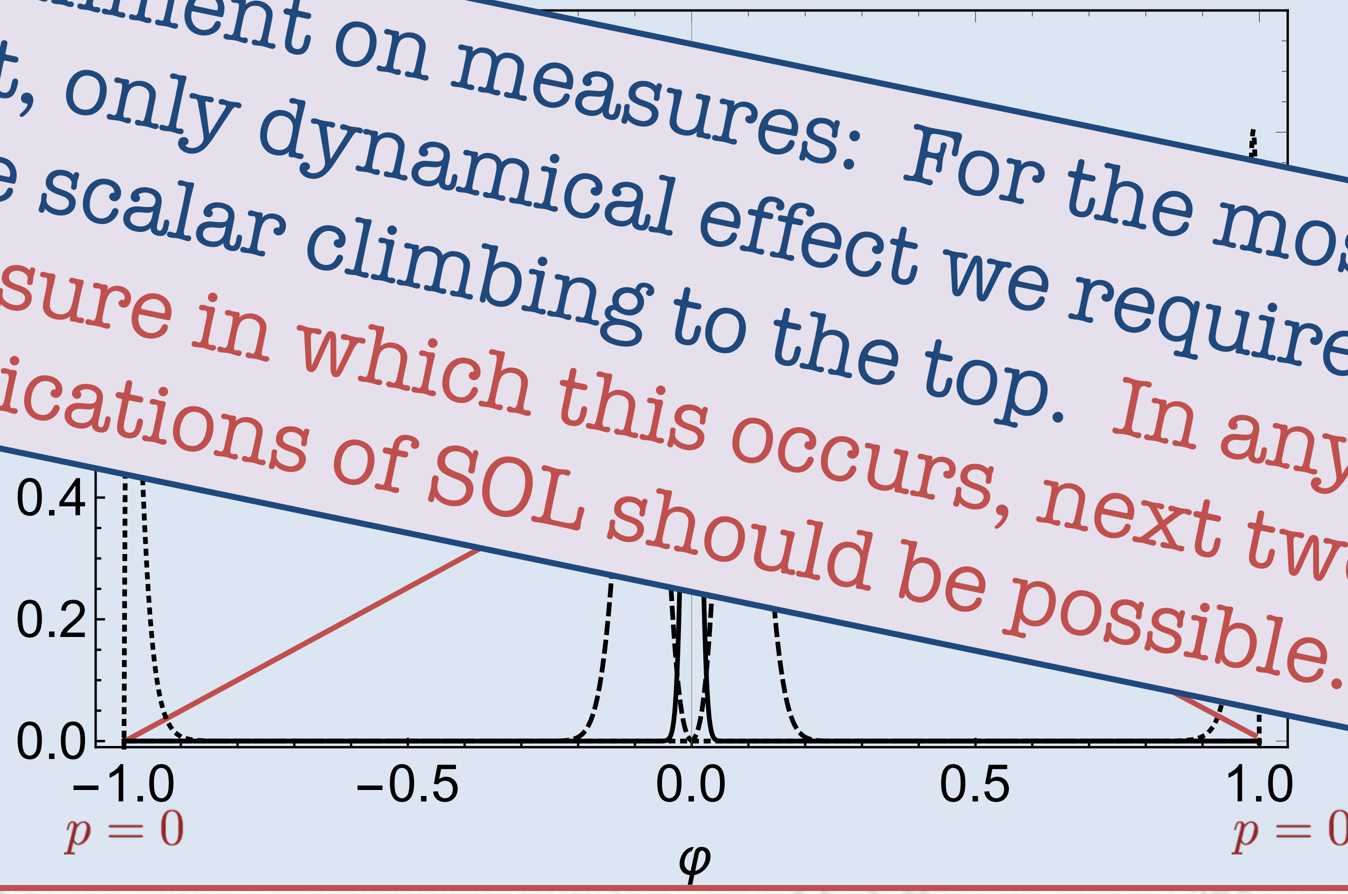


Comment on measures: For the most part, only dynamical effect we require is the scalar climbing to the top. In any measure in which this occurs, next two applications of SOL should be possible.

How long until stasis?

If we are going to reach a relatively stationary state, in order to be independent of boundary conditions we have to wait a time

$$\alpha\beta t_S \qquad , \qquad \alpha\beta < 1$$

in the classical regime, where t_S is the entropy-bound timescale, and

$$\sqrt{\alpha\beta} t_S \qquad , \qquad \alpha\beta > 1$$

in the quantum regime.

Sherpas are eternal...

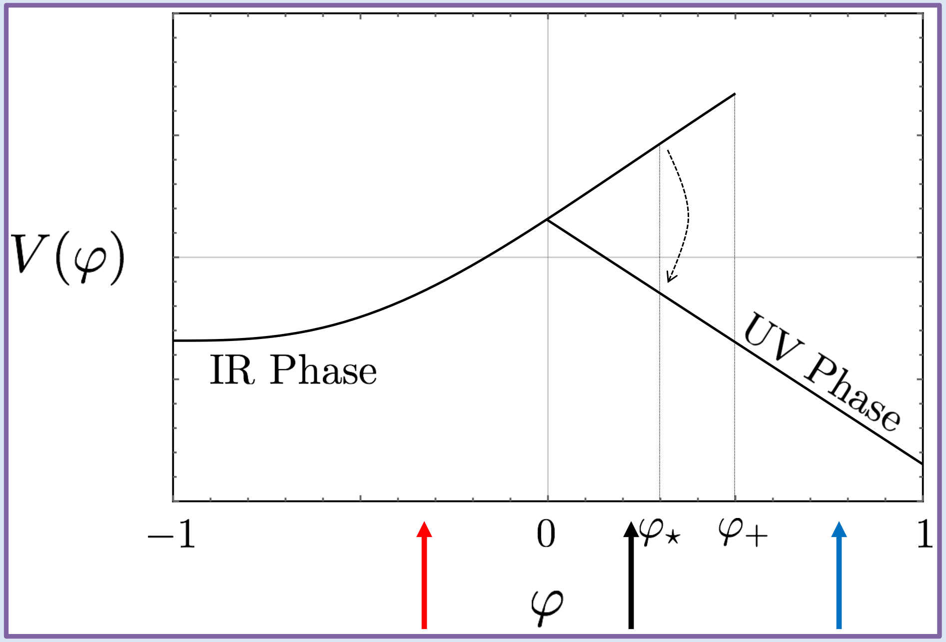
Putting this to use…

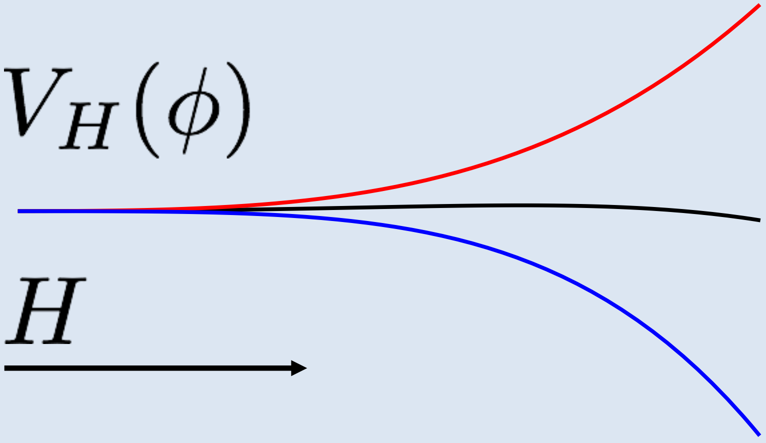
Application: SM Quartic

Suppose the scalar is scanning the SM parameters. In particular, the Higgs quartic, consistent with EFT

$$V(\varphi, h) = \frac{M^4}{g_*^2}\,\omega(\varphi) + \frac{\lambda(\varphi, h)}{4}\left(h^2 - v^2\right)^2$$

Scalar potential is:





