Interpreting the IceCube events by decaying dark matter hints and constraints

Arman Esmaili

Laboratori Nazionali del Gran Sasso (LNGS)
Theory Group

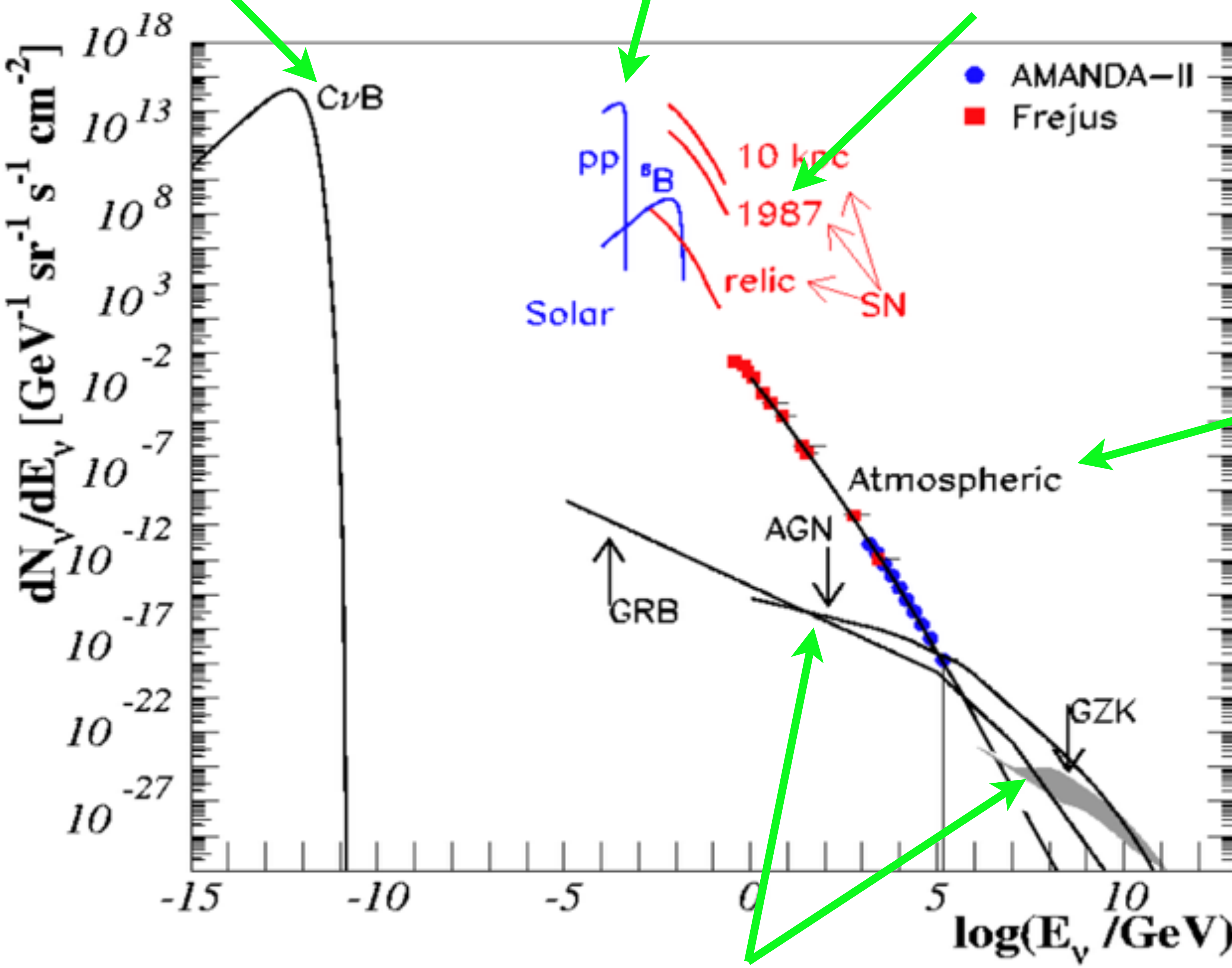
17/Sep/2015

Neutrino Sky

:(

SNO, BOREXINO,....

SK, IMB, ...



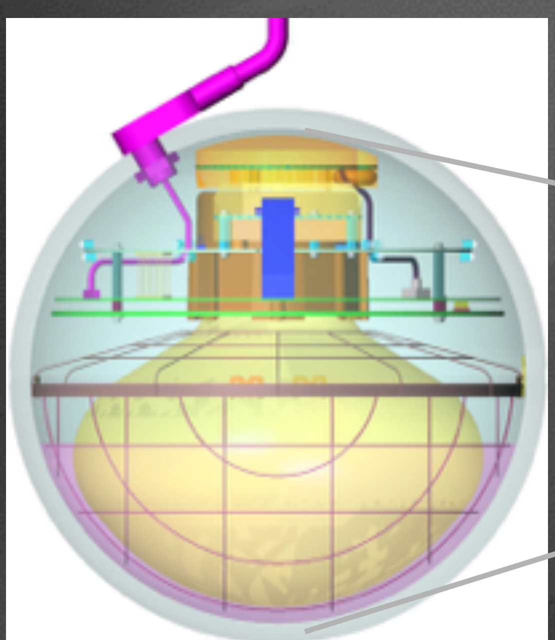
SK, AMANDA, IceCube...

Background for astrophysical neutrinos

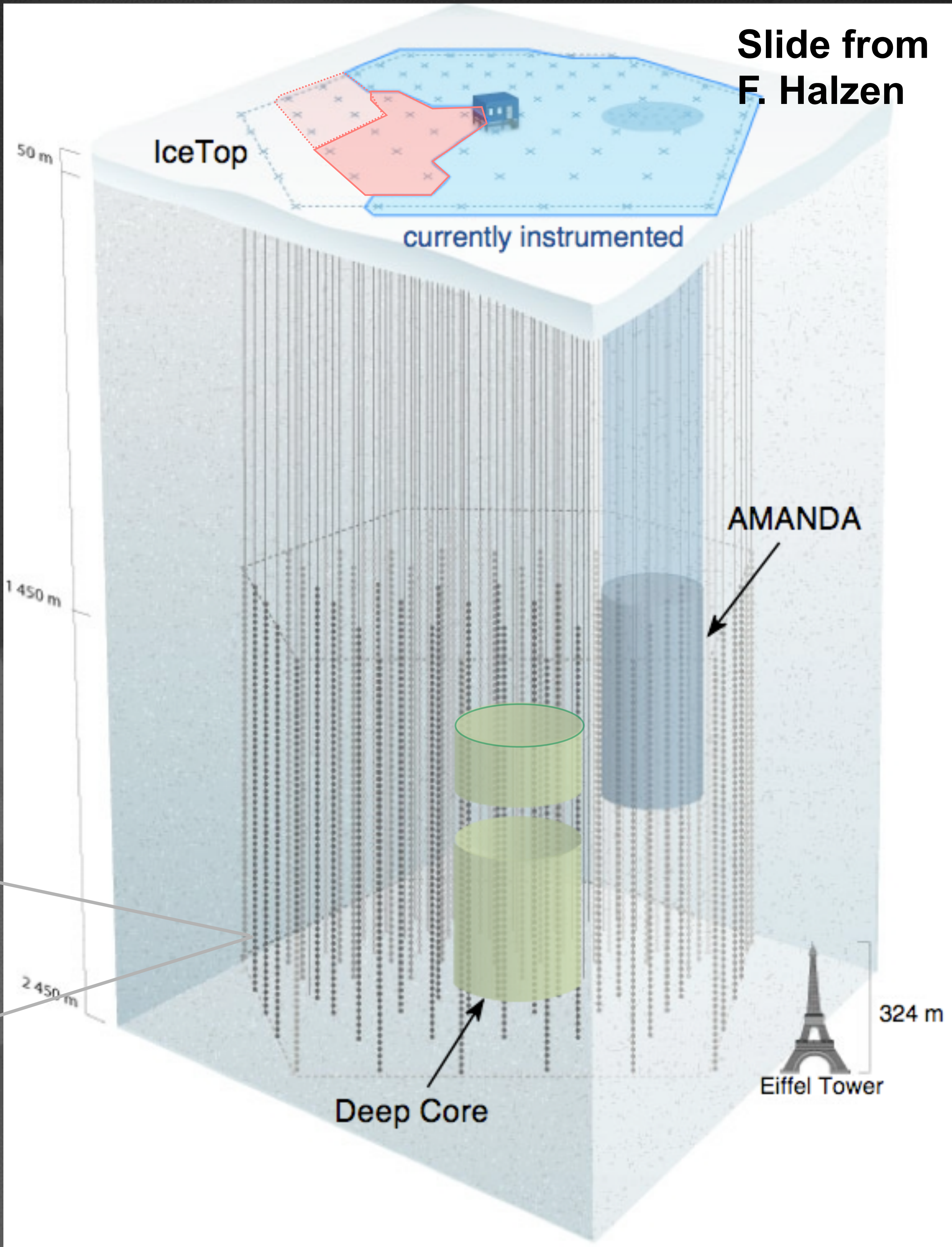
IceCube ?

IceCube / Deep Core

- 5320 optical modules on 86 strings (+ IceTop)
- detects ~220 neutrinos and 1.7x10^{8} muons per day
- threshold 10 GeV
- angular resolution < 1 degree

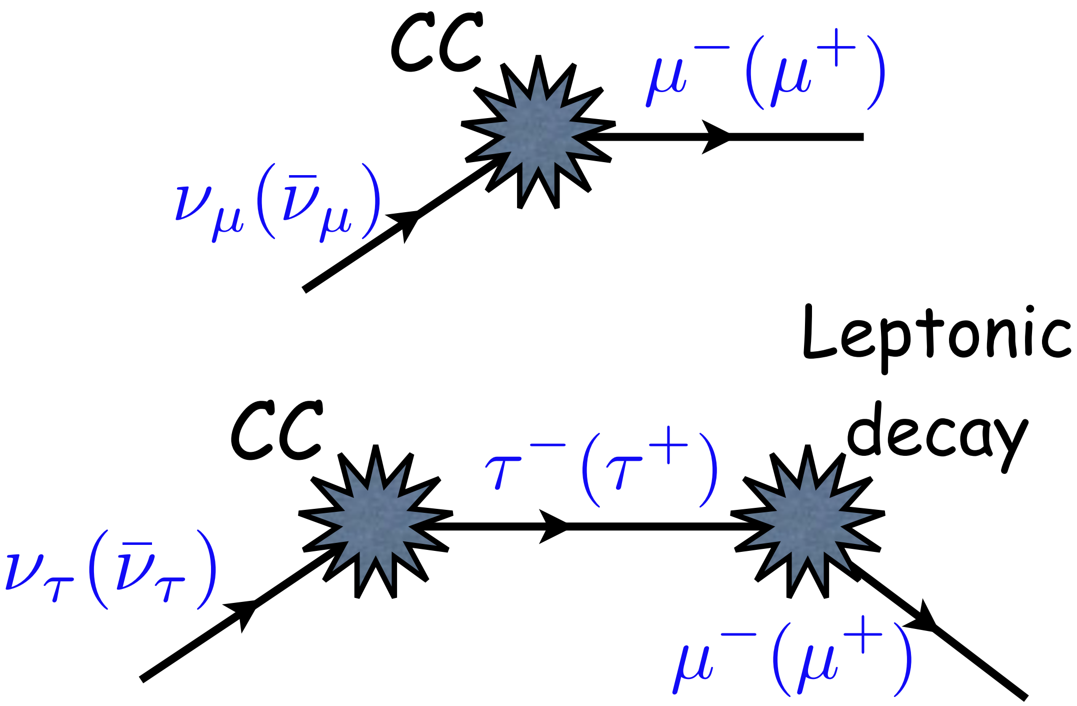


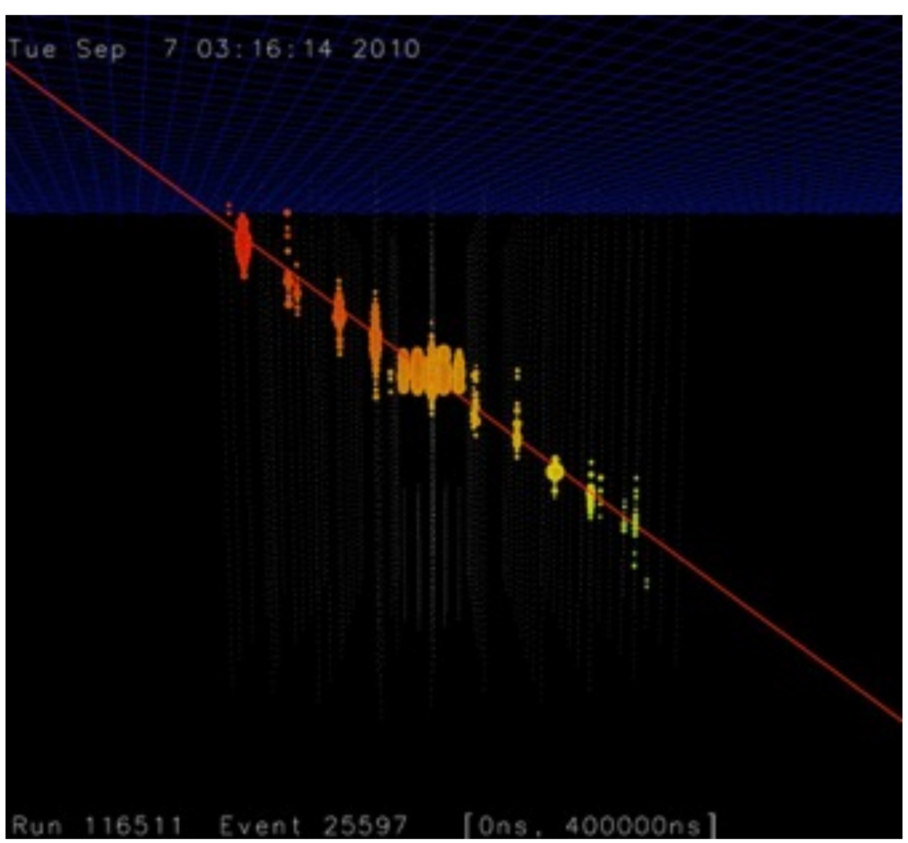
Digital Optical Module (DOM)



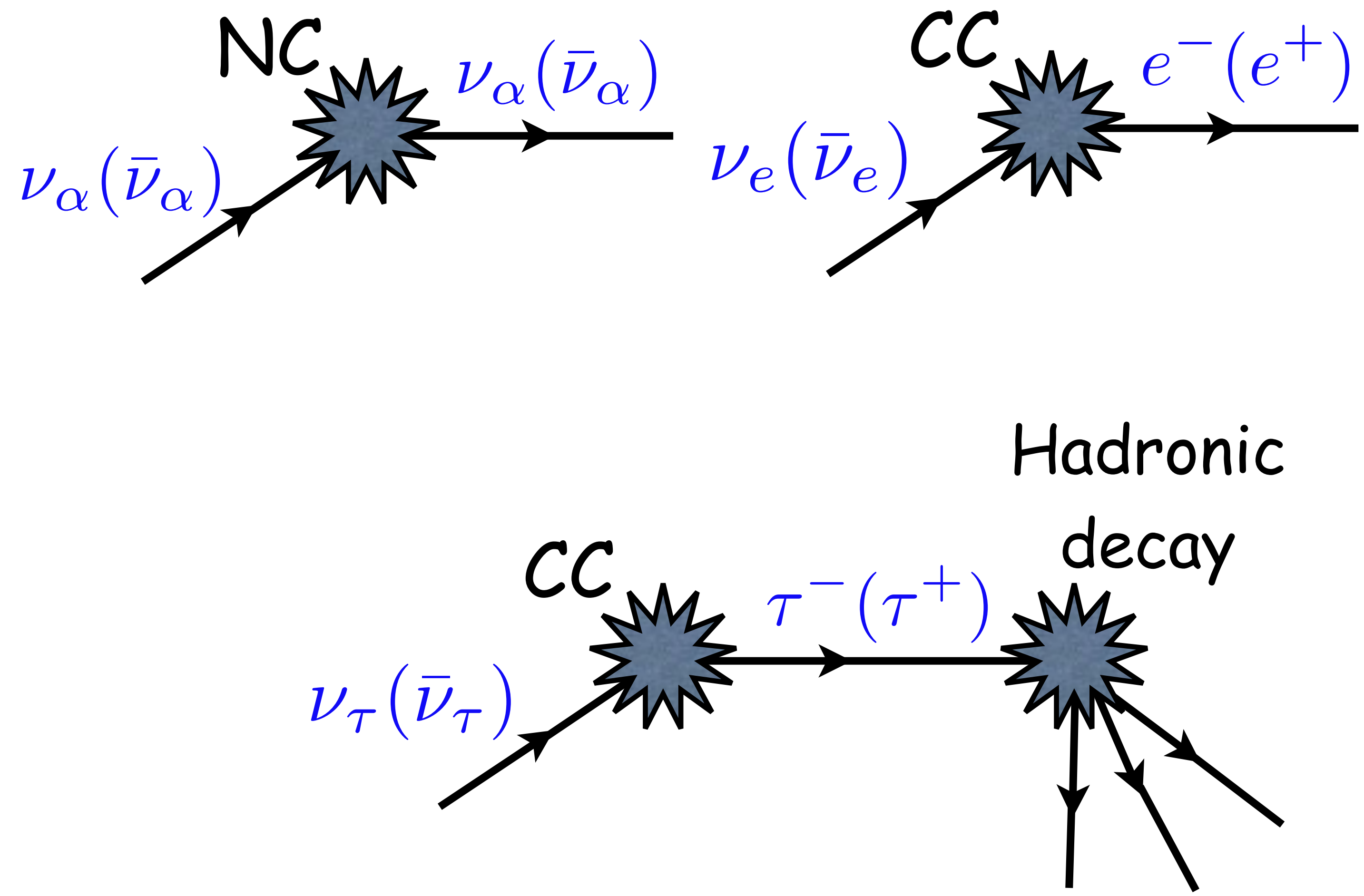
Flavoring at IceCube

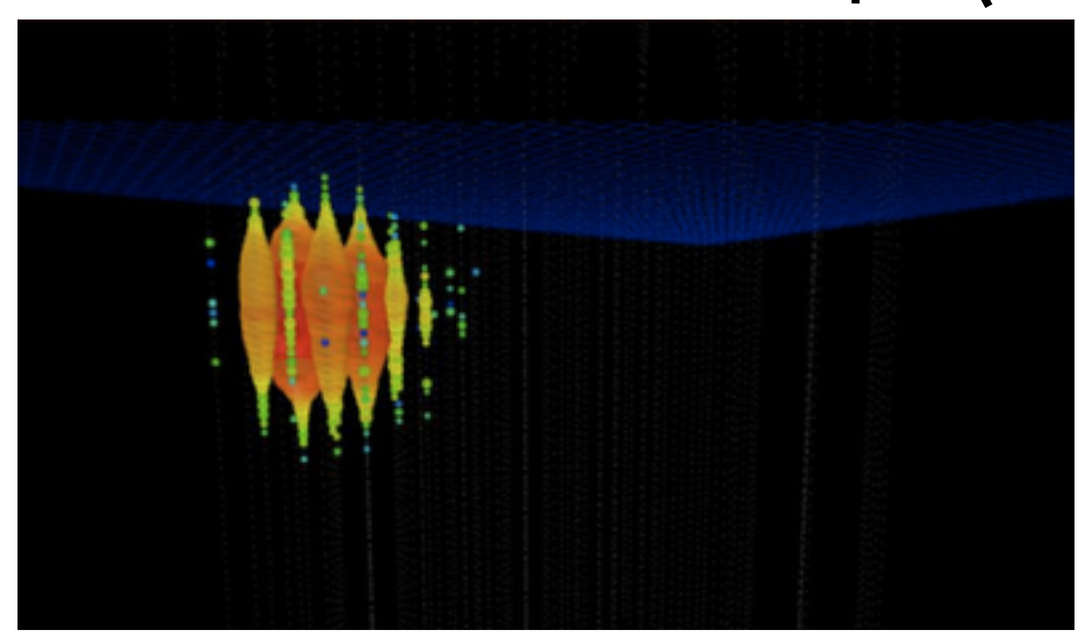
muon-track events





cascade events





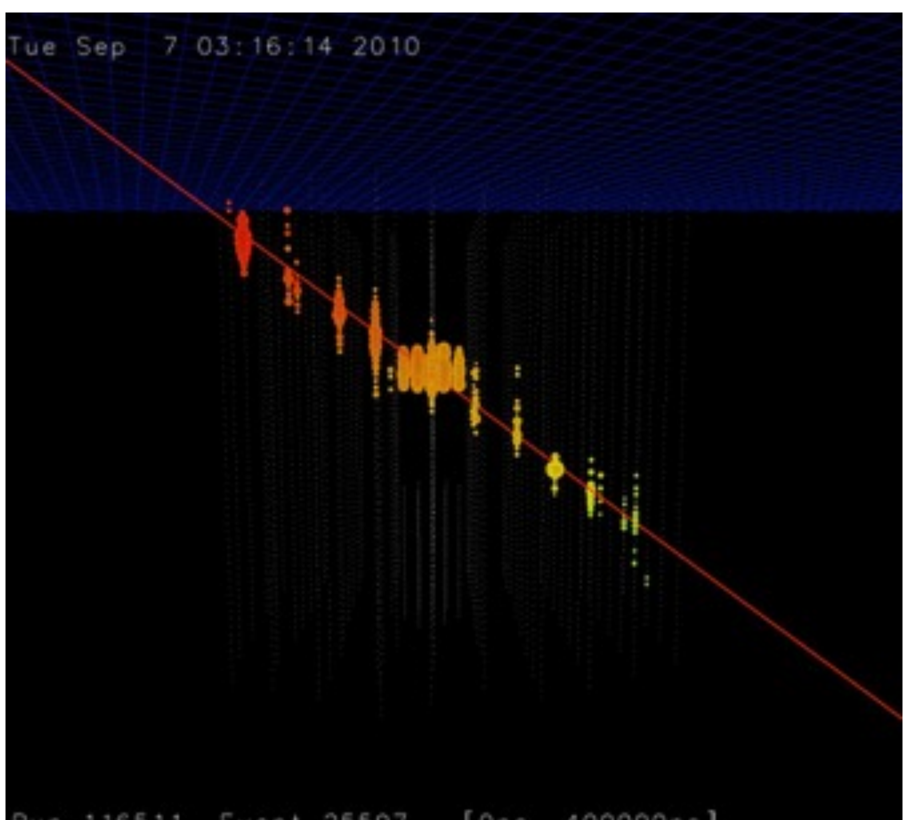
figures from IceCube website

Flavoring at IceCube

muon-track events

great angular resolution (< 1°)

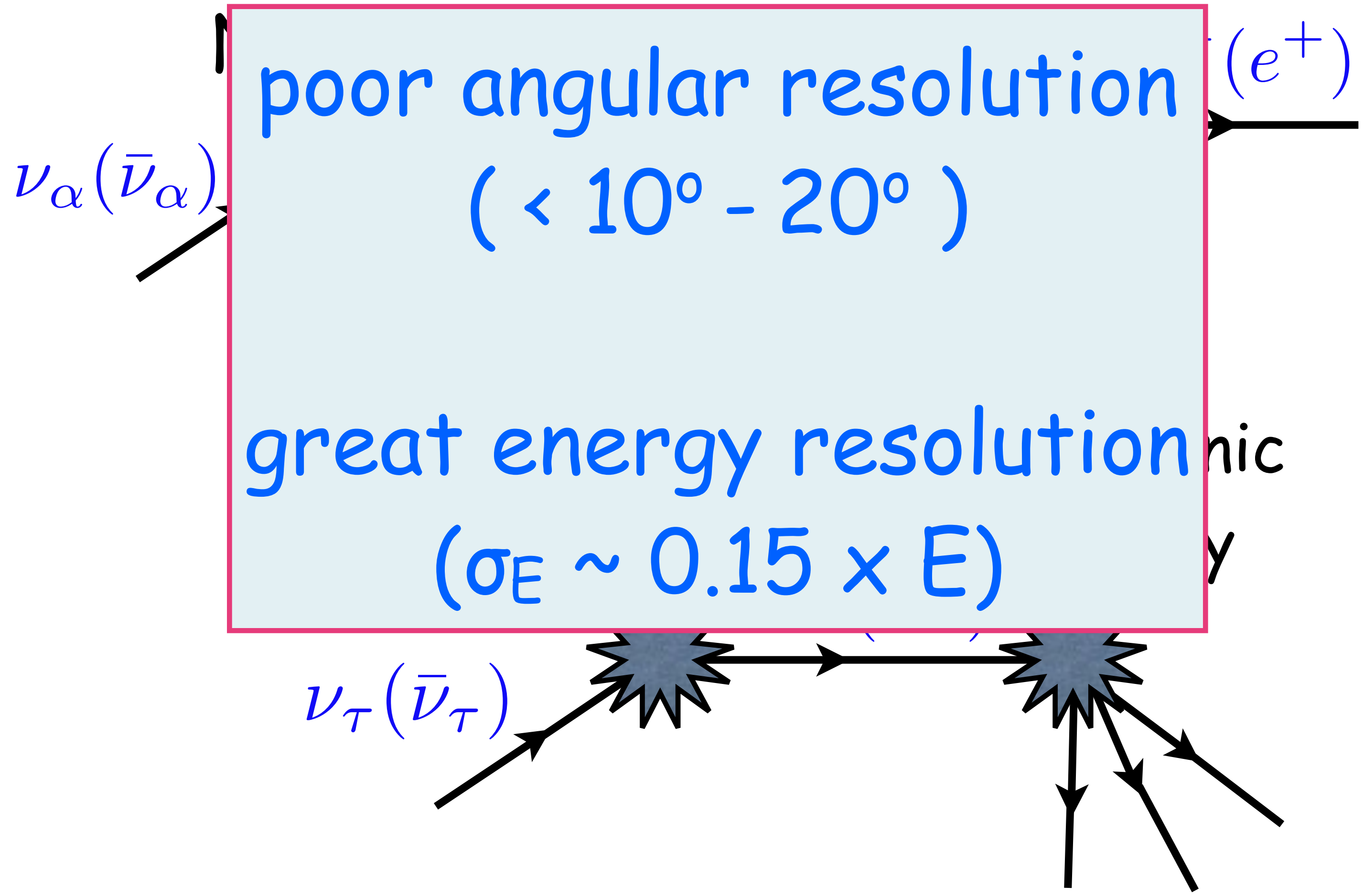
moderate energy resolution ($\sigma_E \sim E$)

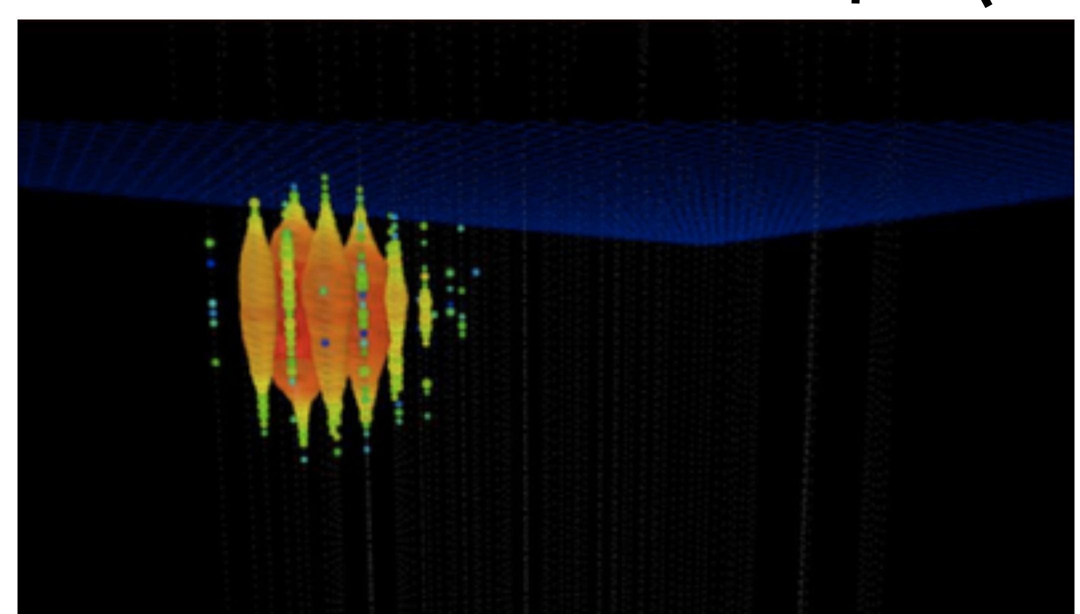


cascade events

poor angular resolution (< 10° - 20°)

great energy resolution ($\sigma_E \sim 0.15 \times E$)

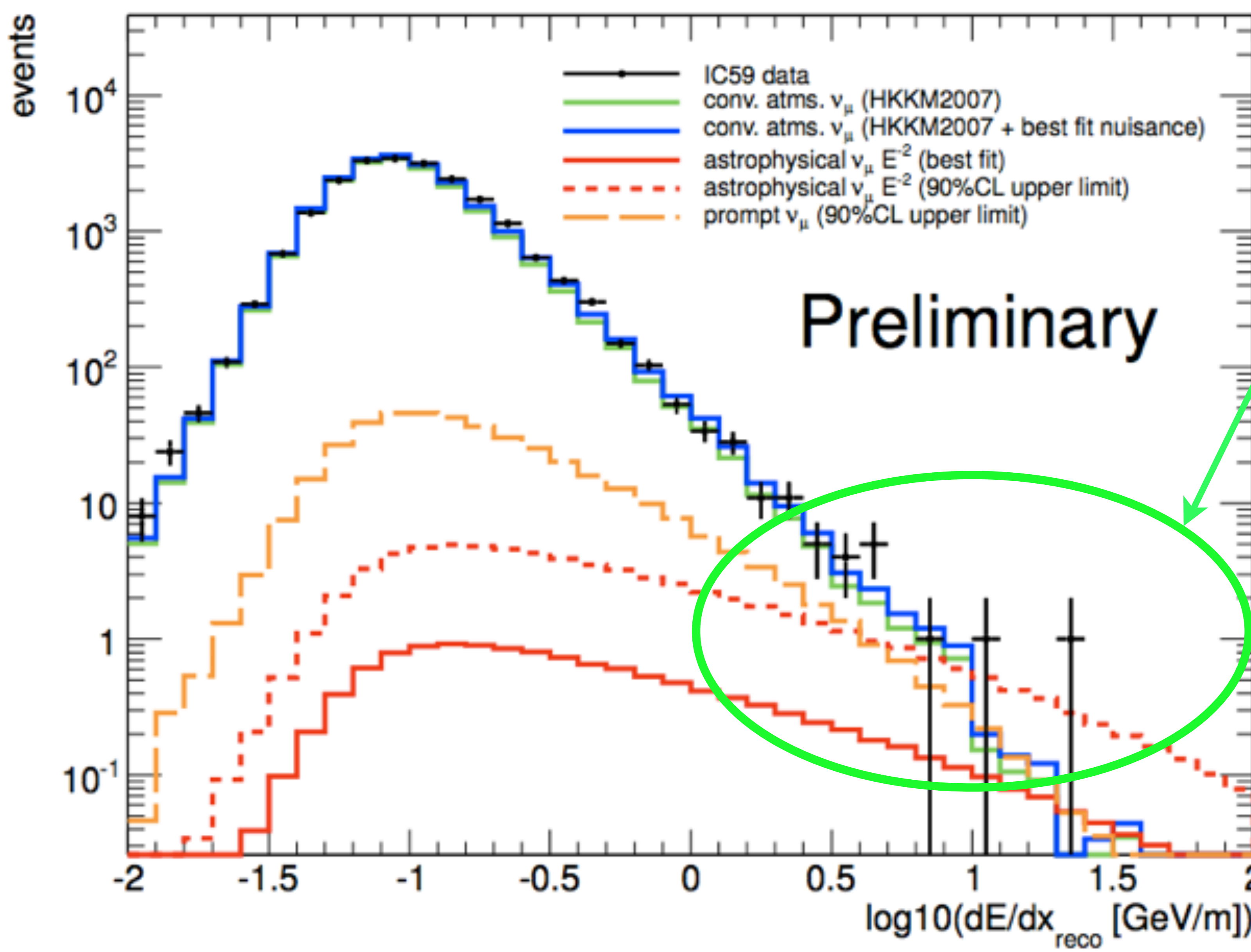




figures from IceCube website

Mission for IceCube began !

✓ muon-track events at IceCube-59, 348 days livetime.



excess in high energy tail (~ 300 TeV)

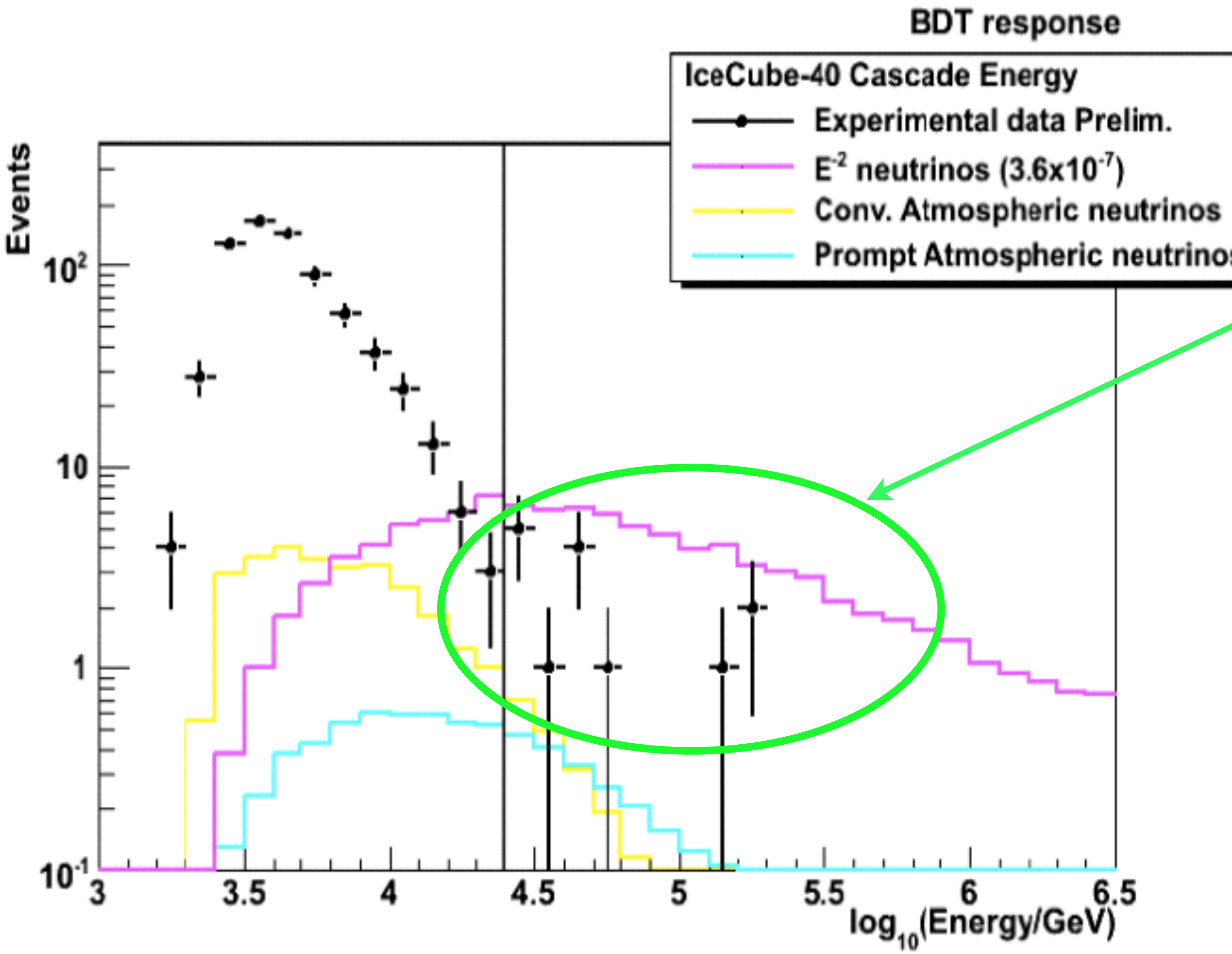
prompt atm neutrinos ?
or
astrophysical neutrinos?

A. Schukraft [IceCube Collaboration]
Nucl. Phys. Proc. Suppl. 266 (2013) [arXiv:1302.0127]

analysis shows preference to astrophysical origin of excess, significance ~ 2.1σ

Mission for IceCube began !

✓ cascade events at IceCube-40, 367 days livetime.



E. Middell, PhD thesis

excess in high energy tail (> 100 TeV)

prompt+conventional ~ 4

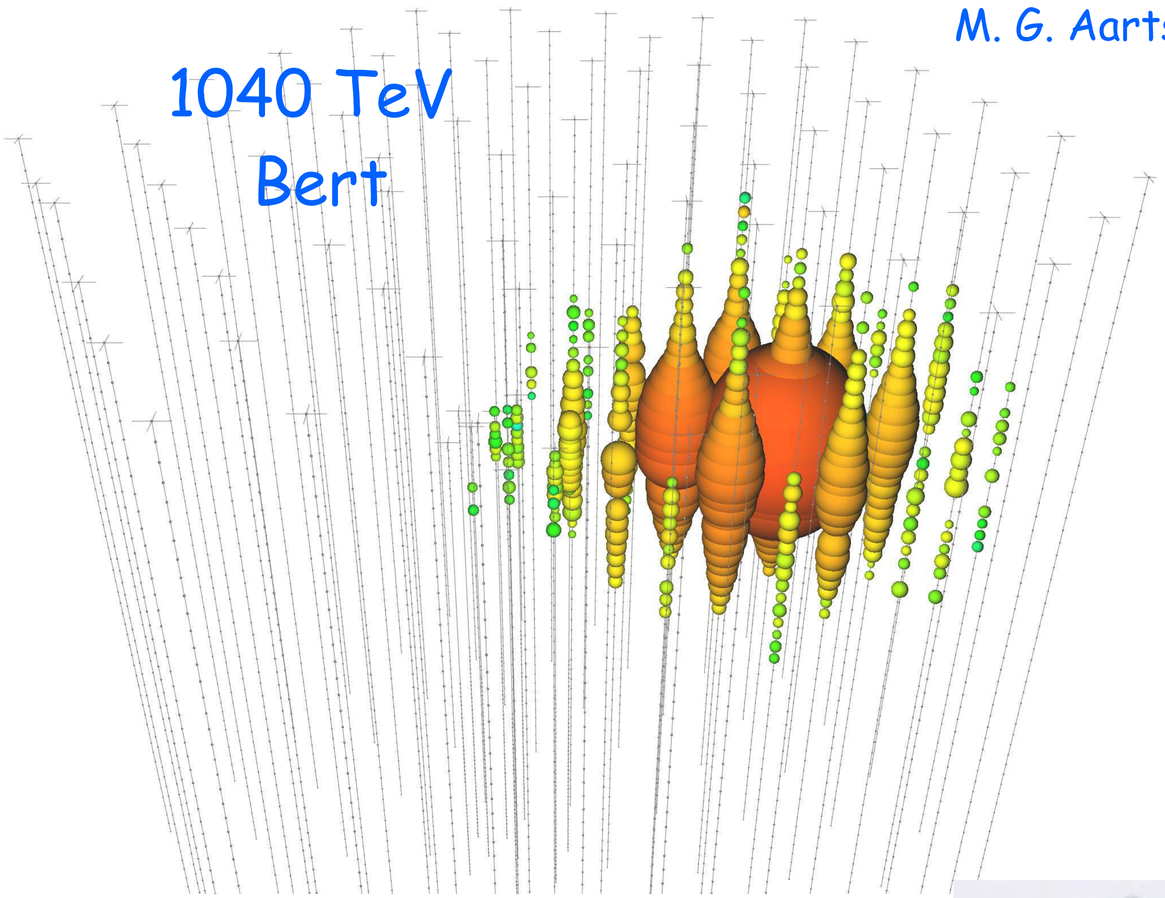
observed ~ 14

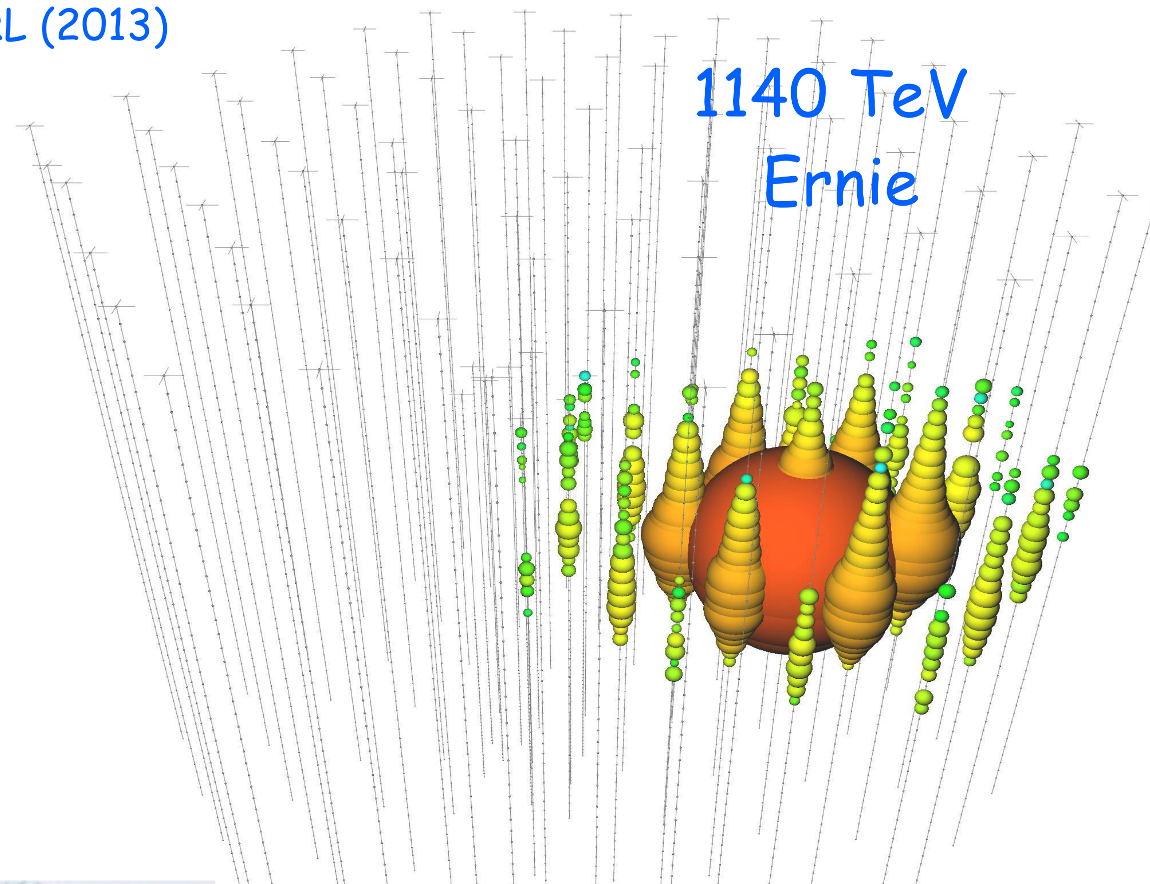
hint on astrophysical flux ~ 2.4σ

Mission for IceCube began !