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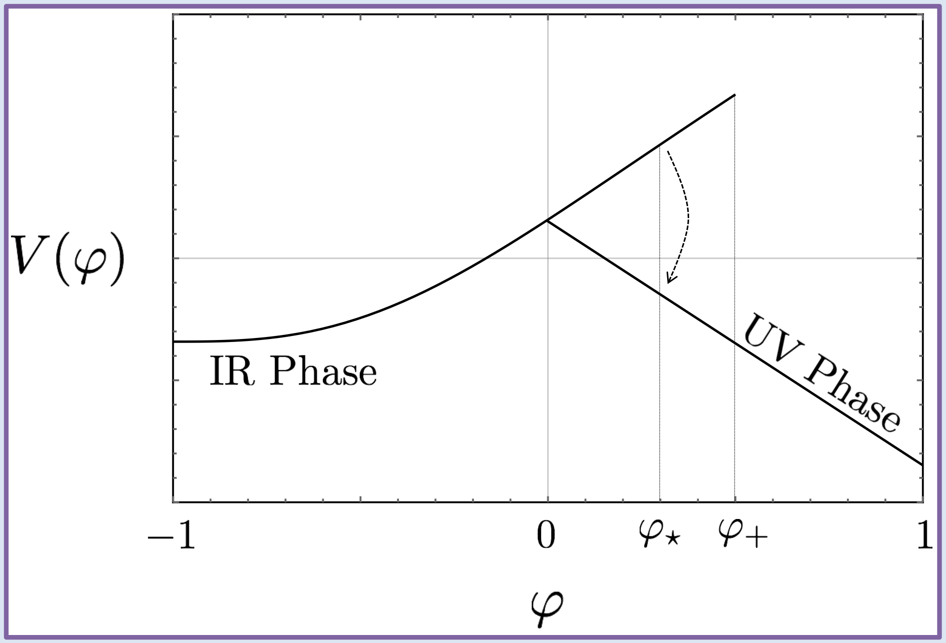

Tunneling from top to bottom is accounted for by having continuity of flux at this point and absorbing boundary conditions at the tunneling point.

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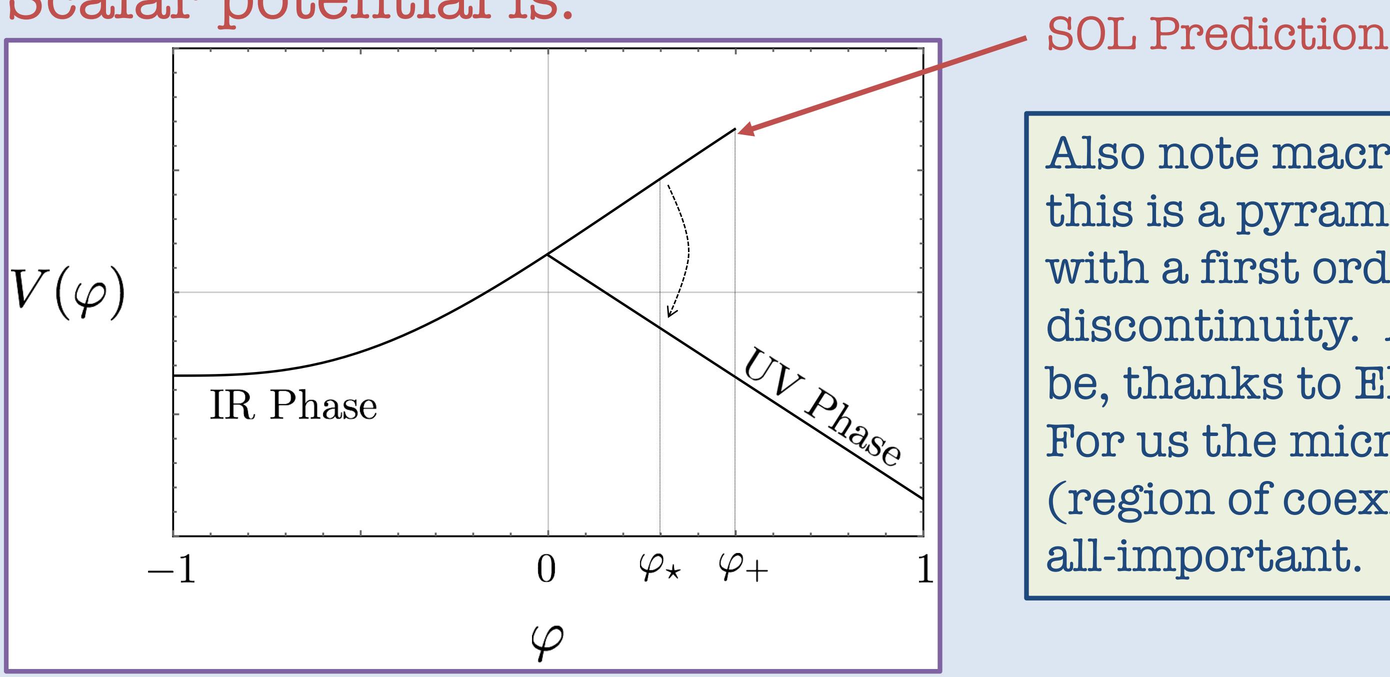
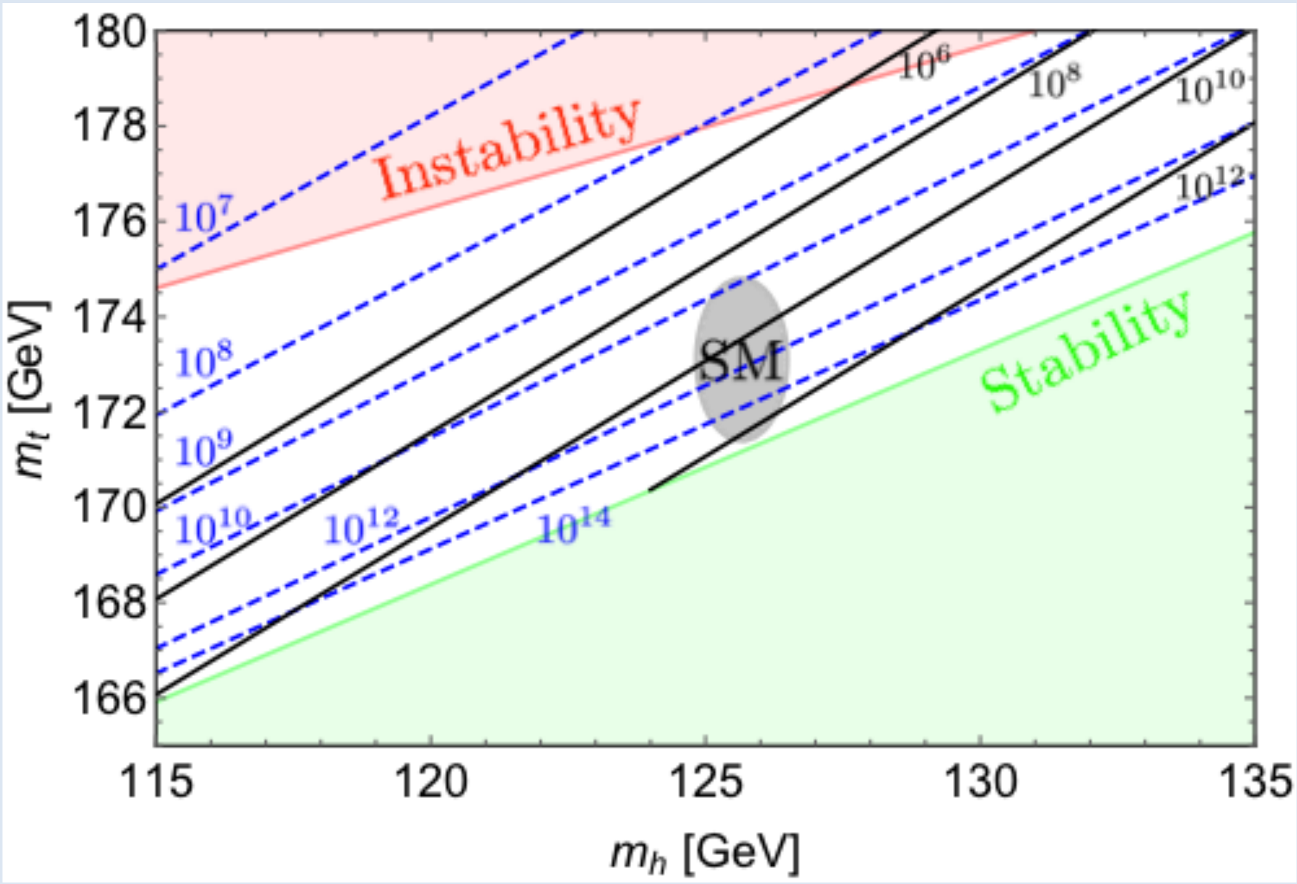


Also note macroscopically this is a pyramid potential with a first order discontinuity. As it should be, thanks to Ehrenfest. For us the microscopics (region of coexistence) is all-important.

Application: SM Quartic

Prediction is metastability region, since top of potential:



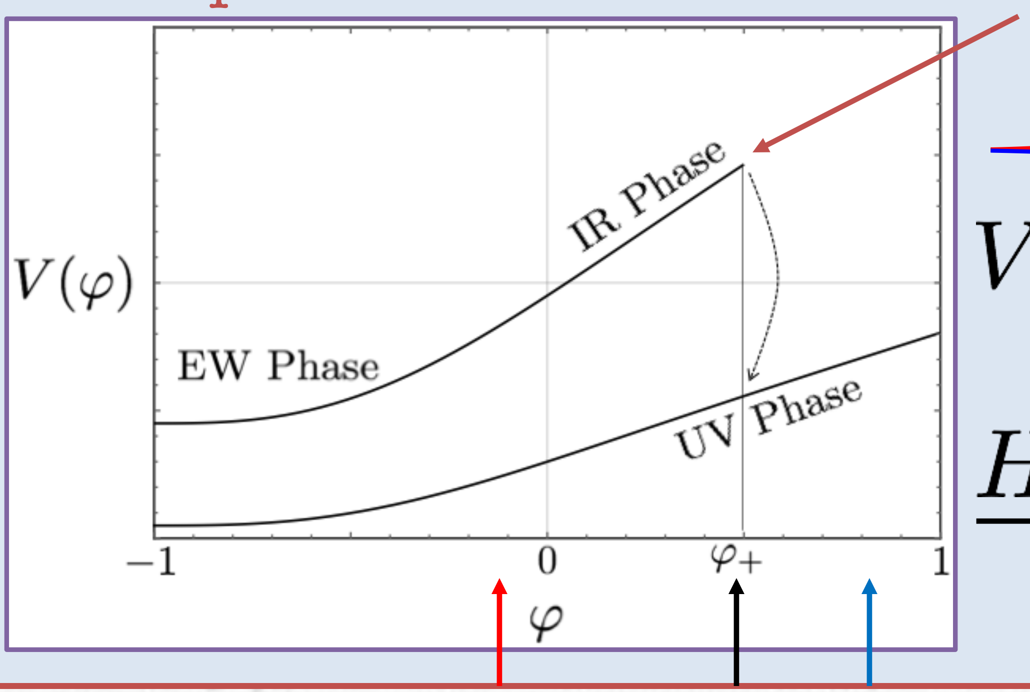
Blue is instability scale, black fixed Hubble contours.

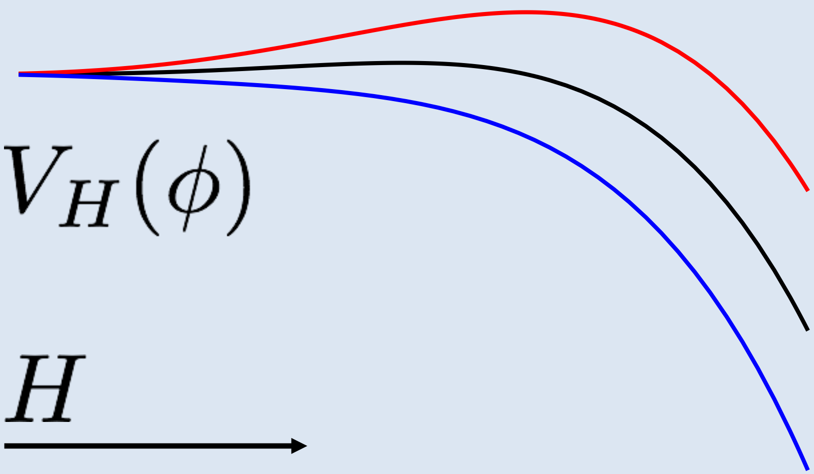
Application: SM Naturalness

Consider the same Higgs instability question, but with a field-dependent mass bilinear:

$$V(\varphi, h) = \frac{M^4}{g_*^2}\,\omega(\varphi) - \frac{\varphi M^2 h^2}{2} + \frac{\lambda(h)\,h^4}{4}$$

Scalar potential is:

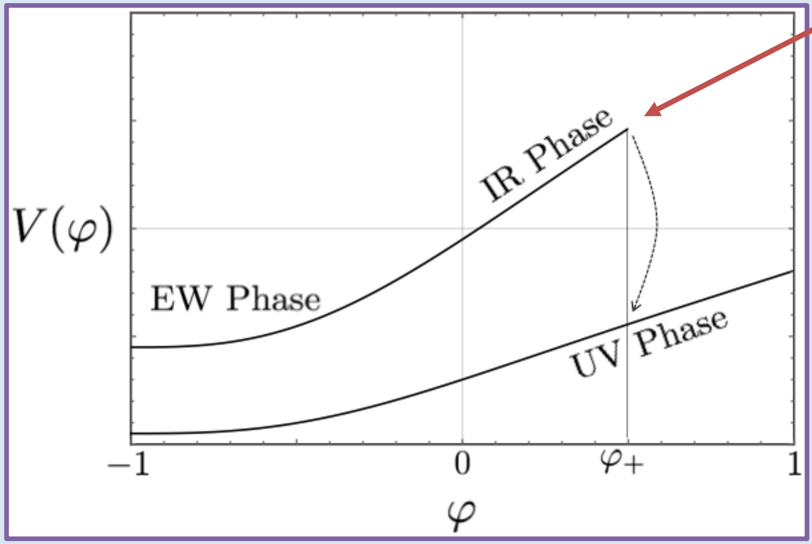




Application: SM Naturalness

Within the SM alone the SOL prediction is the SM instability scale. This is remarkable: Quartic running generates an exponential scale separation between cutoff and instability scale.

SOL Prediction

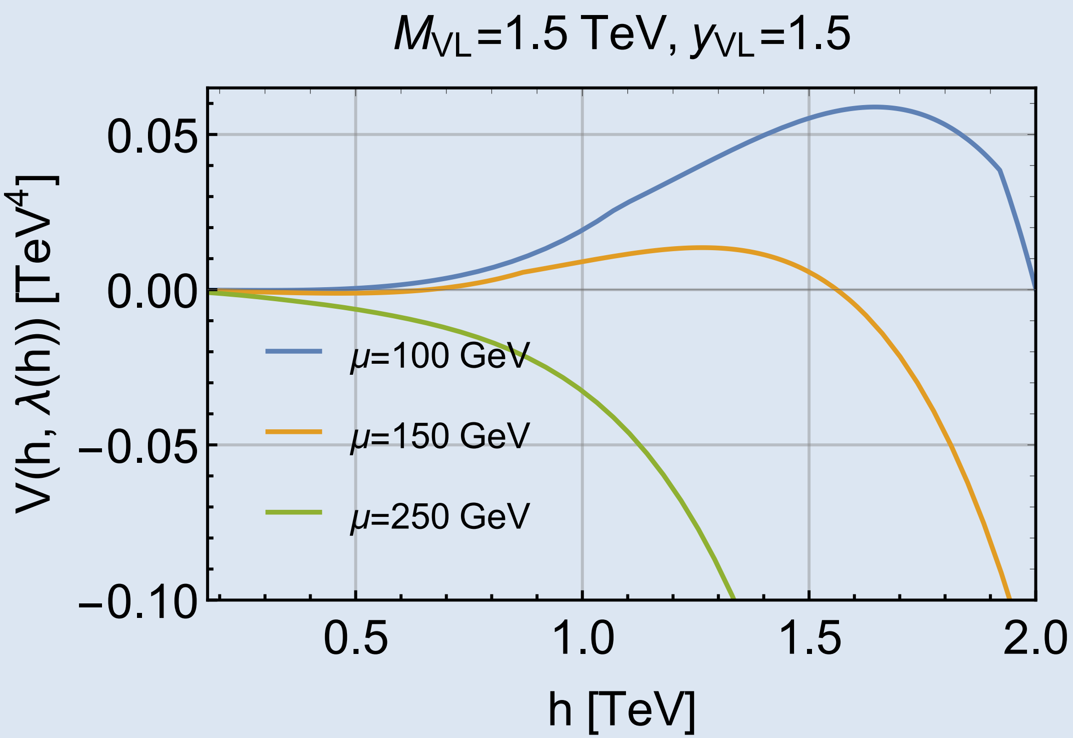


But the instability scale is still way above the weak scale...

$$\Lambda_{\rm Inst} \sim 10^{10}\ {\rm GeV}$$

Junk Food

If we add (naturally) light vector-like-fermions coupled to the Higgs the instability scale can be brought down significantly, just about consistent with SM stability.



Yet another motivation for new states near the TeV scale.