✓ The two PeV cascade events, 616 days livetime

M. G. Aartsen et al, PRL (2013)

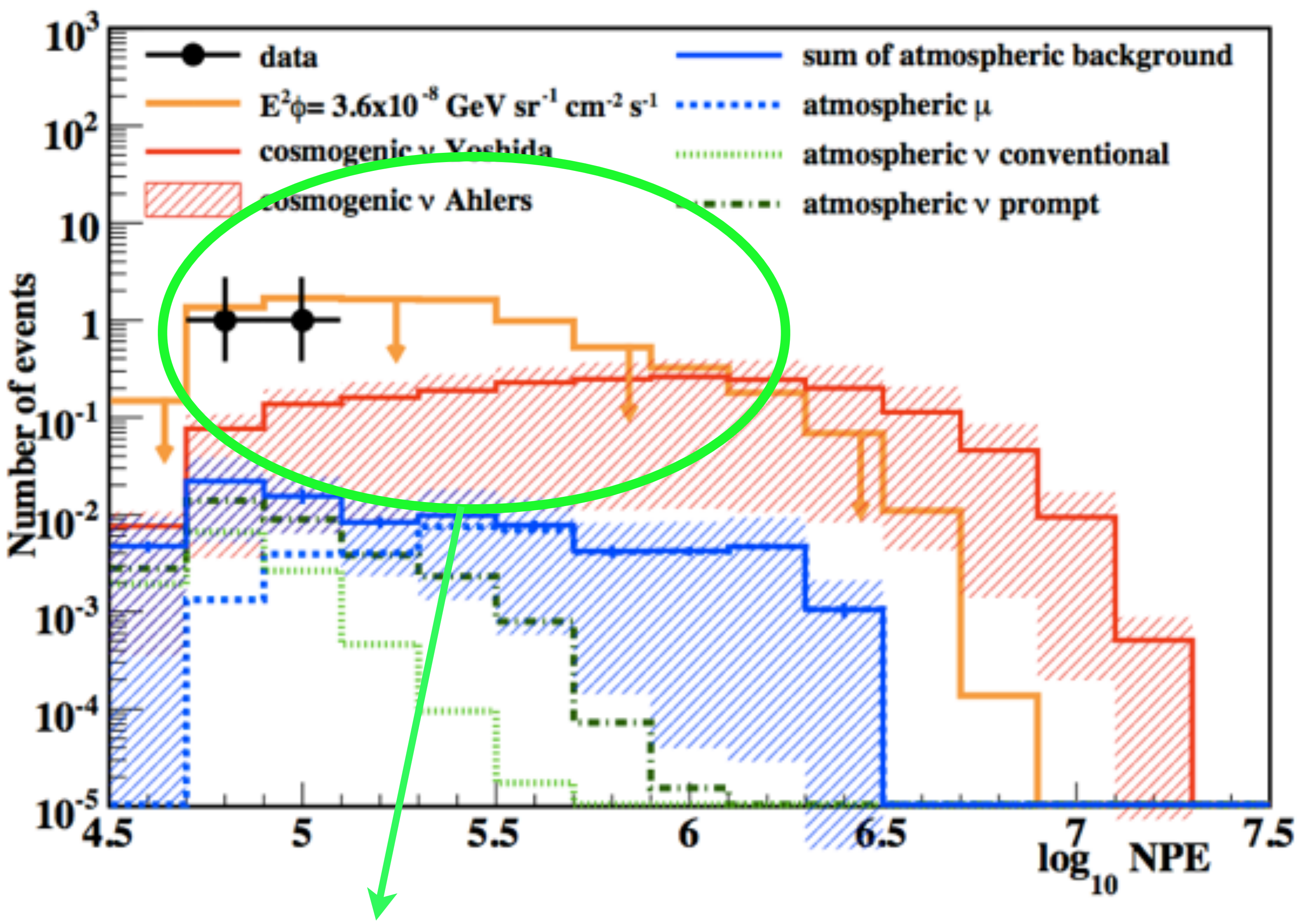




Mission for IceCube began !

✓ The two PeV cascade events, 616 days livetime

M. G. Aartsen et al. [IceCube Collaboration], Phys. Rev. Lett. 111 (2013), [arXiv:1304.5356]



it is like a cut-off at ~ PeV

expected bkg. (conventional+prompt) ~ 0.08(-0.057)(+0.041) sys.

excess of events ~ 2.8σ

GZK ? too low energy, more events should be seen in higher energies

astrophysical ? an E^{-2} spectrum would give ~ 9 more events in higher energies

demands more statistics

flavor composition ?
NC of ν_α or CC of ν_e

isotropy ?

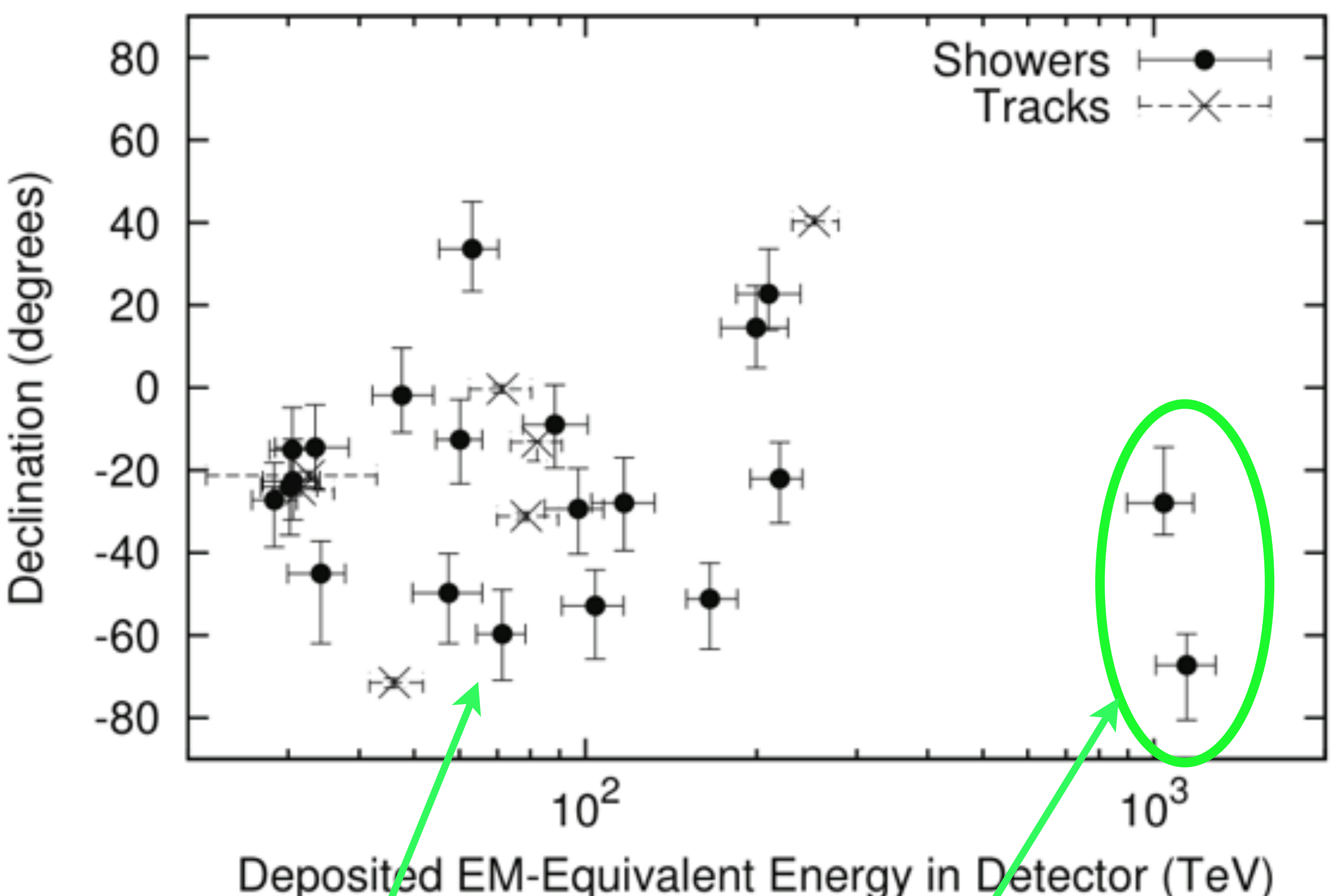
HESE analysis

- ✓ select events interacting inside the detector only
- ✓ no light in the veto region
- ✓ veto for atmospheric muons and neutrinos (which are typically accompanied by muons)
- ✓ energy measurement: total absorption calorimetry

Mission for IceCube began !

✓ Looking for lower energy contained events, 662 days livetime

M. G. Aartsen et al. [IceCube Collaboration], Science 342 (2013), [arXiv:1311.5238]



The whole family!

previous PeV cascade events (Bert and Ernie)

✓ 26 more events

✓ all the new events are lower in energy

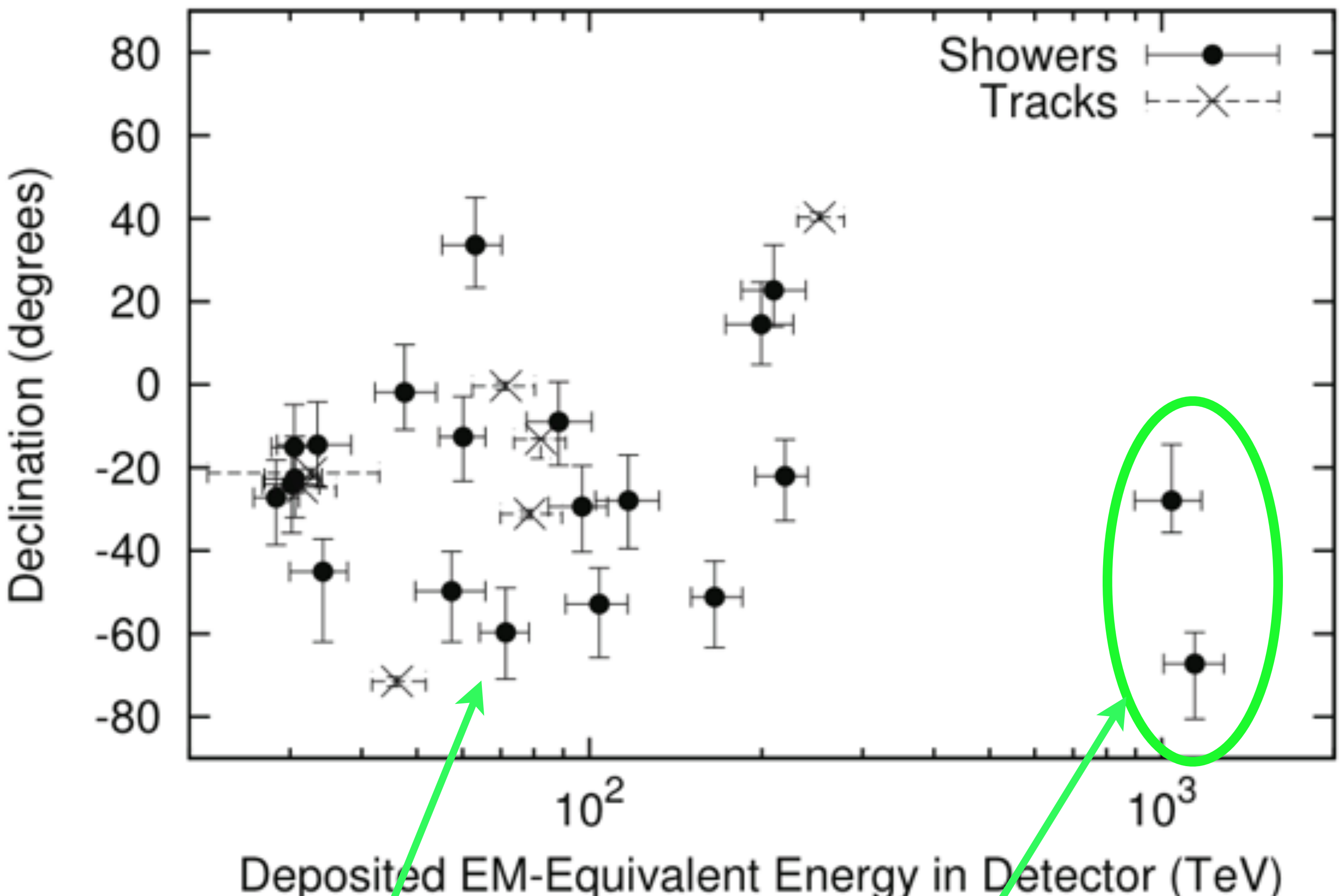
✓ expected bkg. (conventional+prompt) ~ 10.6(+4.5)(-3.5) sys.

excess of events ~ 4.3σ

Mission for IceCube began !

✓ Looking for lower energy contained events, 662 days livetime

M. G. Aartsen et al. [IceCube Collaboration], Science 342 (2013), [arXiv:1311.5238]



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previous PeV cascade events (Bert and Ernie)

✓ 26 more events

✓ all the new events are lower in energy

✓ expected bkg. (conventional+prompt) ~ 10.6(+4.5)(-3.5) sys.

excess of events ~ 4.3σ

which one?

atmospheric ?

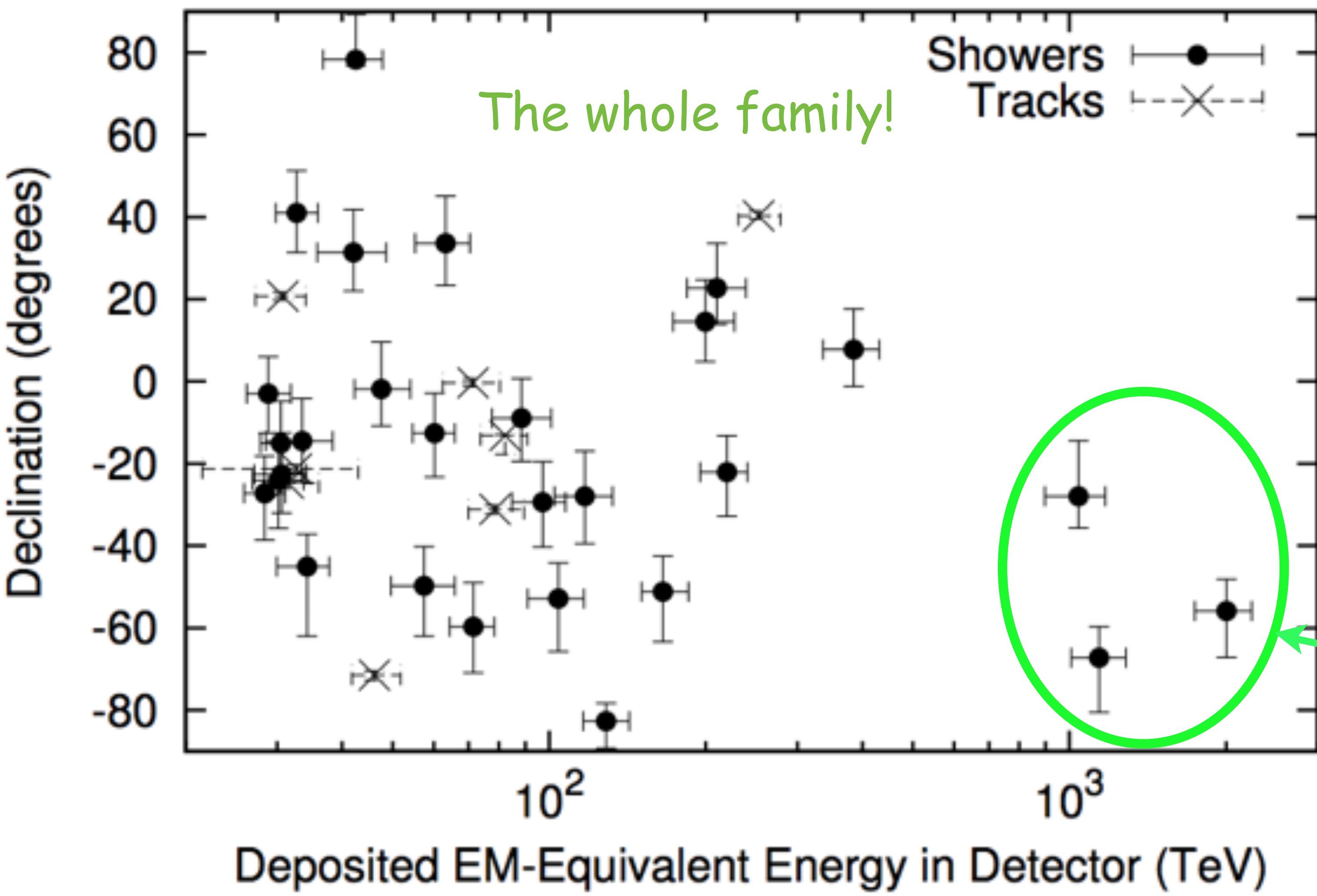
astrophysical ?

or something else ?

Mission for IceCube began !

✓ Looking for lower energy contained events, 988 days livetime

M. G. Aartsen et al. [IceCube Collaboration], PRL 113 (2014), [arXiv:1405.5303]



✓ totally 37 events

✓ three events with energy ~ PeV

✓ expected bkg. (conventional+prompt) ~ 15.6(+10.1)(-5.8) sys.

excess of events ~ 5.7σ

which one?

atmospheric ?

astrophysical ?

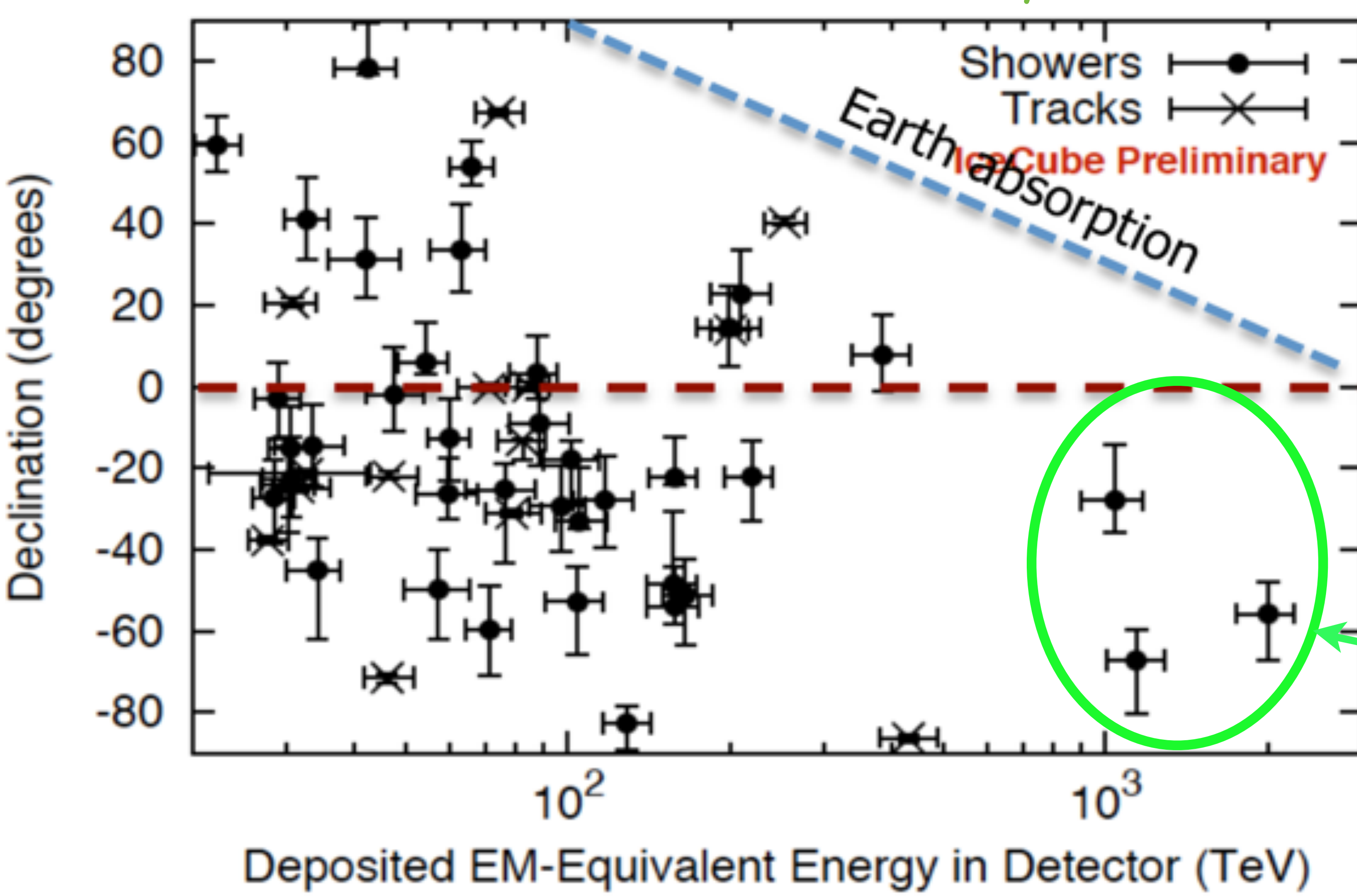
or something else ?

Mission for IceCube began !

✓ Looking for lower energy contained events, 1347 days livetime

IPA 2015

The whole family!



✓ totally 54 events

✓ still three events with energy ~ PeV

4 years of data

excess of events ~ 7σ

which one?

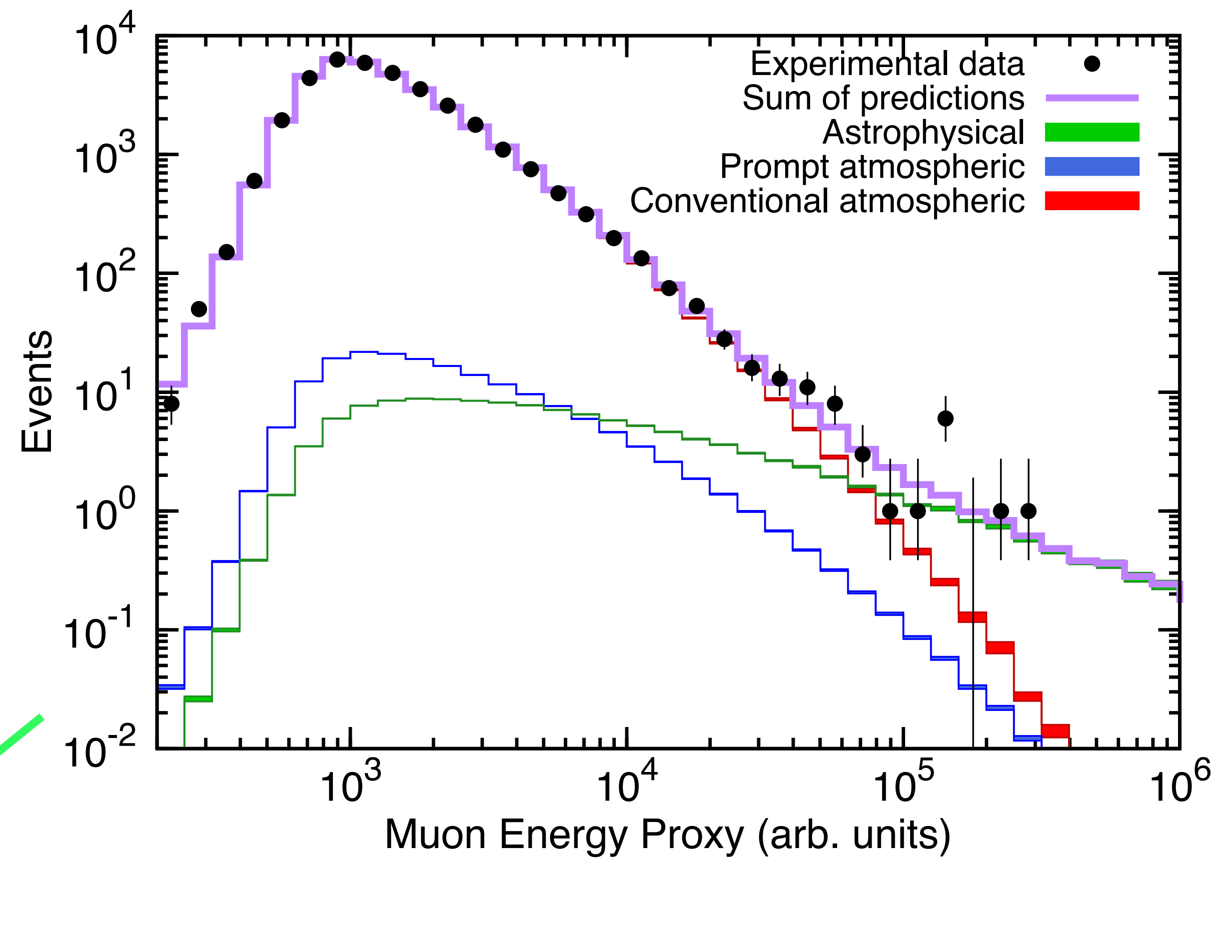
atmospheric ?

astrophysical ?

or something else ?

Mission for IceCube began !

✓ Looking for muon-track events, 660 days livetime



$$\Phi(E_\nu) = 9.9^{+3.9}_{-3.4} \times 10^{-19} \left(\frac{E_\nu}{100\,\mathrm{TeV}}\right)^{-2} \mathrm{GeV}^{-1}\,\mathrm{cm}^{-2}\,\mathrm{sr}^{-1}\,\mathrm{s}^{-1}$$

, fitting the index: 2.2 ± 0.2

lowering the energy threshold of HESE analysis to 1 TeV

$$\Phi(E_\nu) = 2.06^{+0.4}_{-0.3} \times 10^{-18} \left(\frac{E_\nu}{100\,\mathrm{TeV}}\right)^{-2.46\pm0.12} \mathrm{GeV}^{-1}\,\mathrm{cm}^{-2}\,\mathrm{sr}^{-1}\,\mathrm{s}^{-1}$$

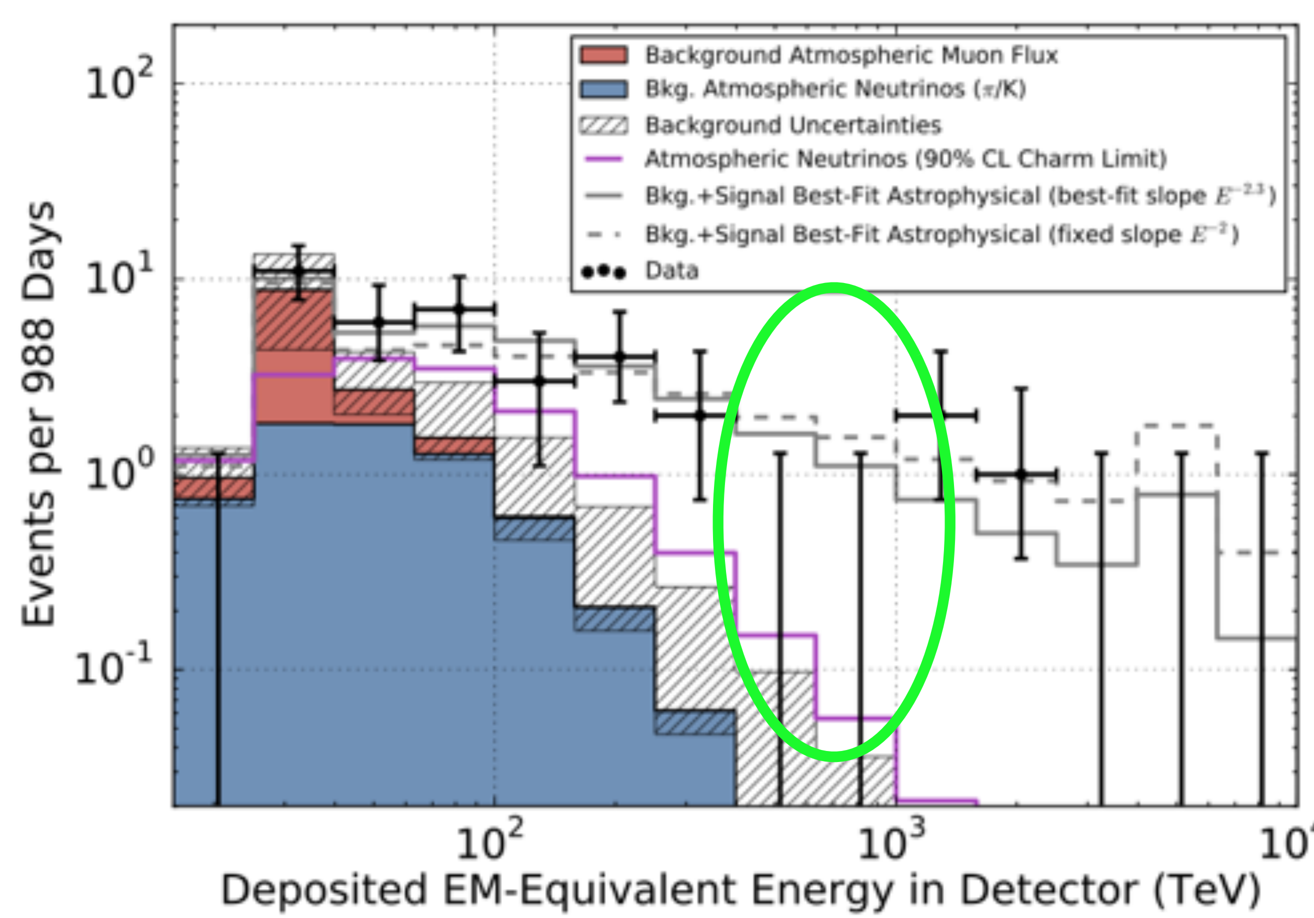
IceCube data

Problems with the astrophysical/atm interpretation of IceCube data

IceCube data

Problems with the astrophysical/atm interpretation of IceCube data

✓ deficit of events in the energy range ~ (400 - 1000) TeV

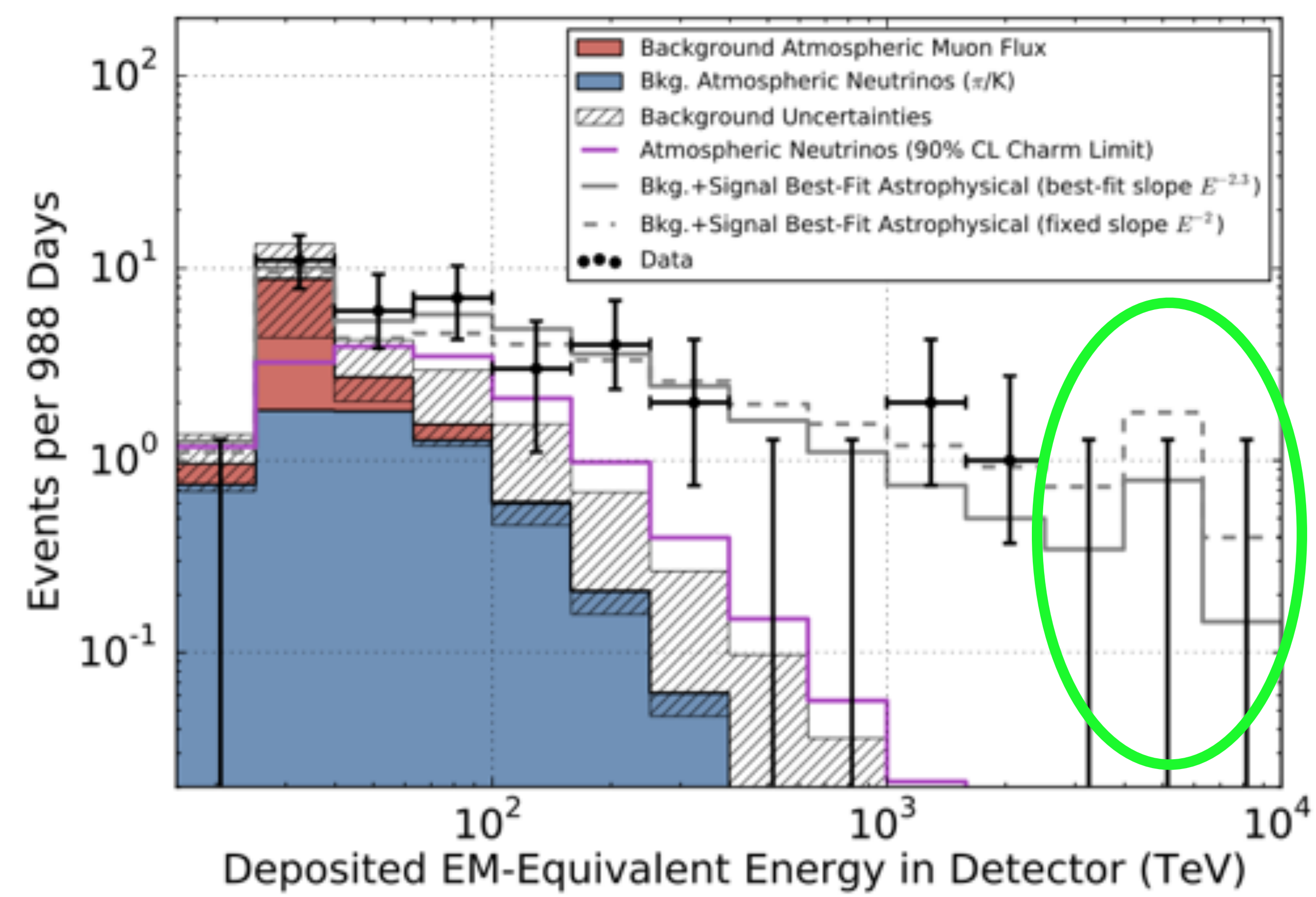


IceCube data

Problems with the astrophysical/atm interpretation of IceCube data

✓ deficit of events in the energy range ~ (400 - 1000) TeV

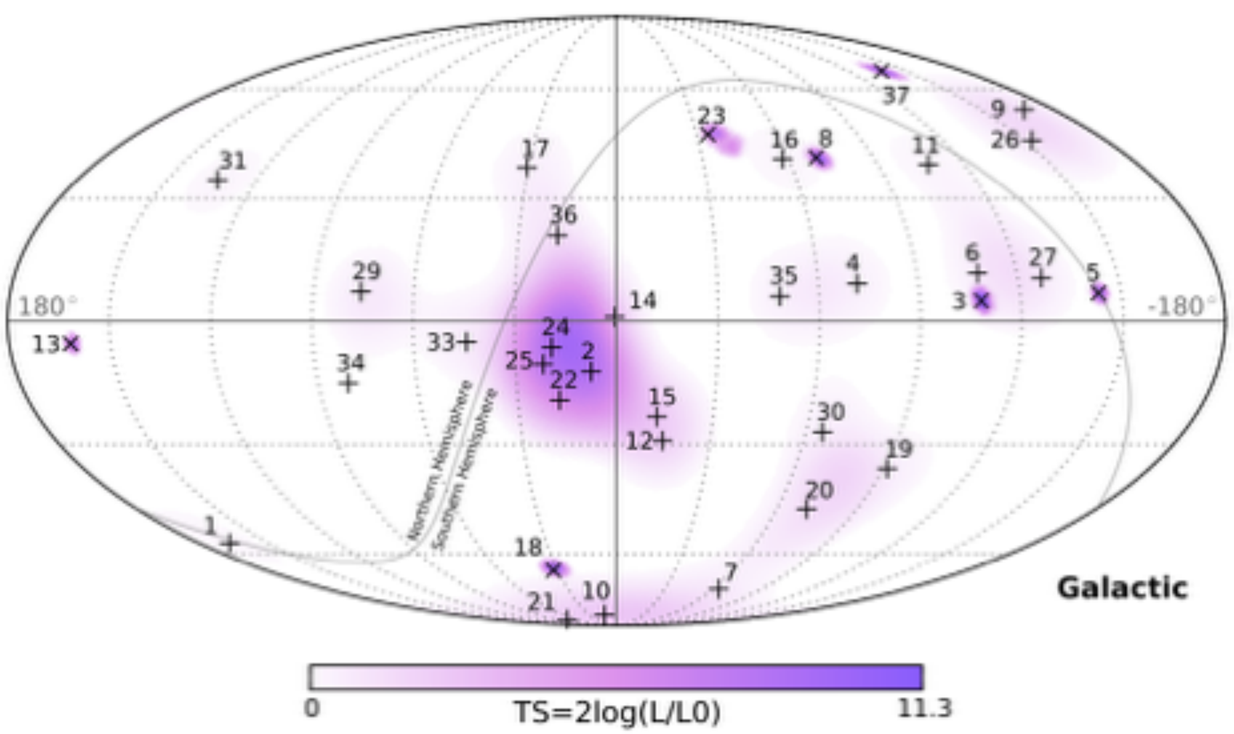
✓ cut-off in events: no events observed with energy > 2 PeV

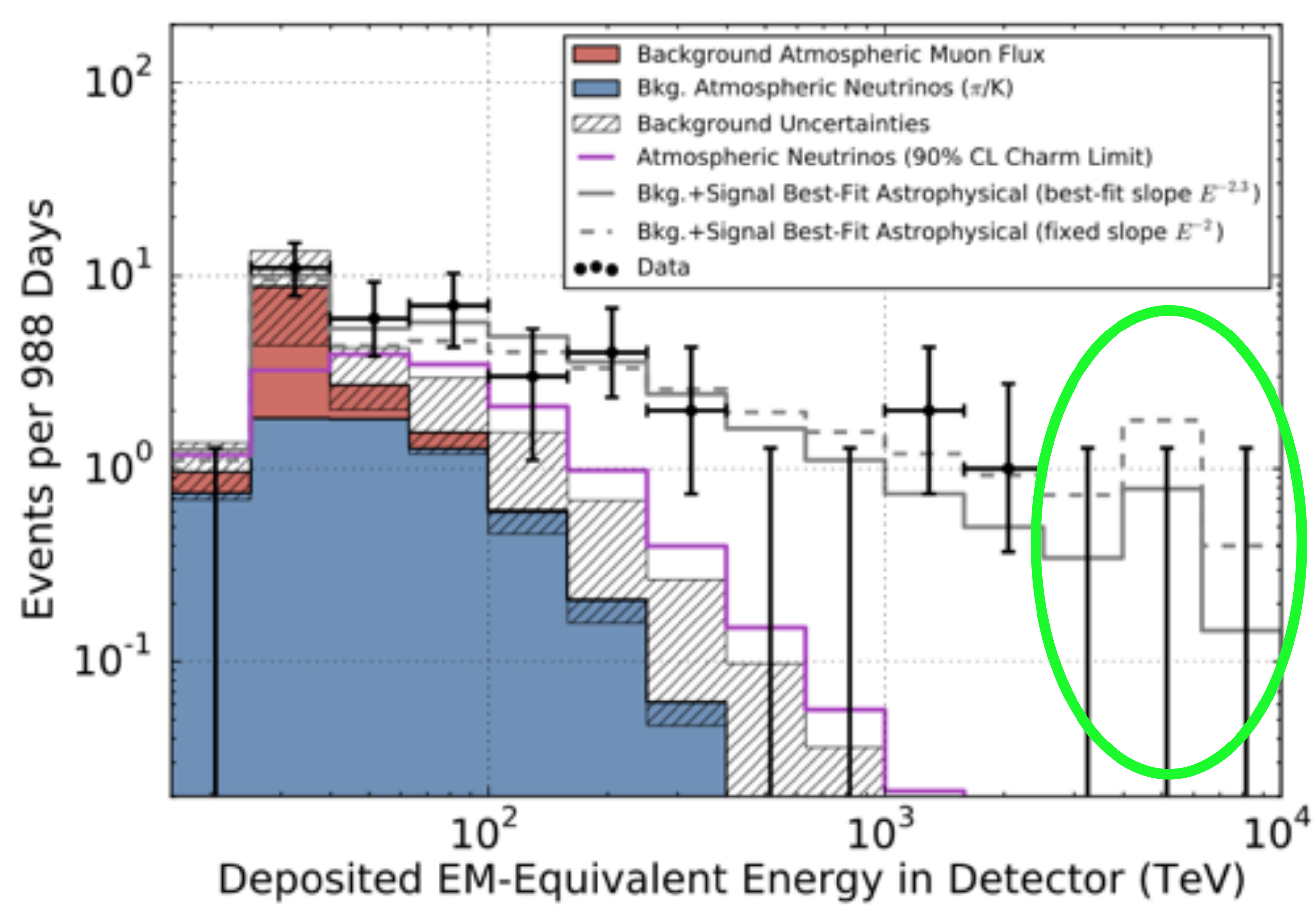


IceCube data

Problems with the astrophysical/atm interpretation of IceCube data

- ✓ deficit of events in the energy range ~ (400 - 1000) TeV
- ✓ cut-off in events: no events observed with energy > 2 PeV
- ✓ angular distribution of events show mild anisotropies (enhanced toward GC)





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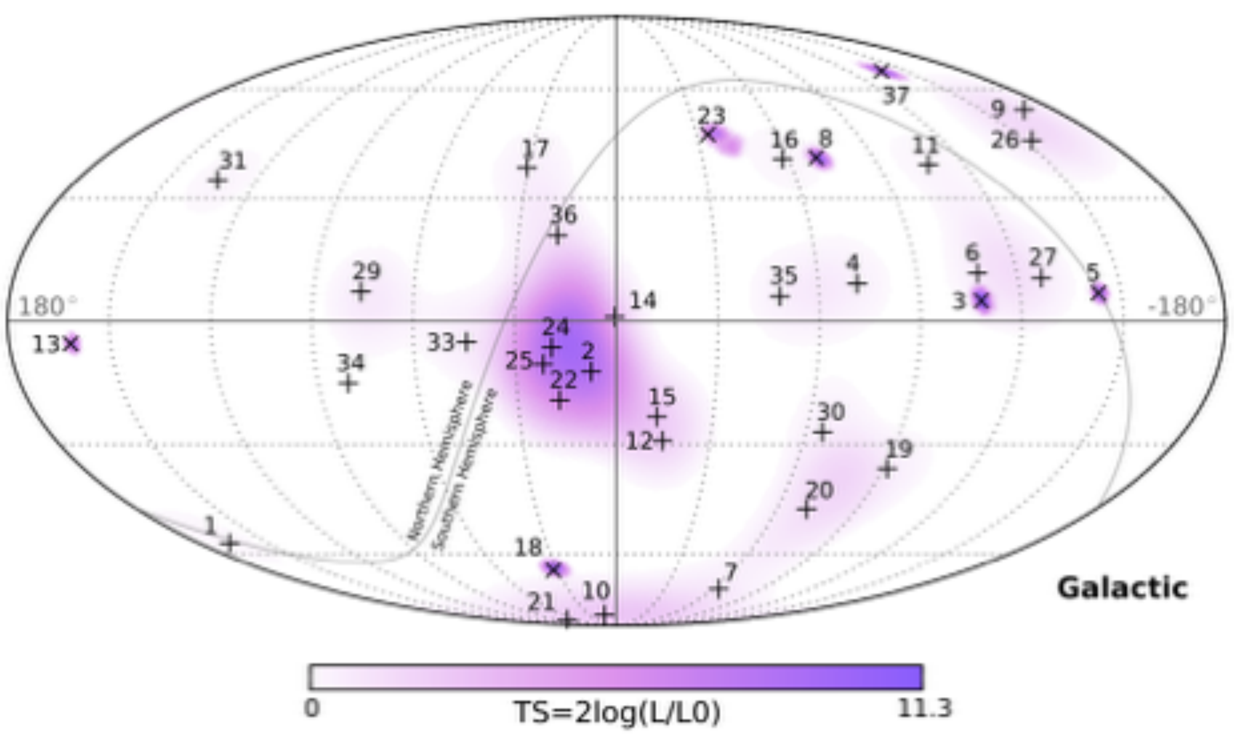
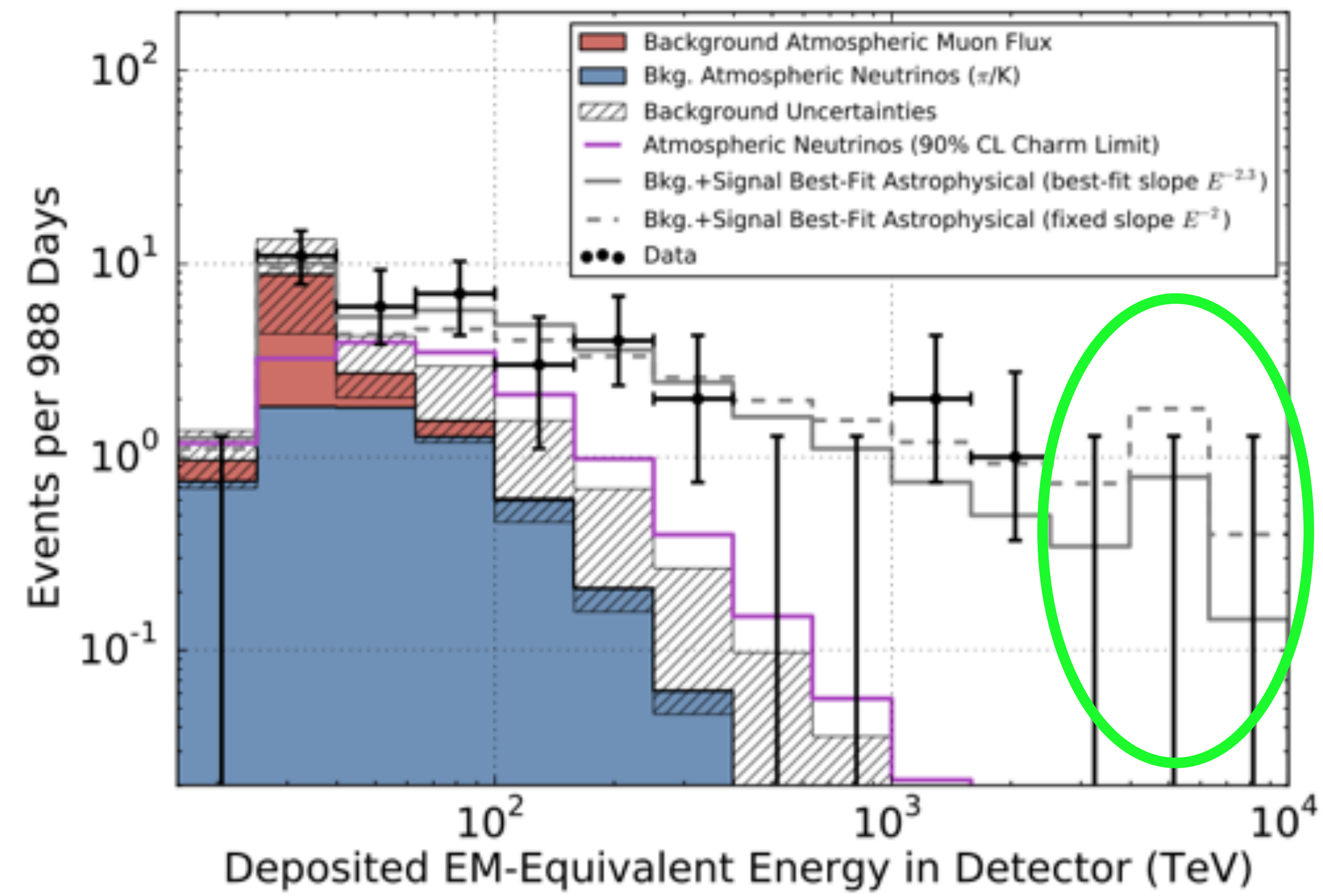
⚠ none of the above-mentioned issues are significant