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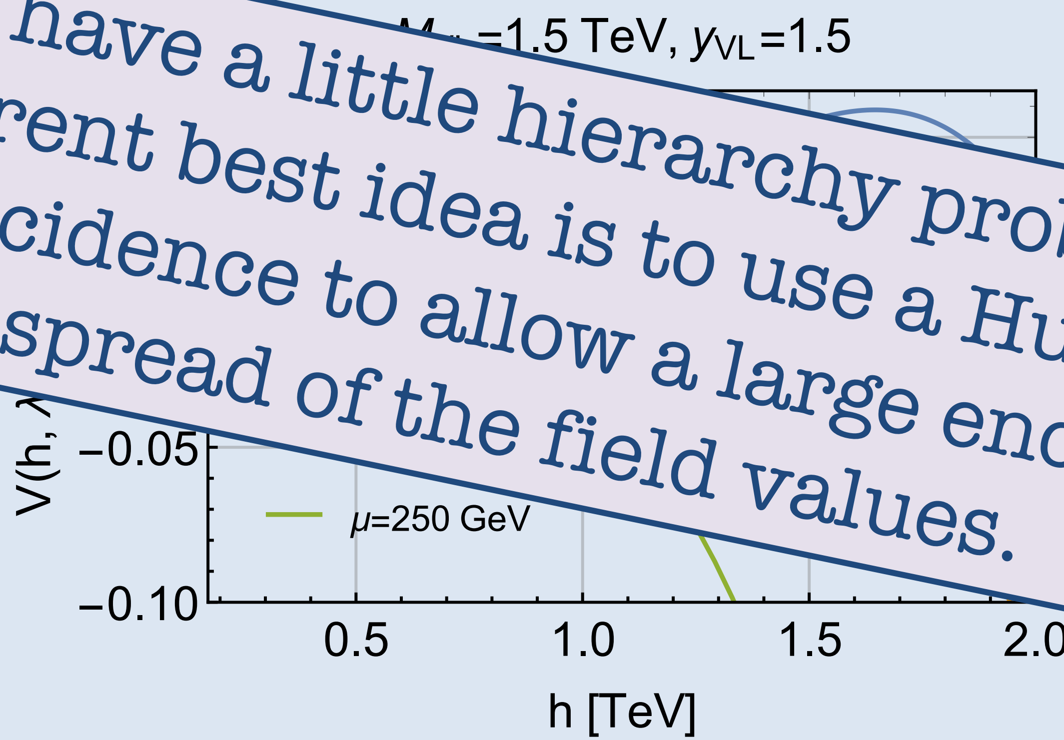


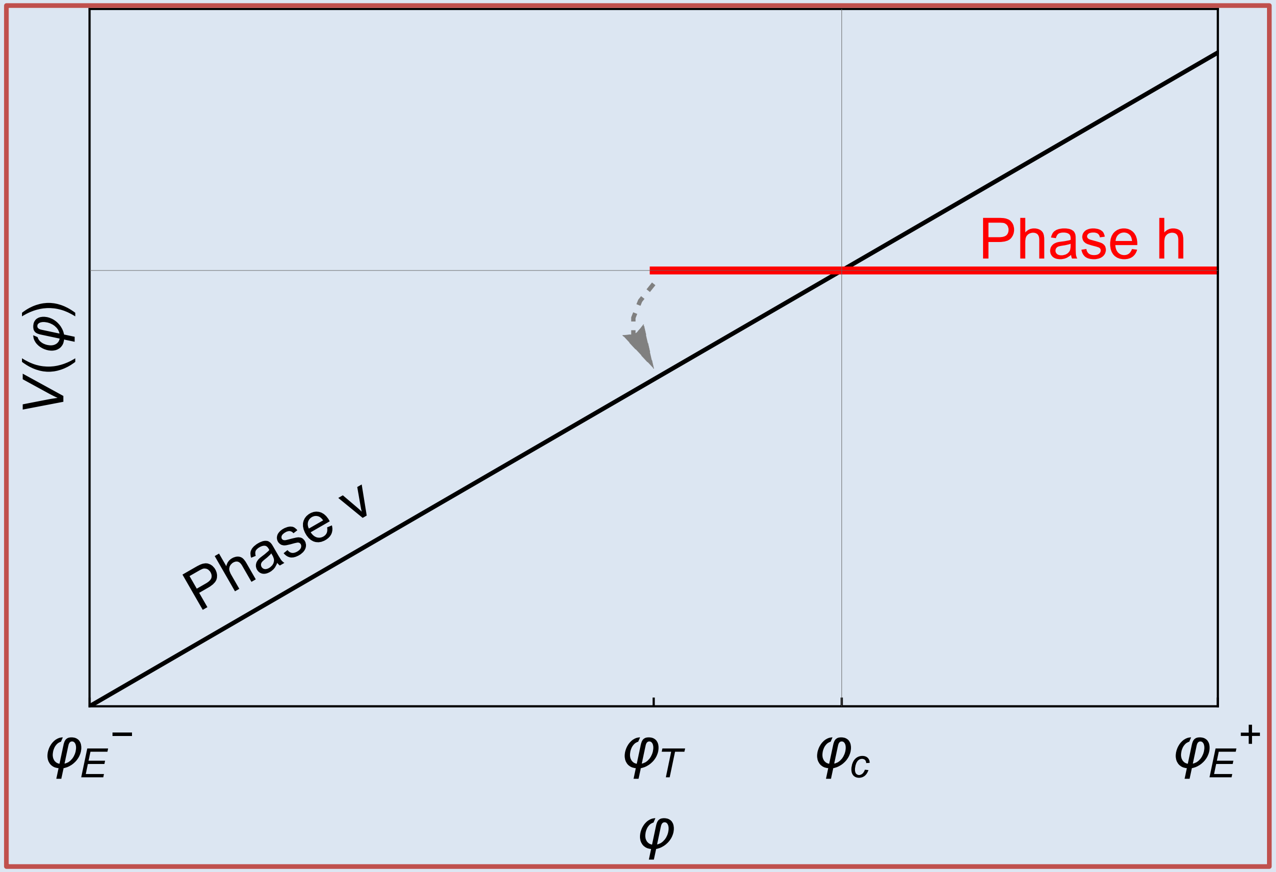
Still have a little hierarchy problem…
Current best idea is to use a Hubble-coincidence to allow a large enough spread of the field values.

Yet another motivation for new states near the TeV scale.

From quite speculative, to spectacularly speculative...

Waterfalls

Consider the following "Waterfall" scalar potential:



What does the field distribution look like in steady state?

Waterfalls

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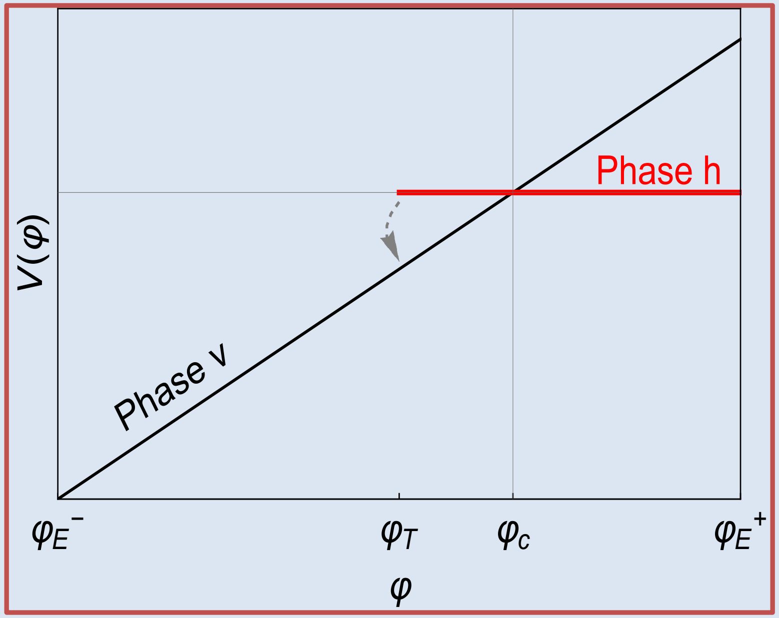