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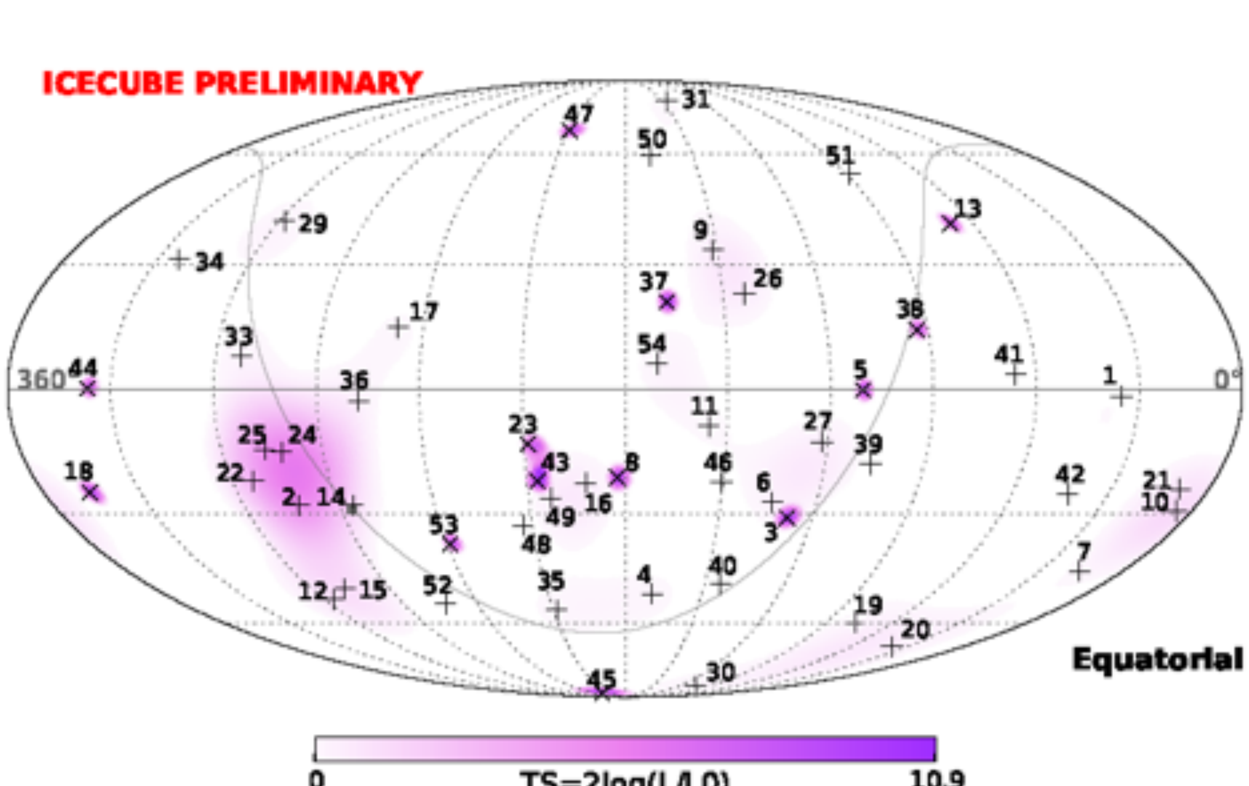

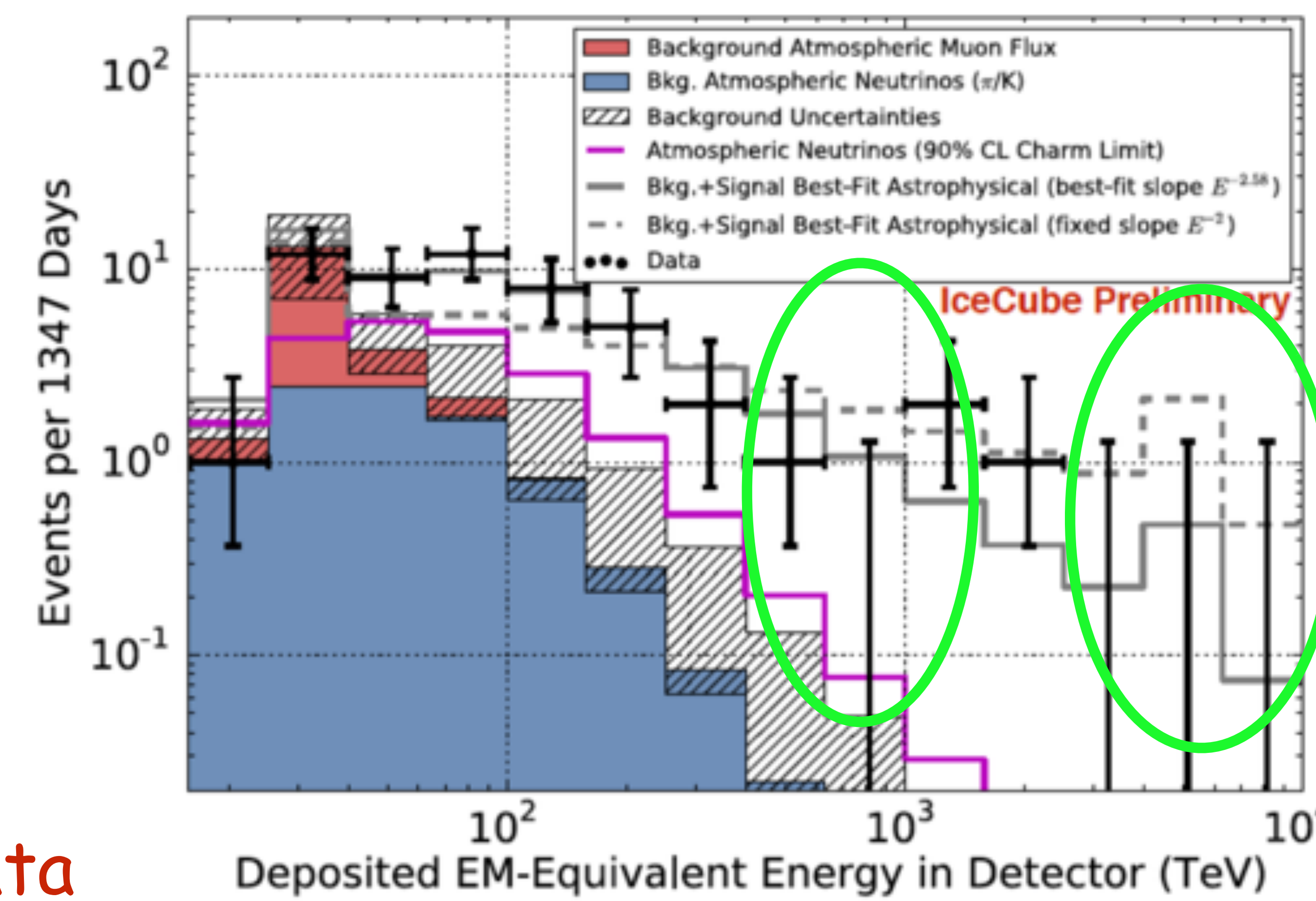

4 years of data

Interpretations of IceCube data

✓ "Conventional" interpretations of IceCube data

Cosmic ray sources

GRBs

Galaxy clusters

Star-forming galaxies

AGNs

Fermi bubbles

Galactic Center activities

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K. Murase, M. Ahlers and B. C. Lacki, Phys. Rev. D 88, no. 12, 121301 (2013) [arXiv:1306.3417 [astro-ph.HE]].
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A. Bhattacharya, R. Enberg, M. H. Reno and I. Sarcevic, arXiv:1407.2985 [astro-ph.HE].
C. Lunardini, S. Razzaque, K. T. Theodoseau and L. Yang, Phys. Rev. D 90, 023016 (2014) [arXiv:1311.7188 [astro-ph.HE]].

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For a review

L. A. Anchordoqui, V. Barger, I. Cholis, H. Goldberg, D. Hooper, A. Kusenko, J. G. Learned and D. Marfatia et al., Journal of High Energy Astrophysics 1-2, 1 (2014) [arXiv:1312.6587 [astro-ph.HE]].

Interpretations of IceCube data

✓ "New Physics" interpretations of IceCube data

Lepto-quarks

Secret neutrino interactions

resonant absorption on cosmic neutrino background

supermassive long-lived particles

Dark matter decay

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Y. Ema, R. Jinno and T. Moroi, Phys. Lett. B 733, 120 (2014) [arXiv:1312.3501 [hep-ph]].
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A. Esmaili and P. D. Serpico, arXiv:1505.06486

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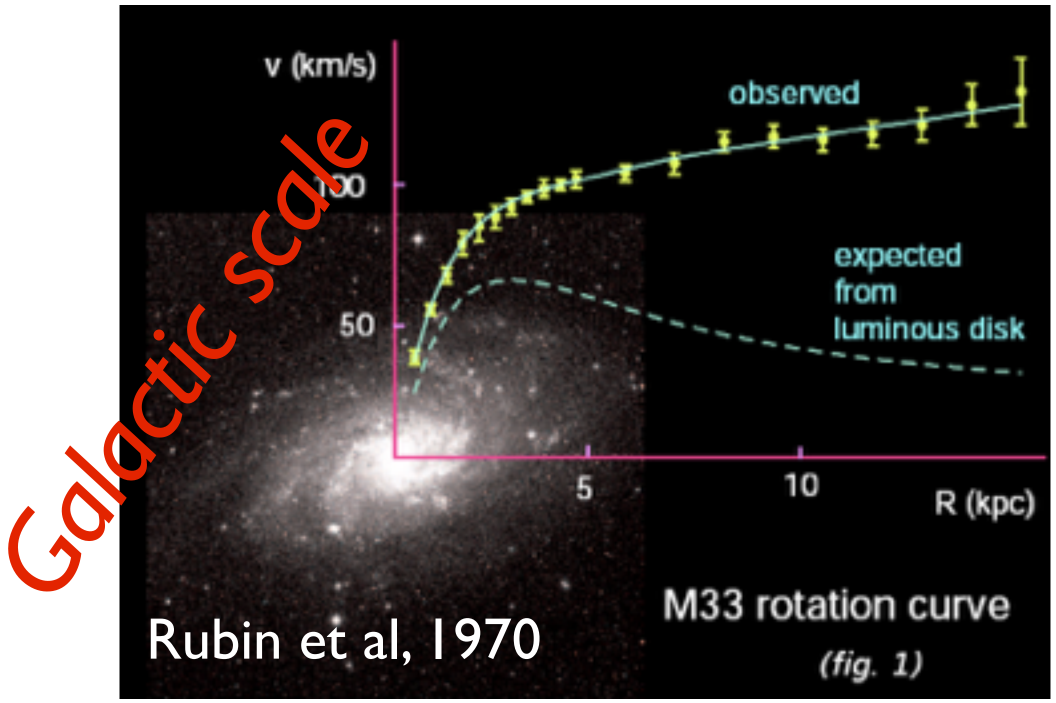
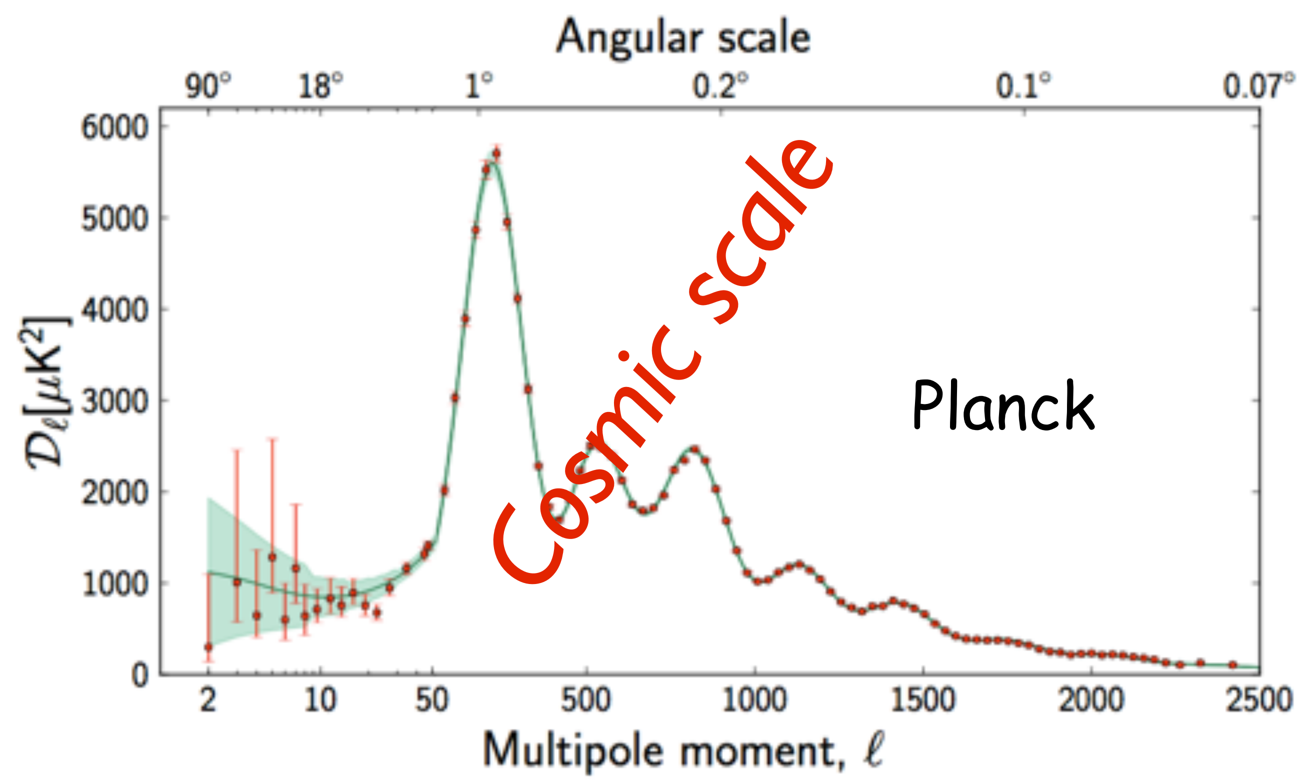
A note on Dark Matter

DM exist!





A note on Dark Matter

DM exist!

What We Know?

- ✓ Non Baryonic
- ✓ No electric and color charges
- ✓ Cold (or perhaps warm)
- ✓ Long lived (not necessarily stable)

All of these come from gravitational effects

A note on Dark Matter

DM exist!

What We Do Not Know?

✓ Mass

axion — Sterile ν — WIMP — Wimpzilla

10^{-6} eV — ~KeV — ~100 GeV — M_{GUT} (10^{16} GeV)

⚠ "WIMP" paradigm ?

Note that WIMP paradigm is a "particle physics" conjecture, needs to be validated at colliders

DM exist!

What We Do Not Know?

✓ Mass

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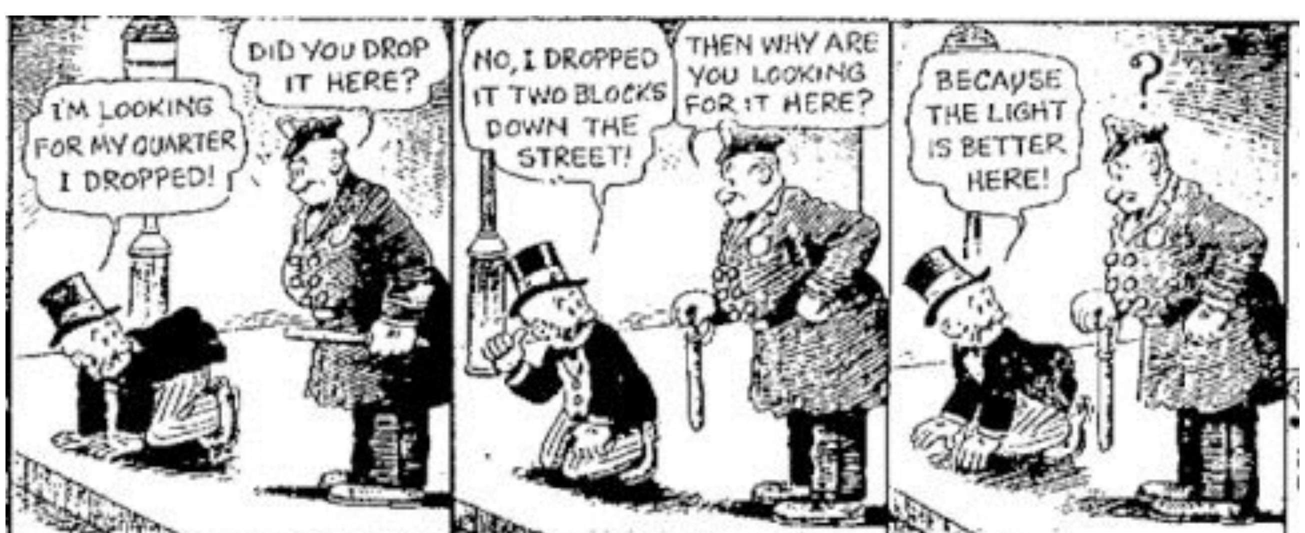
Note that WIMP paradigm is a "particle physics" conjecture, needs to be validated at colliders

caution: streetlight effect



Mulla Nasreddin

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A note on Dark Matter

DM exist!

What We Do Not Know?

✓ Mass

⚠ "WIMP" paradigm ?

Note that WIMP paradigm is a "particle physics" conjecture, needs to be validated at colliders

✓ Lifetime: stable (∞) or

$\tau_{DM} > 4.3 \times 10^{17}$ s (age of Universe)

$\tau_{DM} > 2.2 \times 10^{19}$ s (CMB) Y. Gong and X. Chen, PRD77 (2008), arXiv:0802.2296

✓ Possible decay and/or annihilation channels

✓ ...

Limits on lifetime from neutrino experiments

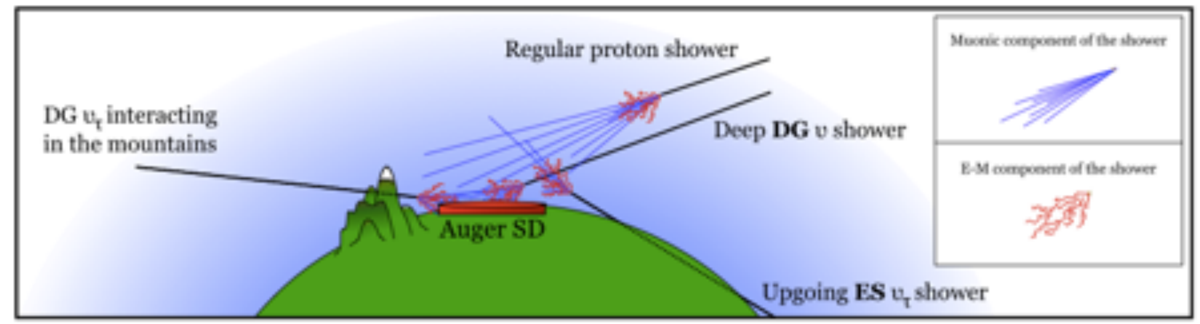
before recent IceCube data

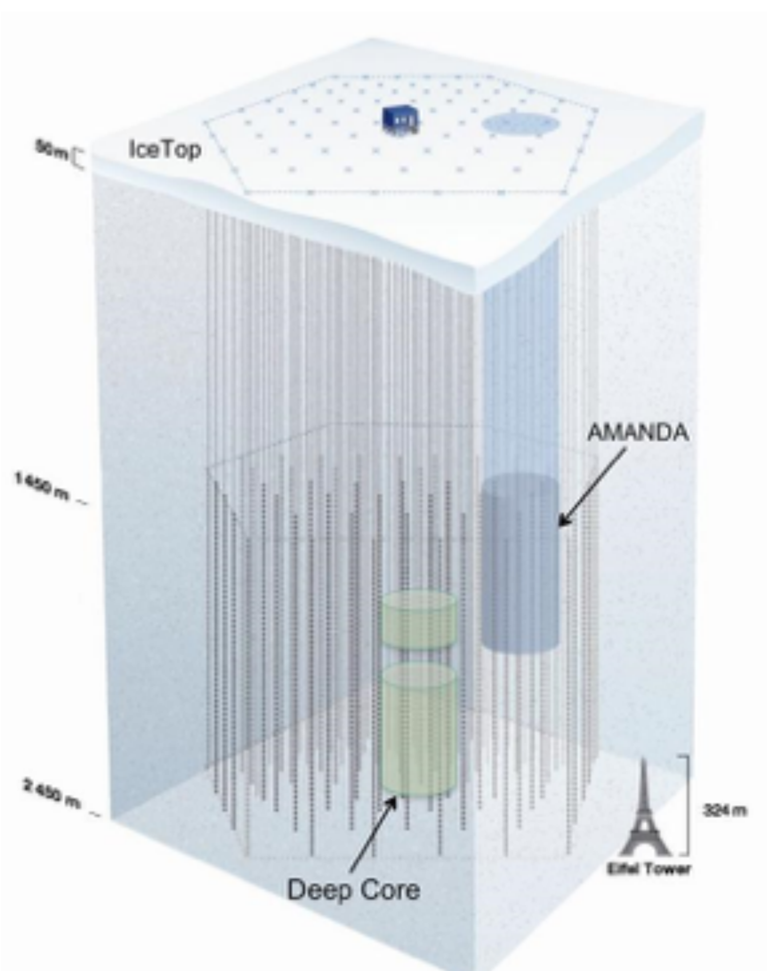
$16 - 2.5 \times 10^3$ TeV
AMANDA

$340 - 2 \times 10^5$ TeV
IceCube-22

$3 \times 10^3 - 6.3 \times 10^6$ TeV
IceCube-40

$10^5 - 10^8$ TeV
Auger

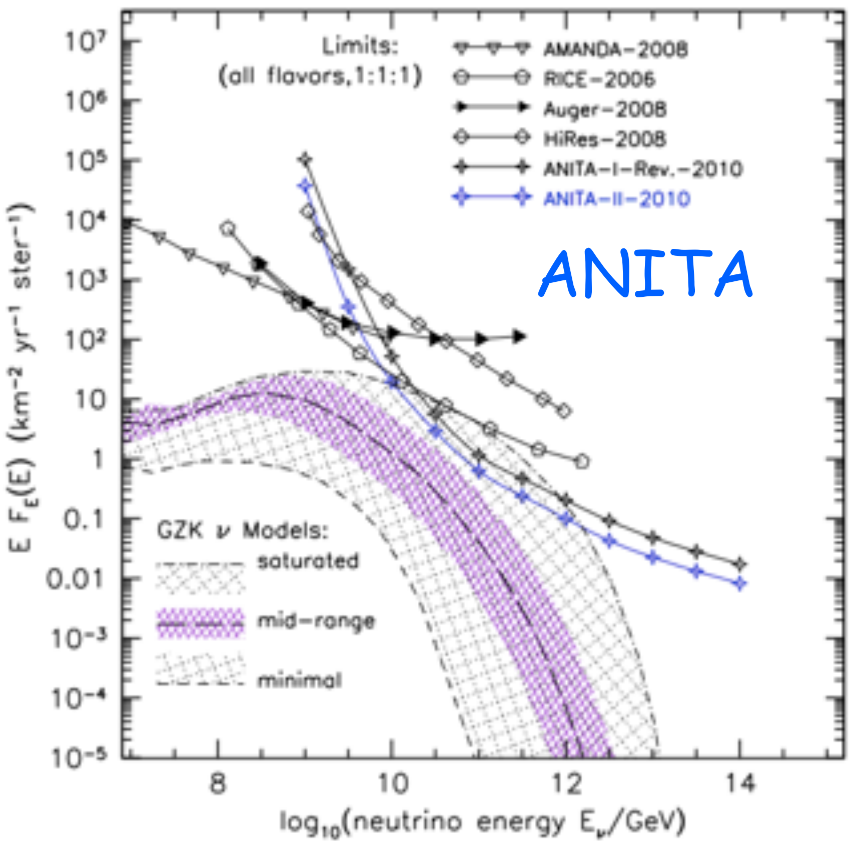




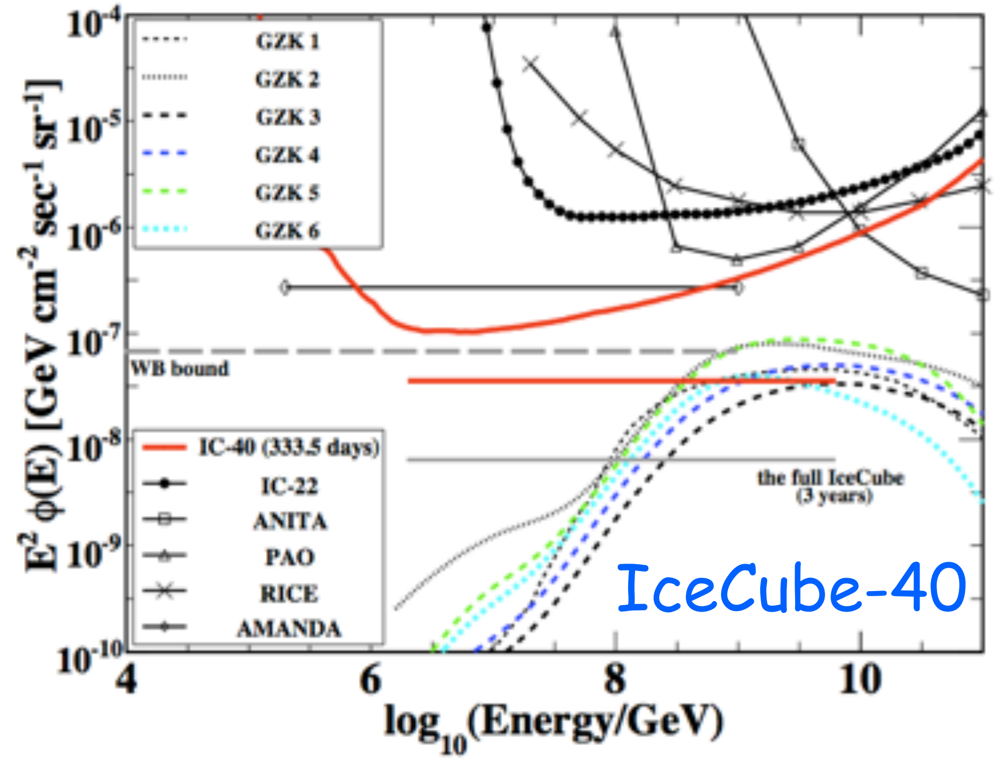
$10^6 - 3.2 \times 10^{11}$ TeV
ANITA

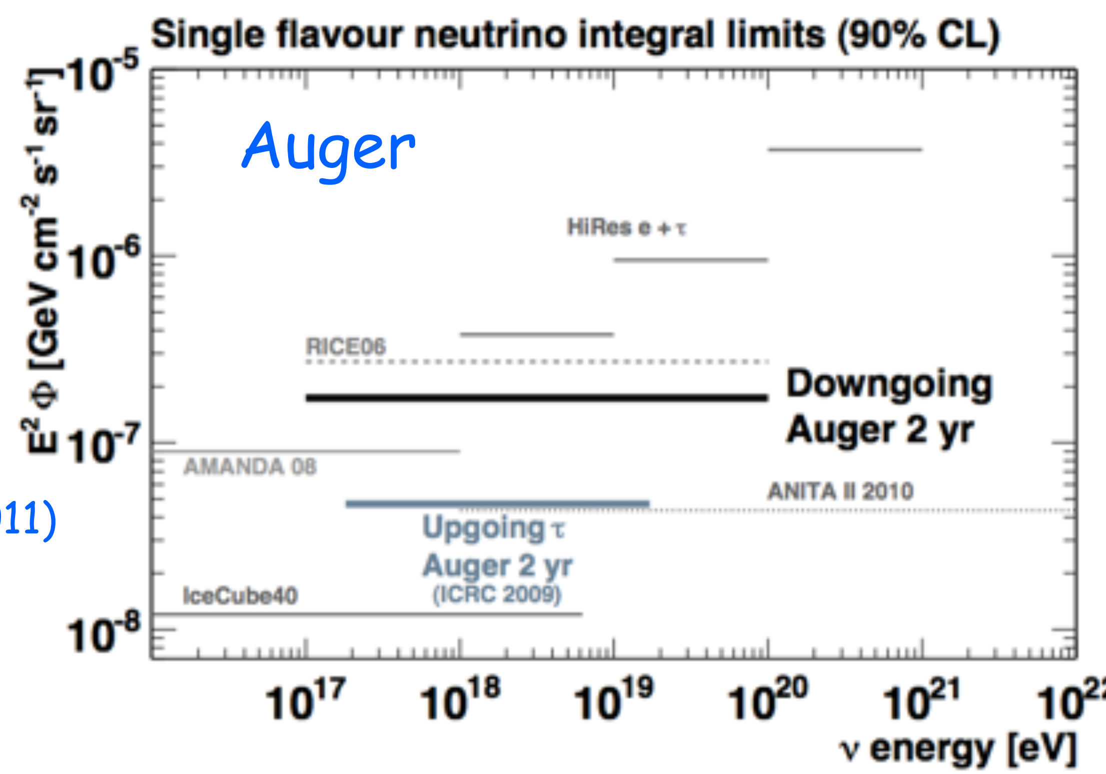
Experiments

P. Gorham et al. [ANITA Collaboration], Phys.Rev.D85 (2012)
arXiv:1011.5004, arXiv:1003.2961



R. Abbasi et al. [IceCube Collaboration], Phys.Rev.D83 (2011)
arXiv:1103.4250





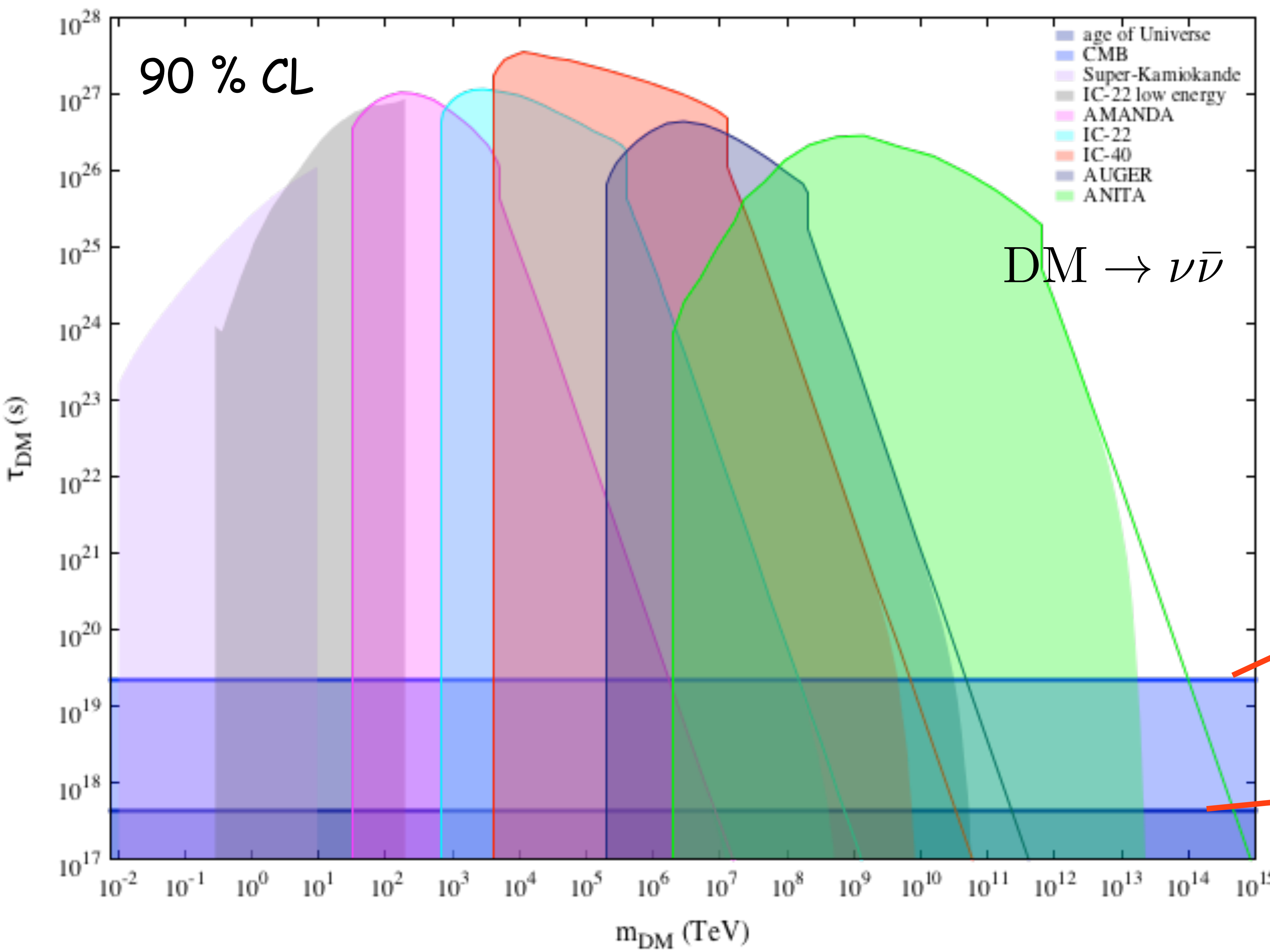
P. Abreu et al. [Pierre Auger Collaboration], Phys.Rev.D84 (2011)
arXiv:1202.1493

Limits on lifetime from neutrino experiments

before recent IceCube data

✓ Lifetime: stable (∞) or

A.E., Alejandro Ibarra and Orlando L. G. Peres
JCAP (2012) [arXiv: 1205.5281]



Using AMANDA, IceCube, Auger and ANITA data

$\tau_{DM} > 2.2 \times 10^{19}$ s (CMB)

$\tau_{DM} > 4.3 \times 10^{17}$ s (age of Universe)

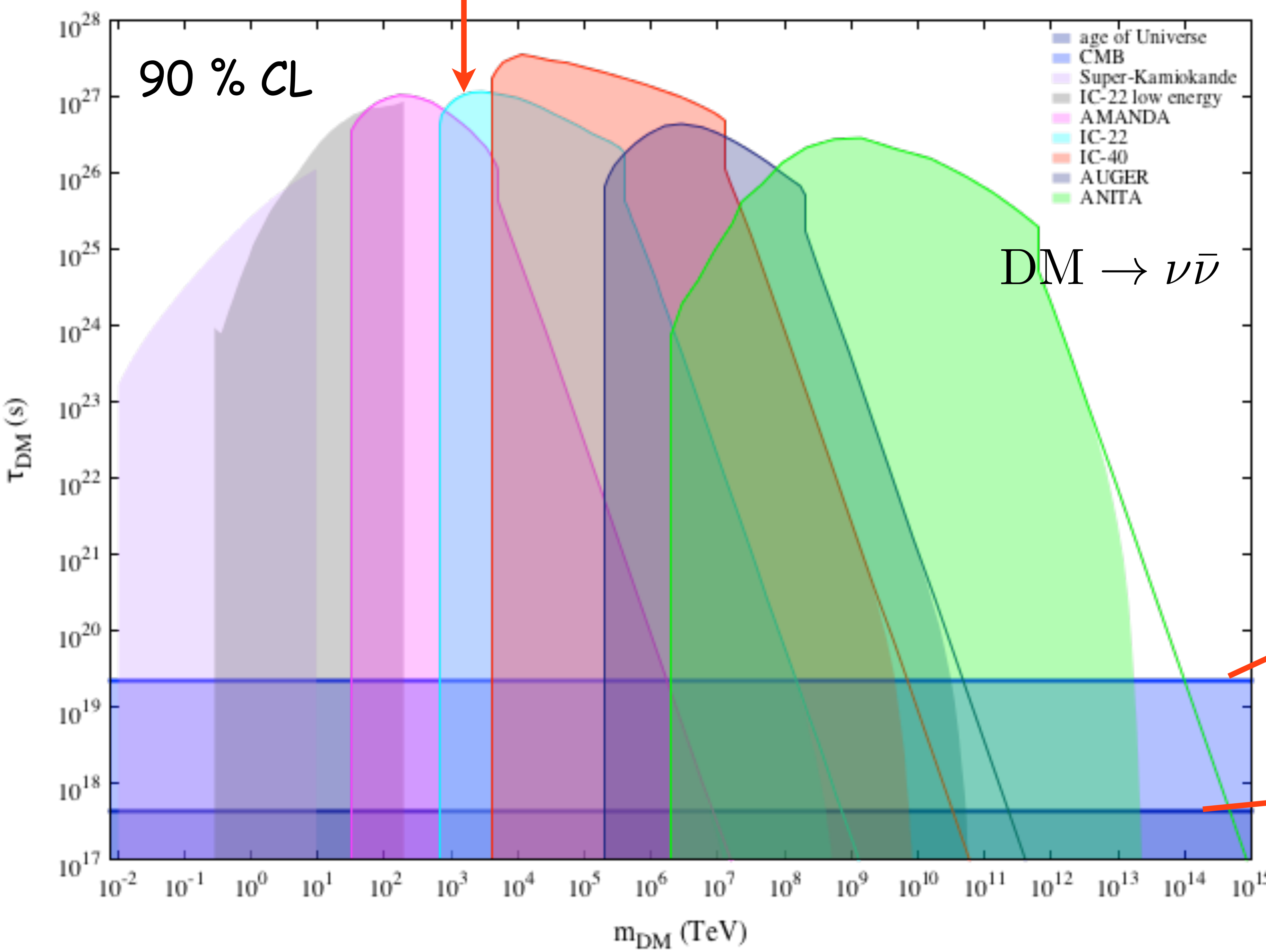
Limits on lifetime from neutrino experiments

before recent IceCube data

✓ Lifetime: stable (∞) or

this talk

A.E., Alejandro Ibarra and Orlando L. G. Peres
JCAP (2012) [arXiv: 1205.5281]



Using AMANDA, IceCube, Auger and ANITA data

$\tau_{DM} > 2.2 \times 10^{19}$ s (CMB)

$\tau_{DM} > 4.3 \times 10^{17}$ s (age of Universe)

Limits on lifetime from neutrino experiments

before recent IceCube data

✓ Lifetime: stable (∞) or

this talk

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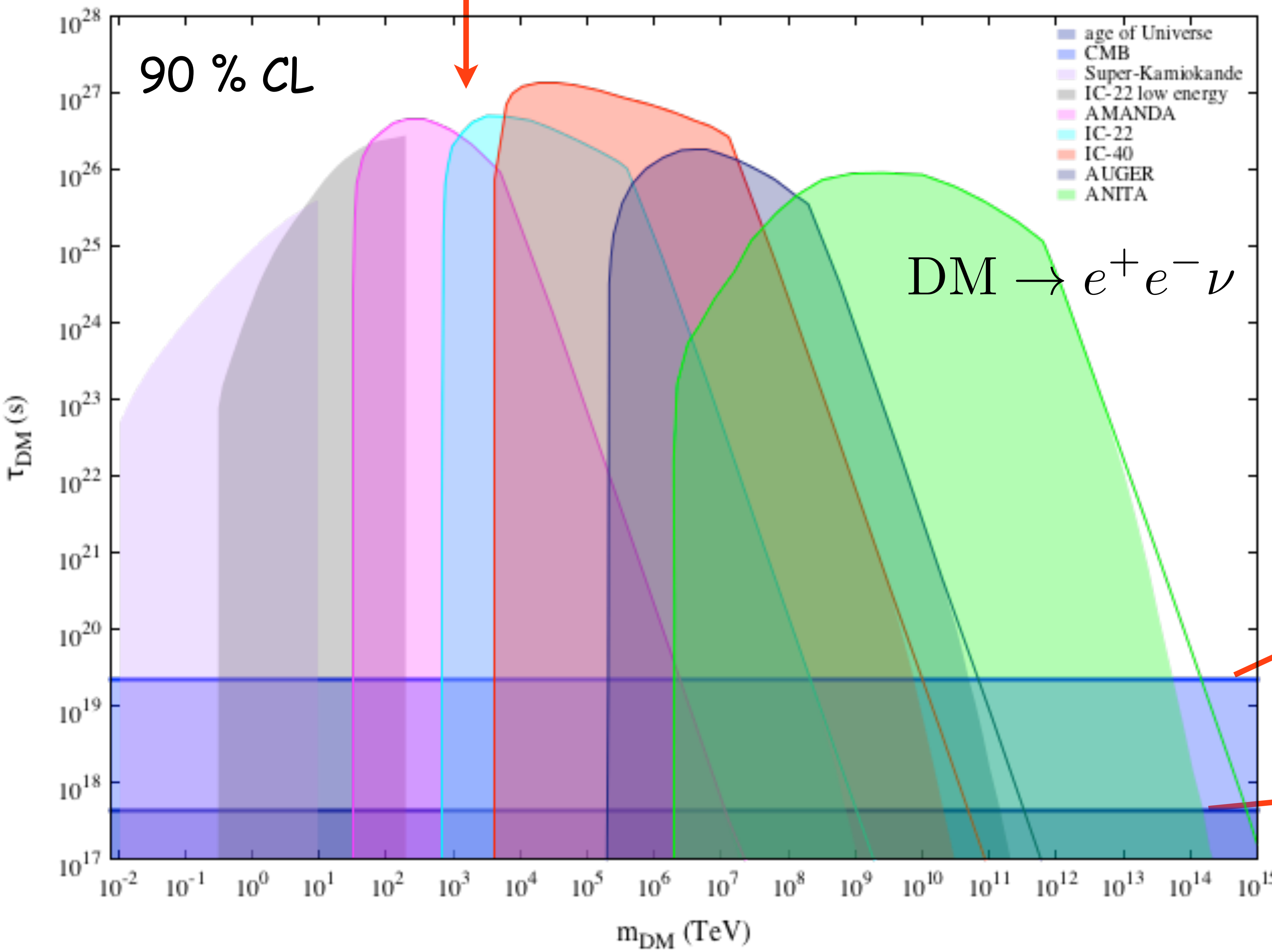

Using AMANDA, IceCube, Auger and ANITA data

$\tau_{DM} > 2.2 \times 10^{19}$ s (CMB)

$\tau_{DM} > 4.3 \times 10^{17}$ s (age of Universe)

Interpreting the IceCube events by decaying dark matter

Two main diagnostics:

✓ Energy distribution

✓ Angular distribution

Energy distribution of neutrinos from decaying DM

✓ Galactic contribution:

NFW $\rho_{\rm halo}(r) \simeq \frac{\rho_h}{r/r_c(1+r/r_c)^2}$

$$\frac{{\rm d}J_{\rm h}}{{\rm d}E_\nu}(l,b) = \frac{1}{4\pi\, m_{\rm DM}\, \tau_{\rm DM}} \frac{{\rm d}N_\nu}{{\rm d}E_\nu} \int_0^\infty {\rm d}s\ \rho_{\rm h}[r(s,l,b)]$$

$r(s,l,b) = \sqrt{s^2 + R_\odot^2 - 2sR_\odot \cos b \cos l}$

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M. Cirelli et. al., JCAP (2011)

energy spectrum of neutrinos
at production point
(including the EW corrections)

quarks

neutrinos,
charged leptons

$$\frac{\mathrm{d}N_\nu}{\mathrm{d}E_\nu} = (1-b_{\rm H}) \left.\frac{\mathrm{d}N_\nu}{\mathrm{d}E_\nu}\right|_{\rm S} + b_{\rm H} \left.\frac{\mathrm{d}N_\nu}{\mathrm{d}E_\nu}\right|_{\rm H}$$

Energy distribution of neutrinos from decaying DM

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at the Earth

$$\begin{pmatrix} J_e \\ J_\mu \\ J_\tau \end{pmatrix} = \begin{pmatrix} P_{ee} & P_{e\mu} & P_{e\tau} \\ & P_{\mu\mu} & P_{\mu\tau} \\ & & P_{\tau\tau} \end{pmatrix} \begin{pmatrix} I_e \\ I_\mu \\ I_\tau \end{pmatrix}$$

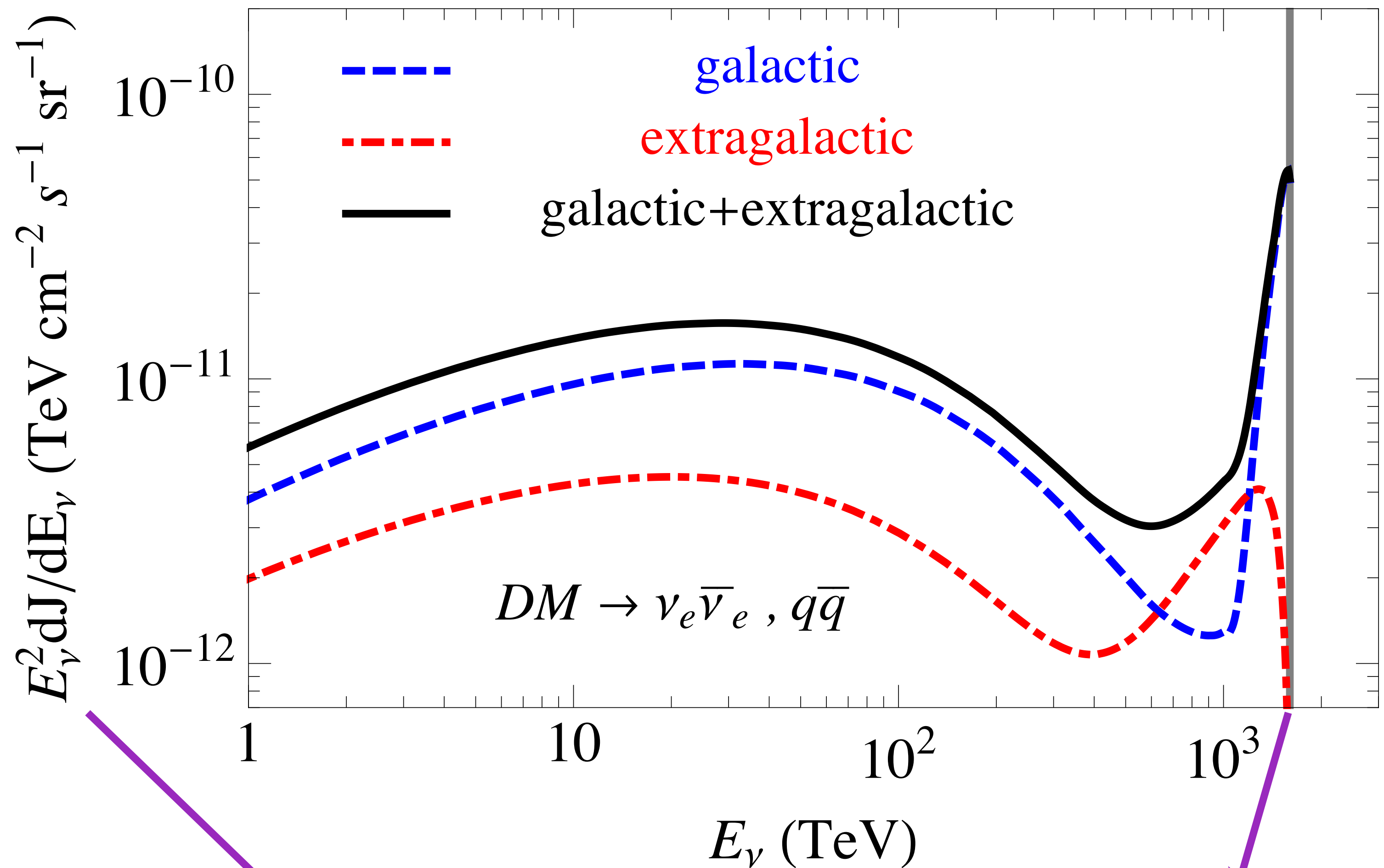
production point

decoherent oscillation

Flux of neutrinos from decaying DM

✓ an example:

A. E., Pasquale D. Serpico,
JCAP (2013) [arXiv:1308.1105]



$(\nu_e + \nu_\mu + \nu_\tau)/3$ $\quad$ $m_{DM}/2 = 1.6$ PeV

$b_H = 0.12$ and $\tau_{DM} = 2 \times 10^{27}$ s

Flux of neutrinos from decaying DM

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intriguing features:

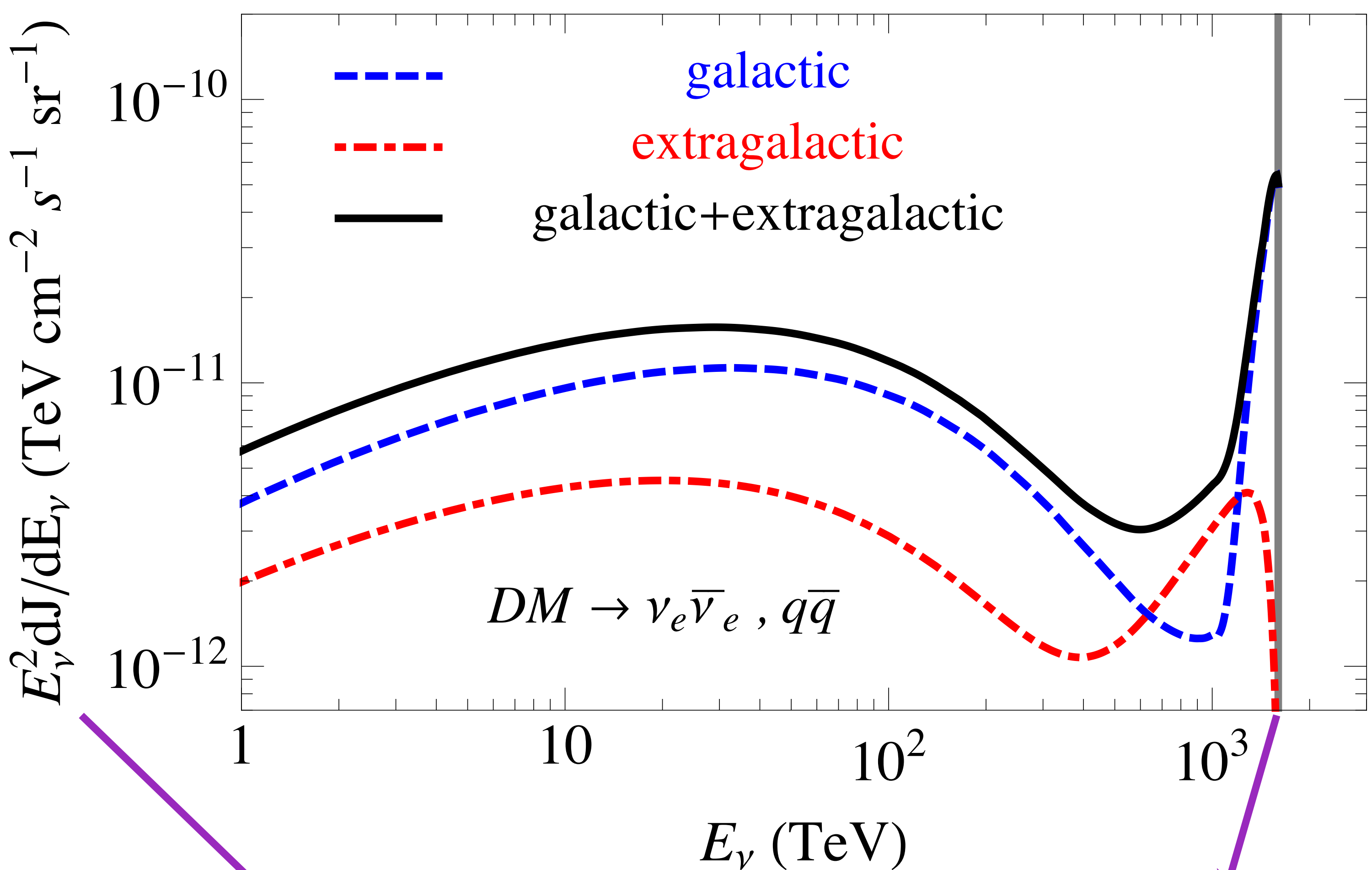
a cut-off at $m_{DM}/2$

a peak in ~ PeV

a dip in ~ (0.4-1) PeV

populated spectrum in < 0.4 PeV

due to soft channel and EW cascades



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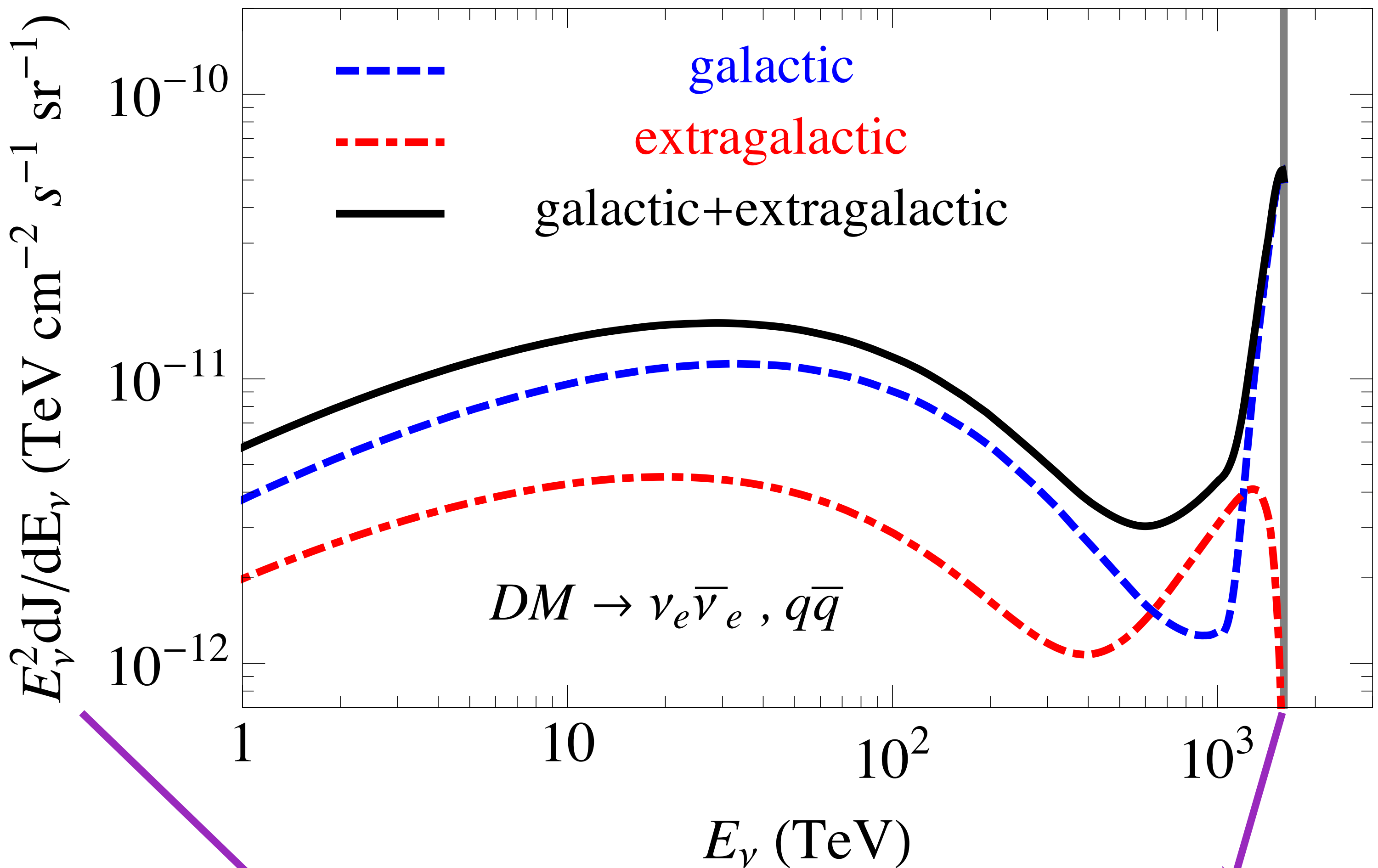
b_H controls the peak height at ~ PeV

τ_{DM} controls the low energy population

A. E., Pasquale D. Serpico, JCAP (2013) [arXiv:1308.1105]



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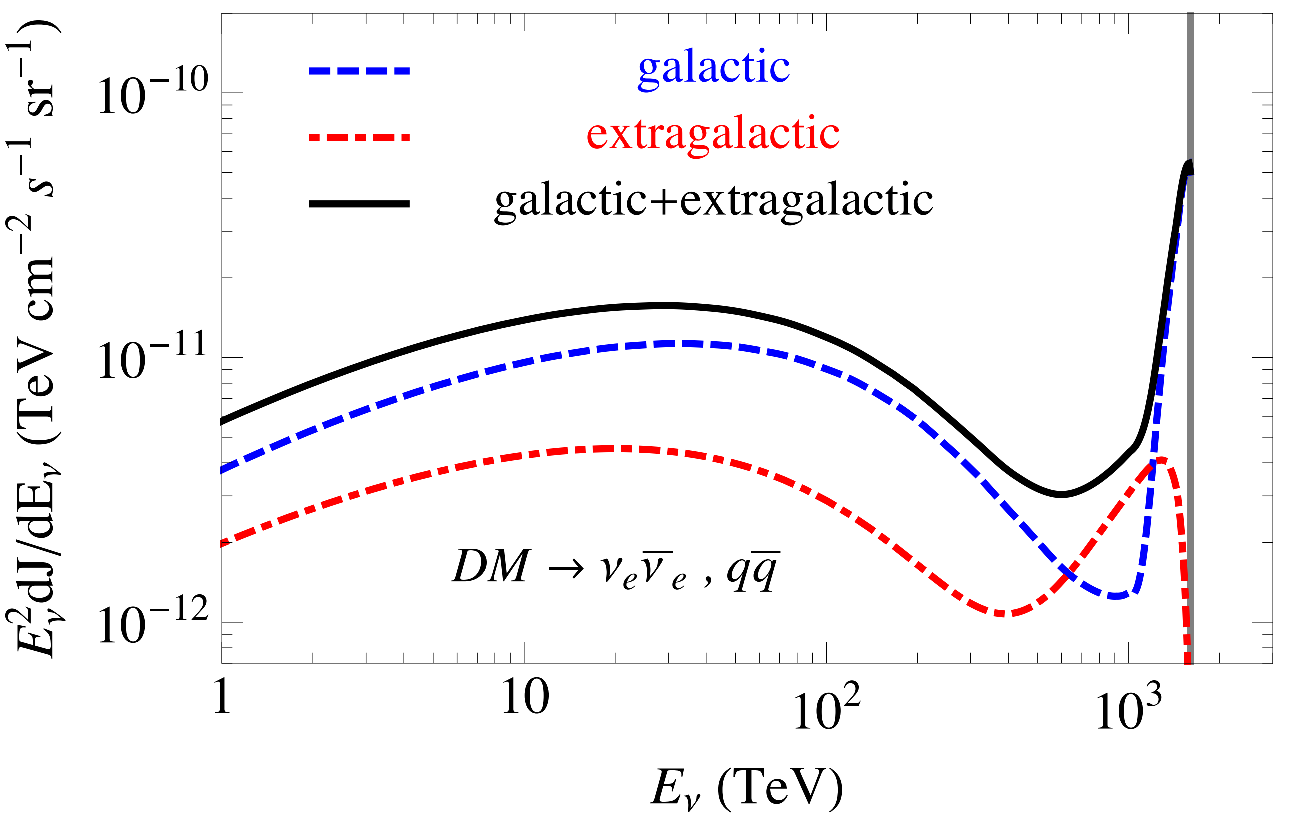
Flux of neutrinos from decaying DM

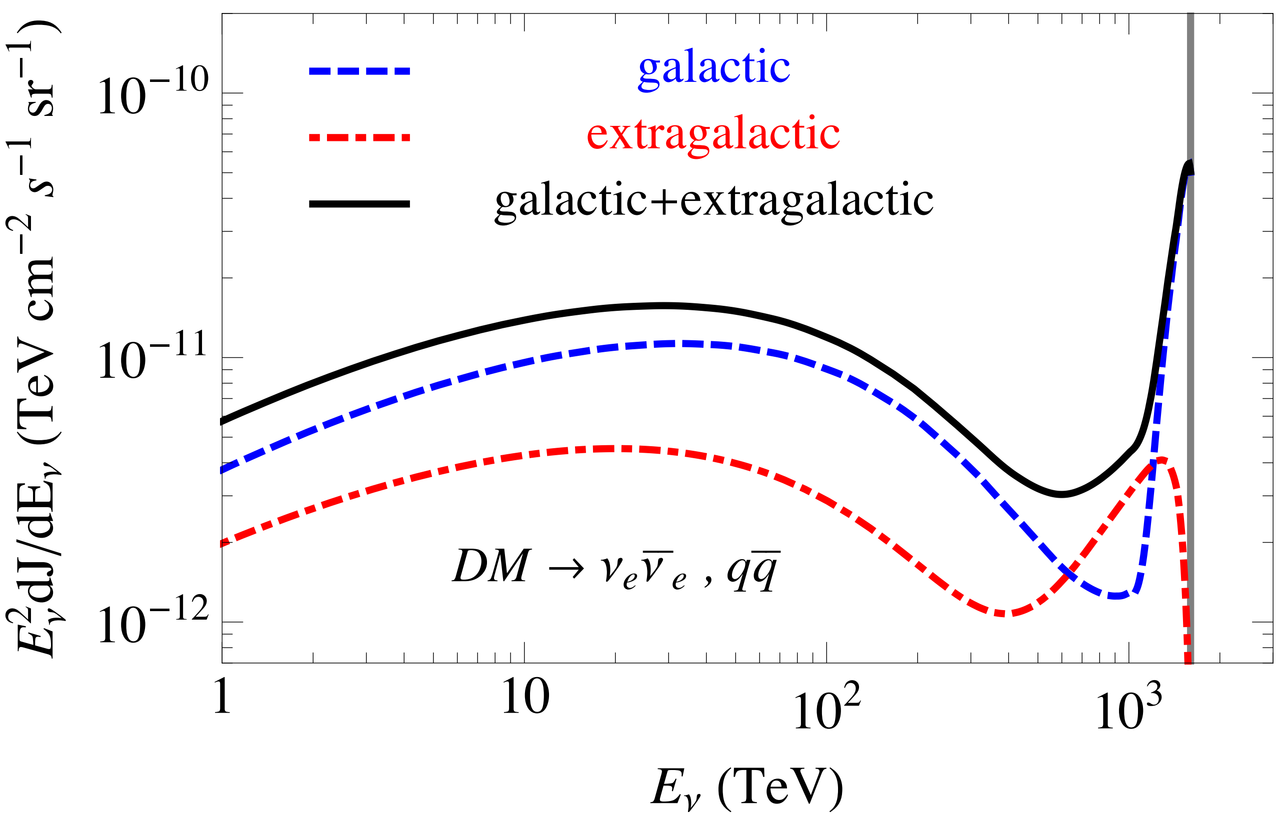
✓ fine-tuned decay channels ?

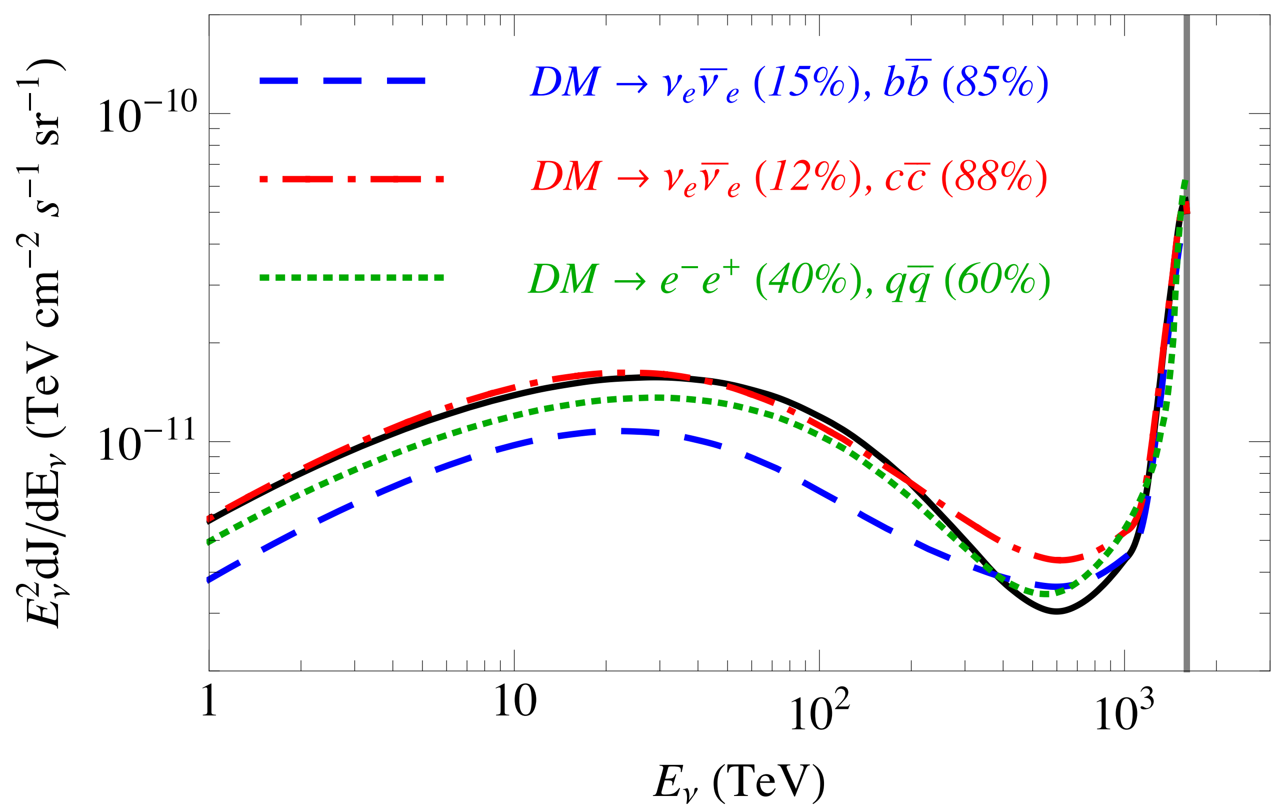


Flux of neutrinos from decaying DM

✓ fine-tuned decay channels ?

τ_{DM} = (1-3) x 10^{27} s





the intriguing features are generic

Flux of neutrinos from decaying DM

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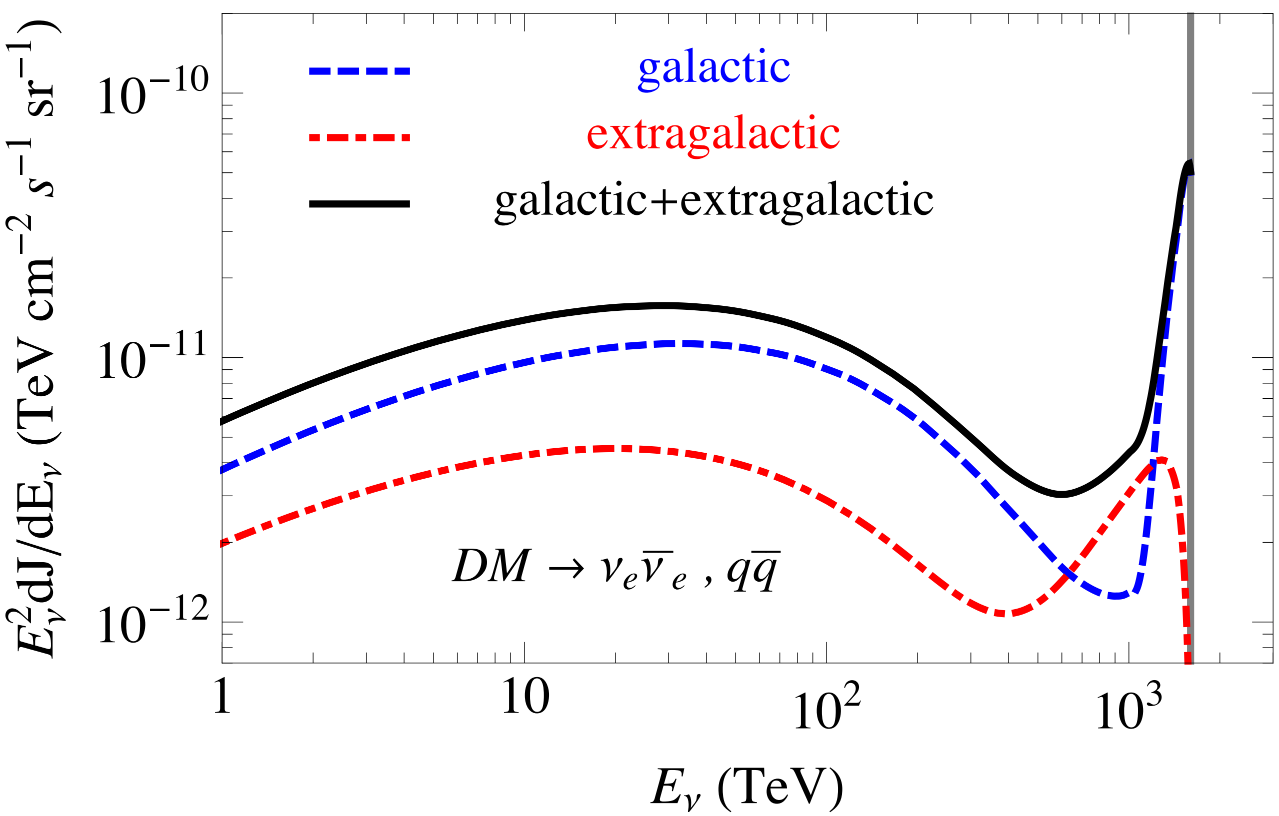
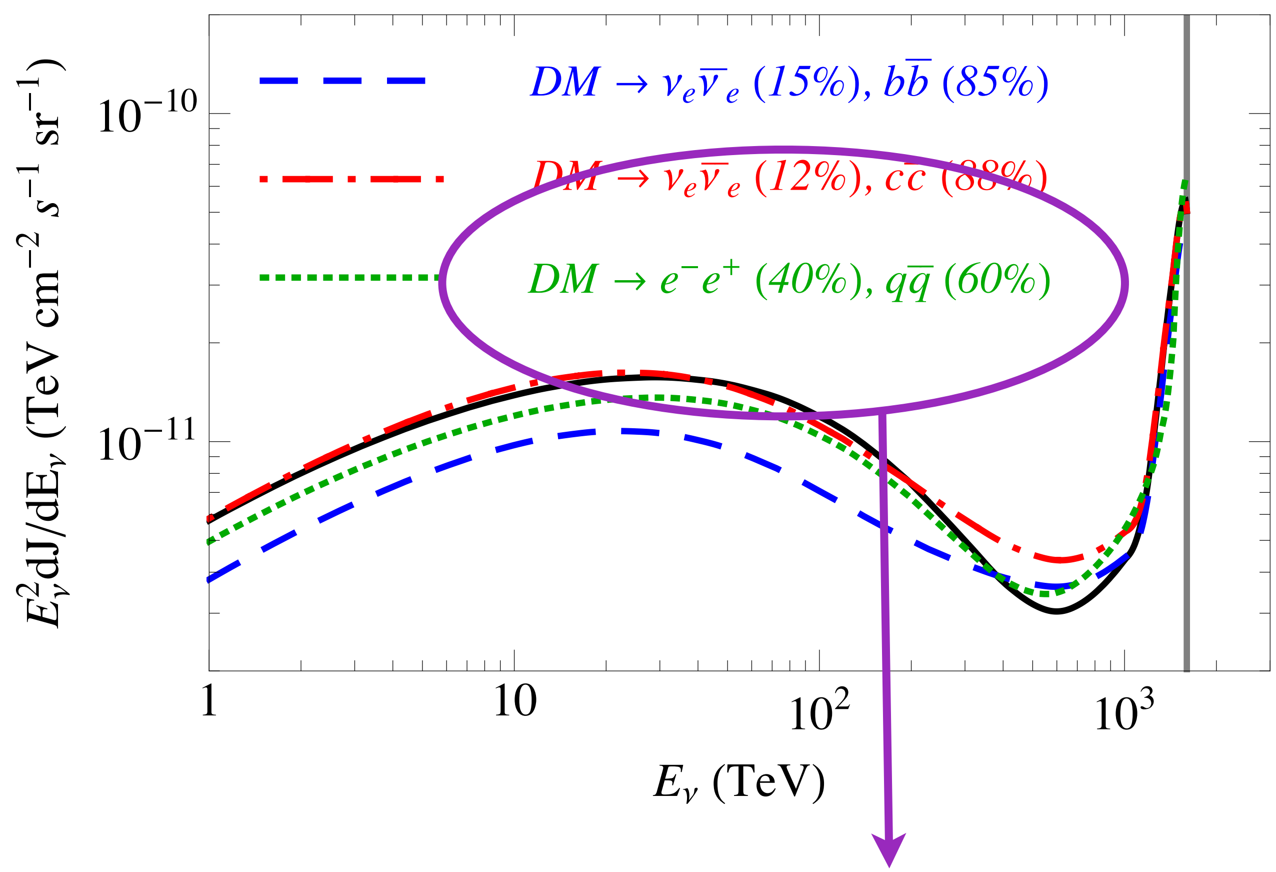




the crucial role of EW cascades

the intriguing features are generic

Confronting with energy distribution of IceCube data

two years data set

✓ branching ratio b_H gives the two PeV events

✓ soft channel and lifetime τ_{DM} gives the upturn in low energy

the value of τ_{DM} is compatible with the bounds derived from neutrinos and gamma rays

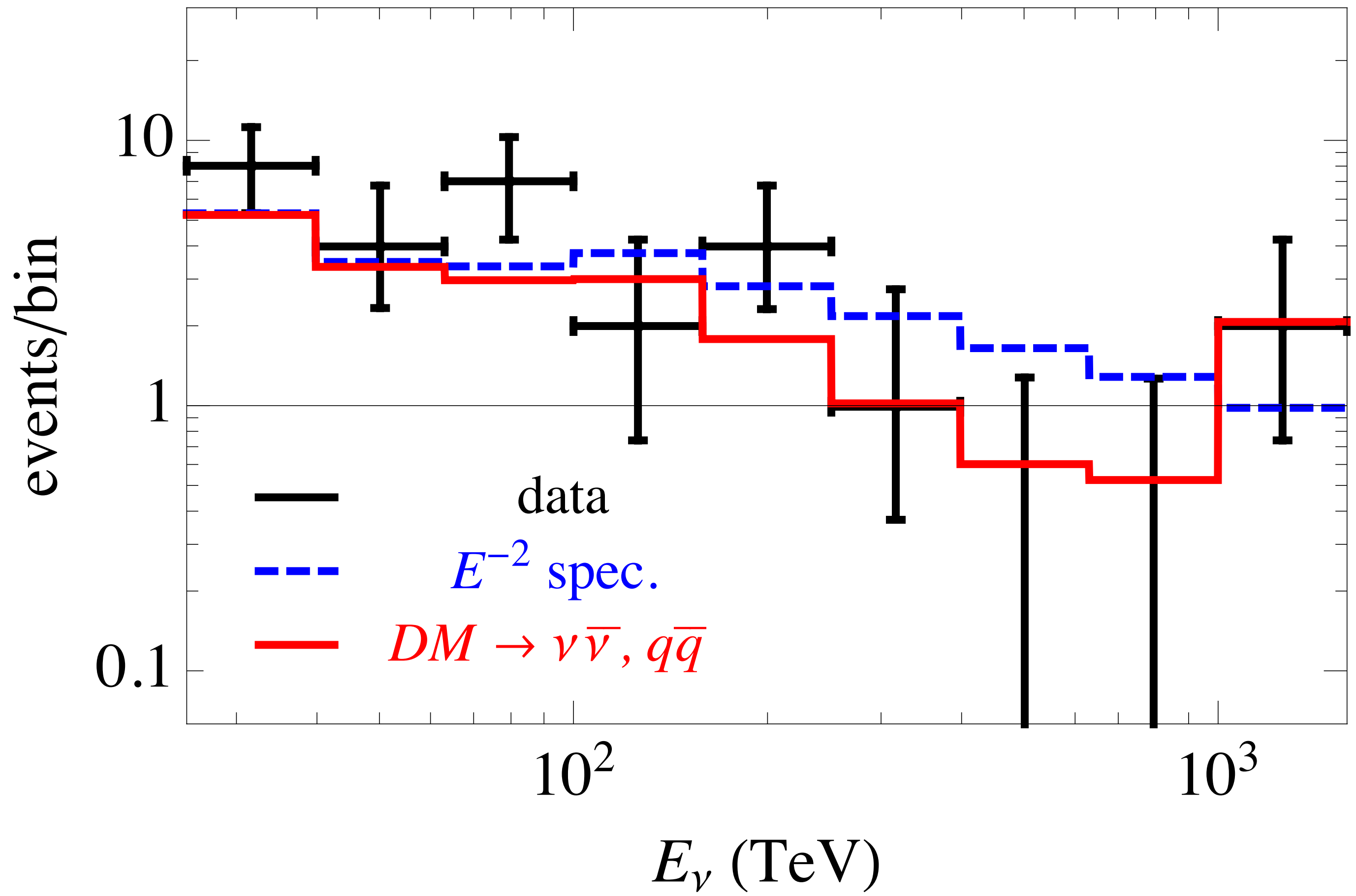
✓ natural explanation for the lack of events > PeV

the value of m_{DM} can be changed within the current uncertainty of the highest energy events

✓ the low energy bins contain large bkg. contribution

the important discriminators of DM vs astrophysical model are high energy bins, where clearly data shows preference to DM model

$b_H = 0.12$ and $\tau_{DM} = 2 \times 10^{27}$ s



✓ different decay channels lead to qualitatively same result

Confronting with energy distribution of IceCube data

three years data set

SM sector ⟷ Dark sector

portal type:

$$\mathcal{L}_{\rm protal} = \frac{\mathcal{O}_{\rm SM}\mathcal{O}_{\rm DM}}{\Lambda^{d-4}}$$

Confronting with energy distribution of IceCube data

three years data set

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"neutrino" portal: $$\mathcal{O}_{\rm SM} \to HL$$

A. Falkowski, J. Juknevich and J. Shelton
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✓ d = 4 : $$\mathcal{O}_{\rm DM} \to N$$

heavy sterile neutrino, DM candidate

T. Higaki, R. Kitano and R. Sato, JHEP (2014)
arXiv:1405.0013

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UV completion: $SU(3)_c \times SU(2)_L \times U(1)_Y \times U(1)_{B-L}$

$m_\phi \sim 10^{13}\ {\rm GeV}$

"Higgs" field ϕ_{B-L} plays the role of inflaton

$T_R \sim 10^7\ {\rm GeV}$

Confronting with energy distribution of IceCube data

three years data set

Leptogenesis: $\phi \to N_2 N_2$ $\quad M_2 \sim 10^{12}\ \text{GeV} \longrightarrow \frac{n_B}{s} \sim 10^{-10}$

DM abundance: $$\Omega_{N_1} \simeq 0.2 \left(\frac{M_1}{4\ \text{PeV}}\right)^3 \left(\frac{T_R}{3 \times 10^7\ \text{GeV}}\right)^{-1}$$

DM lifetime: $$\tau_{N_1} \simeq 8 \times 10^{28}\ \text{s} \left(\frac{M_1}{1\ \text{PeV}}\right)^{-1} \left(\frac{10^{-29}}{|y_N|^2}\right)$$

DM decay channels:

$$\text{Br}(\ell^{\pm} W^{\mp}) = 2\text{Br}(\nu_\ell Z) = 2\text{Br}(\nu_\ell h) = |U_{\ell 1}|^2$$ NH

$$\text{Br}(\ell^{\pm} W^{\mp}) = 2\text{Br}(\nu_\ell Z) = 2\text{Br}(\nu_\ell h) = |U_{\ell 3}|^2$$ IH

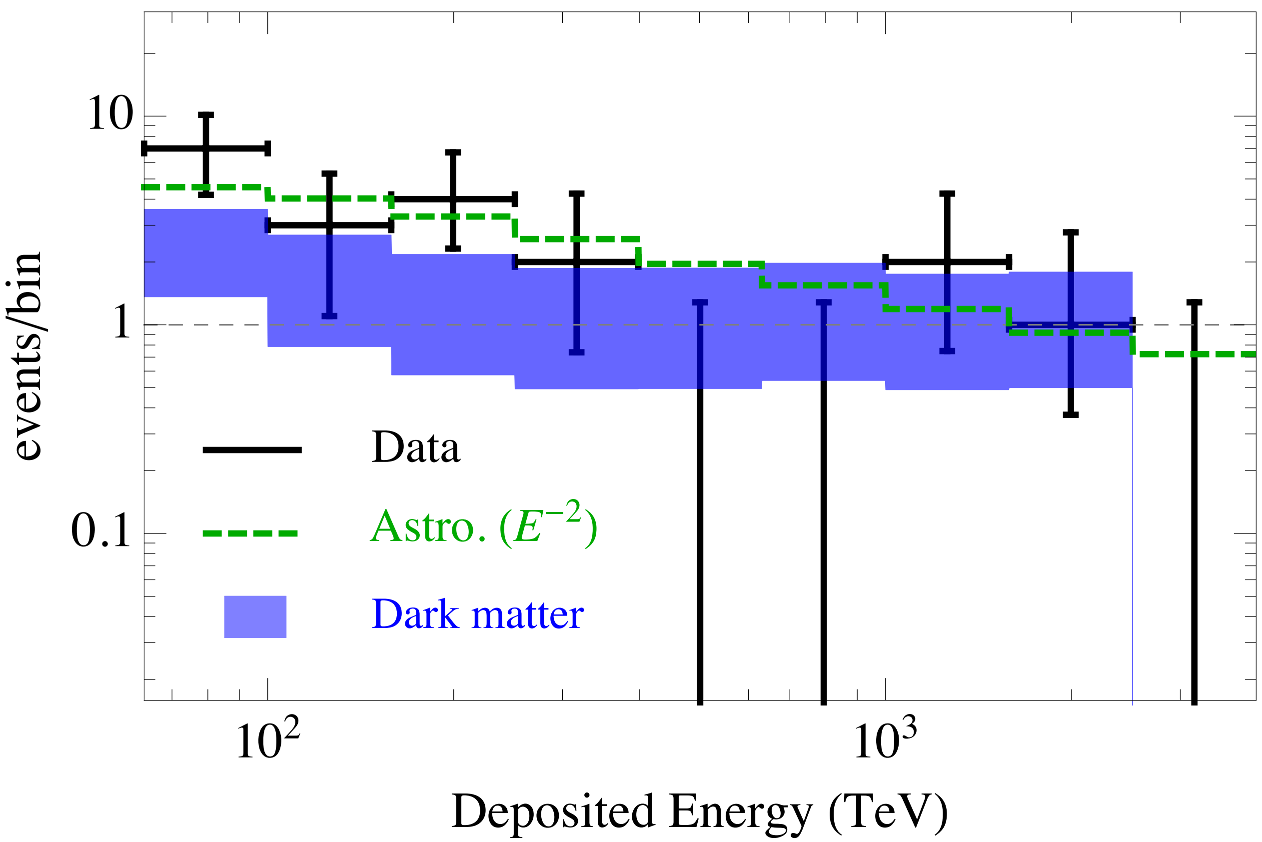
Confronting with energy distribution of IceCube data

three years data set

IH

τ_{DM} = 1.1 x 10^{28} s

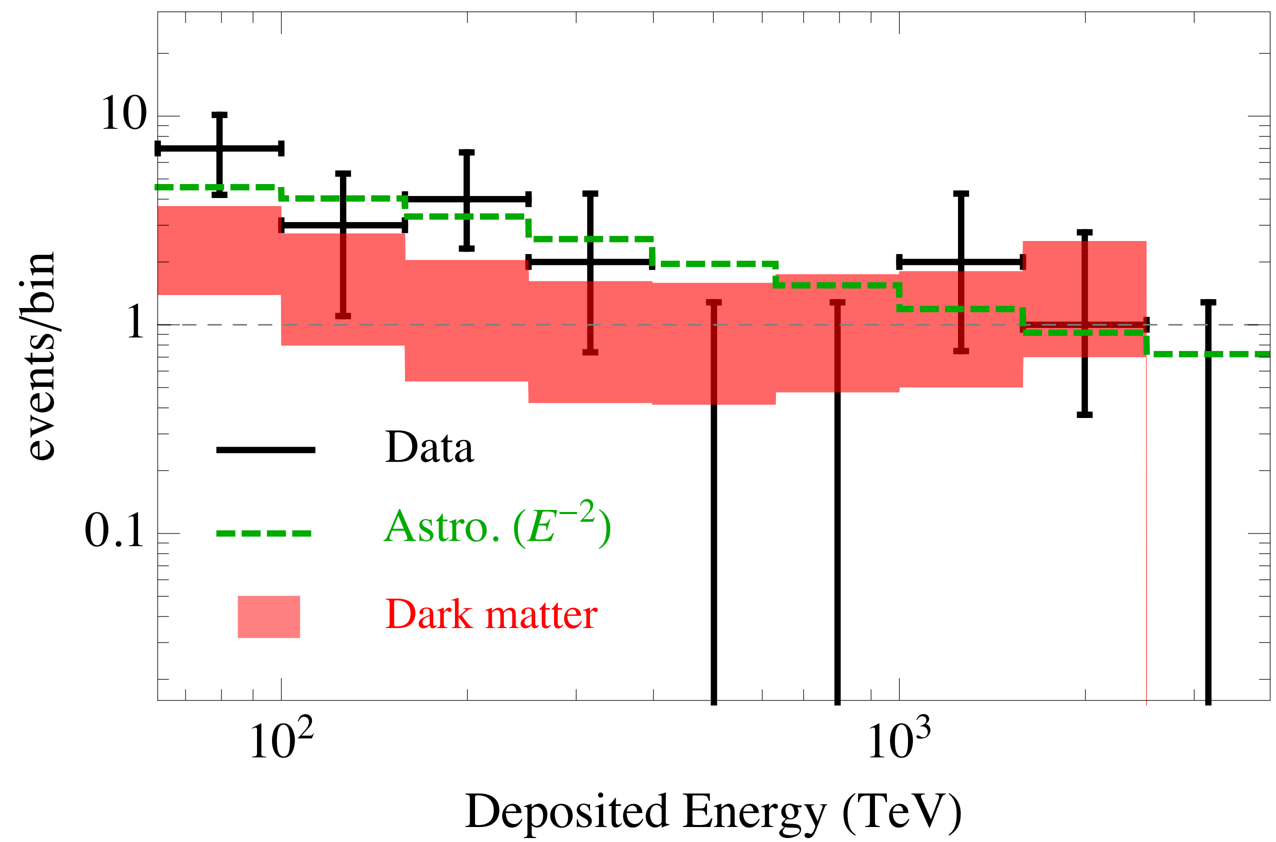
shaded: ±1σ



NH

τ_{DM} = 7.3 x 10^{27} s

shaded: ±1σ



m_{DM} = 4 PeV

Confronting with energy distribution of IceCube data

three years data set

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A. Falkowski, J. Juknevich and J. Shelton
arXiv:0908.1790 [hep-ph].

✓ d =4: $\mathcal{O}_{\rm DM} \to N$

production mechanism:

$m_\phi \gg m_N$ inflaton decay

$m_\phi \ll m_N$ freeze-in

$g\phi NN \; , \; g \simeq 10^{-6}$

Confronting with energy distribution of IceCube data

three years data set

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✓ d = 5 : $\mathcal{O}_{\rm DM} \to \chi\phi$ singlet fermion and scalar (Asymmetric DM)

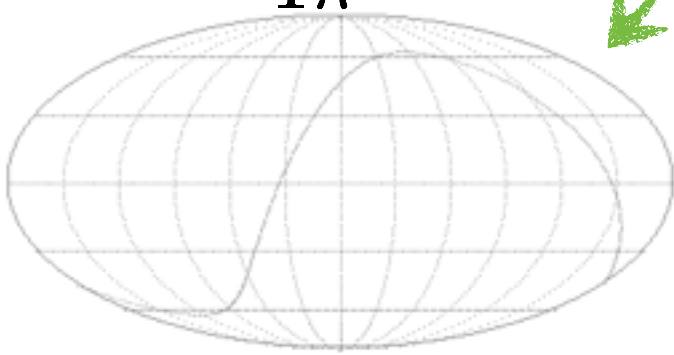
✓ d = 6 : other portals

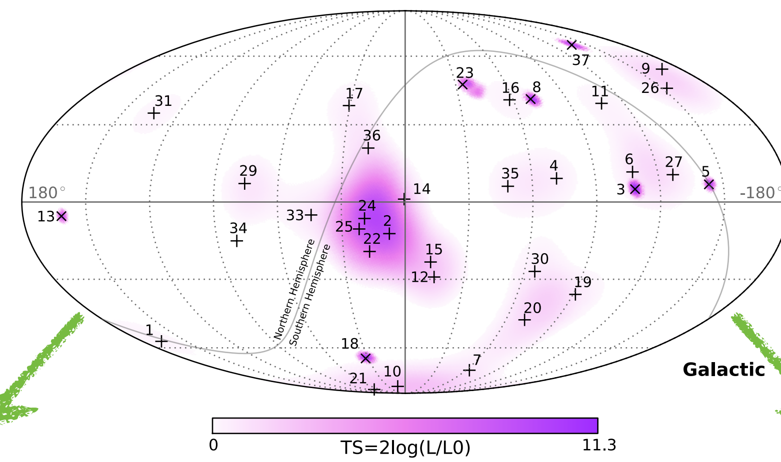
✓ For d > 4 there are more freedom in branching ratios. We have shown that for the most constrained model (d=4) a good fit to the data can be obtained. Obviously better fits can be achieved for d > 4.

Angular distribution of neutrinos from decaying DM

✓ We would compare

$$p^{\rm iso} = \frac{1}{4\pi}$$





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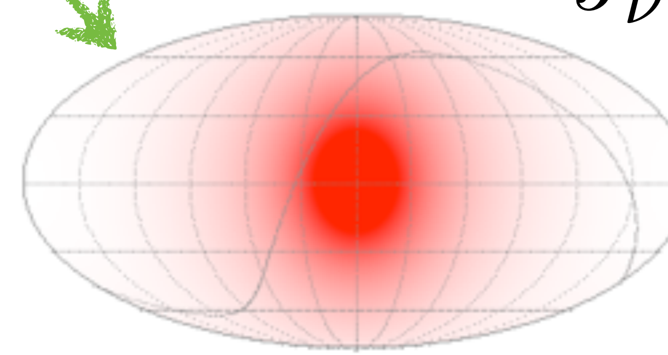

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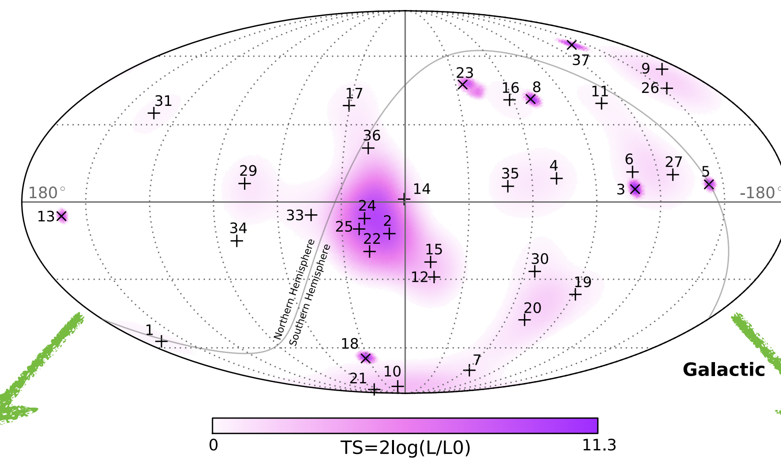

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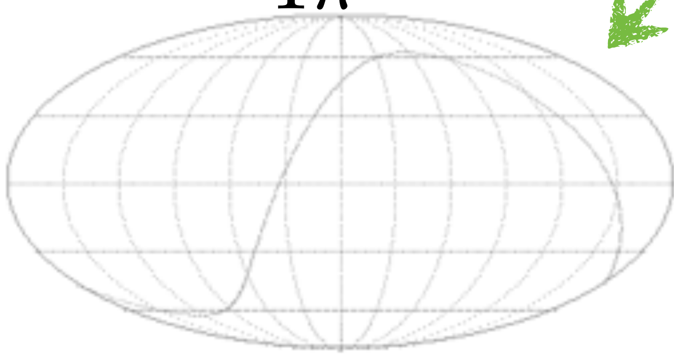

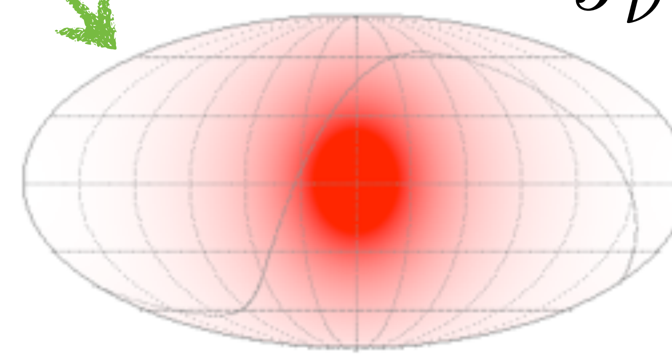

PDF of data

$$p_i(b,l) = \frac{1}{2\pi\sigma_i^2}\exp\left[-\frac{|\vec{x}-\vec{x}_i|^2}{2\sigma_i^2}\right]$$

"flat sky" approximation

PDF of isotropic dis.

$$p^{\rm iso} = \frac{1}{4\pi}$$

PDF of DM

$$p^{\rm DM}(b,l) = \frac{1}{J_\nu}\frac{{\rm d}^2 J_\nu}{{\rm d}b\,{\rm d}l} = \frac{\int_0^\infty \rho[r(s,b,l)]{\rm d}s + \Omega_{\rm DM}\rho_c\beta}{4\pi(\eta+\Omega_{\rm DM}\rho_c\beta)}$$

Angular distribution of neutrinos from decaying DM

✓ Likelihood analysis

Number of signal events

Test Statistics

$$\mathrm{TS}_{\mathrm{like}} = 2\sum_{i=1}^{N}\left(\ln f_i - \ln p_i^{\mathrm{iso}}\right) = 2\ln\left(\prod_{i=1}^{N} f_i\right) - 2N\ln\left(\frac{1}{4\pi}\right)$$

$$f_i = \int p_i(b,l)\, p^{\mathrm{DM}}(b,l)\cos(b)\,\mathrm{d}b\,\mathrm{d}l = \frac{1}{2\pi\sigma_i^2}\int e^{-\frac{|\vec{x}_i-\vec{x}|^2}{2\sigma_i^2}}\, p^{\mathrm{DM}}(b,l)\cos(b)\,\mathrm{d}b\,\mathrm{d}l$$

N = 35 ? too optimistic!

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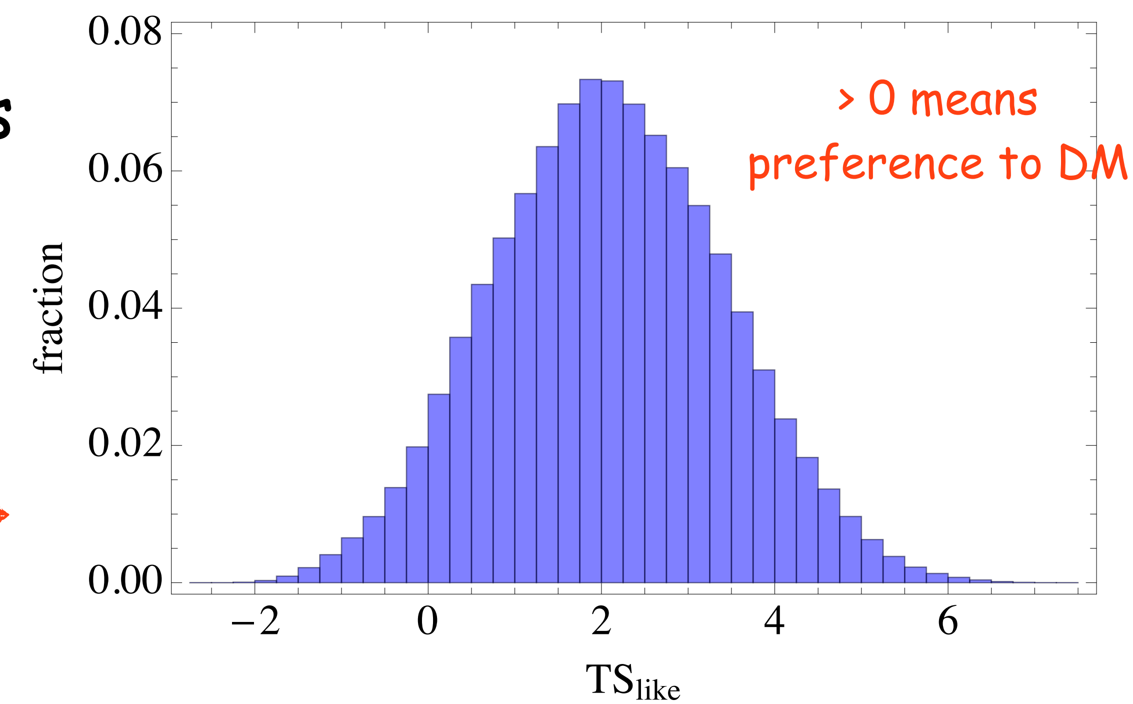
N = 35 ? too optimistic!

✓ let's assume N_b = 15 and all the events with E > 150 TeV are signal events

→ $\binom{26}{15}$ ways of selecting the bkg events among the low energy events

Distribution of $\mathrm{TS}_{\mathrm{like}}$ for all these realizations (mean value = 2.1) →



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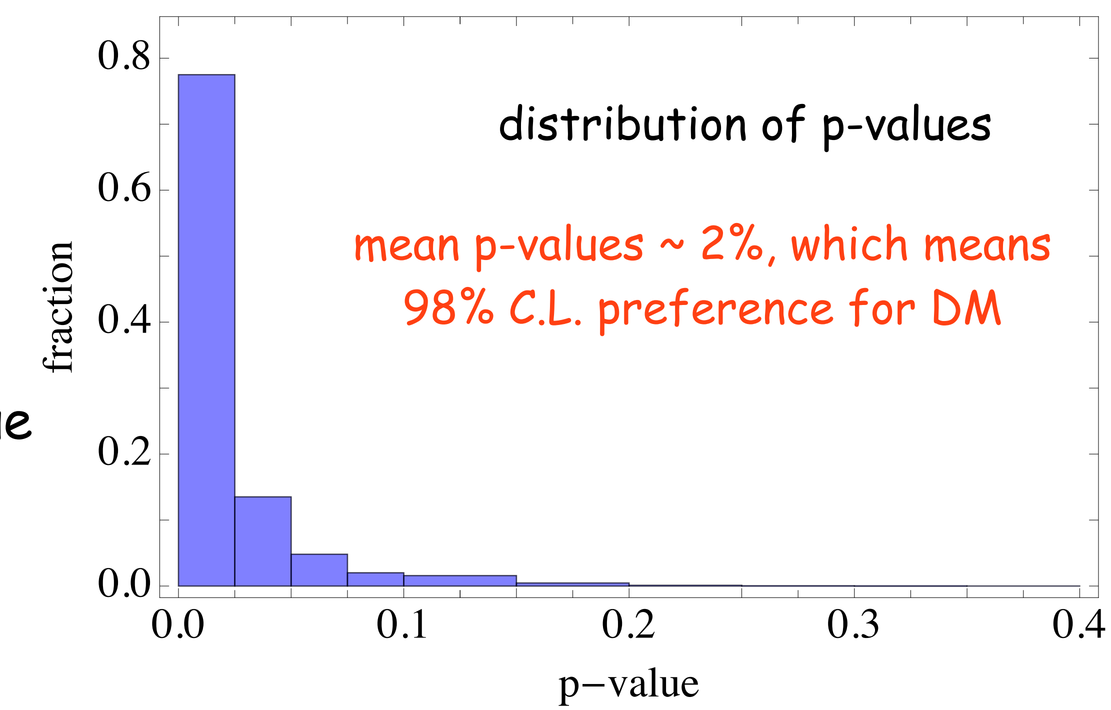
Quantifying the preference

generating a sample (10^5) of isotropically distributed set of 20 events

↓

for each realization of bkg choosing, p-value is the fraction of generated events which have smaller $\mathrm{TS}_{\mathrm{like}}$ than the one computed by observed data

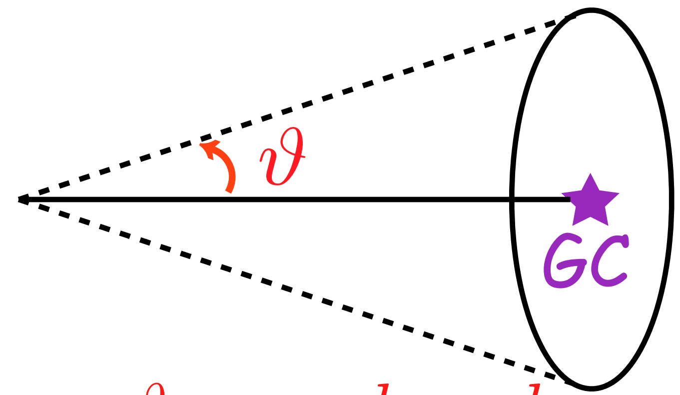
p-value →



Angular distribution of neutrinos from decaying DM

✓ **Kolmogorov-Smirnov test:** a powerful non-parametric test

The 2-dim KS test have some ambiguities



$$p^{\rm iso}(\vartheta) = \int_0^{2\pi} p^{\rm iso}(\vartheta,\varphi)\,{\rm d}\varphi = \int_0^{2\pi} \frac{1}{4\pi}\,{\rm d}\varphi = \frac{1}{2}$$

$$p^{\rm DM}(\vartheta) = \int_0^{2\pi} p^{\rm DM}(\vartheta,\varphi)\,{\rm d}\varphi = \frac{\int_0^\infty \rho[r(s,\vartheta)]{\rm d}s + \Omega_{\rm DM}\rho_c\beta}{2(\eta + \Omega_{\rm DM}\rho_c\beta)}$$

Angular distribution of neutrinos from decaying DM

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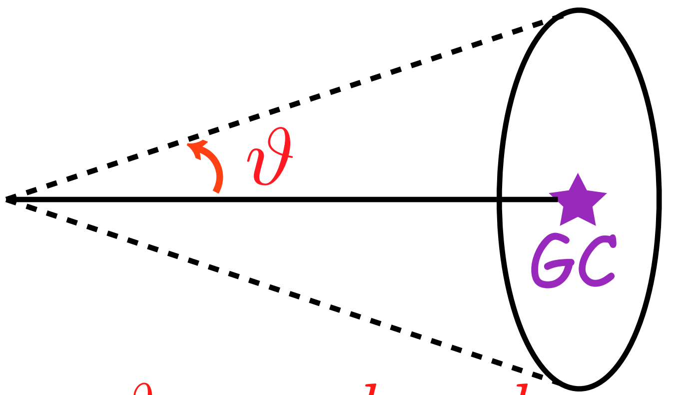

$\cos\vartheta = \cos b \cos l$

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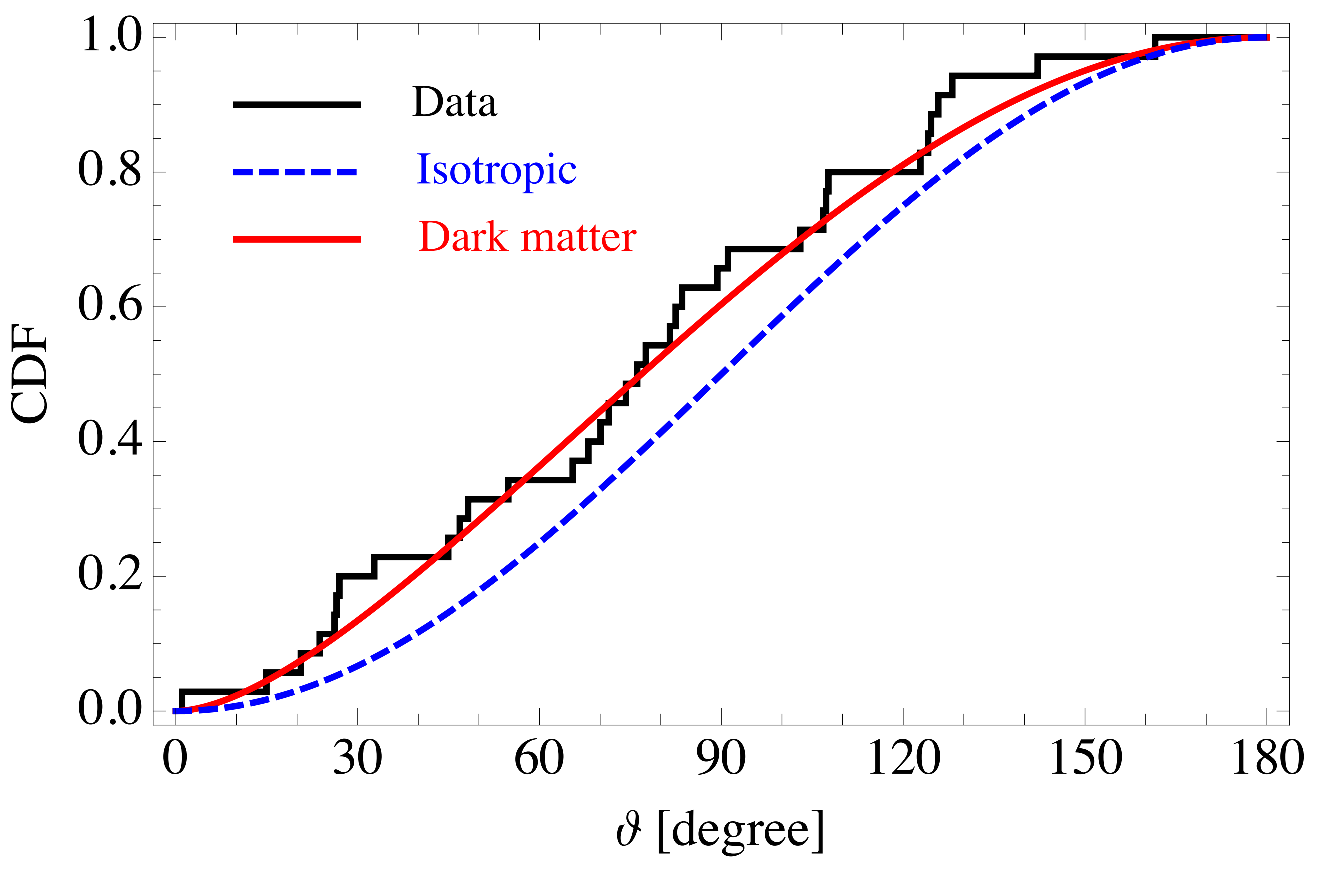

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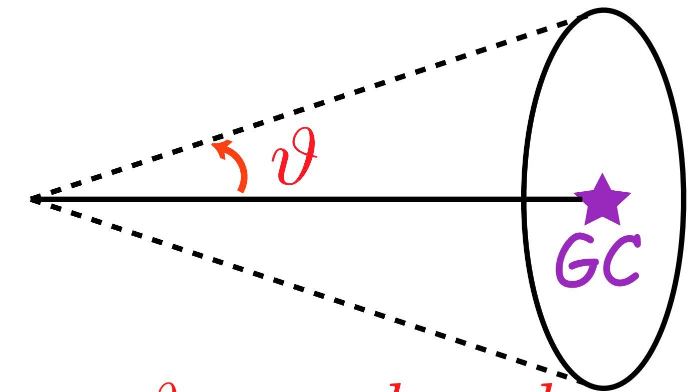

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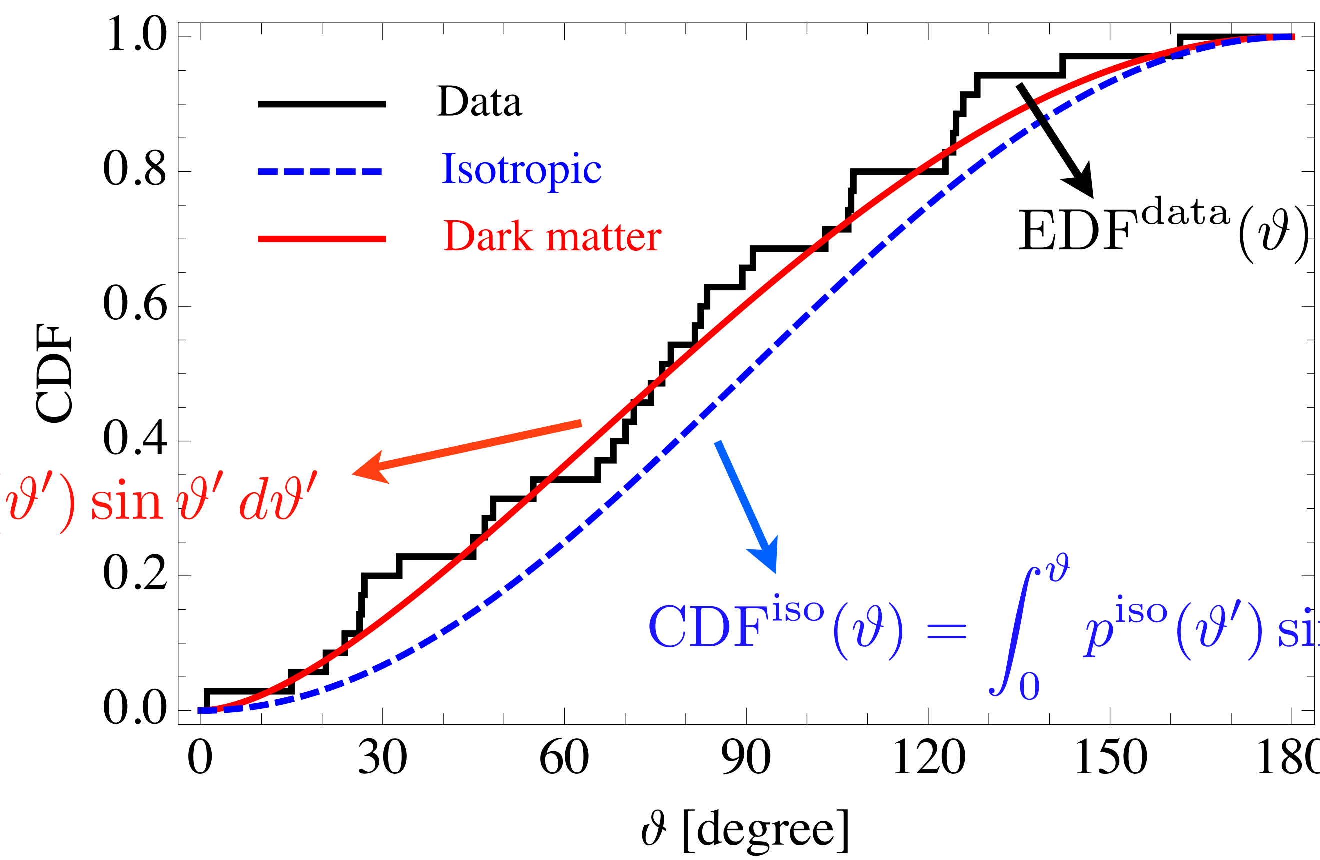
$$p^{\rm DM}(\vartheta) = \int_0^{2\pi} p^{\rm DM}(\vartheta,\varphi)\,{\rm d}\varphi = \frac{\int_0^\infty \rho[r(s,\vartheta)]{\rm d}s + \Omega_{\rm DM}\rho_c\beta}{2(\eta + \Omega_{\rm DM}\rho_c\beta)}$$



$$\mathrm{EDF}^{\rm data}(\vartheta) = \frac{1}{N}\sum_{i=1}^{N}\Theta(\vartheta - \vartheta_i)$$

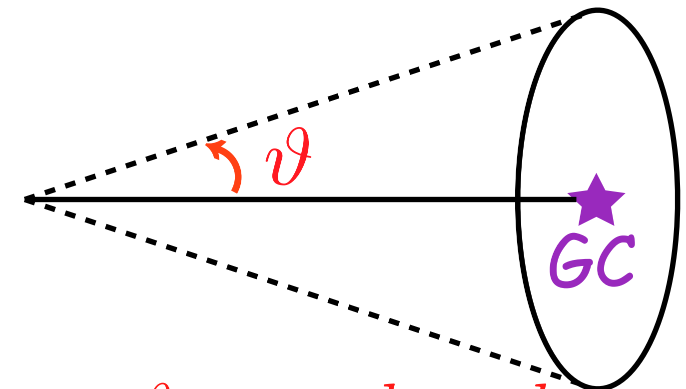
$$\mathrm{CDF}^{\rm DM}(\vartheta) = \int_0^\vartheta p^{\rm DM}(\vartheta')\sin\vartheta'\,d\vartheta'$$

$$\mathrm{CDF}^{\rm iso}(\vartheta) = \int_0^\vartheta p^{\rm iso}(\vartheta')\sin\vartheta'\,d\vartheta' = \frac{1-\cos\vartheta}{2}$$

Angular distribution of neutrinos from decaying DM

✓ **Kolmogorov-Smirnov test:** a powerful non-parametric test

The 2-dim KS test have some ambiguities



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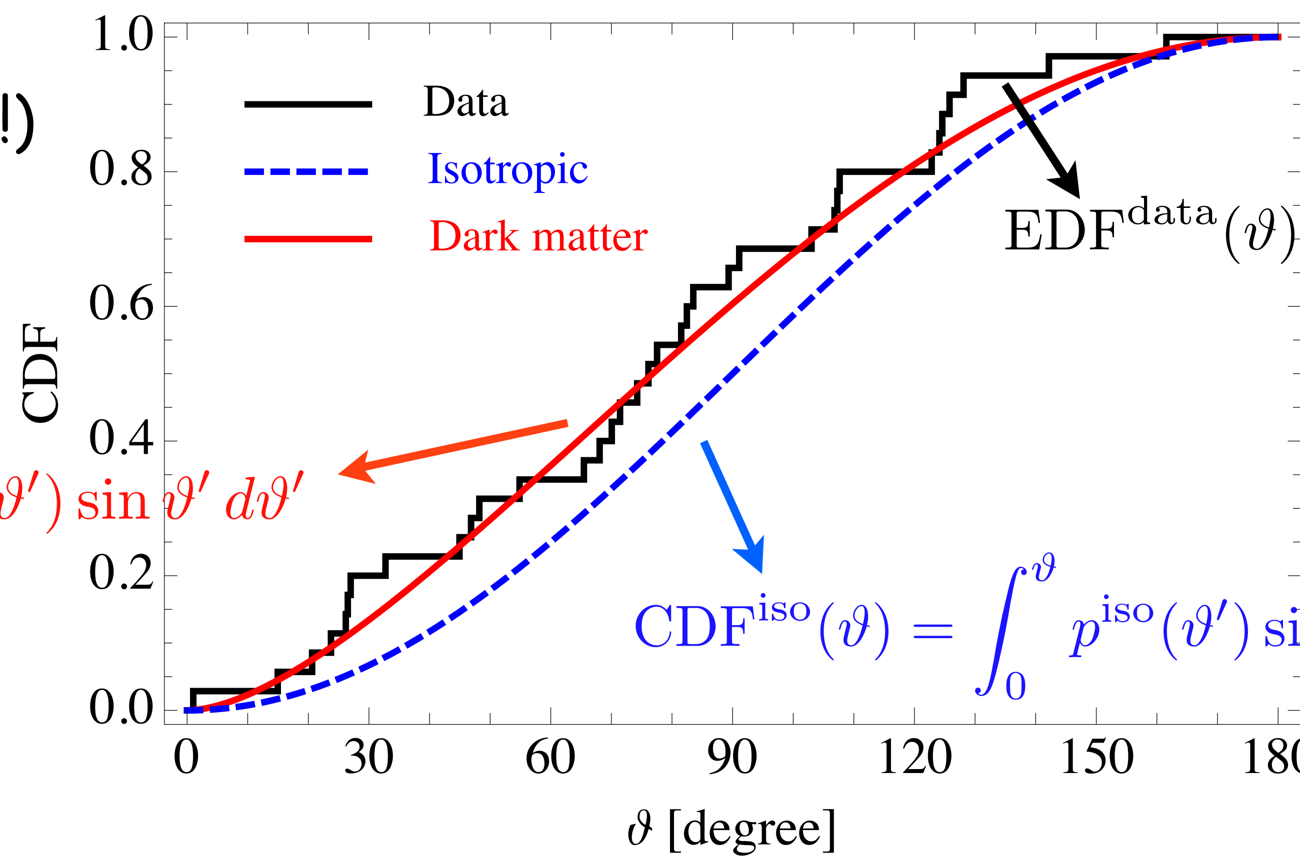
N = 35 (too optimistic!)



$$\mathrm{EDF}^{\rm data}(\vartheta) = \frac{1}{N}\sum_{i=1}^{N}\Theta(\vartheta-\vartheta_i)$$

$$\mathrm{CDF}^{\rm DM}(\vartheta) = \int_0^{\vartheta} p^{\rm DM}(\vartheta')\sin\vartheta'\,d\vartheta'$$

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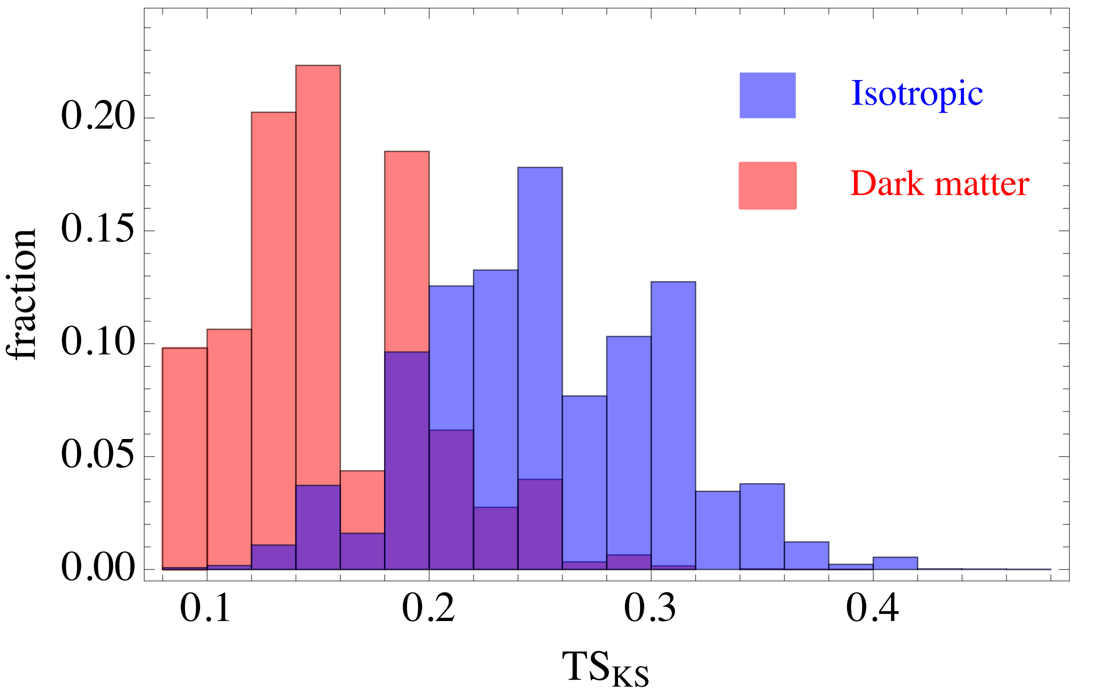
Angular distribution of neutrinos from decaying DM

✓ Kolmogorov-Smirnov test:

statistically larger TS for isotropic distribution



Test Statistics

$$\mathrm{TS_{KS}} = \max_{1\le i\le N}\left\{\mathrm{CDF^{DM}}(\vartheta_i) - \frac{i-1}{N}, \frac{i}{N} - \mathrm{CDF^{DM}}(\vartheta_i)\right\}$$

Angular distribution of neutrinos from decaying DM

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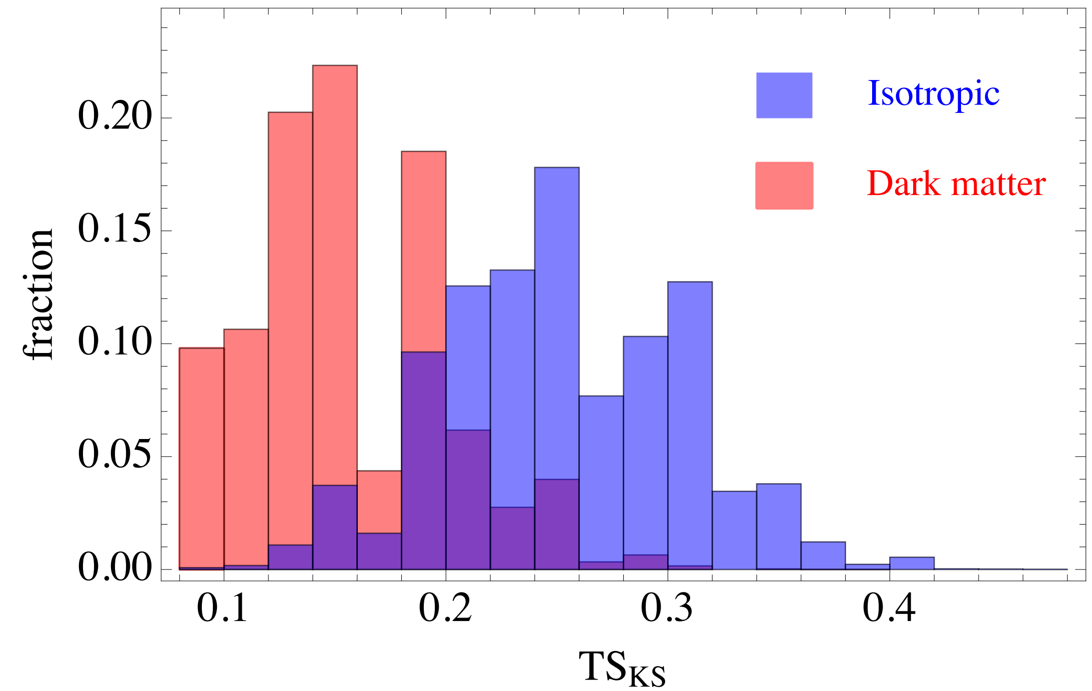


Test Statistics

$$TS_{KS} = \max_{1\leq i\leq N} \left\{ \mathrm{CDF}^{\mathrm{DM}}(\vartheta_i) - \frac{i-1}{N}, \frac{i}{N} - \mathrm{CDF}^{\mathrm{DM}}(\vartheta_i) \right\}$$

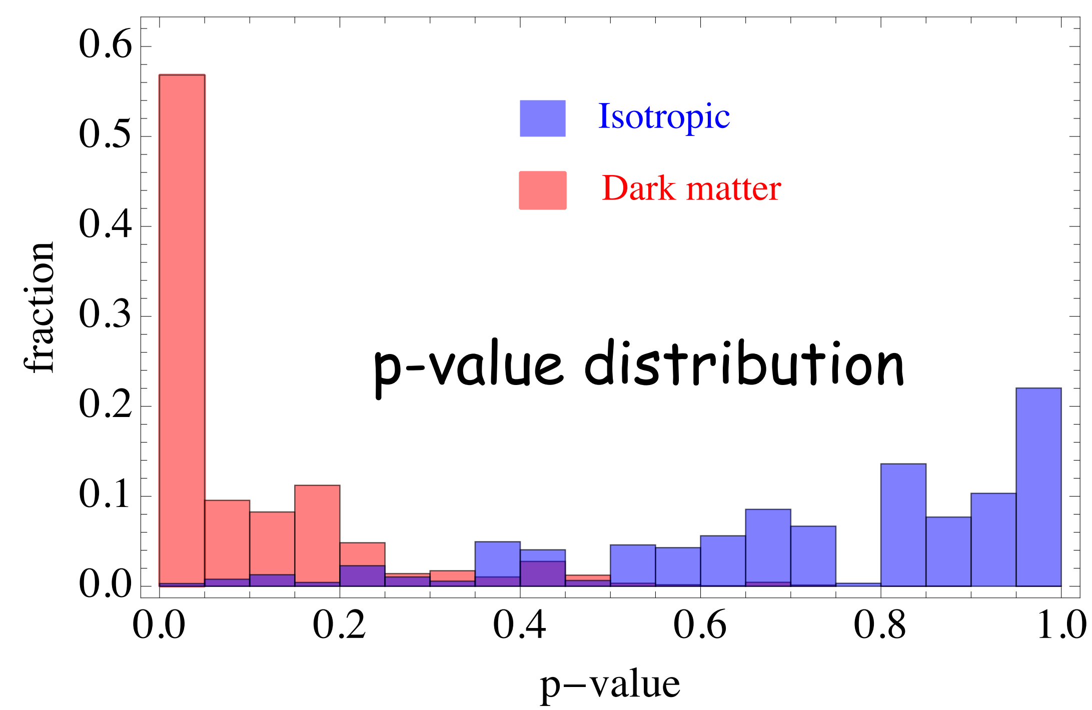
again, generating a sample (10^5) of isotropically distributed set of 20 events

↓

on the average, 10% of generated isotropic sample have smaller TS_{KS} than the values obtained for data vs DM dis.
for data vs isotropic dis. it is 73%

↓

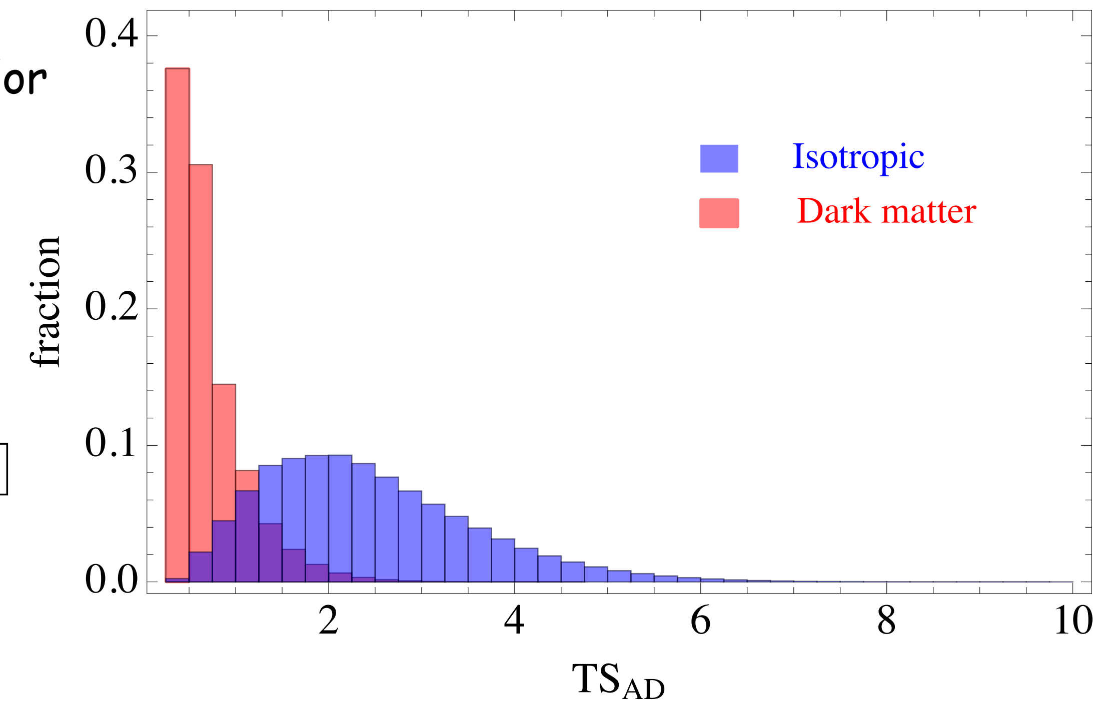
less than 2σ preference for DM dis.



Angular distribution of neutrinos from decaying DM

✓ **Anderson-Darling test:** a powerful non-parametric test, especially sensitive to the end points

statistically larger TS for isotropic distribution



Test Statistics

$$\mathrm{TS}_{\mathrm{AD}} = -N - \frac{1}{N}\sum_{i=1}^{N}(2i-1)\left[\ln\left(\mathrm{CDF}^{\mathrm{DM}}(\vartheta_i)\right) + \ln\left(1-\mathrm{CDF}^{\mathrm{DM}}(\vartheta_{N+1-i})\right)\right]$$

Angular distribution of neutrinos from decaying DM

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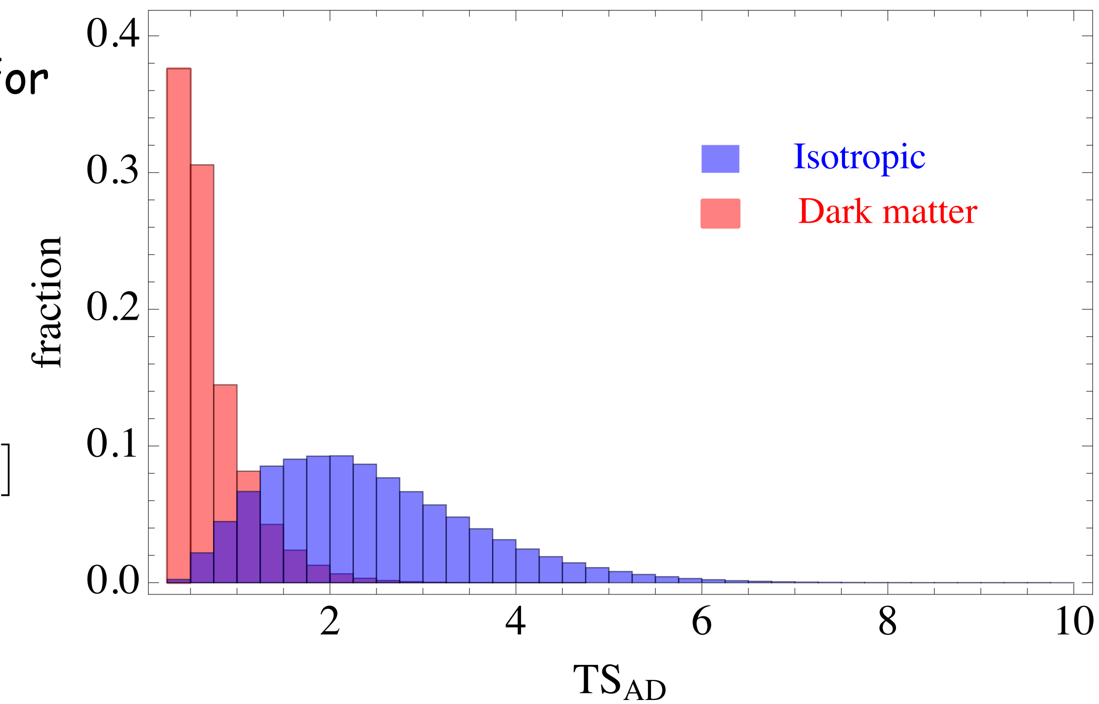

Test Statistics

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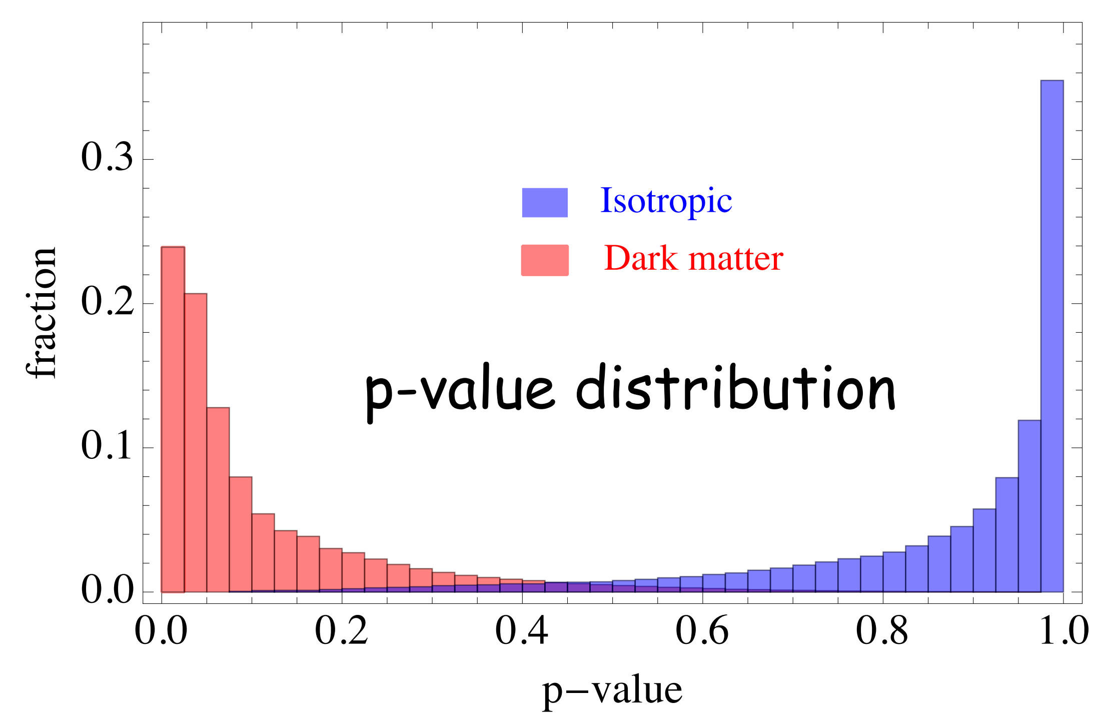
again, generating a sample (10^5) of isotropically distributed set of 20 events

↓

on the average, 11% of generated isotropic sample have smaller TS_{KS} than the values obtained for data vs DM dis.
for data vs isotropic dis. it is 86%

↓

less than 2σ preference for DM dis.