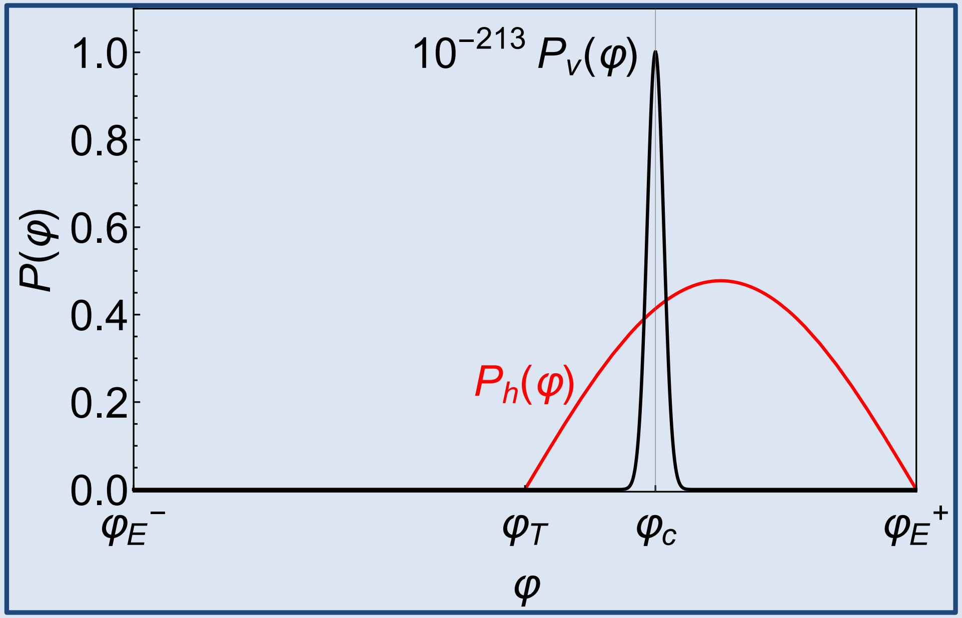

In steady state the solution on the “v” branch is localised at the point where it crosses the same height at the “h” branch, even though tunneling far away.

Waterfalls

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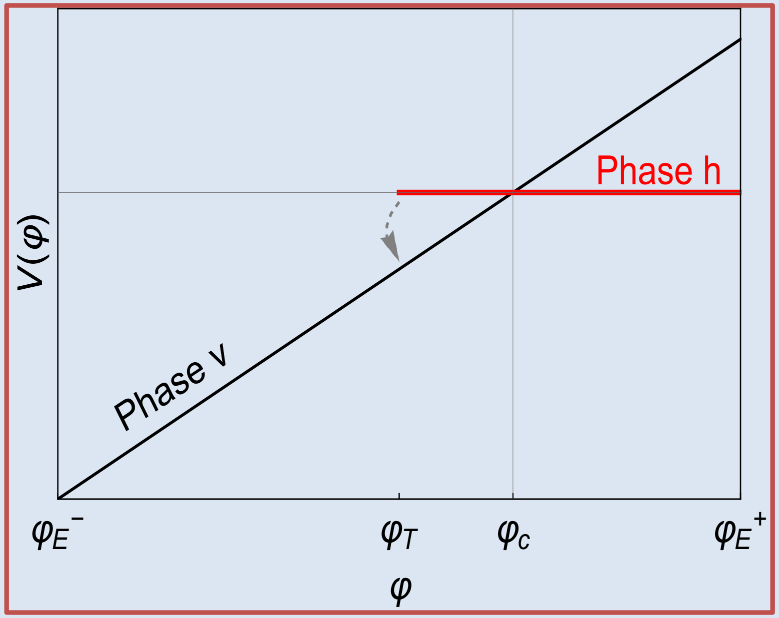
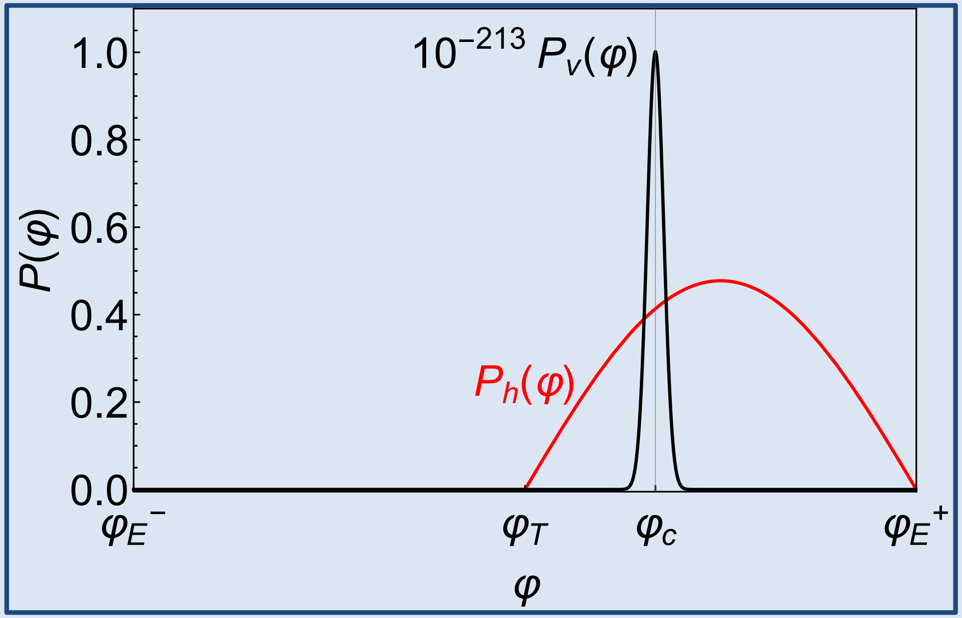
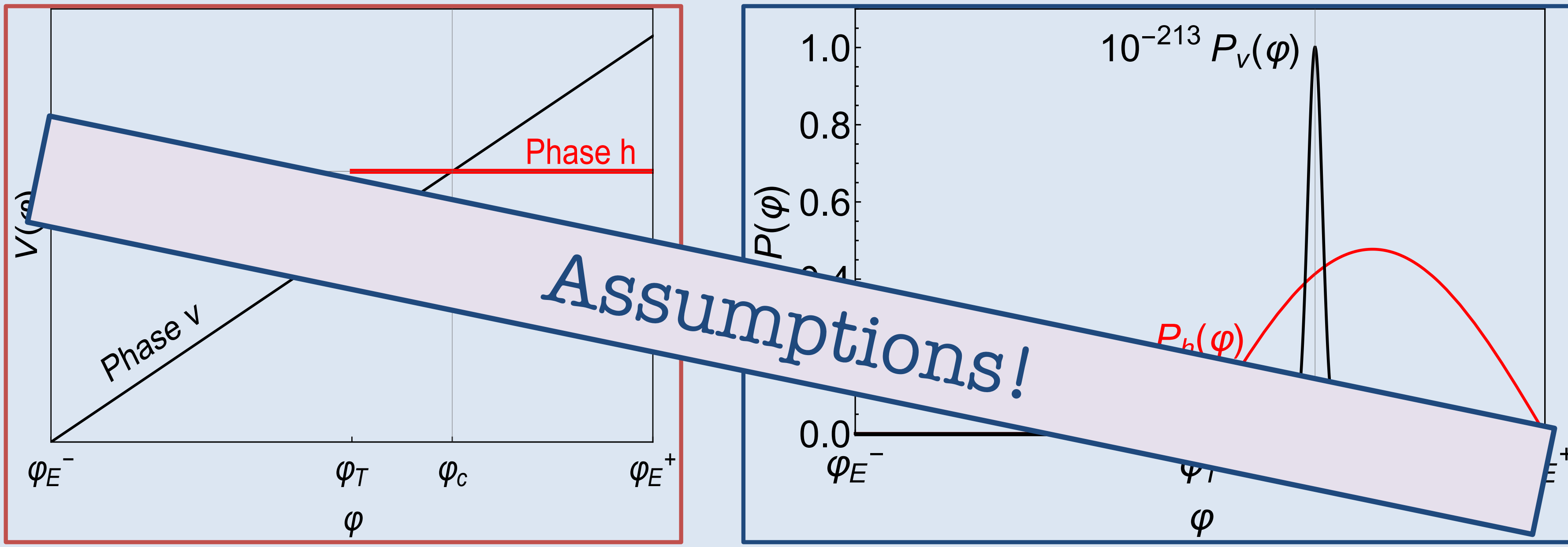




Reason for this is simple. To have non-zero solutions on both branches in steady state, they must both, on average, inflate at the same rate = same Hubble.

Waterfalls

Consider the following "Waterfall" scalar potential:



Reason for this is simple. To have non-zero solutions on both branches in steady state, they must both, on average, inflate at the same rate = same Hubble.

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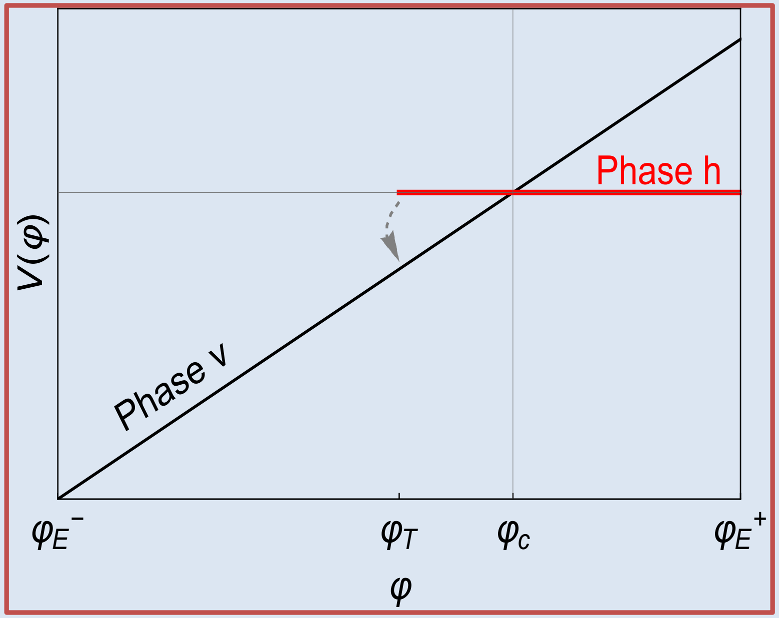

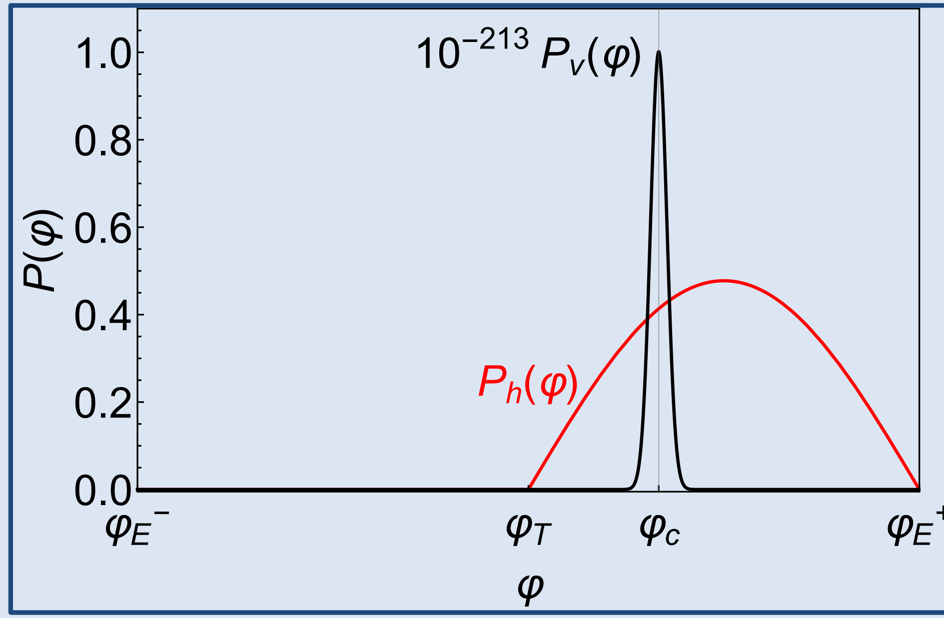

- "v" has non-zero flux injected at the top.
- "h" branch also has flux.
- (No significant sensitivity to these BCs though.)

Waterfalls

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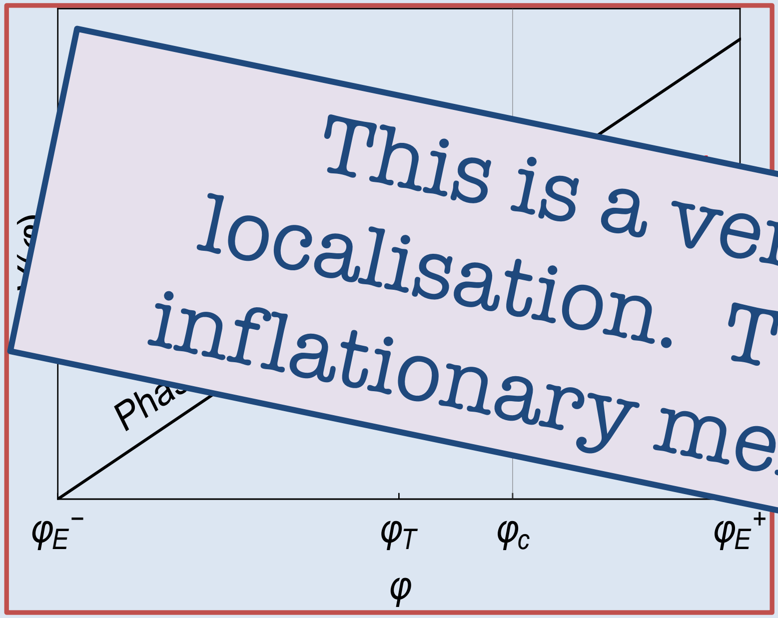
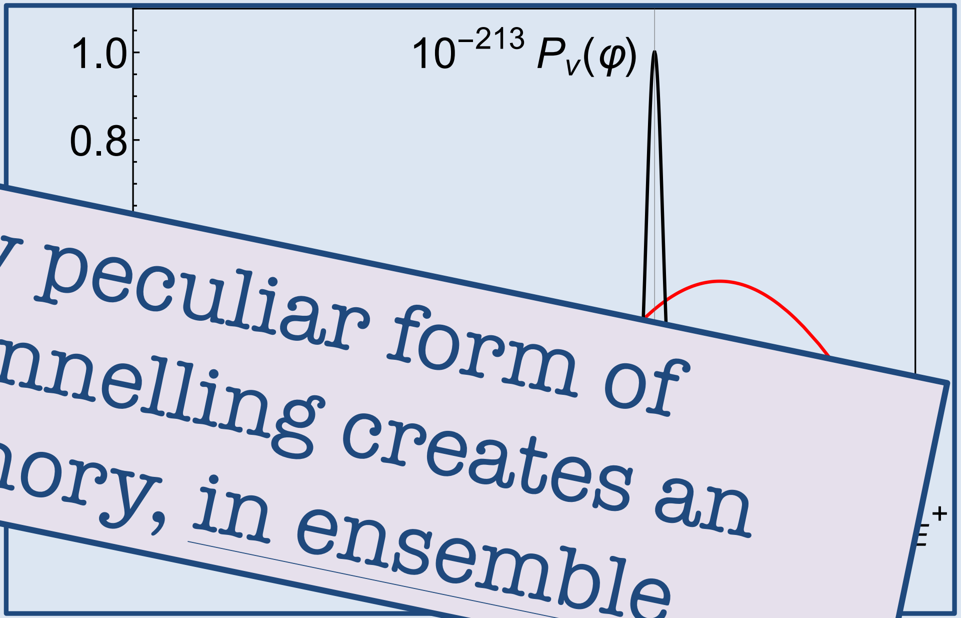




This is a very peculiar form of localisation. Tunnelling creates an inflationary memory, in ensemble....

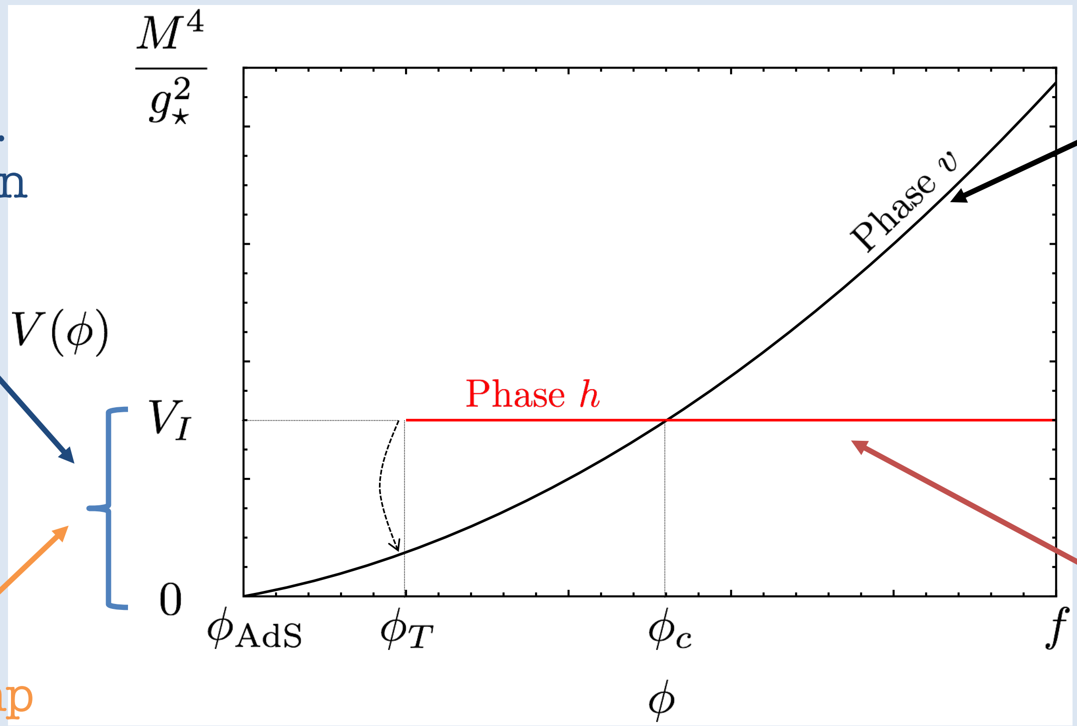
- "v" has non-zero flux injected at the top.
- "h" branch also has flux.
- (No significant sensitivity to these BCs though.)