Angular distribution of neutrinos from decaying DM

✓ Some issues:

Angular resolution in KS and AD tests? → even after shifting all the events to higher ϑ values still the preference for DM persist

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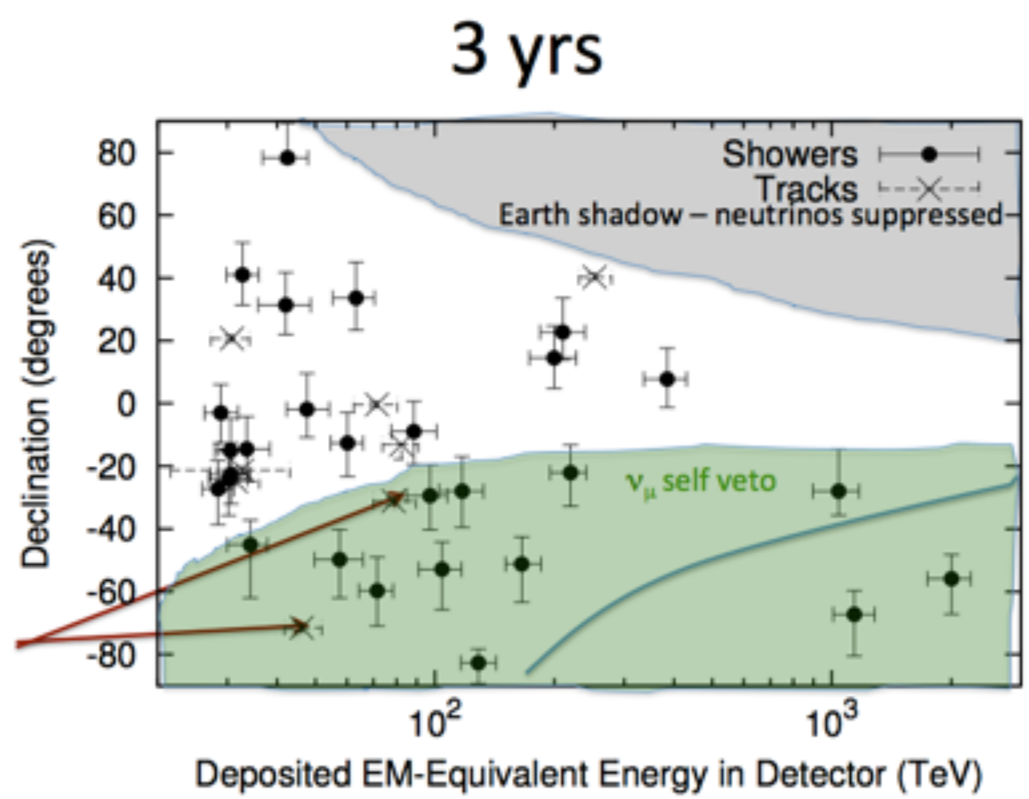

Figure from T. Gaisser

Background rejection? →

Angular distribution of neutrinos from decaying DM

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Angular resolution in KS and AD tests? → even after shifting all the events to higher ϑ values still the preference for DM persist

3 yrs

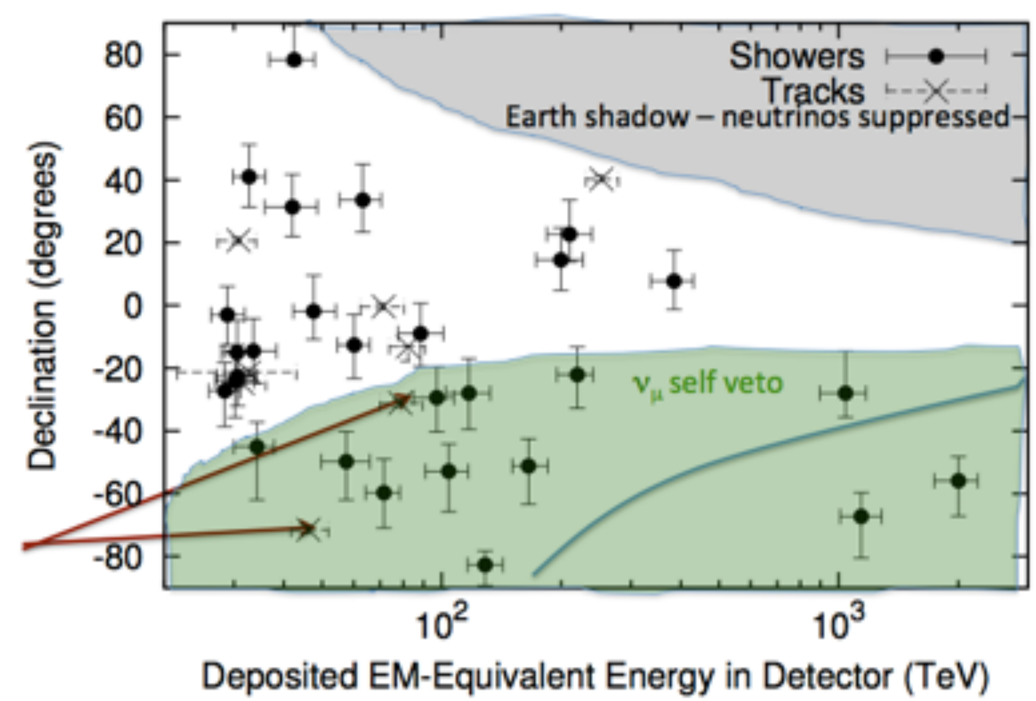
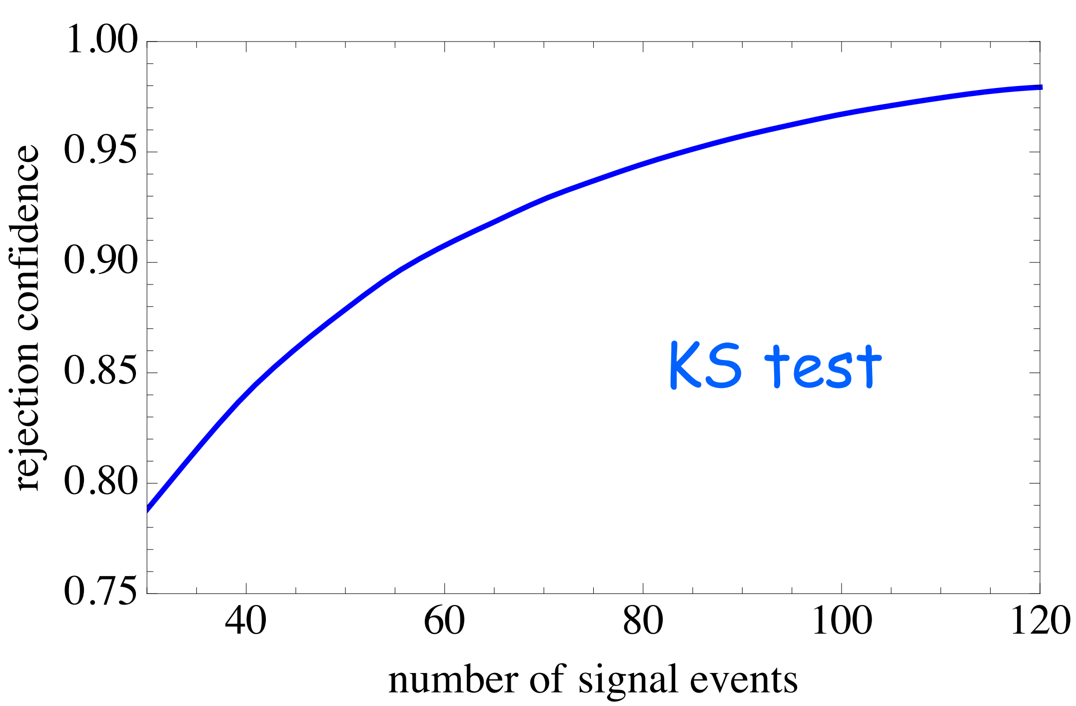


Figure from T. Gaisser

Background rejection? →



How many events are needed for a 3σ discrimination?

Gamma ray bounds

Universe is opaque for gamma-rays with E > 1 TeV → cascades develop: gamma-ray interaction with interstellar radiation field and CMB → gamma-rays populate at lower energies < $10^{(2-3)}$ GeV

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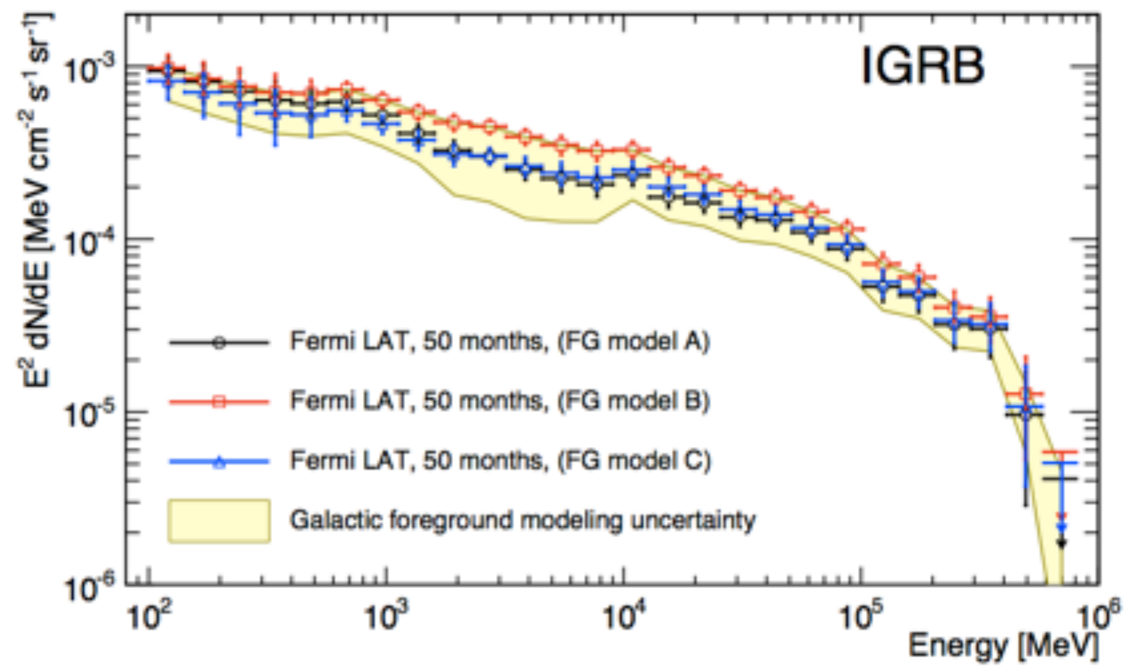
✓ Isotropic diffuse gamma-ray background by Fermi-LAT

M. Ackermann et al. [The Fermi LAT Collaboration], arXiv:1410.3696 [astro-ph.HE].



integrated energy density

$$\omega_\gamma = \frac{4\pi}{c}\int_{E_1}^{E_2} E_\gamma \frac{\mathrm{d}\varphi_\gamma}{\mathrm{d}E_\gamma}\,\mathrm{d}E_\gamma \lesssim 4.4\times 10^{-7}\,\mathrm{eV/cm^3}$$

$E_1 \sim \mathcal{O}(1)$ GeV $\qquad E_2 \sim \mathcal{O}(100)$ GeV

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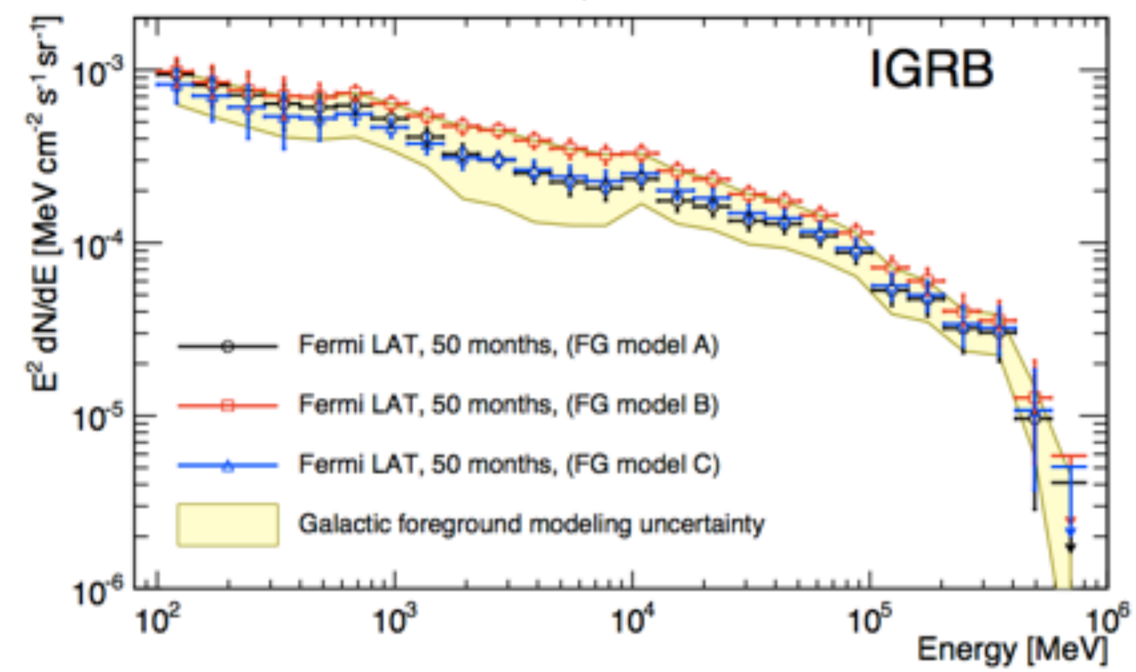

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$$E_1 \sim \mathcal{O}(1)\ \mathrm{GeV} \qquad E_2 \sim \mathcal{O}(100)\ \mathrm{GeV}$$

total electromagnetic energy budget (NH case)

$$\frac{4\pi}{c}\int \sum_{\mathrm{i=gal,extragal}} \left[E_\gamma \left.\frac{\mathrm{d}\varphi_\gamma}{\mathrm{d}E_\gamma}\right)^{\mathrm{i}} + E_e \left.\frac{\mathrm{d}\varphi_{e^\pm}}{\mathrm{d}E_e}\right)^{\mathrm{i}} \right] \mathrm{d}E \simeq 5.2\times 10^{-8}\ \mathrm{eV/cm^3}$$

✓

Gamma ray bounds

Universe is opaque for gamma-rays with E > 1 TeV → cascades develop: gamma-ray interaction with interstellar radiation field and CMB → gamma-rays populate at lower energies < $10^{(2-3)}$ GeV

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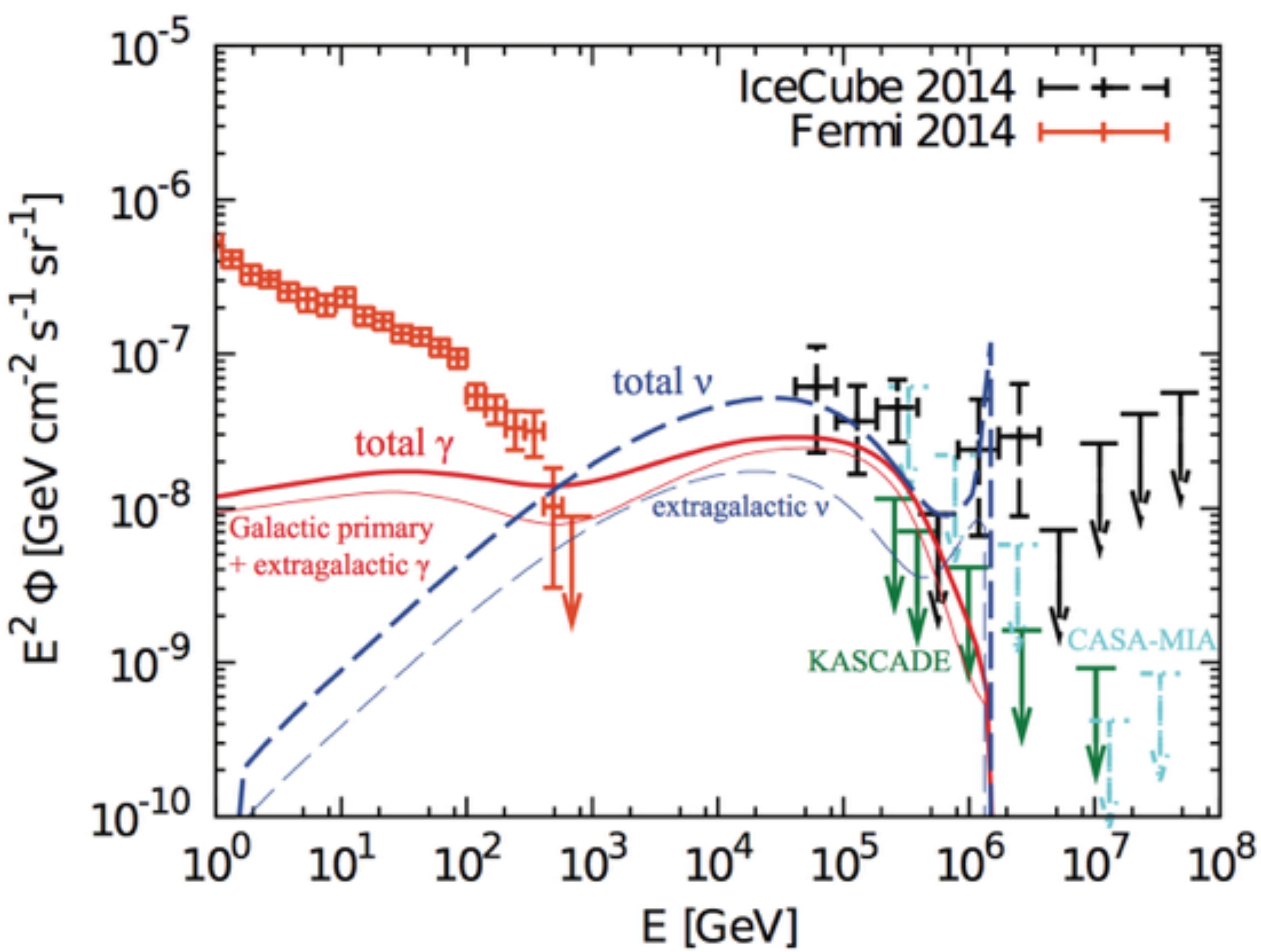
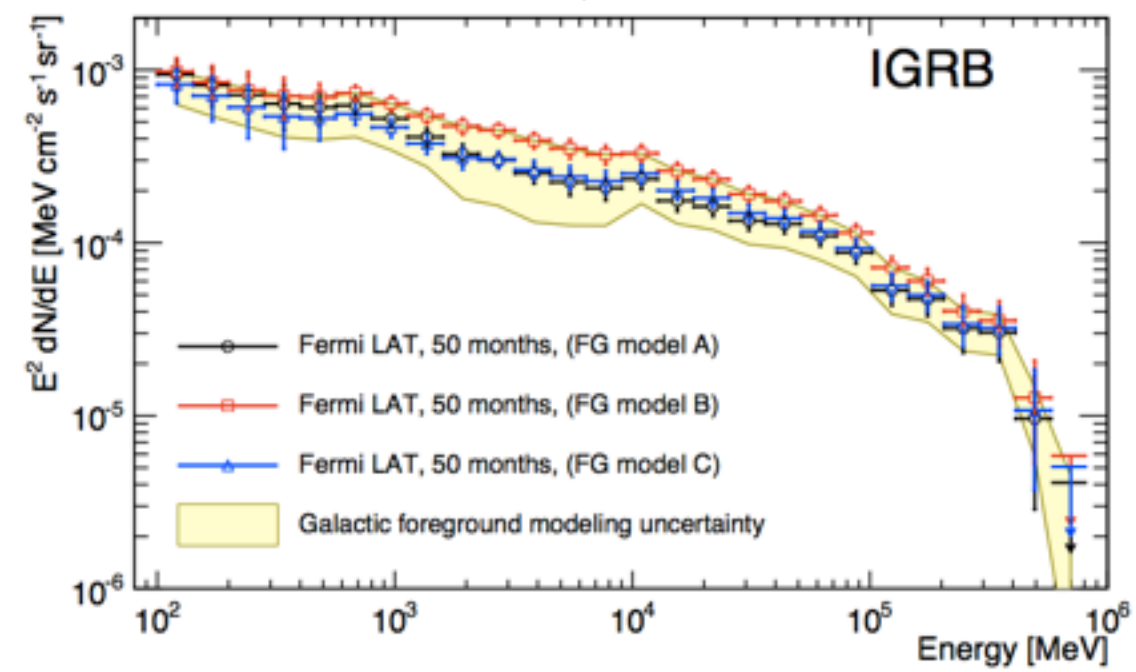


Murase, Laha, Ando, Ahlers, arXiv:1503.04663

M. Ackermann et al. [The Fermi LAT Collaboration], arXiv:1410.3696 [astro-ph.HE].



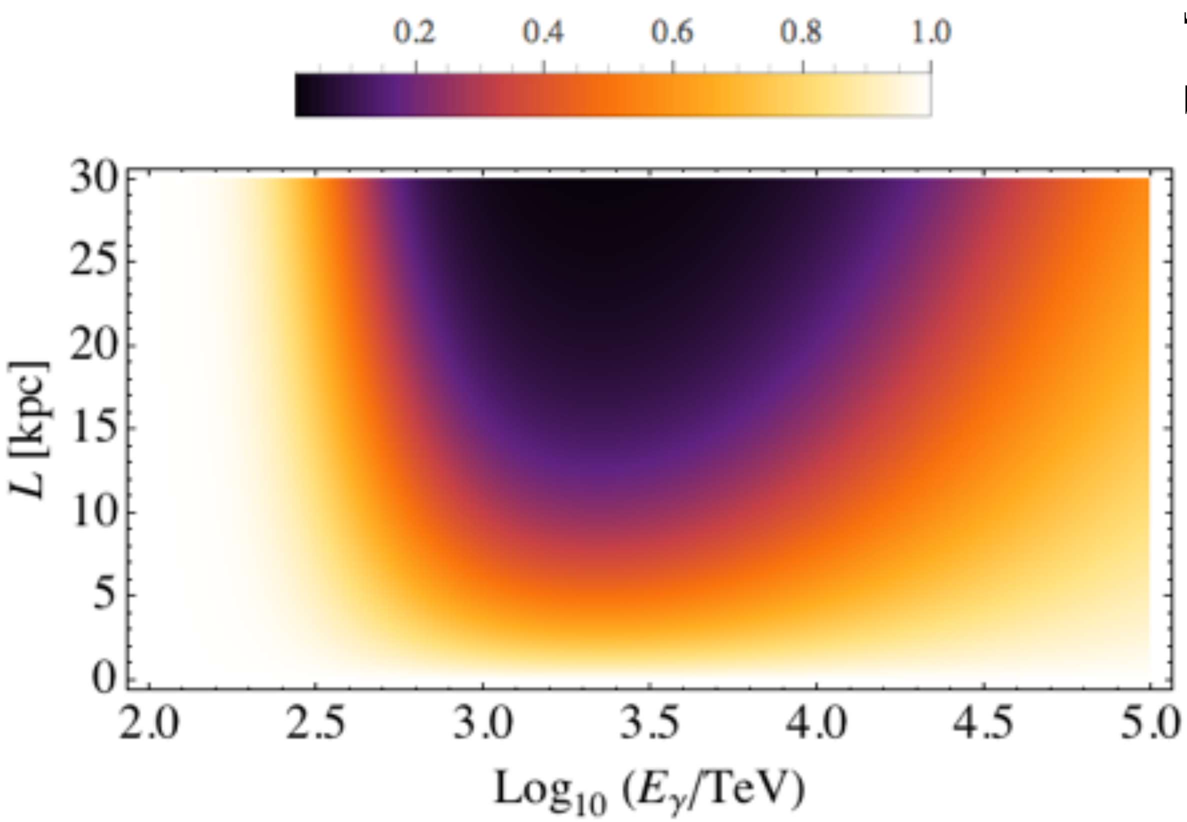
Gamma ray bounds

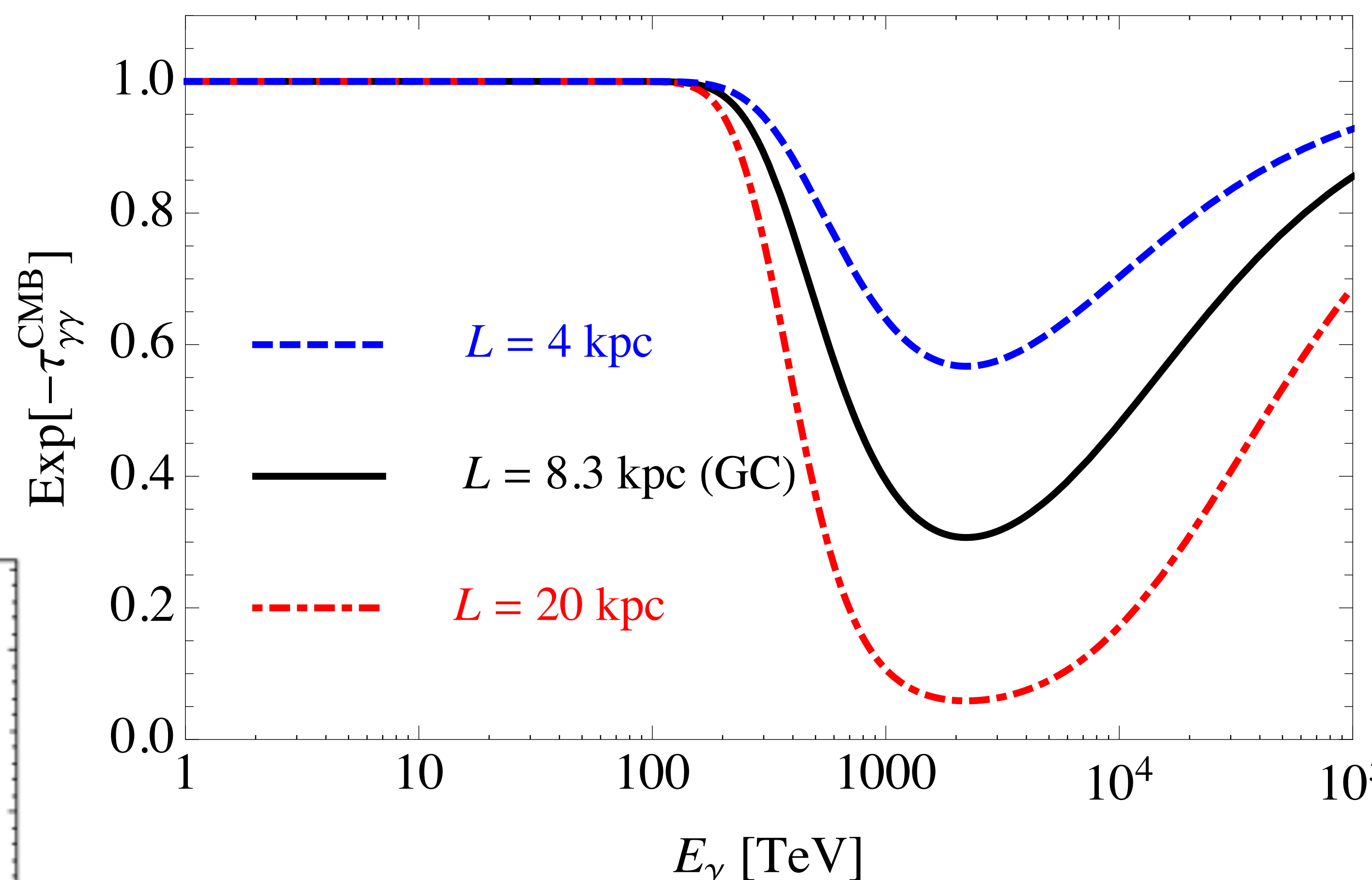
✓ Galactic component

at ~ PeV, the absorption length of gamma-rays are comparable to Galactic distances → neither full absorption or cascade development, nor full transparency

A. E. and P. Serpico, arXiv:1505.06486

Absorption due to pair production on CMB photons





Absorption at ~ PeV

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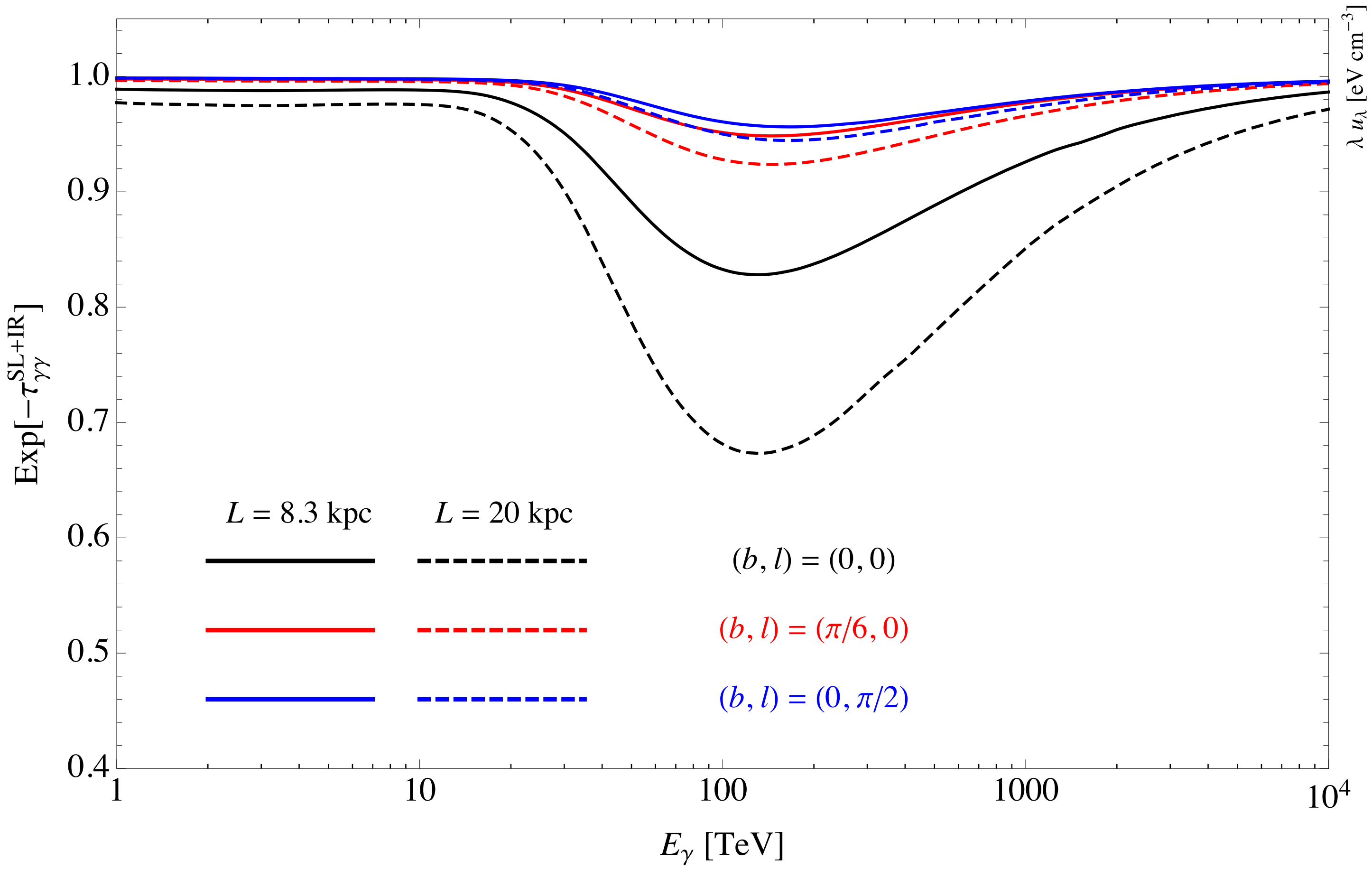
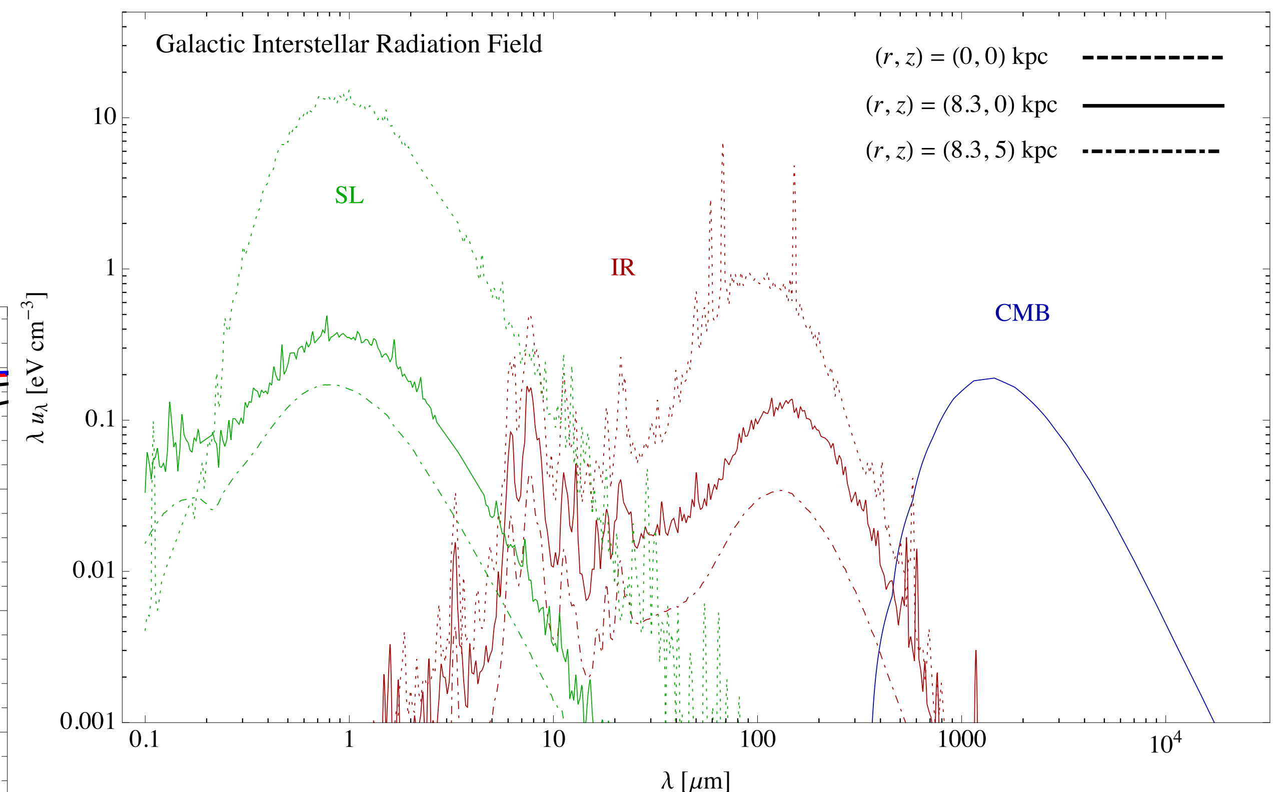
Absorption due to pair production on SL+IR photons





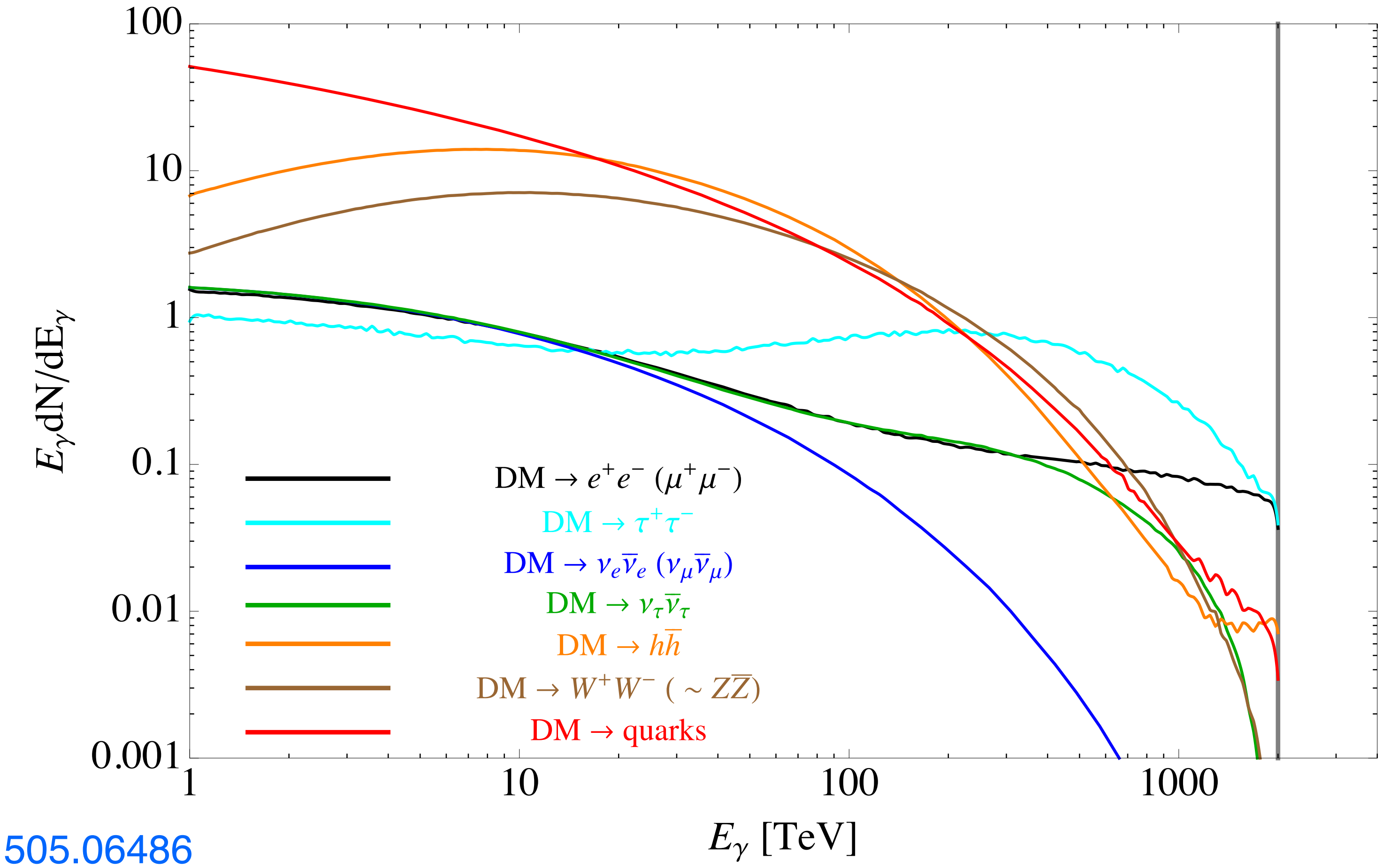
Absorption at ~ 100 TeV

Gamma ray bounds

✔ Galactic component

at ~ PeV, the absorption length of gamma-rays are comparable to Galactic distances → neither full absorption or cascade development, nor full transparency