Application…

Waterfalls

Consider the following "Waterfall" scalar potential in a SUSY setup:

Inflationary sector secluded. this contribution vanishes after reheating.



SUSY and R-symmetry broken, large symmetry breaking.

SUSY vacuum, scalar a pNGB with small explicit breaking. No R-breaking.

For C regime max reheat temp is 25 MeV or so.

Waterfalls

Consider the following "Waterfall" scalar potential in a SUSY setup:

Inf
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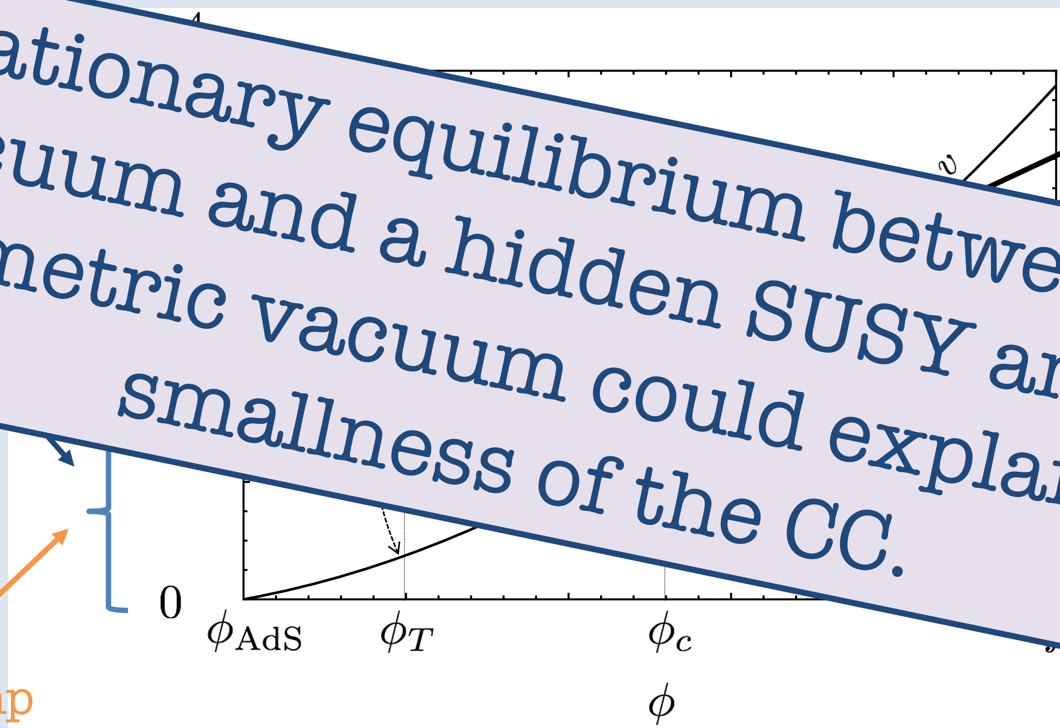
v
n



SUSY and R-symmetry broken, large symmetry

For C regime max reheat temp is 25 MeV or so.

n,
B
explicit breaking. No R-breaking.

Inflationary equilibrium between our vacuum and a hidden SUSY and R-symmetric vacuum could explain the smallness of the CC.

Life after SOL...

Post-Reheating Dynamics

For these models on their own, the scalar will continue rolling after inflation. Many potential effects, but limits on the evolution of dark energy imply:

$$\alpha^2 \beta > \left(\frac{\hbar\, H_0^4}{M_P H_{\mathrm{now}} \Lambda^2} \right)^2 = \left(\frac{H_0}{2 \times 10^{-3}\ \mathrm{eV}} \right)^8$$

Unless the Hubble scale is extremely low during inflation, this constraint forces us towards the Q regime for most applications.

Alternatively, could have some post-inflationary trapping etc, but we opt for simplicity.

Experimental Predictions

Details differ in each implementation, however a common feature is a very very light scalar field which would still be rolling down its scalar potential.

Predicts that dark energy is not a exactly a constant and evidence for $w=-1$ being violated would provide support for SOL.

Summary

Whatever your thoughts on eternal inflation.

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Whatever your thoughts on eternal inflation.

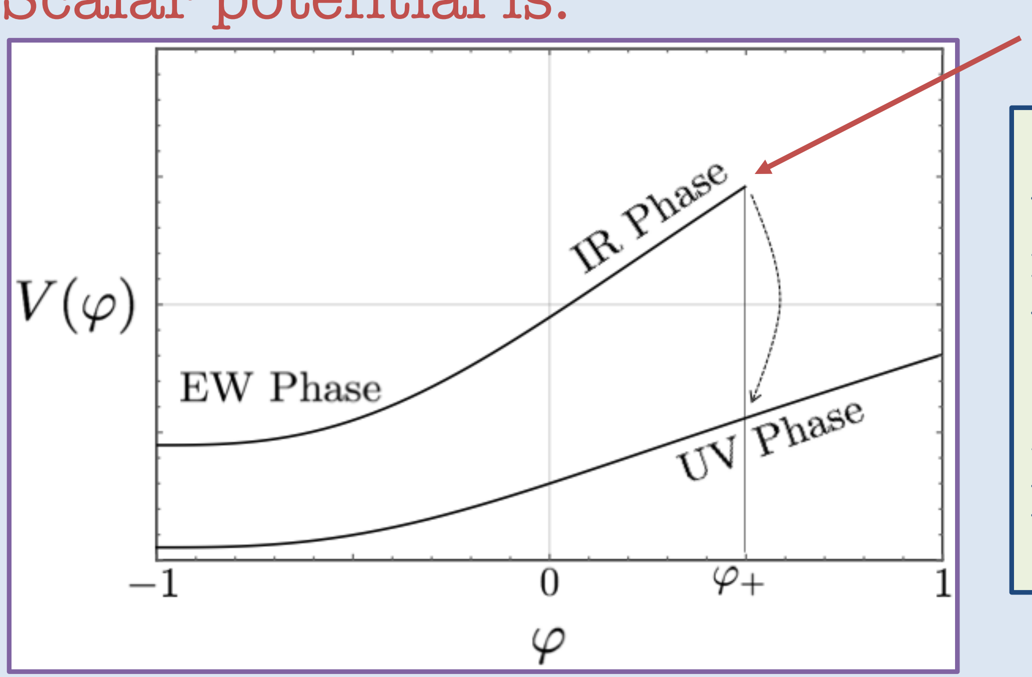
SOL could offer answers as to why our Universe is so determinedly critical, despite Wilsonian arguments to the contrary.

Application: SM Naturalness

Consider the same Higgs instability question, but with a field-dependent mass bilinear:

$$V(\varphi, h) = \frac{M^4}{g_*^2}\,\omega(\varphi) - \frac{\varphi M^2 h^2}{2} + \frac{\lambda(h)\, h^4}{4}$$

Scalar potential is:



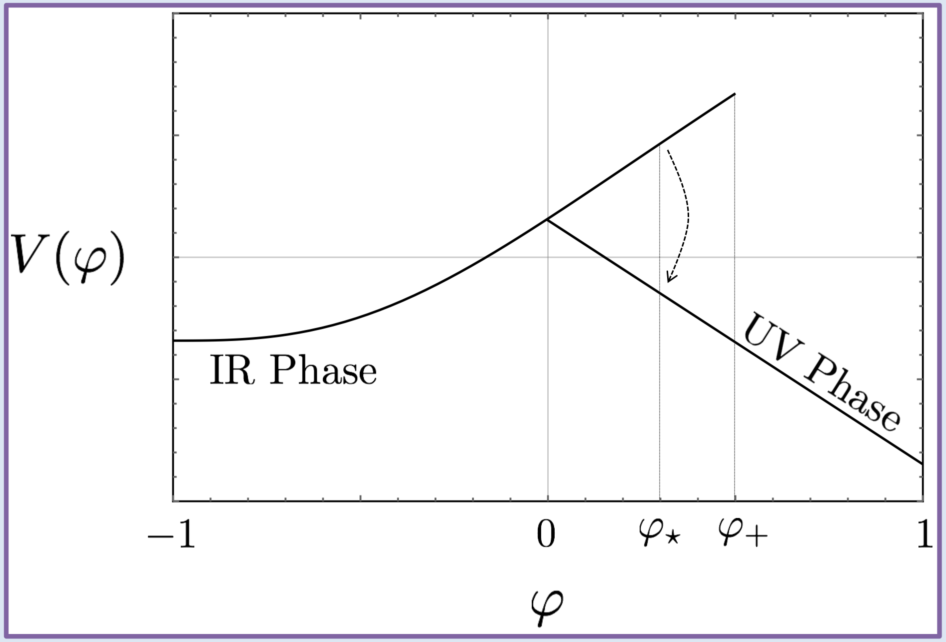
This essentially follows from the fact that for a negative mass-squared greater than the instability scale there is only the UV phase. Whereas for masses below it and positive mass-squared both can coexist.