Prompt component
$$\frac{d\Phi_\gamma}{dE_\gamma}(E_\gamma, b, l) = \frac{1}{4\pi\, m_{\rm DM}\, \tau_{\rm DM}} \frac{dN_\gamma}{dE_\gamma}(E_\gamma) \int_0^\infty \rho_{\rm h}[\varrho(s,b,l)]\, e^{-\tau_{\gamma\gamma}(E_\gamma,s,b,l)}\, ds$$



calculated by PYTHIA 8.2

A. E. and P. Serpico, arXiv:1505.06486

Gamma ray bounds

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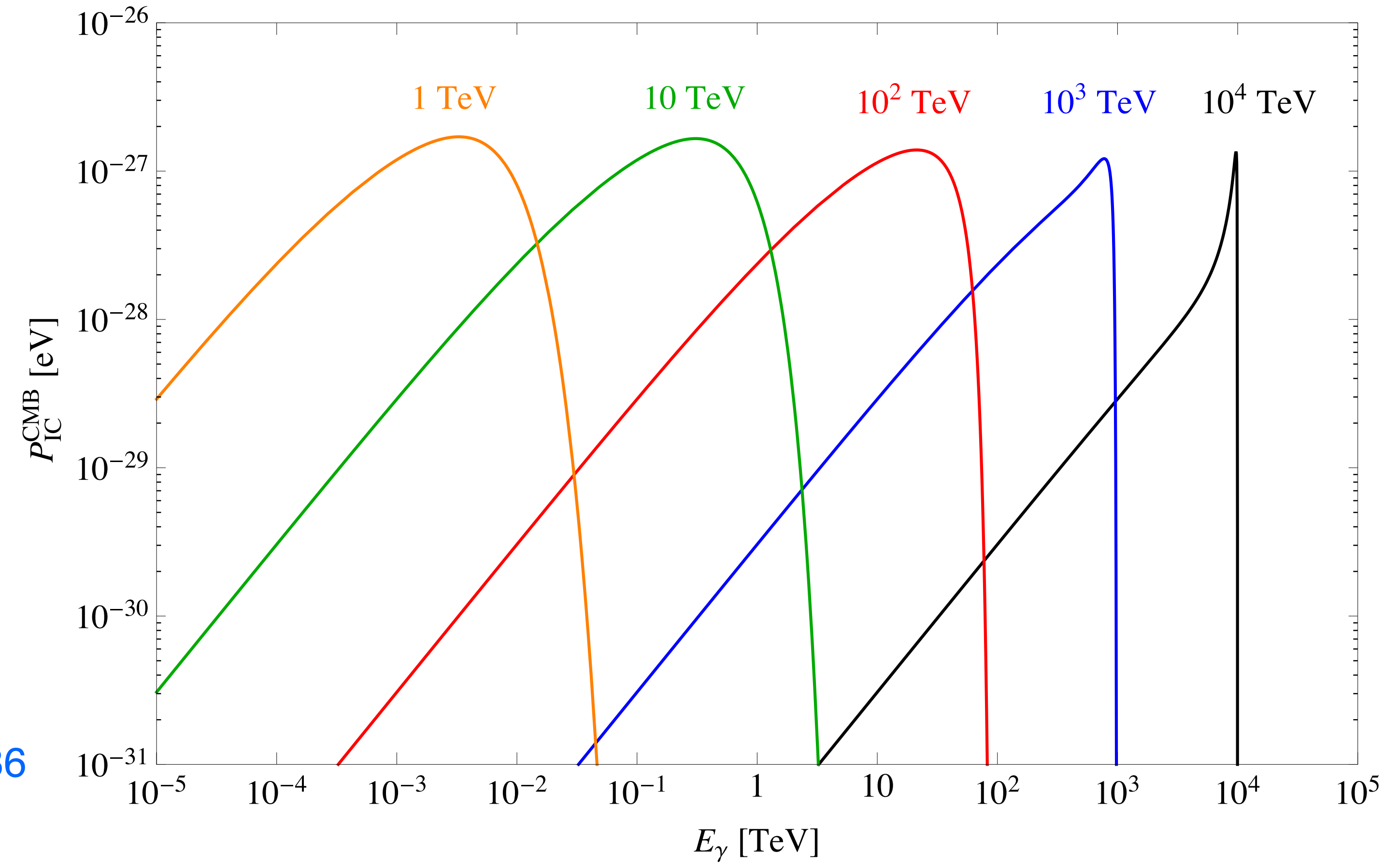
inverse-Compton component

$$\frac{\mathrm{d}\Phi_{\rm IC}}{\mathrm{d}E_\gamma}(E_\gamma,b,l)=\frac{1}{4\pi E_\gamma}\int_0^\infty \mathrm{d}s\, e^{-\tau_{\gamma\gamma}(E_\gamma,s,b,l)}\int_{m_e}^{m_{\rm DM}/2}\mathrm{d}E_e\frac{\mathrm{d}n_e}{\mathrm{d}E_e}(E_e,\varrho)\,P_{\rm IC}(E_e,E_\gamma,\varrho)$$

$$P_{\rm IC}=P_{\rm IC}^{\rm CMB}+P_{\rm IC}^{\rm SL+IR}$$

A. E. and P. Serpico, arXiv:1505.06486



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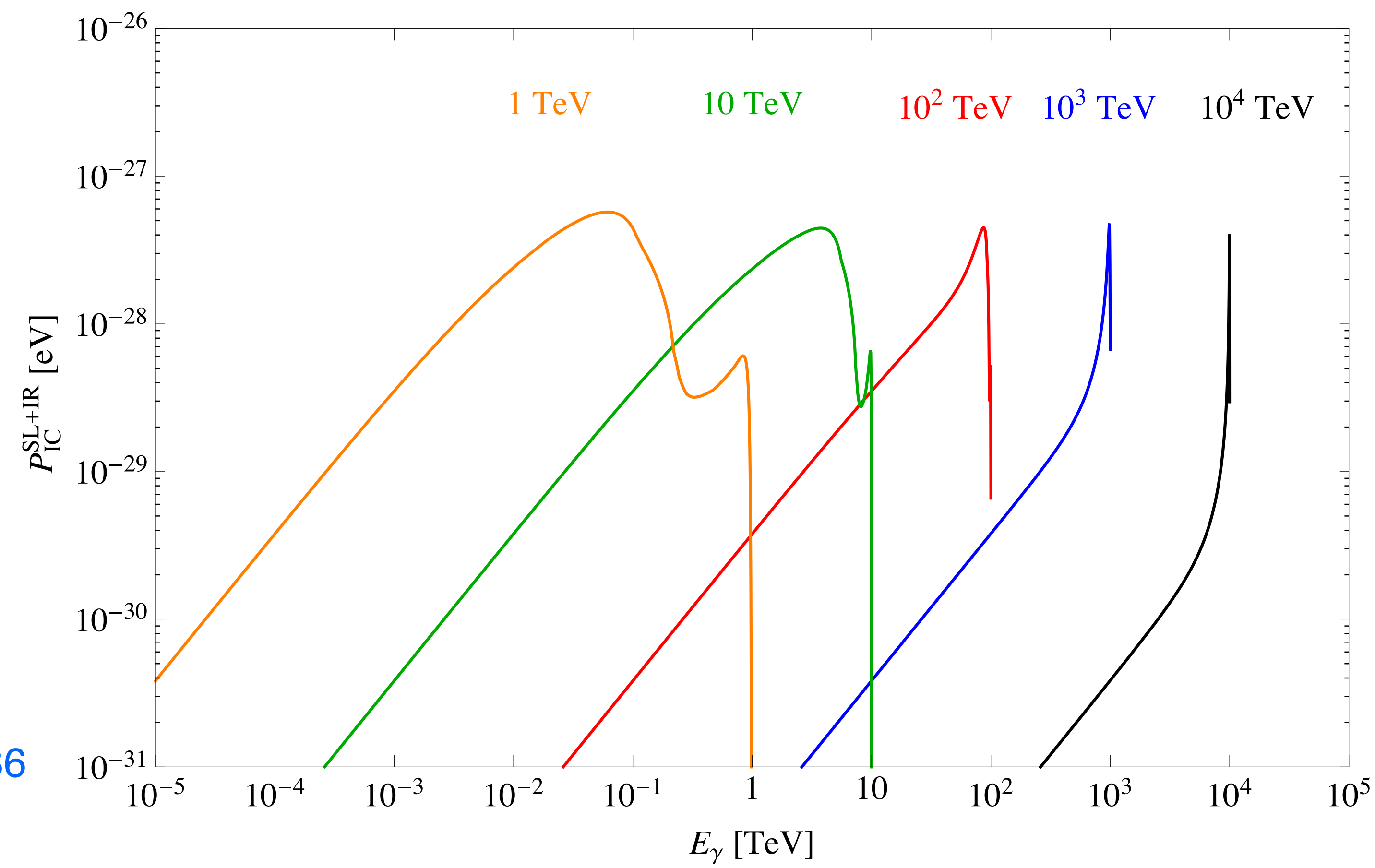
$$P_{IC} = P_{IC}^{CMB} + P_{IC}^{SL+IR}$$



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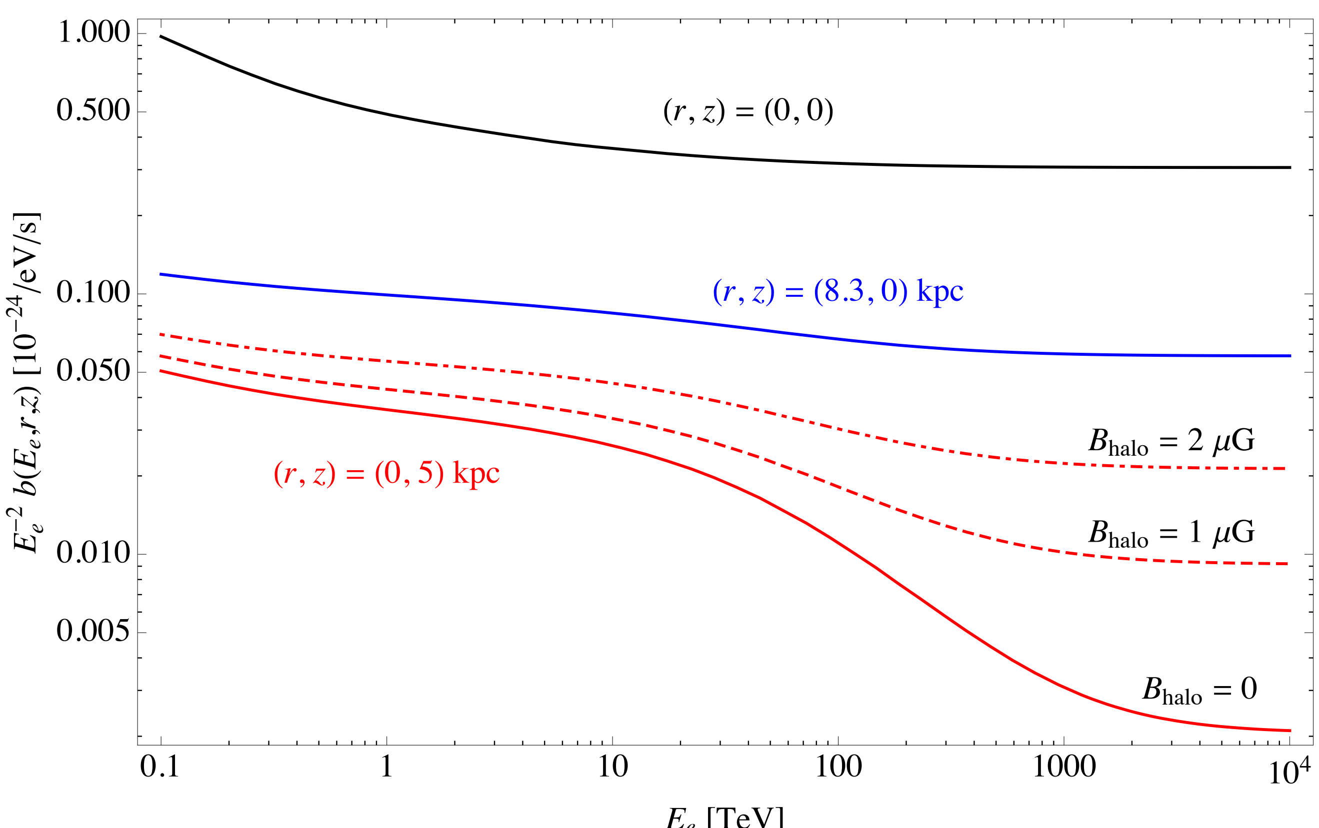
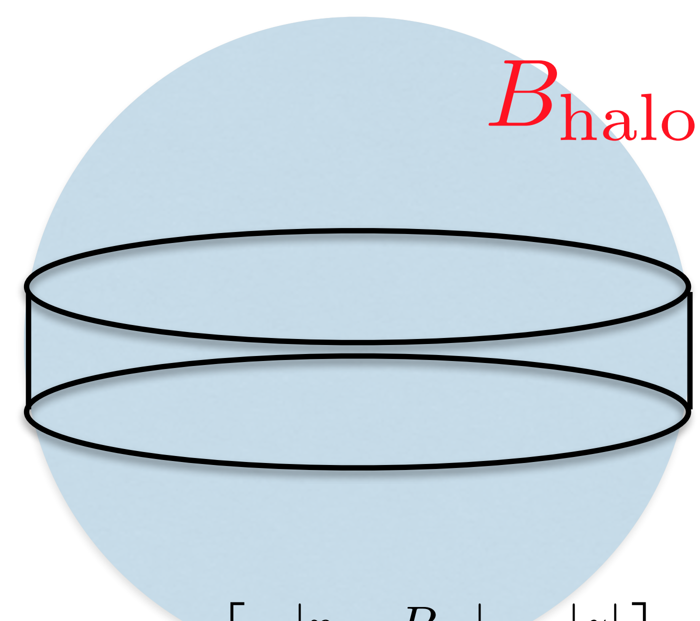
$$\frac{dn_e}{dE_e}(E_e,\vec{x})=\frac{1}{m_{DM}\,\tau_{DM}}\frac{\rho_h(\vec{x})}{b(E_e,\vec{x})}\int_{E_e}^{m_{DM}/2}\frac{dN_e}{dE'_e}(E'_e)\, I_{diff}(E_e,E'_e,\vec{x})\, dE'_e$$



$$b(E_e,\vec{x})\equiv-\frac{dE_e}{dt}=b_{IC}(E_e,\vec{x})+b_{syn}(E_e,\vec{x})$$



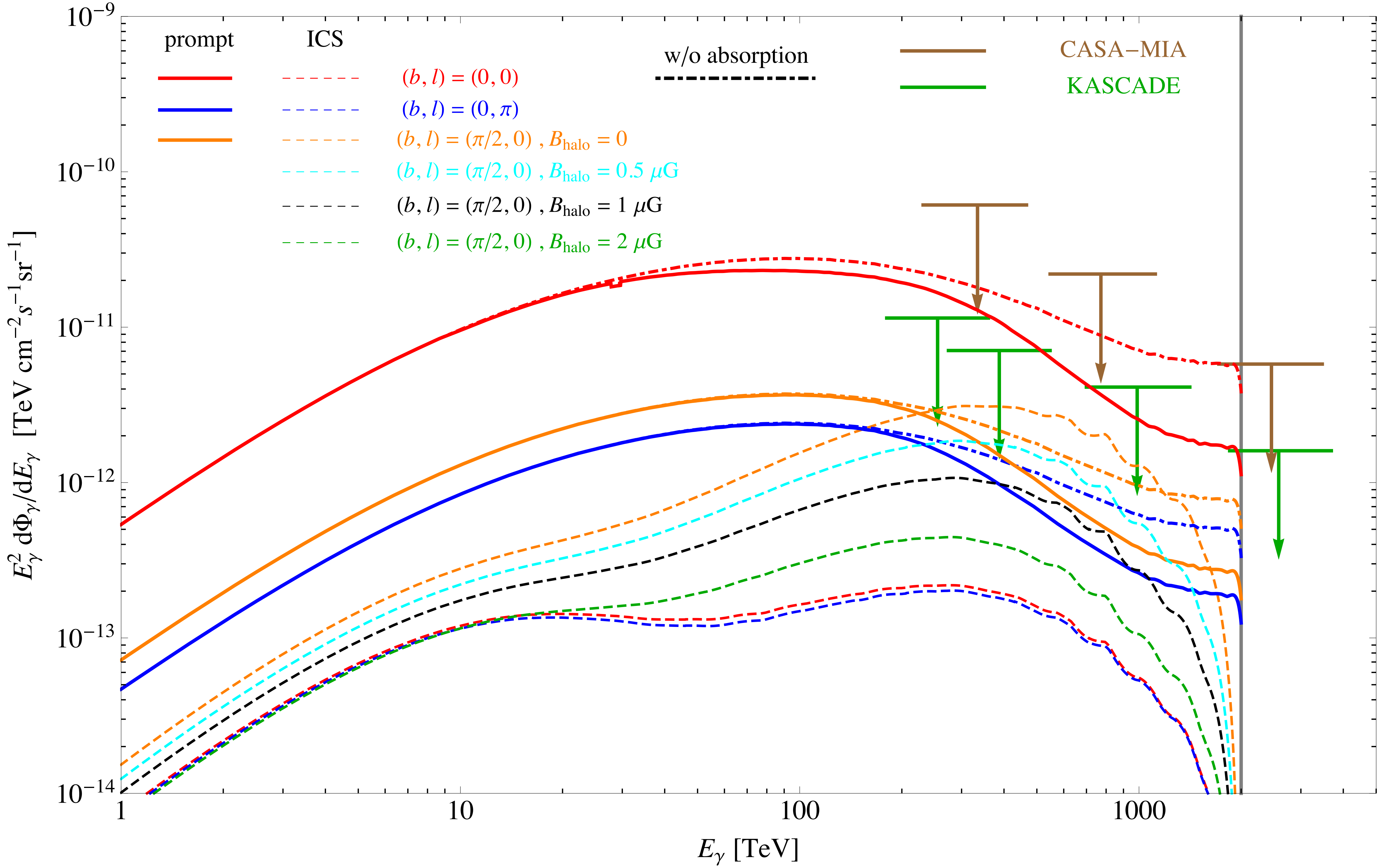
$$B_{reg}(\vec{x})=B_0\ \exp\left[-\frac{|r-R_\odot|}{r_B}-\frac{|z|}{z_B}\right]$$

A. E. and P. Serpico, arXiv:1505.06486

Gamma ray bounds

✓ Galactic component

$\tau_{DM} = 10^{28}$ s and $m_{DM} = 4$ PeV



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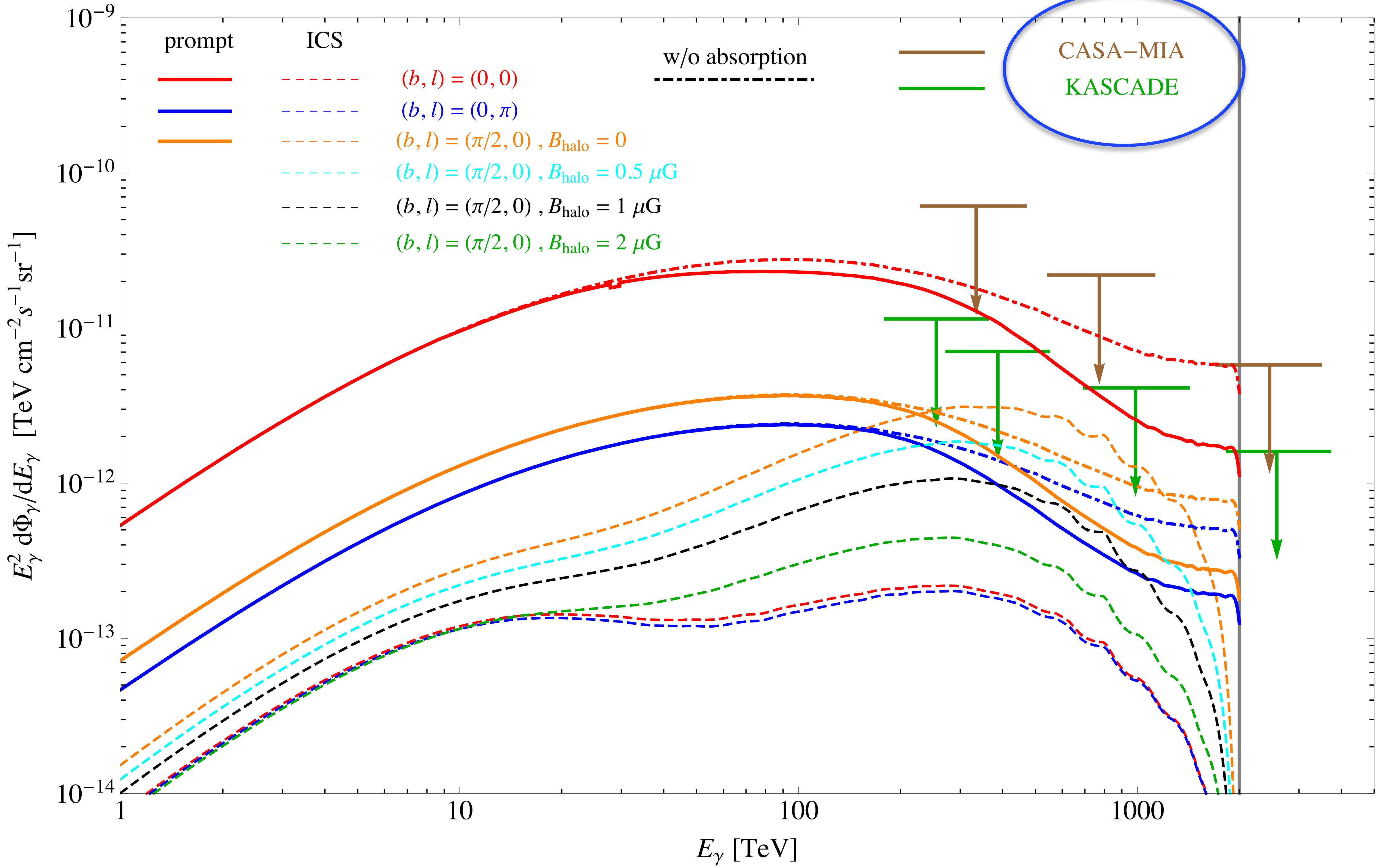
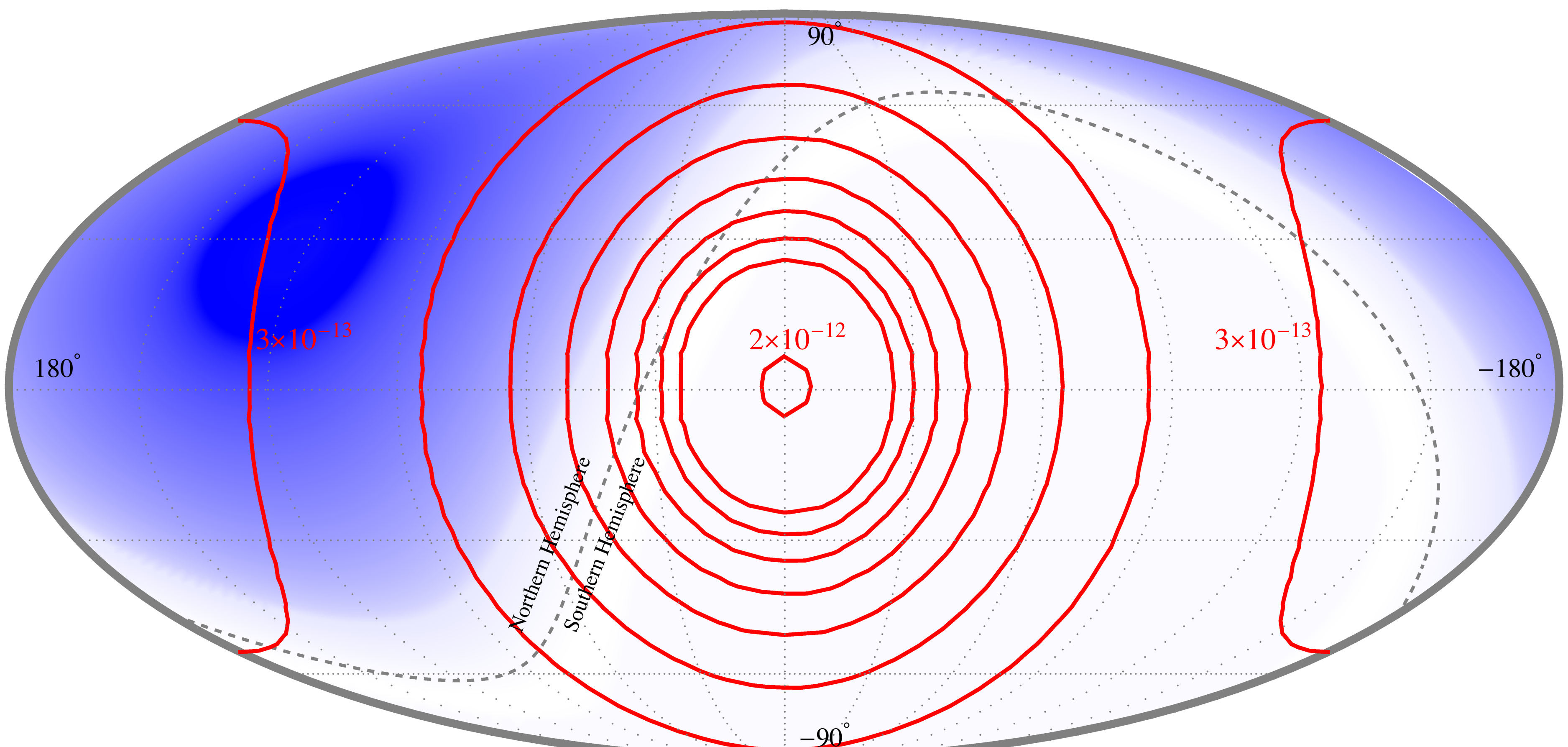


A. E. and P. Serpico, arXiv:1505.06486

Gamma ray bounds

✓ Galactic component

KASCADE

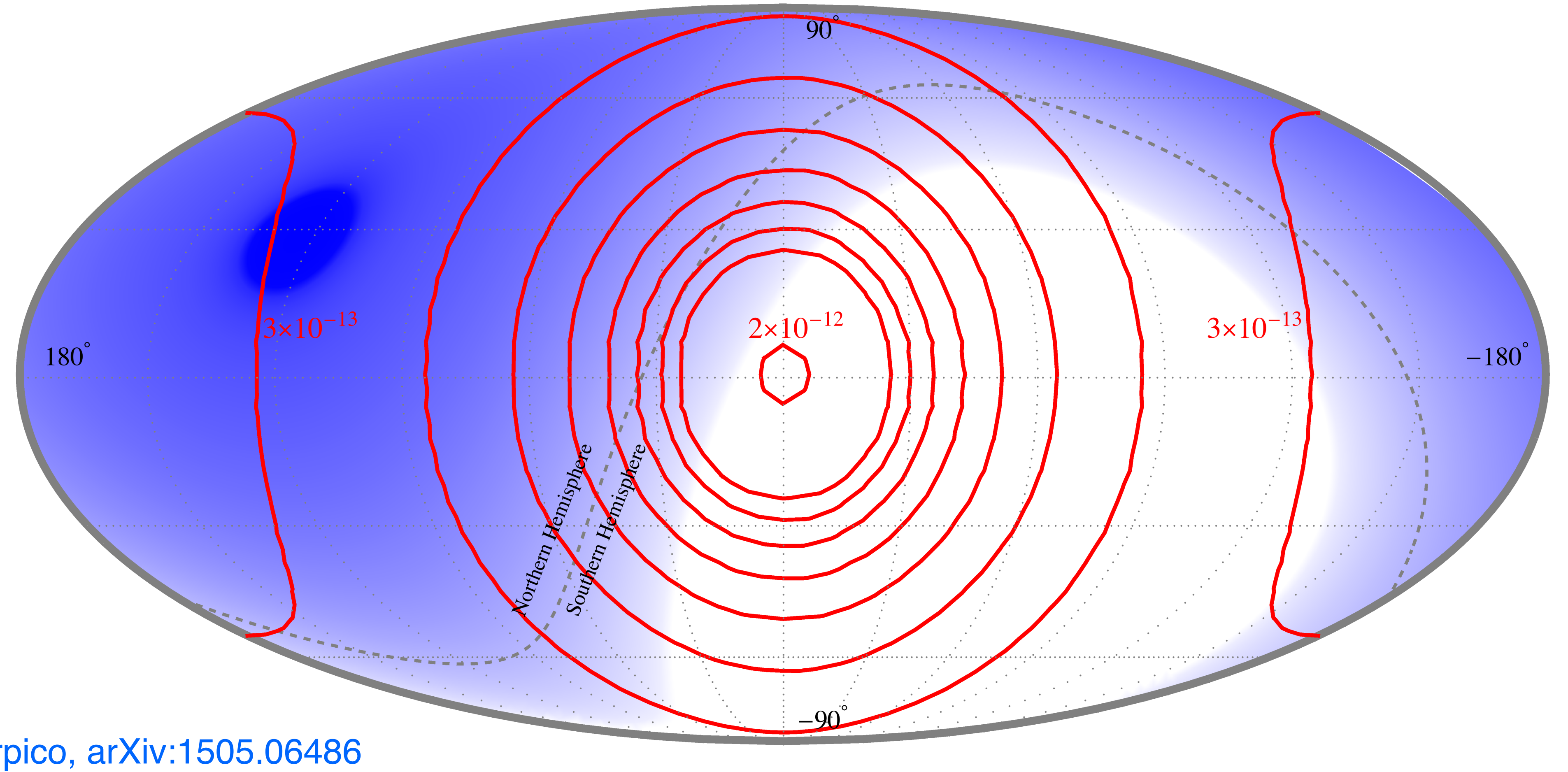


A. E. and P. Serpico, arXiv:1505.06486

Gamma ray bounds

✓ Galactic component

CASA-MIA



A. E. and P. Serpico, arXiv:1505.06486

Gamma ray bounds

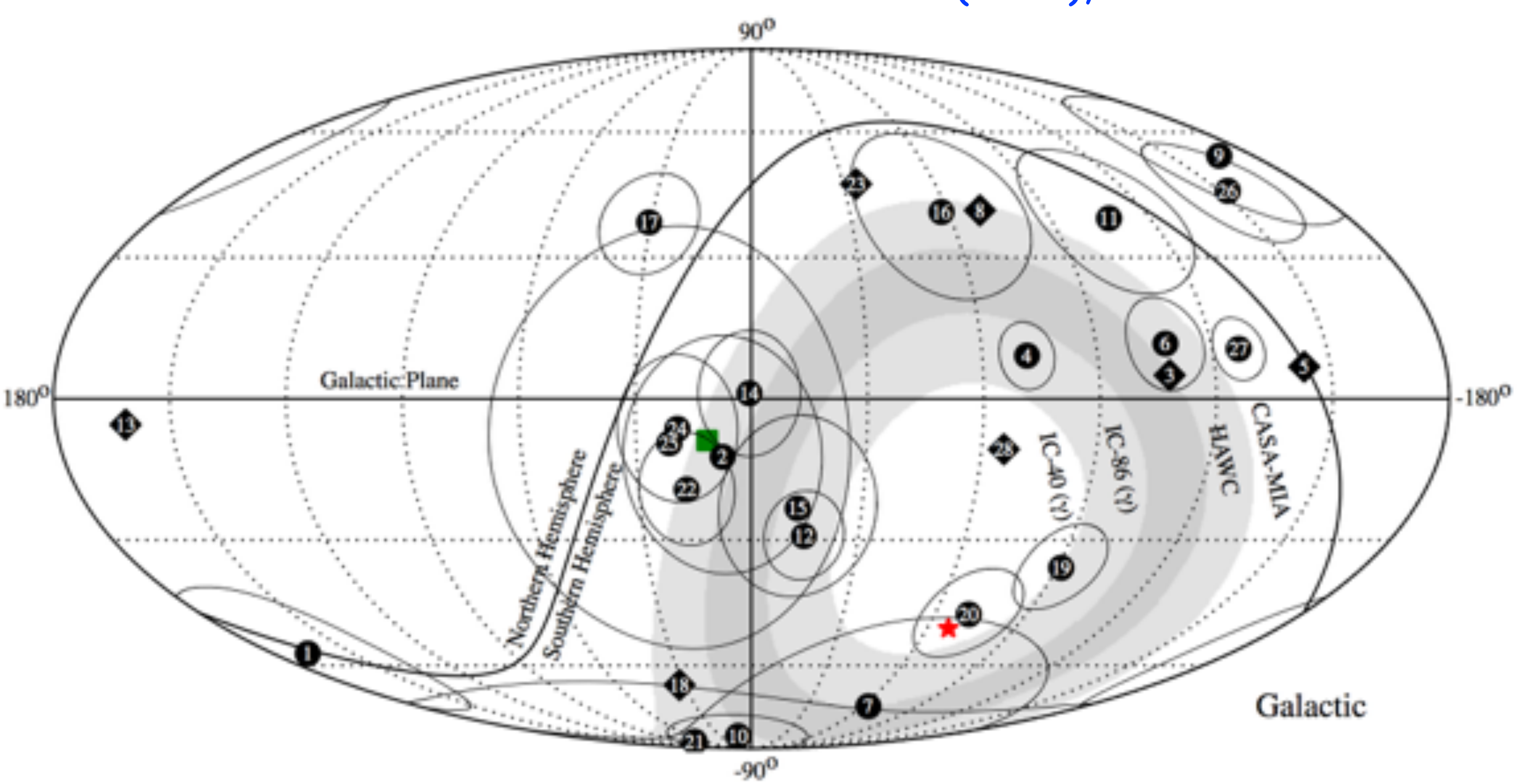
✔ Galactic component

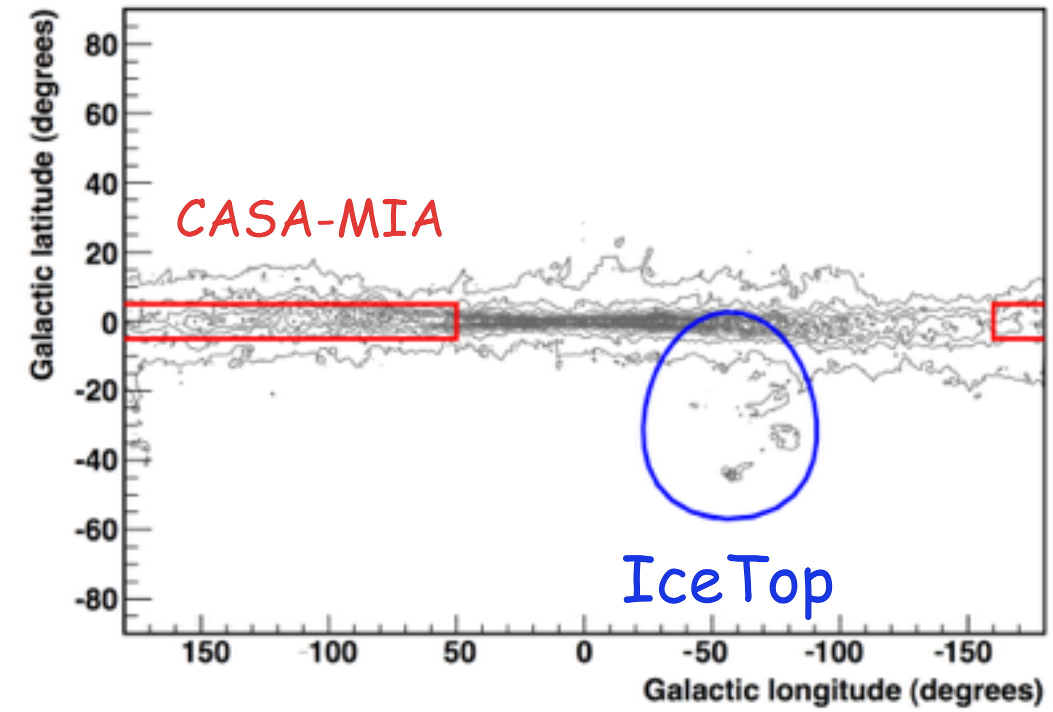
Future experiments

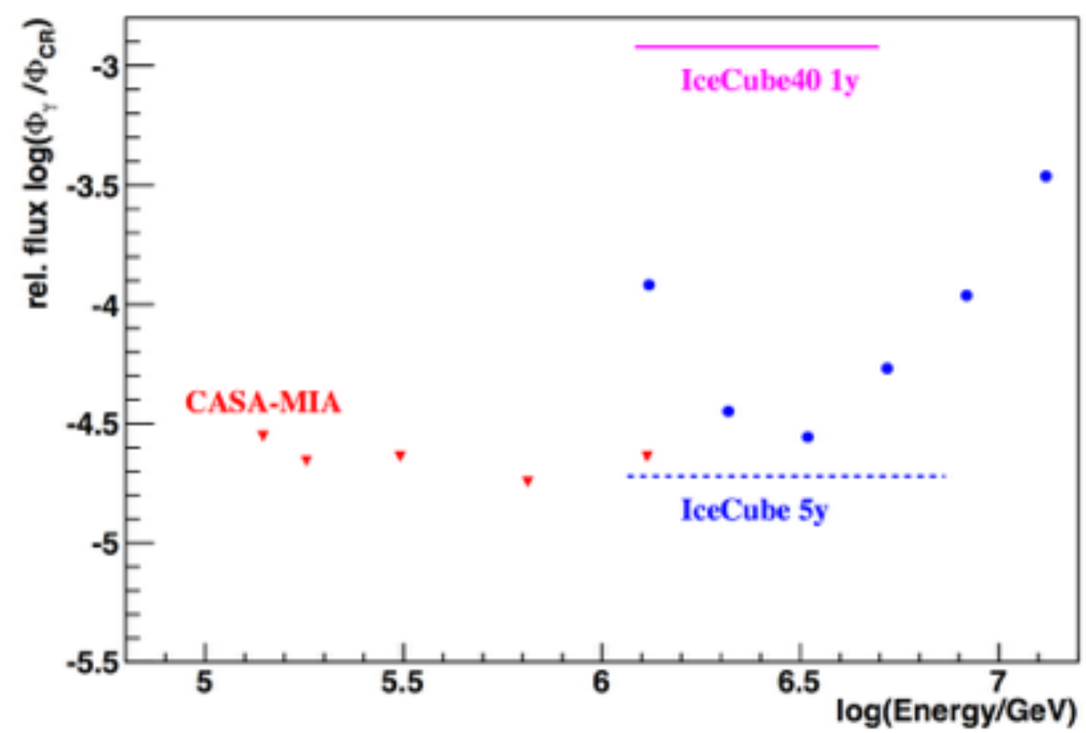
IceTop

HAWC

M. Ahlers and K. Murase, PRD (2014), arXiv:1309.4077







Gamma ray bounds

✓ Galactic component

Anisotropy

$$a_\gamma = \frac{\left.\frac{d\Phi_\gamma}{dE_\gamma}\right|_{GC} - \left.\frac{d\Phi_\gamma}{dE_\gamma}\right|_{anti-GC}}{\frac{d\Phi_{CR}}{dE}}$$

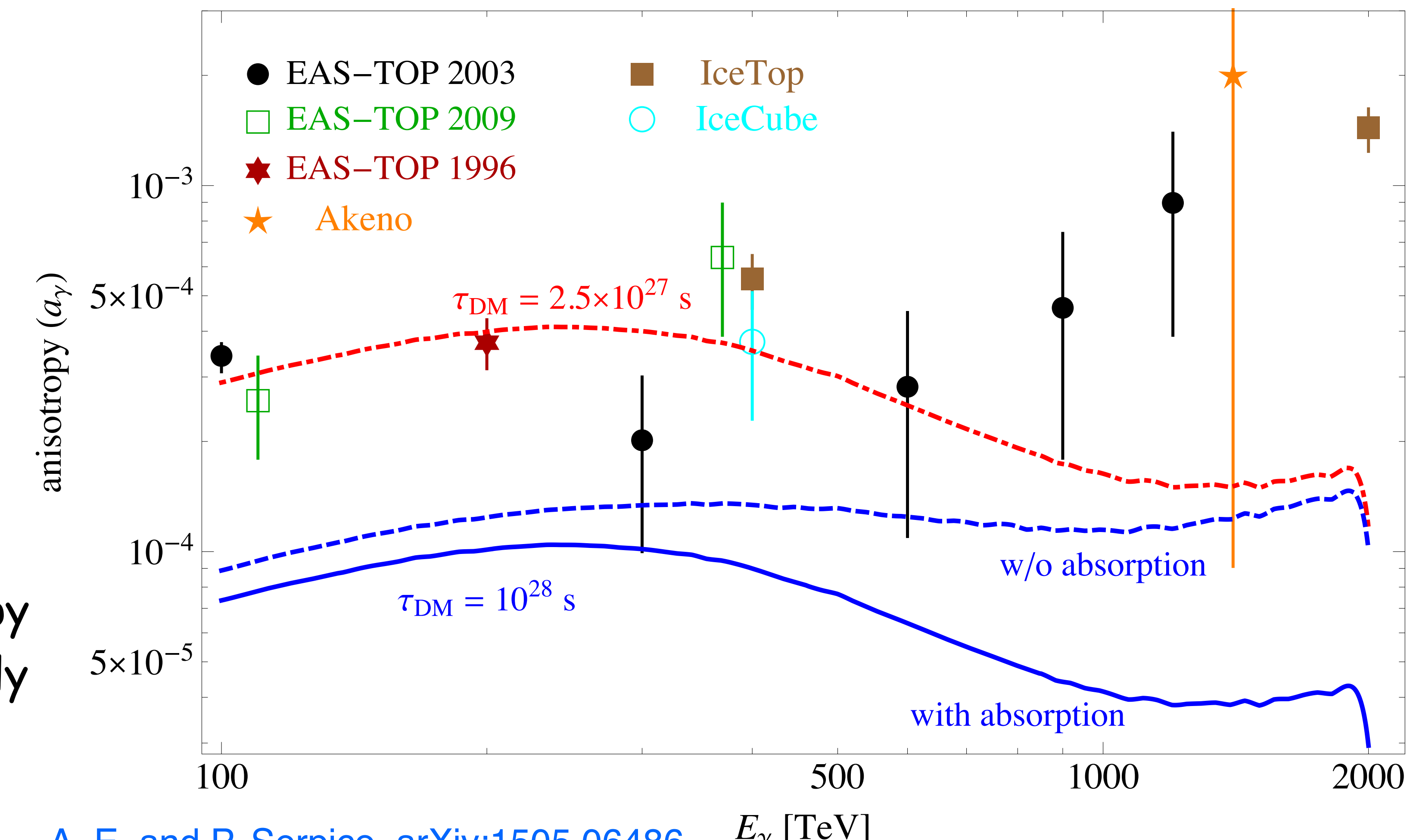
Total CR flux

✓ No need to γ/hadron discrimination

✓ Absorption suppress the anisotropy

✓ The bound 2.5×10^{27} s can be set

✓ Adding the phase info of anisotropy would improve the limits significantly



A. E. and P. Serpico, arXiv:1505.06486

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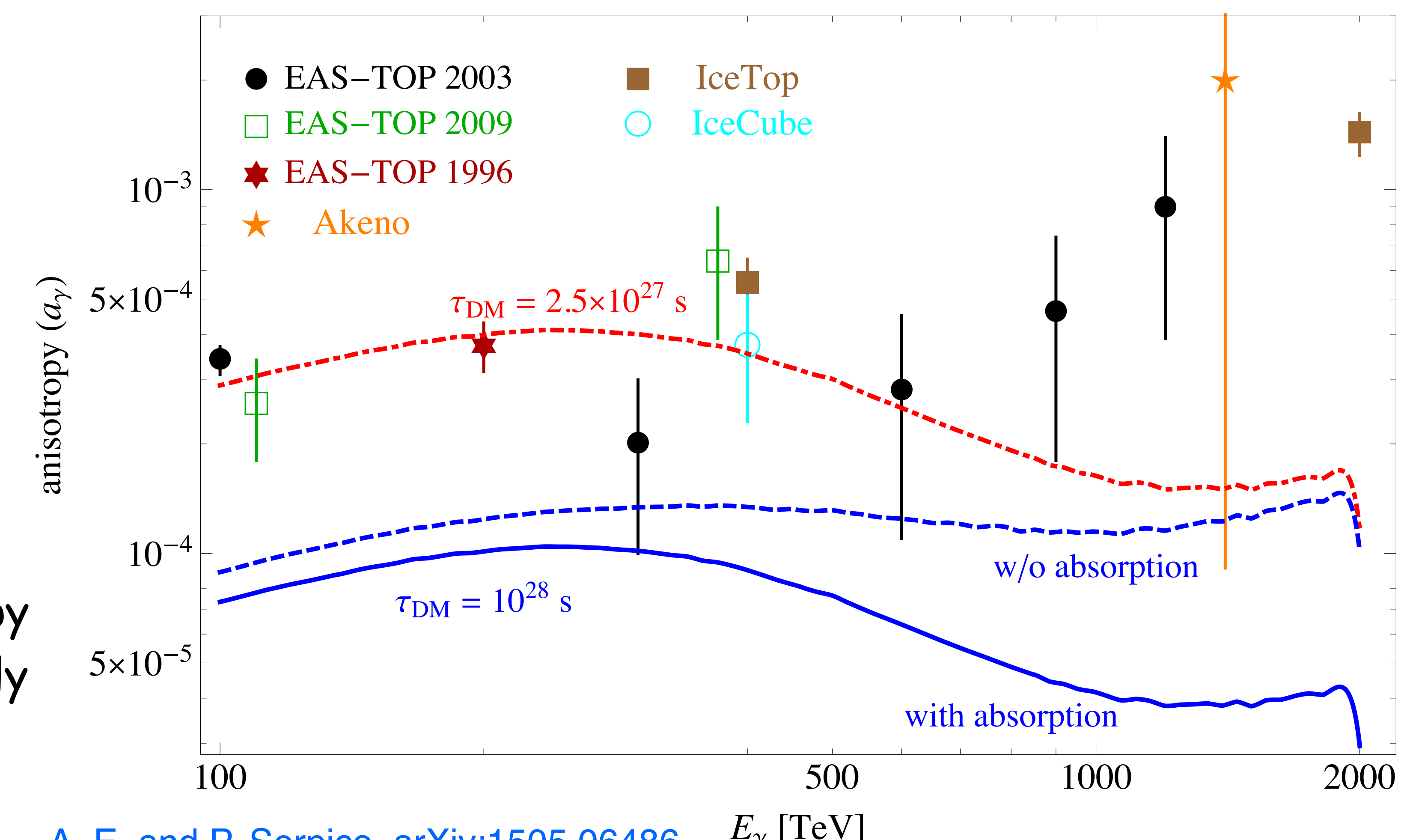


A. E. and P. Serpico, arXiv:1505.06486

What can we learn about DM if the IceCube events originate from conventional astrophysical flux?

constraining:

✓ DM lifetime

✓ annihilation cross section

Constraining DM properties

✓ DM lifetime

contribution of DM to the events in each bin should be smaller than N_{limit}

bin #	$\log_{10}(E_\nu/\text{TeV})$	$N_{\text{astro}}(E_\nu^{-2} \div E_\nu^{-2.3})$	N_{data}	$N_{\text{limit}}\ (E_\nu^{-2} \div E_\nu^{-2.3})$	N_{limit}
#1	1.4 – 1.6	9.46 ÷ 10	11	7.8 ÷ 7.46	16.6
#2	1.6 – 1.8	4.31 ÷ 5.3	6	6.53 ÷ 5.87	10.5
#3	1.8 – 2.0	4.55 ÷ 5.68	7	7.41 ÷ 6.58	11.8
#4	2.0 – 2.2	3.97 ÷ 4.82	3	3.98 ÷ 3.73	6.68
#5	2.2 – 2.4	3.32 ÷ 3.56	4	5.15 ÷ 5.01	8.00
#6	2.4 – 2.6	2.59 ÷ 2.42	2	3.65 ÷ 3.71	5.32
#7	2.6 – 2.8	1.96 ÷ 1.62	0	2.3 ÷ 2.3	2.3
#8	2.8 – 3.0	1.55 ÷ 1.1	0	2.3 ÷ 2.3	2.3
#9	3.0 – 3.2	1.2 ÷ 0.74	2	4.31 ÷ 4.64	5.32
#10	3.2 – 3.4	0.92 ÷ 0.5	1	3.3 ÷ 3.51	3.89
#11	3.4 – 3.6	0.73 ÷ 0.35	0	2.3 ÷ 2.3	2.3
#12	3.6 – 3.8	1.72 ÷ 0.76	0	2.3 ÷ 2.3	2.3

Poisson statistics:

at q% C.L.

$$\frac{q}{100} = \frac{\int_0^{N^i_{\text{limit}}} L(N^i_{\text{data}}, N)\ \mathrm{d}N}{\int_0^{\infty} L(N^i_{\text{data}}, N)\ \mathrm{d}N}$$

$$L(N^i_{\text{data}}, N) = \frac{(N + N^i_{\text{astro}})^{N^i_{\text{data}}}}{N^i_{\text{data}}!} e^{-(N+N^i_{\text{astro}})}$$

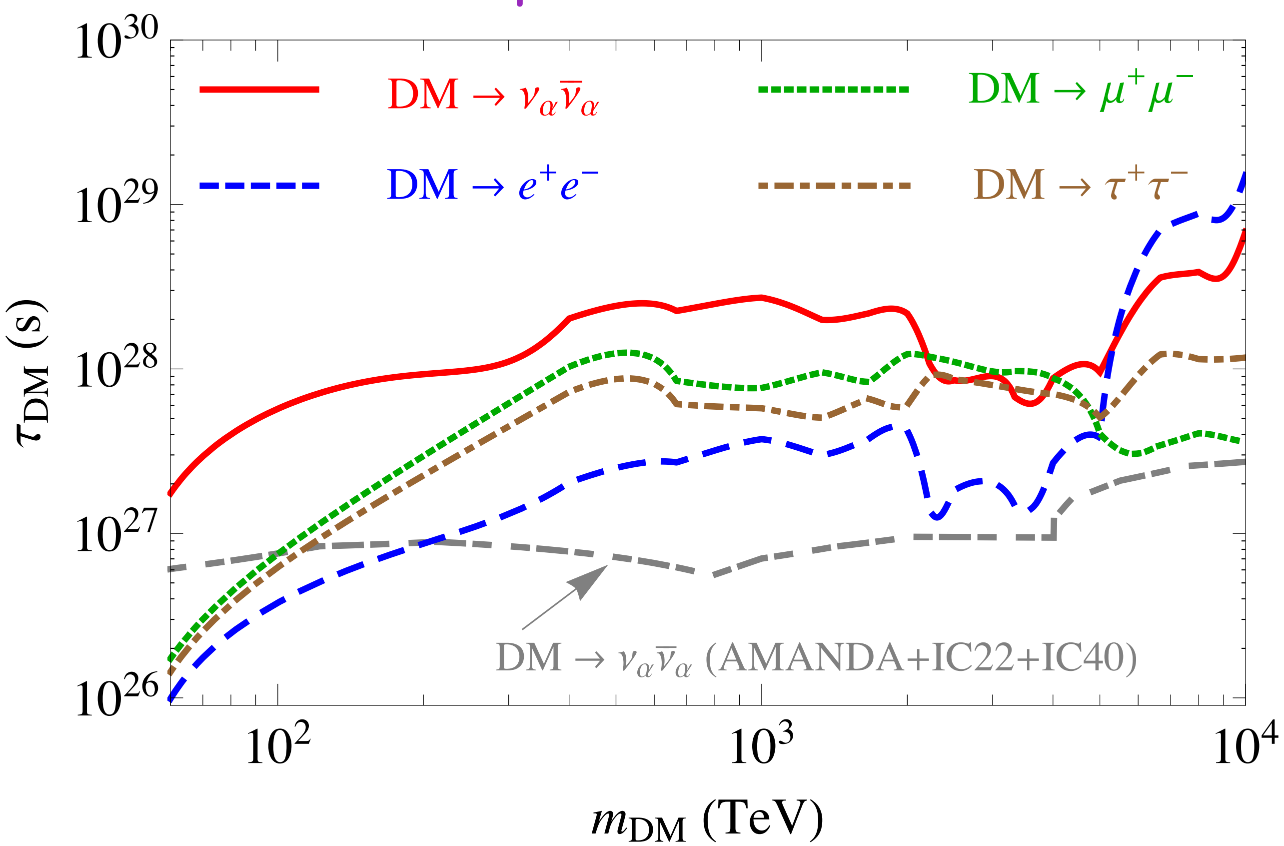
or

$$L(N^i_{\text{data}}, N) = \frac{(N)^{N^i_{\text{data}}}}{N^i_{\text{data}}!} e^{-N}$$

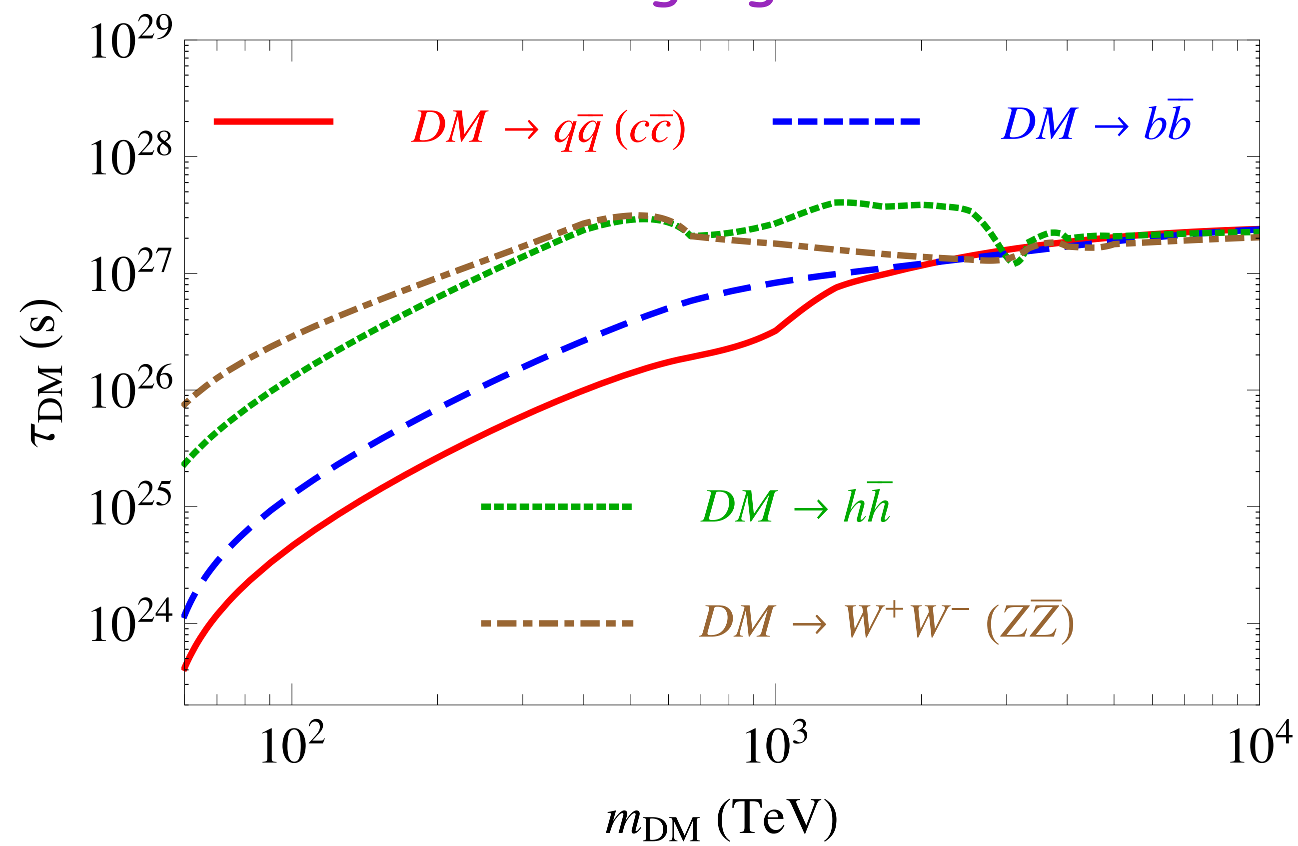
Constraining DM properties

✓ limits on DM lifetime (90% C.L.)

leptonic channels



hadronic/gauge channels



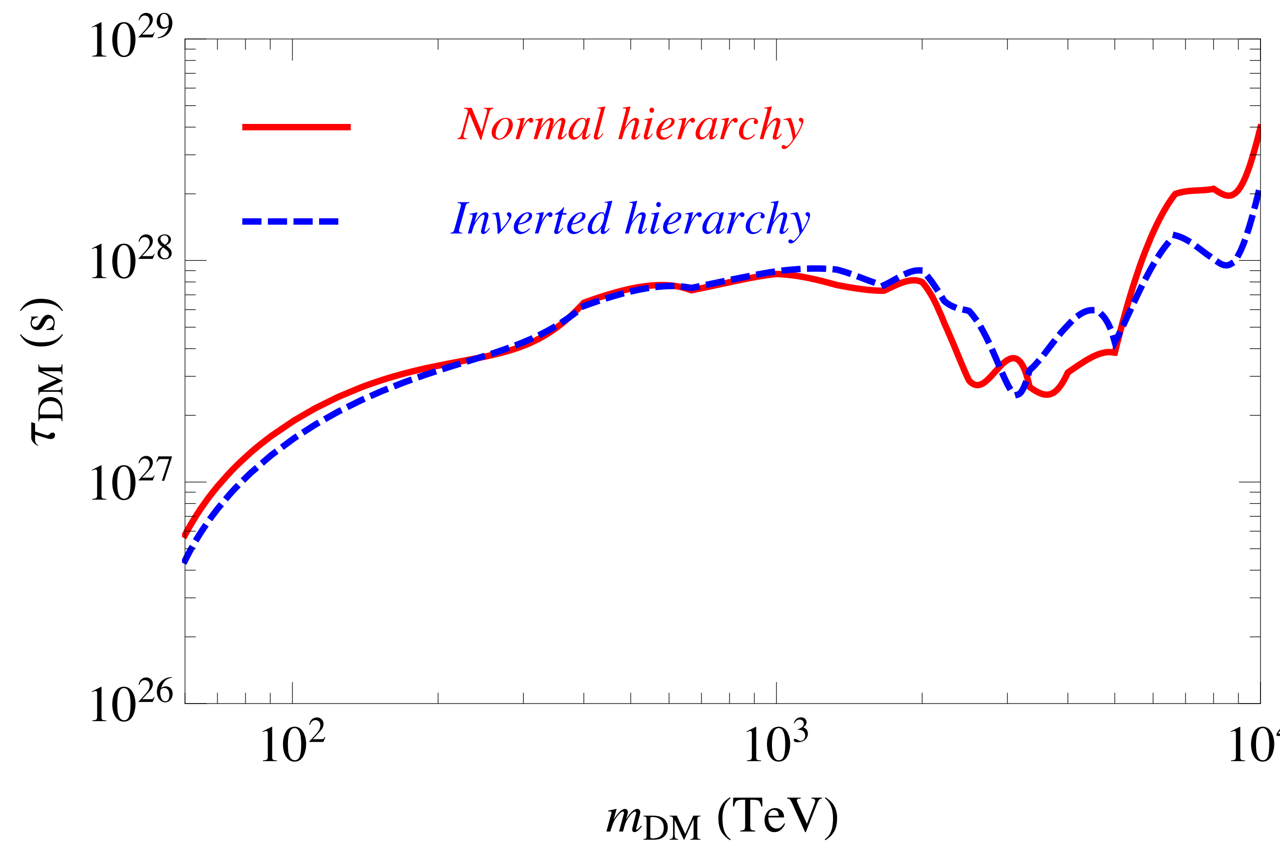
✓ at least one order of magnitude stronger lower limit on the DM lifetime, in the relevant DM mass range

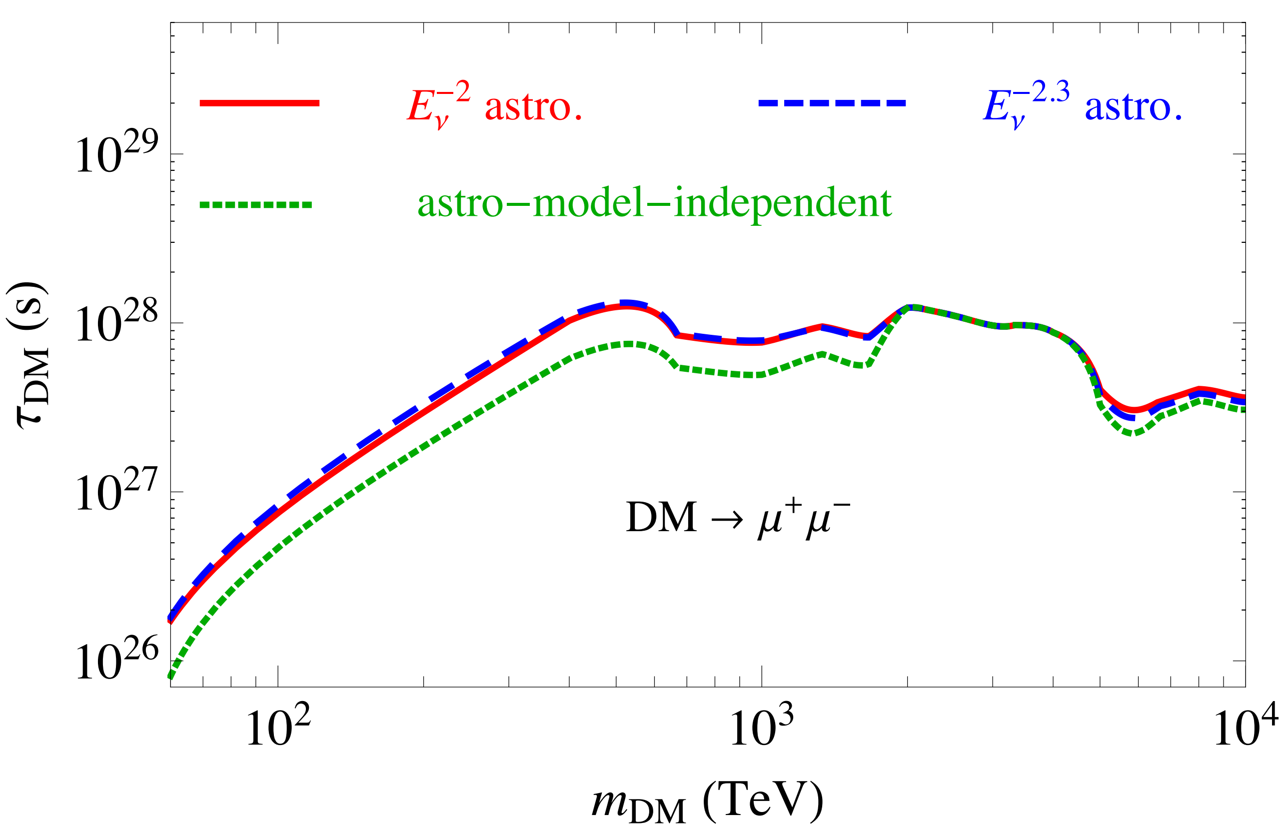
✓ for a specific model, different channels should be scaled according to the corresponding branching ratios

Constraining DM properties

✓ limits on DM lifetime (90% C.L.)

NH and IH cases →





← dependence on the astro. model?

Constraining DM properties

✓ Annihilation cross section

The lower part (< 100 TeV) of the observed spectrum can be used to probe <σv>

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The isotropic components of neutrino flux from DM annihilation:

The residual isotropic flux from the Galactic halo (anti-GC direction)

$$\frac{\mathrm{d}J_{\mathrm{iso}}^{\mathrm{ann}}}{\mathrm{d}E_\nu} = \frac{\langle \sigma v \rangle}{2} \frac{1}{4\pi m_{\mathrm{DM}}^2} \frac{\mathrm{d}N}{\mathrm{d}E_\nu} (\mathrm{l.o.s.})_{\mathrm{anti-GC}} \quad \text{where} \quad (\mathrm{l.o.s.})_{\mathrm{anti-GC}} = \int_0^\infty \rho^2[r(s, b=0, l=\pi)]\, \mathrm{d}s$$

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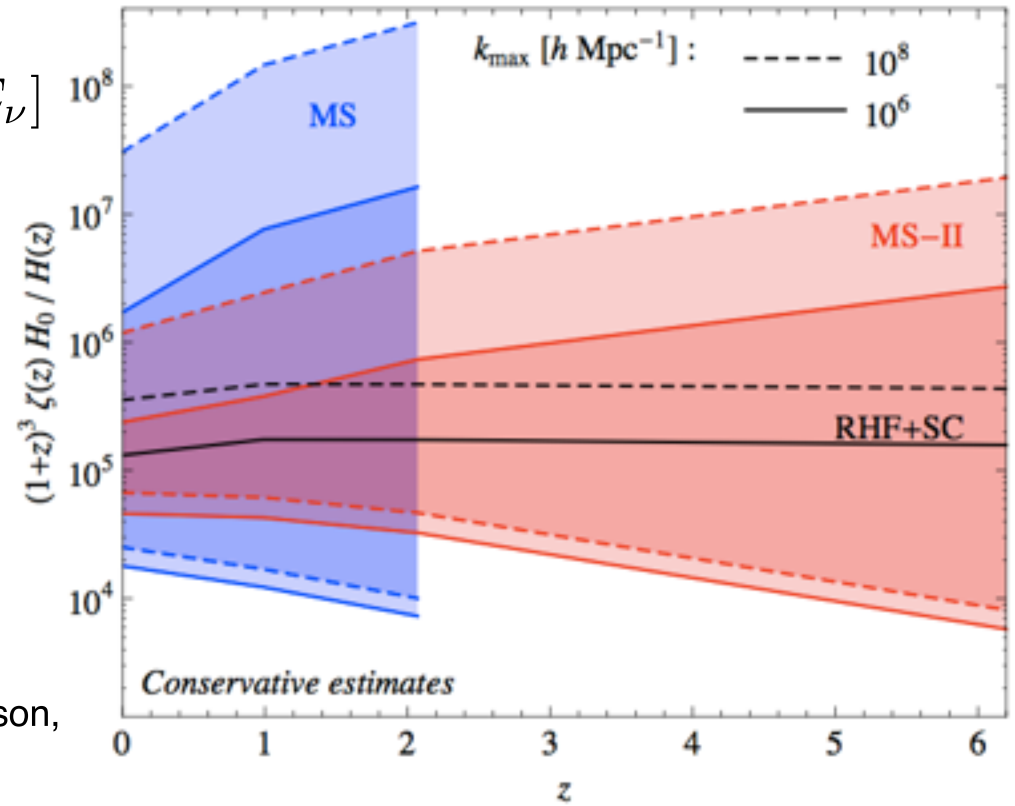
The cosmic flux from all redshift

$$\frac{\mathrm{d}J_{\mathrm{cos}}^{\mathrm{ann}}}{\mathrm{d}E_\nu} = \frac{\langle\sigma v\rangle}{2}\frac{\Omega_{\mathrm{DM}}^2\rho_c^2}{4\pi m_{\mathrm{DM}}^2}\frac{c}{H_0}\int_0^\infty \frac{(1+z)^3\zeta(z)\mathrm{d}z}{\sqrt{\Omega_m(1+z)^3+\Omega_\Lambda}}\frac{\mathrm{d}N}{\mathrm{d}E_\nu}[(1+z)\,E_\nu]$$

$\zeta(z)$ flux multiplier (DM clustering)



E. Sefusatti, G. Zaharijas, P. D. Serpico, D. Theurel and M. Gustafsson, Mon. Not. Roy. Astron. Soc. (2014) [arXiv:1401.2117].

Constraining DM properties

✓ upper limits on annihilation cross section <σv> (90% C.L.)

minimum ÷ maximum value used for ζ(z) unit of <σv> is 10^{-22} cm^3s^{-1}

DM + DM → \ m_{DM}	100 TeV	50 TeV	30 TeV
$\nu_\alpha\bar{\nu}_\alpha$	1.39 ÷ 0.22	1.21 ÷ 0.36	2.44 ÷ 0.88
$q\bar{q}$	489 ÷ 84.5	1427 ÷ 299	9934 ÷ 4603
$b\bar{b}$	185 ÷ 30.4	517 ÷ 106	3514 ÷ 1621
$c\bar{c}$	592 ÷ 100	1708 ÷ 348	11218 ÷ 5215
e^+e^-	14.7 ÷ 2.38	17.8 ÷ 5.06	41.3 ÷ 14.2
$\mu^+\mu^-$	4.47 ÷ 0.65	9.06 ÷ 1.6	23.7 ÷ 9.23
$\tau^+\tau^-$	5.84 ÷ 0.93	10.9 ÷ 2.3	28.5 ÷ 10.8
$h\bar{h}$	21.2 ÷ 3.36	53.4 ÷ 9.49	177 ÷ 76.5
$Z\bar{Z}$	11.9 ÷ 2.05	18.1 ÷ 4.09	40.7 ÷ 16.3
W^+W^-	14.4 ÷ 2.4	23.7 ÷ 4.96	54.5 ÷ 22.3

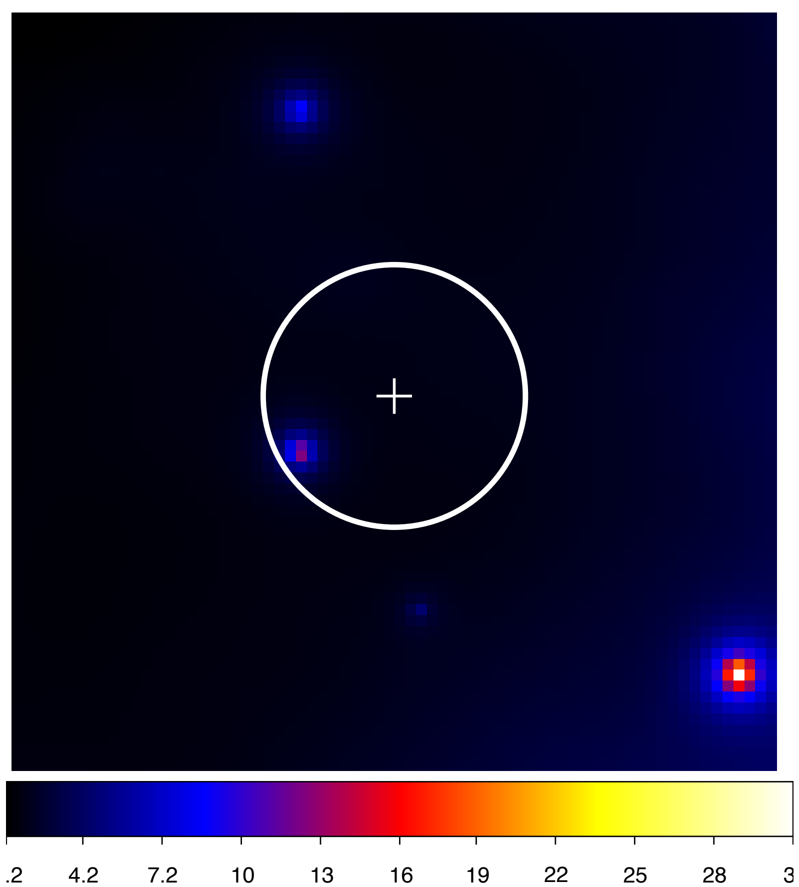
✓ for some final states (neutrinos, charged leptons) the limit is a bit stronger than the unitary bound

IceCube / Fermi-LAT cross correlation

✓ in the 4-yrs data set, 14 muon-tracks observed

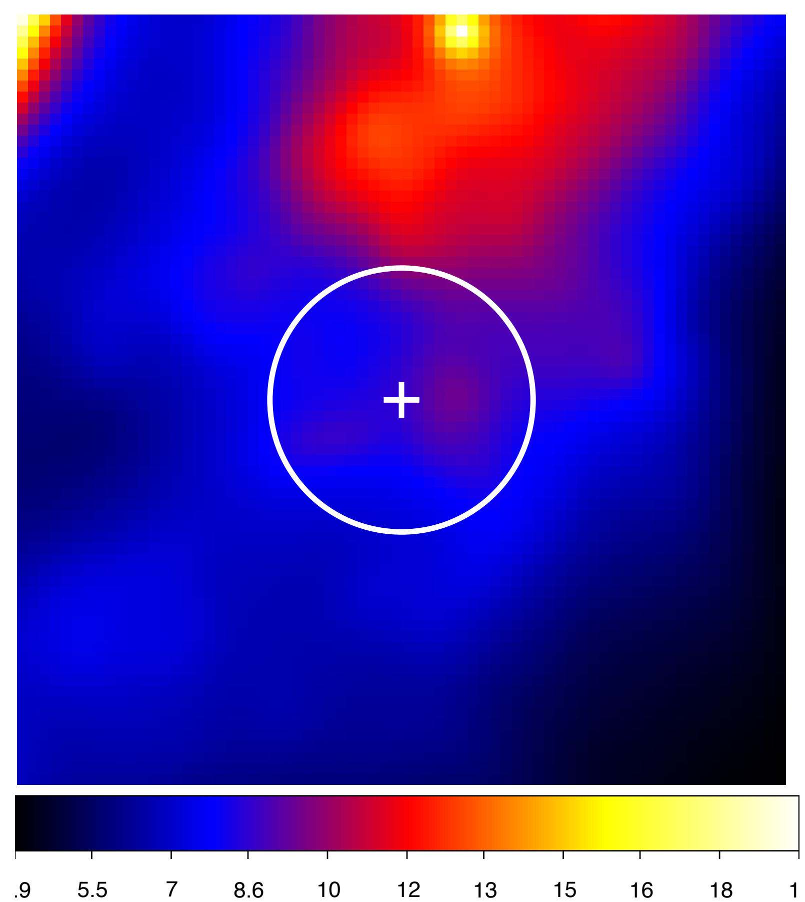
✓ let's look at the Fermi-LAT skymap at the position of these events

3FGL J0725.8-0054



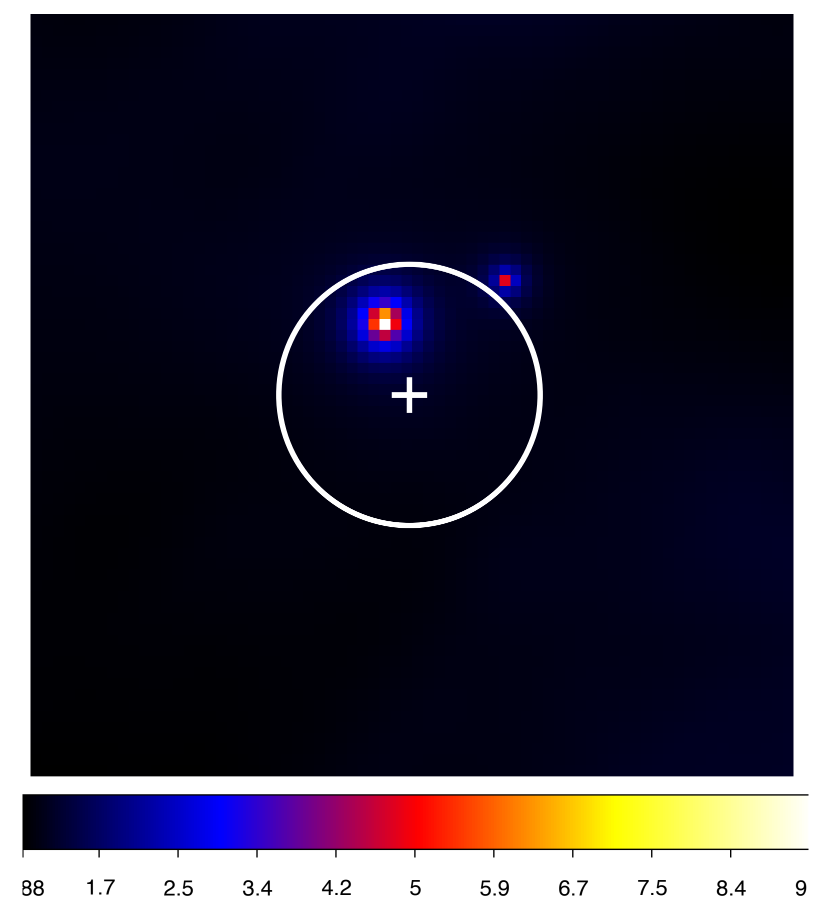
5

IC 443, SNR



38

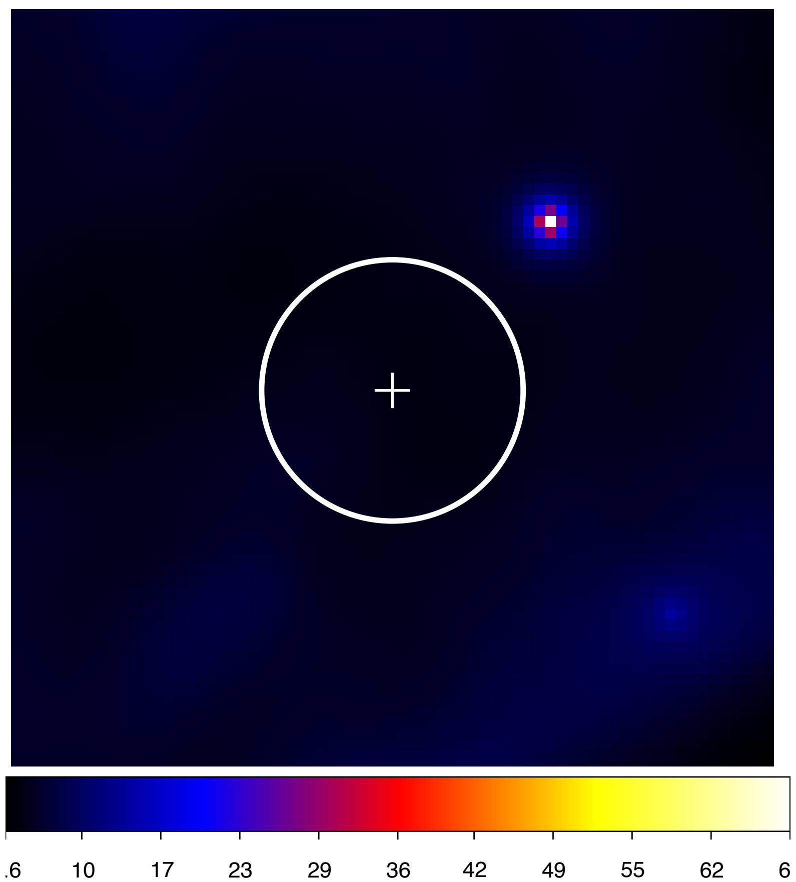
3FGL J2227.8+0040



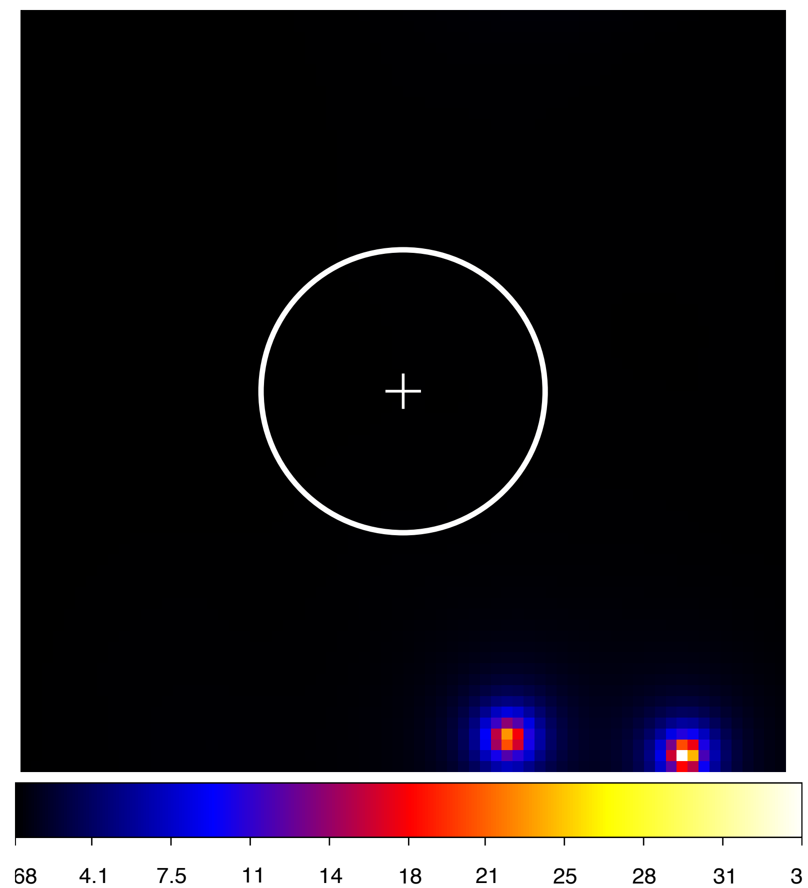
44

work in progress!!!

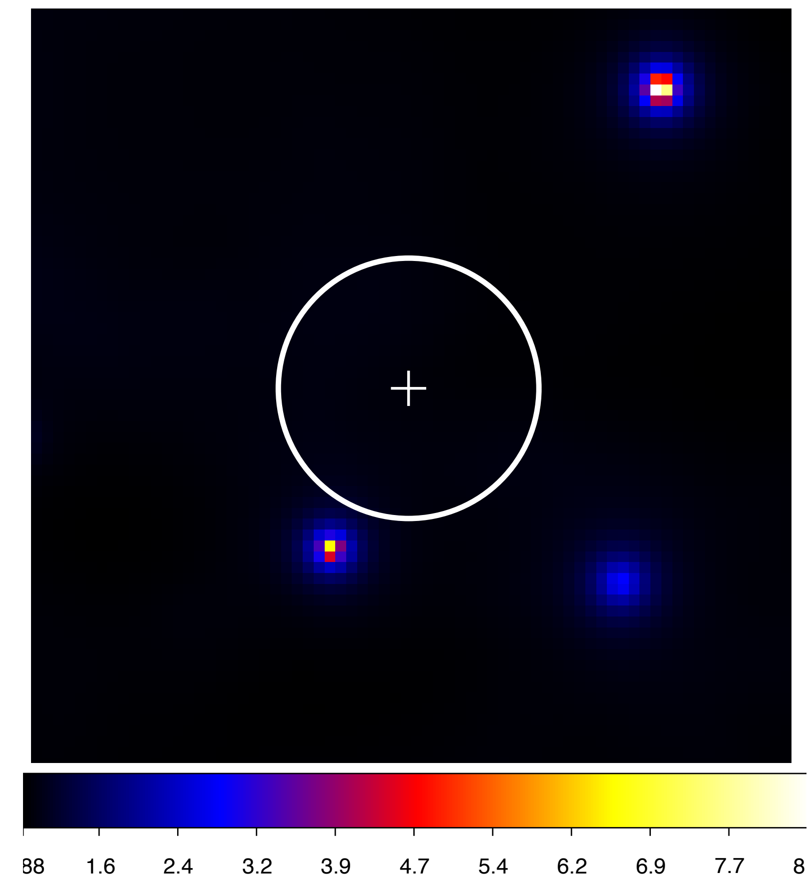
IceCube / Fermi-LAT cross correlation



13



18



47

work in progress!!!

conclusions

- ✓ The excess of events observed by IceCube in the energy range ~ 30 TeV - 2 PeV is an evidence for astrophysical flux or other "New Physics" induced fluxes

- ✓ Several features of the observed events motivate us for a DM interpretation: cut-off at ~ 2 PeV, a mild dip in the (400 - 1000) TeV and anisotropy.

- ✓ We argued that a PeV-scale decaying DM, with generic decay channels, can naturally explain these features. The required lifetime is allowed by the current limits. Both the energy and angular distributions mildly prefer DM interpretation.

- ✓ With more statistics in the next few years, the DM interpretation of IceCube events can be tested. The gamma-ray flux expected in this scenario can be detected by the next generation of EAS detectors. Also, anisotropy measurements in the CR flux would be constraining.

Thank you !

Decaying Dark Matter

$$\Phi = kE_\nu^{-2}$$

	k	E^{min}-E^{max} (TeV)	N_{bg}	N_{sig}	N_{limit}
AMANDA	7.4×10^{-8}	$16 - 2.5\times10^{3}$	6	7	5.4
IceCube-22	1.6×10^{-7}	$340 - 2\times10^{5}$	0.6	3	6.1
IceCube-40	3.6×10^{-8}	$2\times10^{3} - 6.3\times10^{6}$	0.1	0	2.3
Auger	1.7×10^{-7}	$10^{5} - 10^{8}$	0	0	2.3
ANITA	1.3×10^{-7}	$10^{6} - 3.2\times10^{11}$	0.97	1	3.3

$$N_{\rm exp} = T\Delta\Omega \sum_\alpha \left[\int_{E_\nu^{\rm min}}^{E_\nu^{\rm max}} \Phi_{\nu_\alpha+\bar{\nu}_\alpha} A_{\rm eff}^\alpha(E_\nu) dE_\nu \right]$$

$$q/100 = \frac{\int_0^{N_{\rm limit}} L(N_{\rm sig}|N)dN}{\int_0^\infty L(N_{\rm sig}|N)dN}$$ **where** $$L(N_{\rm sig}|N) = \frac{(N+N_{\rm bg})^{N_{\rm sig}}}{N_{\rm sig}!} e^{-(N+N_{\rm bg})}$$

Decaying Dark Matter

✓ extragalactic contribution:

$$H(z) = H_0 \sqrt{\Omega_\Lambda + \Omega_{\rm m}(1+z)^3}$$

$$\frac{{\rm d}J_{\rm eg}}{{\rm d}E_\nu} = \frac{\Omega_{\rm DM}\rho_{\rm c}}{4\pi m_{\rm DM}\tau_{\rm DM}} \int_0^\infty {\rm d}z \, \frac{1}{H(z)} \frac{{\rm d}N_\nu}{{\rm d}E_\nu}\left[(1+z)E_\nu\right] e^{-s_\nu(E_\nu,z)}$$

Opacity

$$s_\nu(E_\nu, z) = \begin{cases} 7.4\times 10^{-17}(1+z)^{7/2}(E_\nu/\,{\rm TeV}), & \text{for } 1 \ll z < z_{\rm eq} \\ 1.7\times 10^{-14}(1+z)^{3}(E_\nu/\,{\rm TeV}), & \text{for } z \gg z_{\rm eq} \end{cases}$$

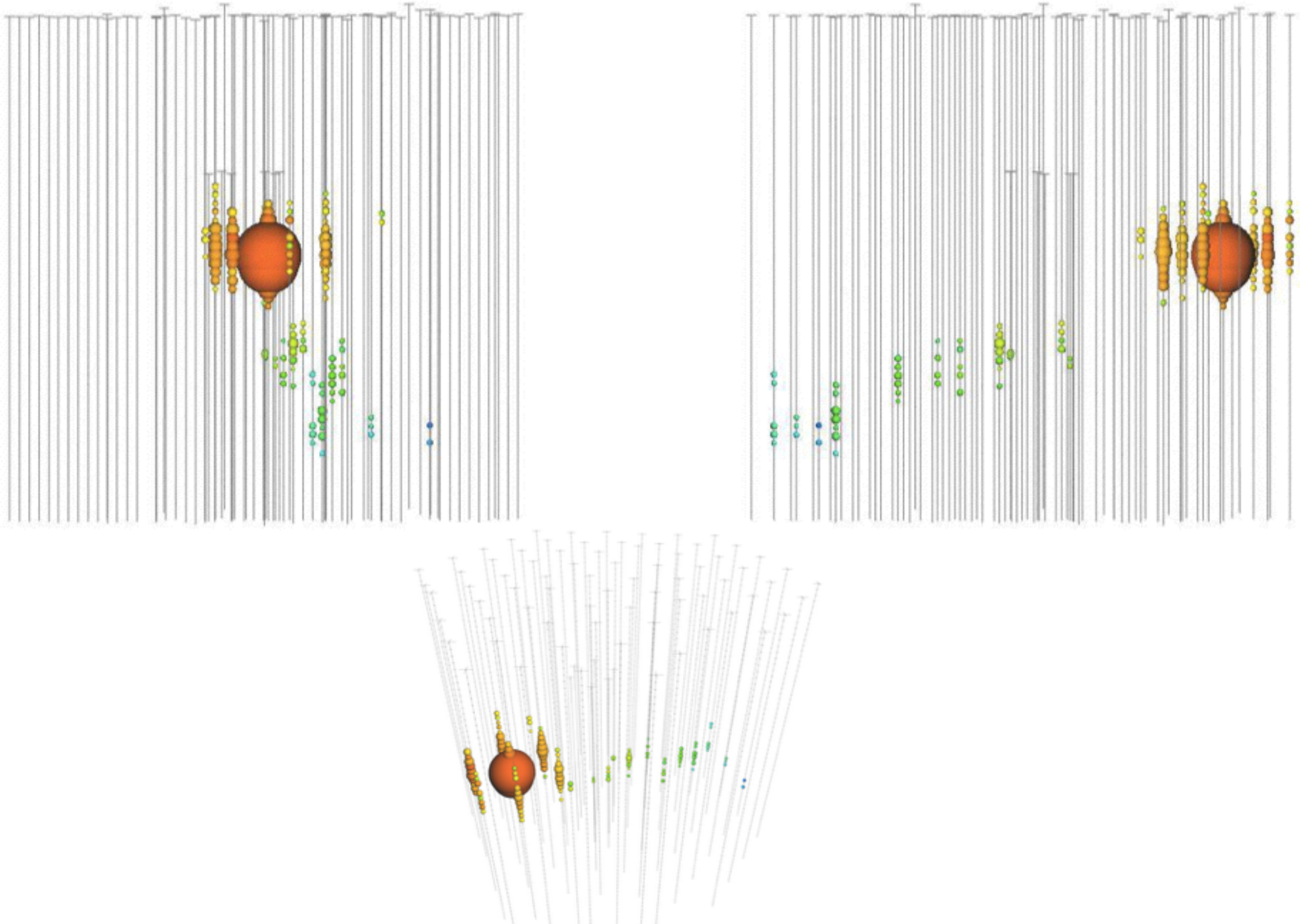
at the Earth

$$\begin{pmatrix} J_e \\ J_\mu \\ J_\tau \end{pmatrix} = \begin{pmatrix} P_{ee} & P_{e\mu} & P_{e\tau} \\ & P_{\mu\mu} & P_{\mu\tau} \\ & & P_{\tau\tau} \end{pmatrix} \begin{pmatrix} I_e \\ I_\mu \\ I_\tau \end{pmatrix}$$

production point

decoherent oscillation

Event 8 (Track)



Gamma ray bounds

✓ Galactic component

at ~ PeV, the absorption length of gamma-rays are comparable to Galactic distances → neither full absorption or cascade development, nor full transparency

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the integrated flux of γ-ray above 330 TeV must be below 1.0×10^{-13} cm^{-2}s^{-1}sr^{-1}

the integrated flux of γ-ray above 775 TeV must be below 2.6×10^{-14} cm^{-2}s^{-1}sr^{-1}

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For DM decay (NH case) unattenuated flux →

the integrated flux of γ-ray above 330 TeV is 1.2×10^{-14} cm^{-2}s^{-1}sr^{-1}

the integrated flux of γ-ray above 775 TeV is 8.8×10^{-15} cm^{-2}s^{-1}sr^{-1}

✓