Application: SM Quartic

Suppose the scalar is scanning the SM parameters. In particular, the Higgs quartic, consistent with EFT

$$V(\varphi, h) = \frac{M^4}{g_*^2}\,\omega(\varphi) + \frac{\lambda(\varphi, h)}{4}\left(h^2 - v^2\right)^2$$

Scalar potential is:



This essentially follows from the fact that for the UV quartic above a zero at the cutoff it remains positive always, so no UV phase. If it is below zero it will cross zero and become positive at some point, so both phases, or negative always, so only UV phase.

Inflation

The theory of inflation has been remarkably successful at solving puzzles:

- Horizon
- Flatness
- …

And at making detailed predictions for our Hubble patch:

- Structures from gravitational collapse of quantum mechanical fluctuations
- Almost scale-invariant spectrum…

In general inflation doing pretty well…

The Good

Th… stochastic approach offers possibility of … acroscopic parameters.

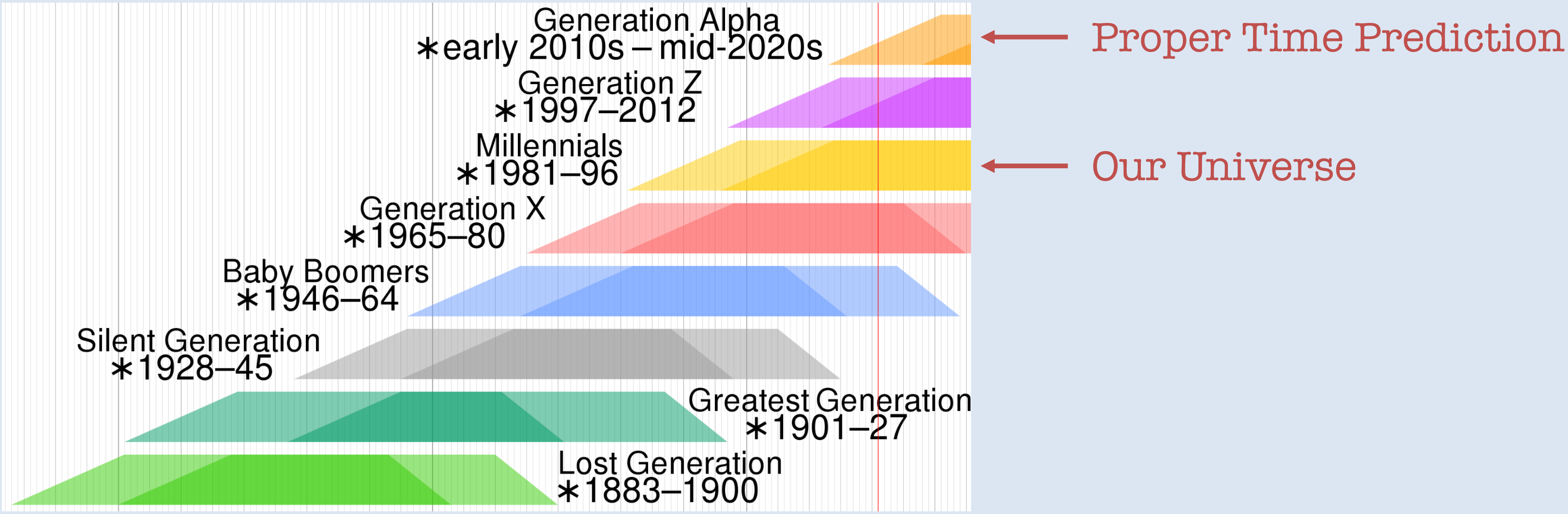
The connection with statistical mechanics is not just coincidence…

"In statistical mechanics, the Fokker–Planck equation is a partial differential equation that describes the time evolution of the probability density function of the velocity of a particle under the influence of drag forces and random forces, as in Brownian motion." Wikipedia

Scalar field fo… trajectory. Average of traj… described by a Fokker-Planck equation.

The Worst

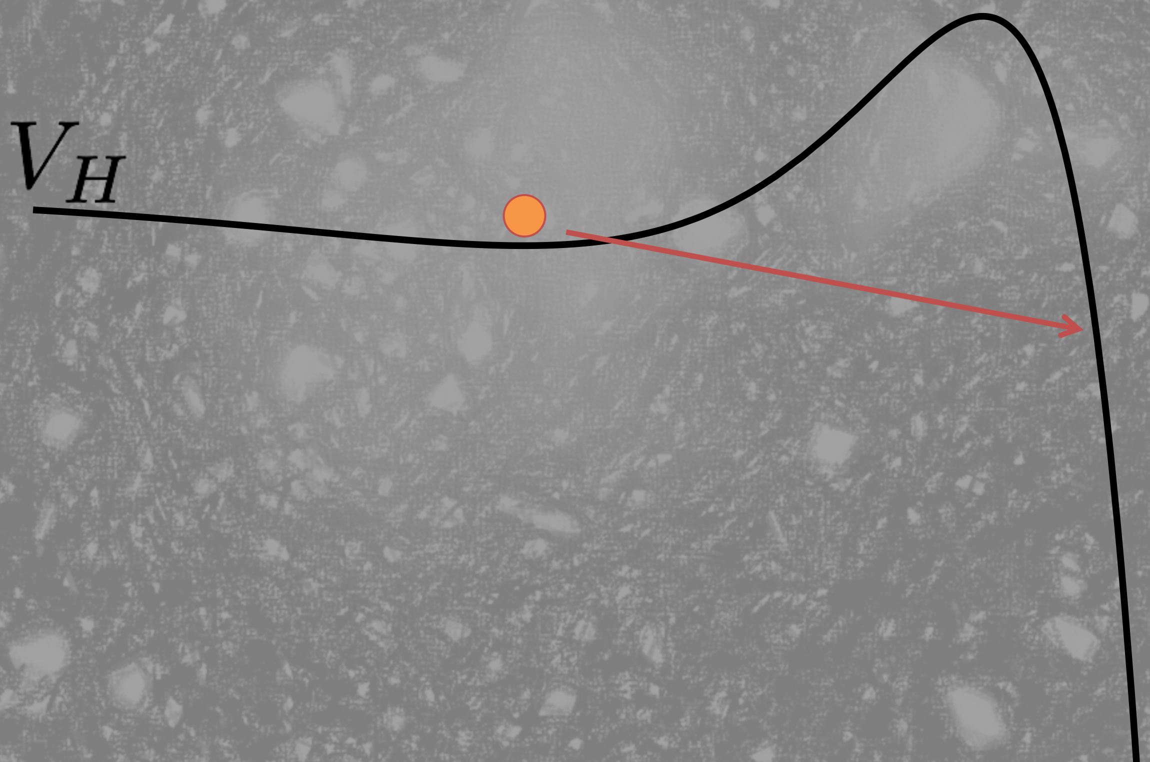
The youngness paradox is much more severe. Emphasised to us by Andrei Linde.



Scale-factor cutoff has no youngness paradox. We are currently unsure how the youngness paradox applies to our models, due to the criticality aspect things could be different.

A Metastable Universe?

This means that the vacuum we are in, as in the Mexican hat pictures, is just local, but there is a deeper one out at large field values.



The Elephant in the Room

Ginzburg-Landau is just a phenomenological model, with no explanation of parameters. The macroscopic parameters follow from the detailed microscopic BCS theory and there are no big surprises.

The order parameter at generic temperatures is of the typical scale associated with underlying microscopic parameters. Criticality means fine-tuning against the fundamental scale.